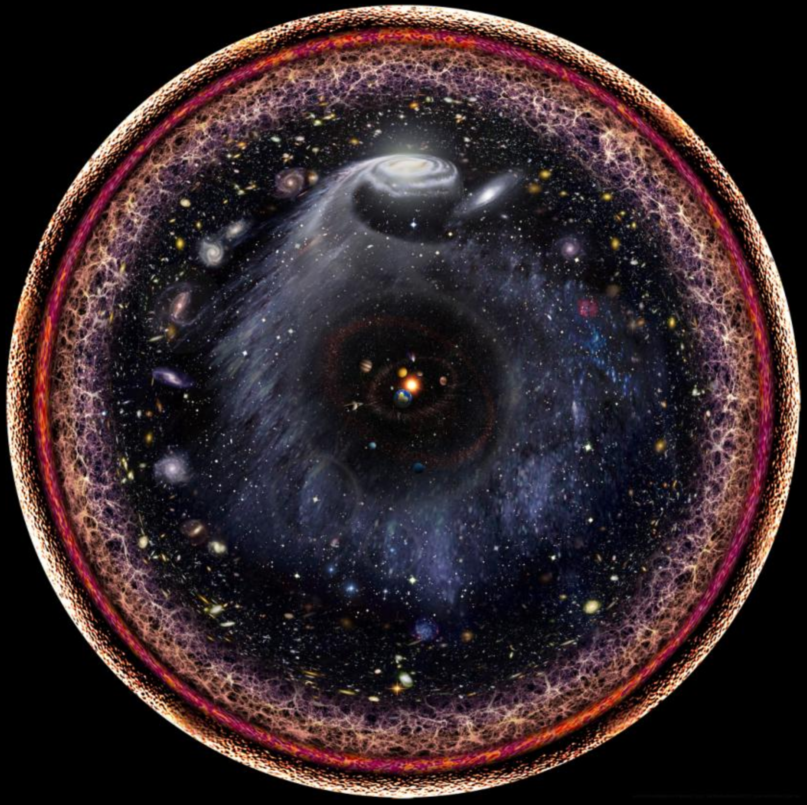
Cosmological searches for dark matter-baryon interactions



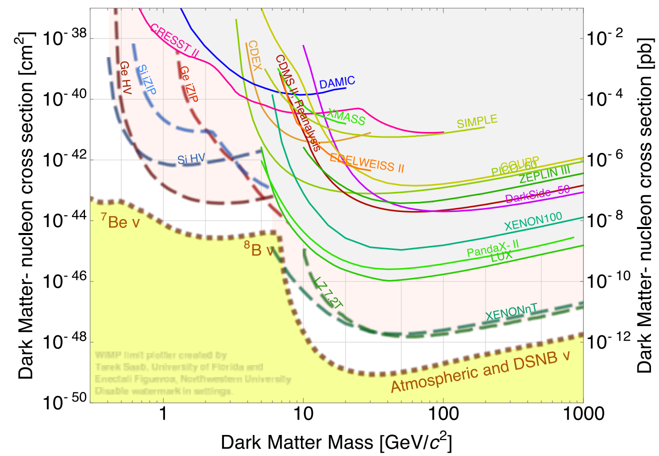
Vera Gluscevic
Eric Schmidt Fellow, Institute for Advanced Study

Image credit: Budassi

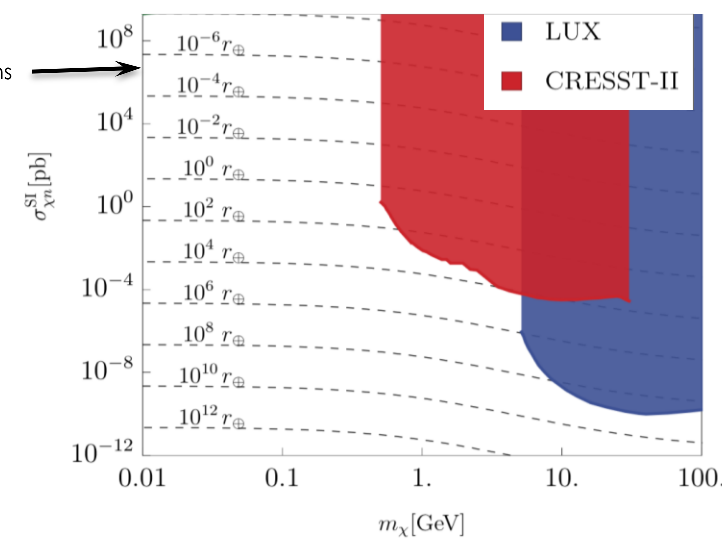
KITP 2018

Caveats of low-energy searches

- Assumptions about <u>local</u> astrophysics of DM. [cf. talk by Nassim]
- Limited ability to pin down nature of the interaction. [see arXiv 1506.04454]
- DD has a ceiling.
- Masses << GeV are still poorly explored.
- Vast sub-GeV lanscape!

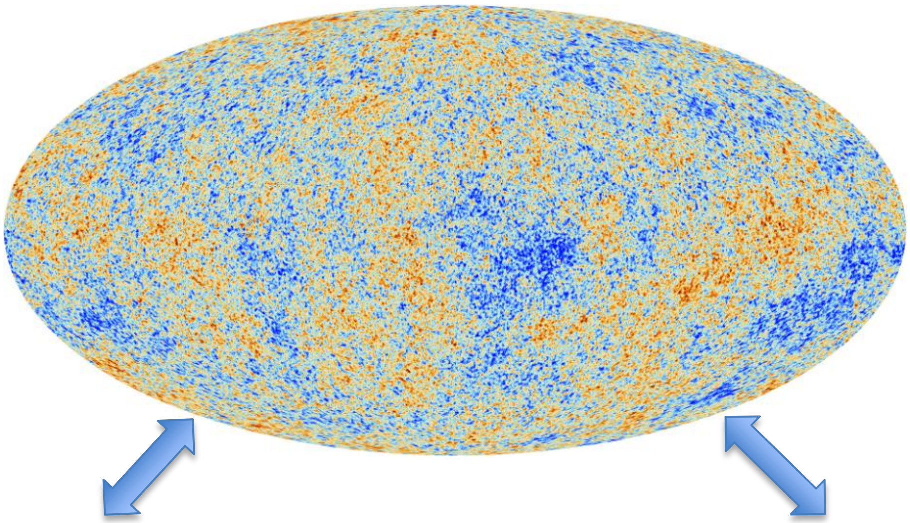


@ 10^{-29} cm^2:
1000 interactions
before target

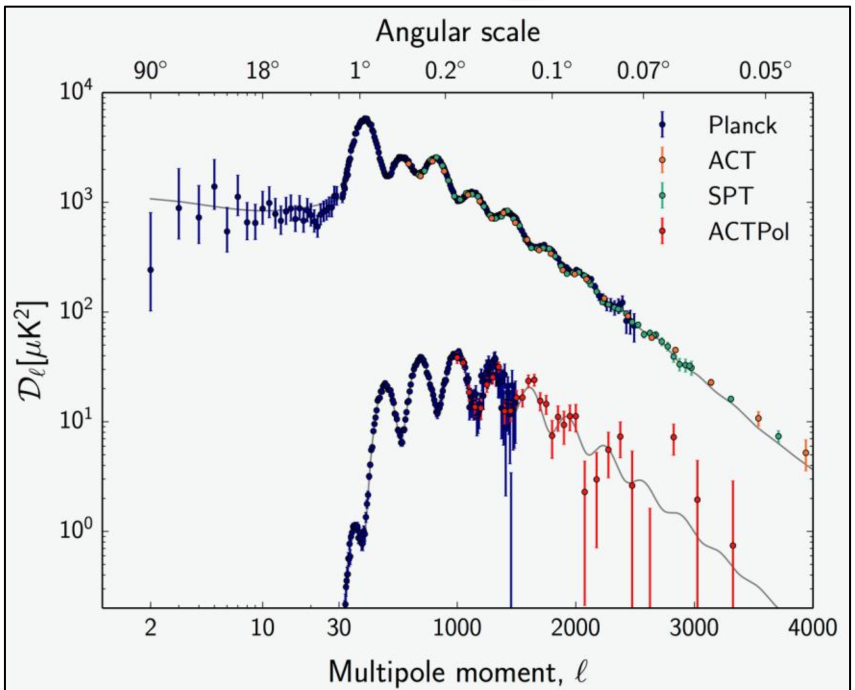


Emken and Kouvaris, 2017

Cosmic microwave background [CMB]

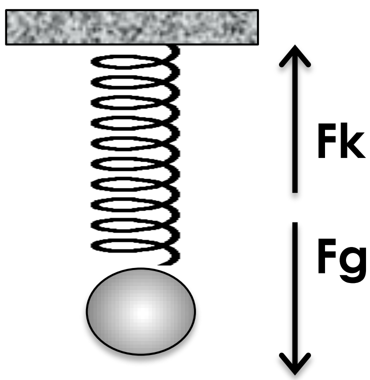


Power spectrum



Plot by E. Calabrese
(for ACTPol)

Acoustic oscillations



CMB on Dark Matter

DM = gravitating fluid (old and uninteresting).

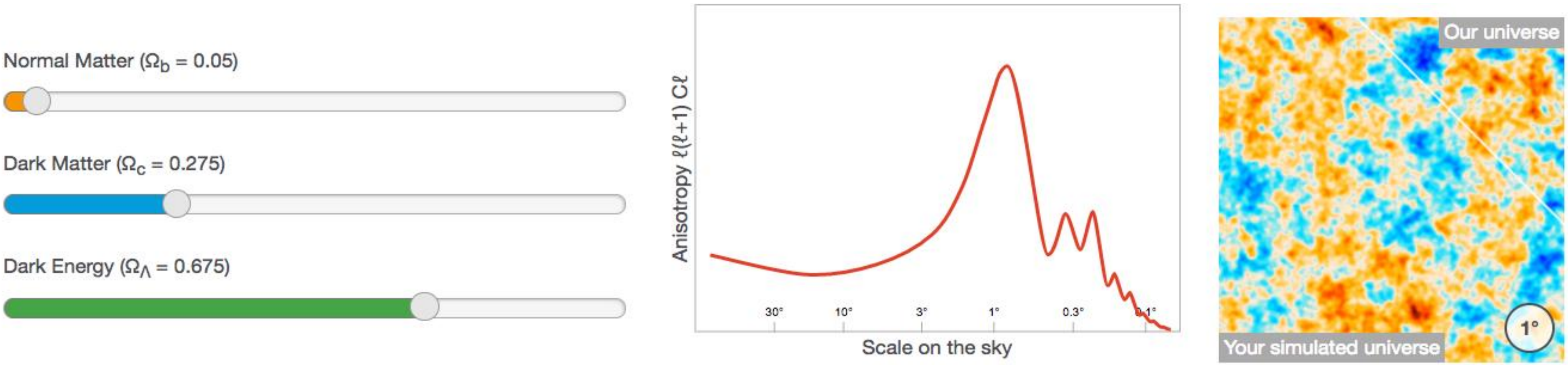
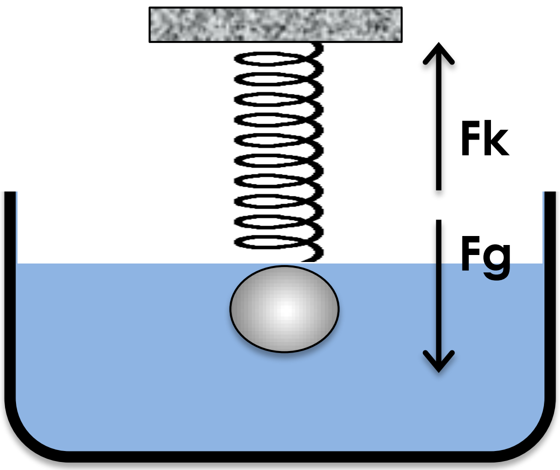
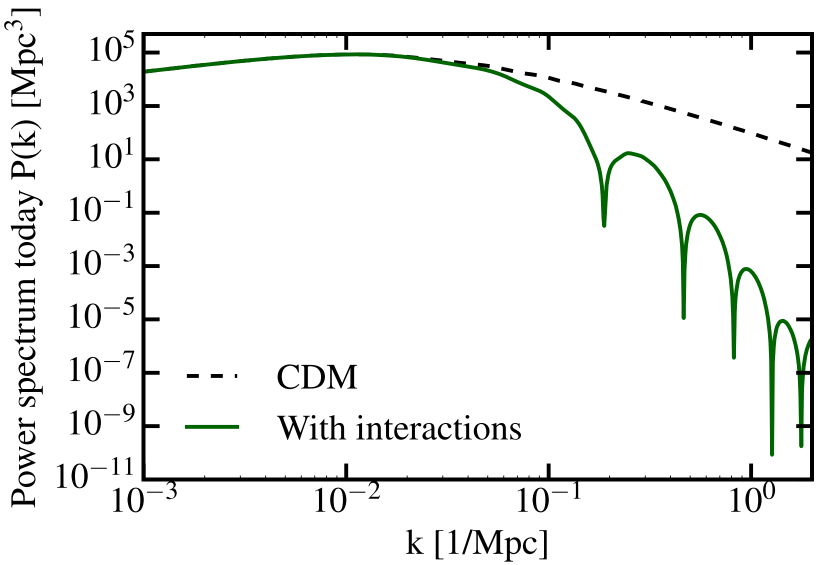


Image credits: Amanda Yoho, Planck CMB simulator at http://strudel.org.uk/planck/#; ESA/Planck

With dark matter-proton scattering:

scattering →drag force→ suppression of small scales

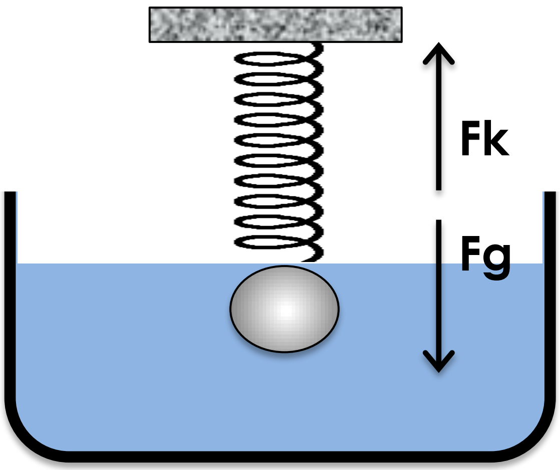




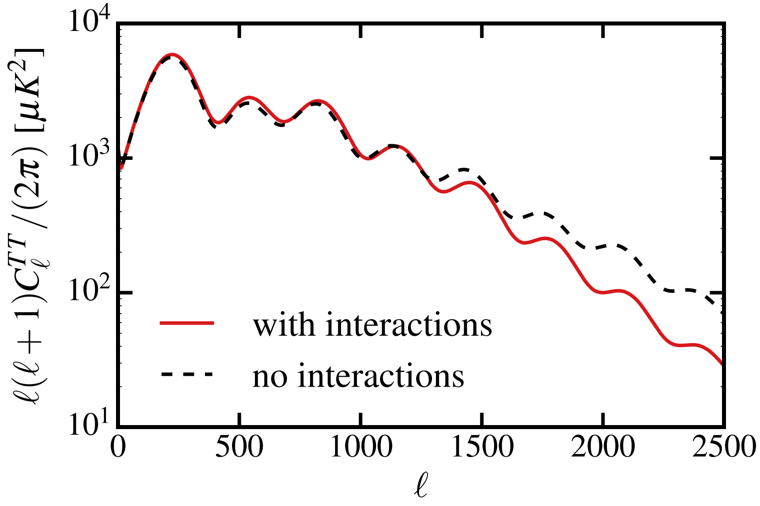
Boddy and VG (2018).

With dark matter-proton scattering:

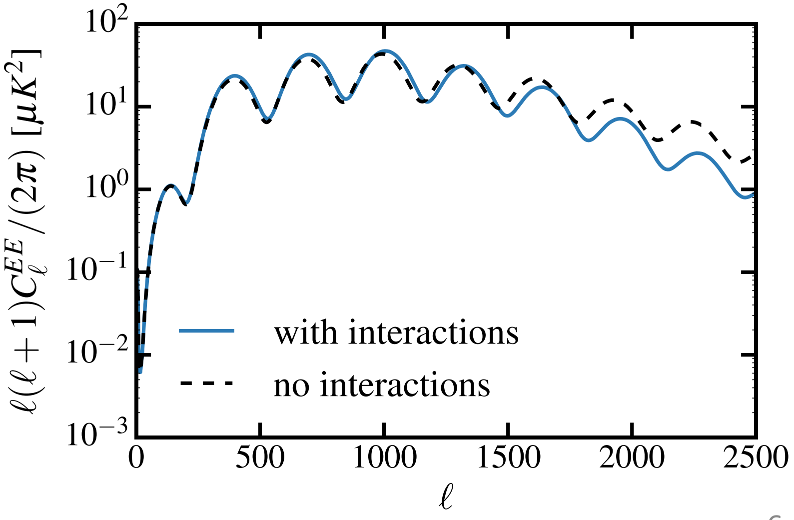
scattering →drag force→ suppression of small scales



Temperature



Polarization



VG and Boddy (2017), Boddy and VG (2018); Previous work: Chen et al (2002), Sigurdson et al (2004); Dvorkin et al (2014); et

Scattering in the early universe

Momentum transfer between baryon-photon fluid and DM

$$\dot{\delta}_\chi = -\theta_\chi - \frac{\dot{h}}{2}, \qquad \dot{\delta}_b = -\theta_b - \frac{\dot{h}}{2},$$

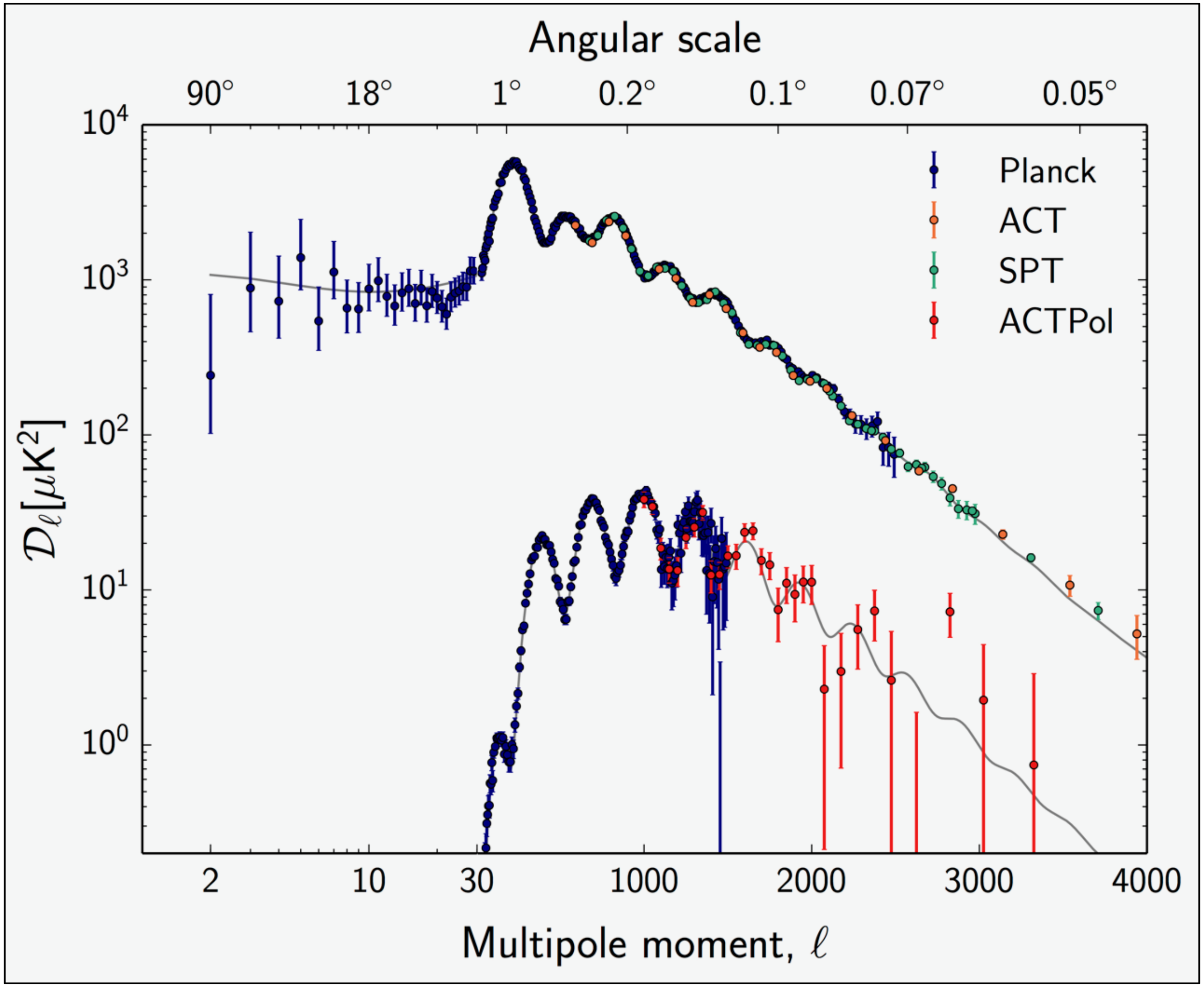
$$\dot{\theta}_\chi = -\frac{\dot{a}}{a}\theta_\chi + c_\chi^2 k^2 \delta_\chi + \underline{R_\chi (\theta_b - \theta_\chi)},$$

$$\dot{\theta}_b = -\frac{\dot{a}}{a}\theta_b + c_b^2 k^2 \delta_b + R_\gamma (\theta_\gamma - \theta_b) + \underline{\frac{\rho_\chi}{\rho_b} R_\chi (\theta_\chi - \theta_b)}$$

Gluscevic and Boddy (2017), Boddy and Gluscevic (2018), Chen et al (2002), Sigurdson et al (2004); Dvorkin et al (2014); etc.

$$R_\chi = \frac{a\, c_n\, \rho_b\, \sigma_0}{m_\chi + m_H} \left(\frac{T_b}{m_H} + \frac{T_\chi}{m_\chi} \right)^{\frac{n+1}{2}} F_{He}$$

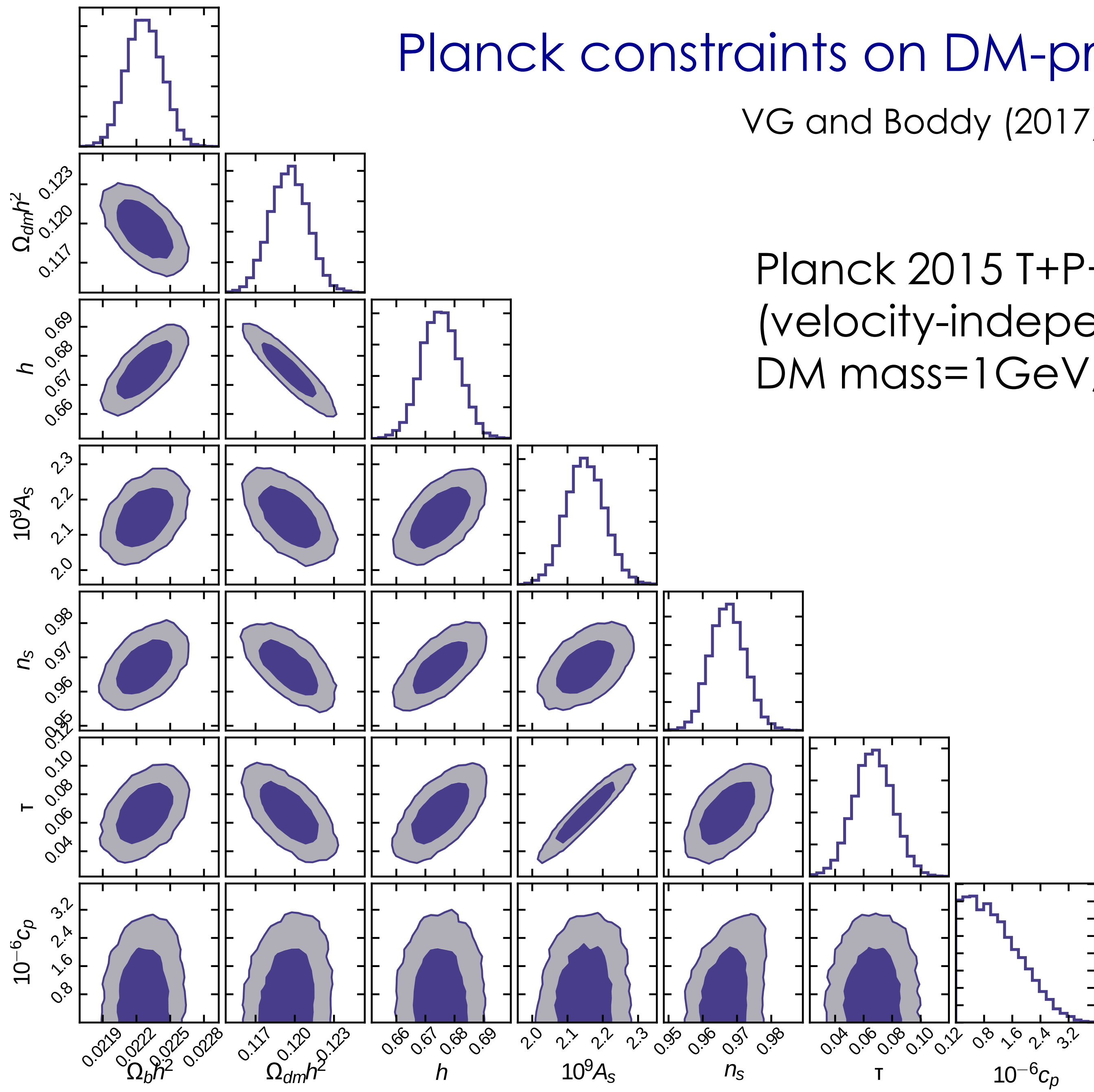
Data



Plot by E. Calabrese [for ACTPol]

Planck constraints on DM-proton scattering

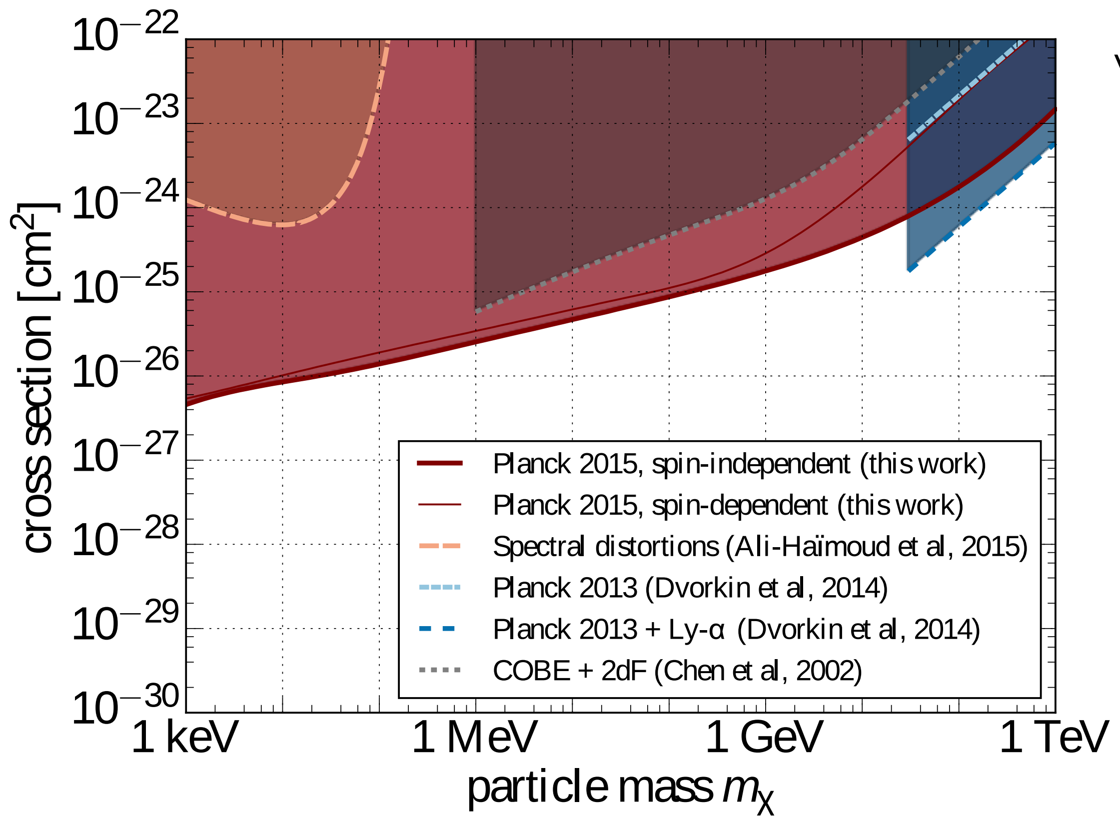
VG and Boddy (2017), Boddy and VG (2018)



Planck 2015 T+P+lensing
(velocity-independent scattering,
DM mass=1GeV, include He):

Cosmological exclusion curves

v-independent DM scattering with proton: 95% confidence upper limit



VG and Boddy (2017)

High cross sections, down to mass ~keV!

Non-relativistic EFT

[Fan et al, 2010; Fitzpatrick et al, 2012; Anand et al, 2013]

$$\mathcal{O}_1 = 1_\chi 1_N$$

$$\mathcal{O}_3 = \vec{S}_N \cdot \left(\frac{i\vec{q}}{m_N} \times \vec{v}^\perp\right)$$

$$\mathcal{O}_4 = \vec{S}_\chi \times \vec{S}_N$$

$$\mathcal{O}_5 = \vec{S}_\chi \cdot \left(\frac{i\vec{q}}{m_N} \times \vec{v}^\perp\right)$$

$$\mathcal{O}_6 = -\left(\vec{S}_\chi \cdot \frac{i\vec{q}}{m_N}\right)\left(\vec{S}_N \cdot \frac{i\vec{q}}{m_N}\right)$$

$$\mathcal{O}_7 = \vec{S}_N \cdot \vec{v}^\perp$$

$$\mathcal{O}_8 = \vec{S}_\chi \cdot \vec{v}^\perp$$

$$\mathcal{O}_9 = \vec{S}_\chi \cdot \left(\vec{S}_N \times \frac{i\vec{q}}{m_N}\right)$$

$$\mathcal{O}_{10} = \vec{S}_N \cdot \frac{i\vec{q}}{m_N}$$

$$\mathcal{O}_{11} = \vec{S}_\chi \cdot \frac{i\vec{q}}{m_N}$$

$$\mathcal{O}_{12} = \vec{S}_\chi \cdot \left(\vec{S}_N \times \vec{v}^\perp\right)$$

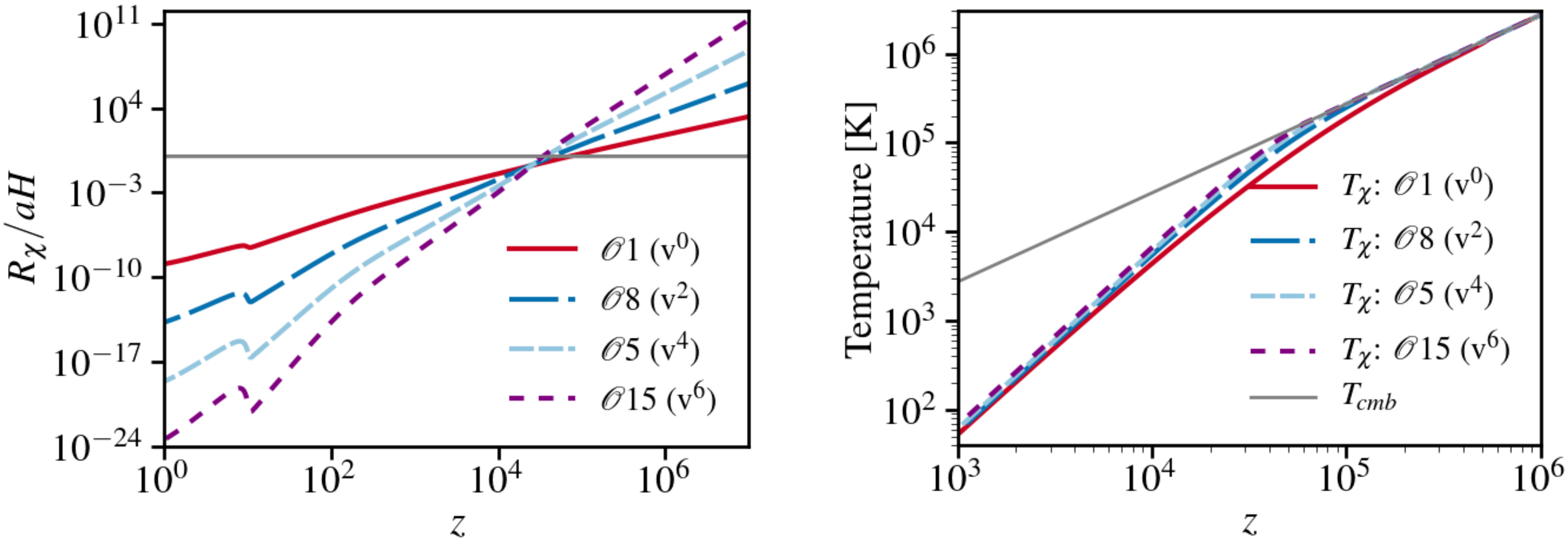
$$\mathcal{O}_{13} = \left(\vec{S}_\chi \cdot \vec{v}^\perp\right)\left(\vec{S}_N \cdot \frac{i\vec{q}}{m_N}\right)$$

$$\mathcal{O}_{14} = \left(\vec{S}_\chi \cdot \frac{i\vec{q}}{m_N}\right)\left(\vec{S}_N \cdot \vec{v}^\perp\right)$$

$$\mathcal{O}_{15} = \left(\vec{S}_\chi \cdot \frac{i\vec{q}}{m_N}\right)\left[\left(\vec{S}_N \times \vec{v}^\perp\right) \cdot \frac{i\vec{q}}{m_N}\right]$$

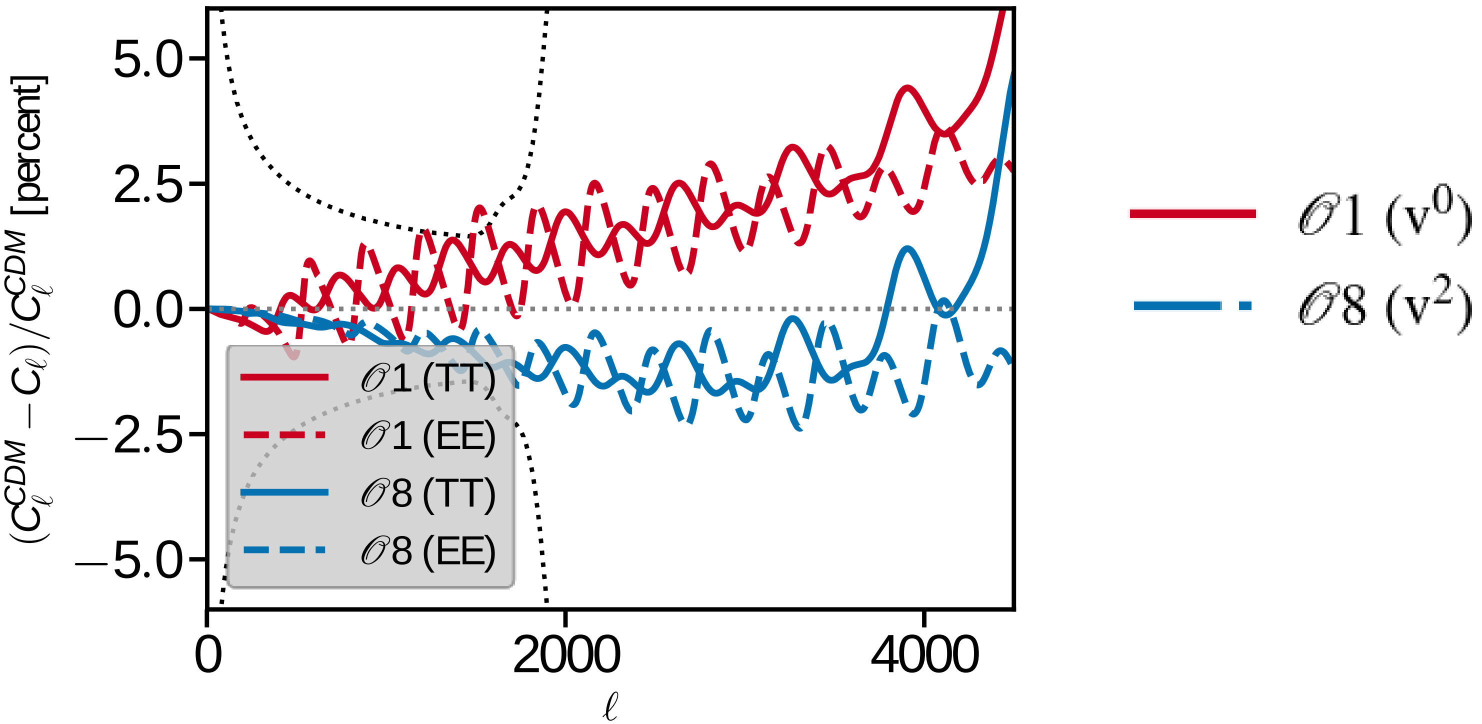
Momentum transfer

✧ Each operator -> cross section with a different dependence on relative particle velocity, different thermal history:

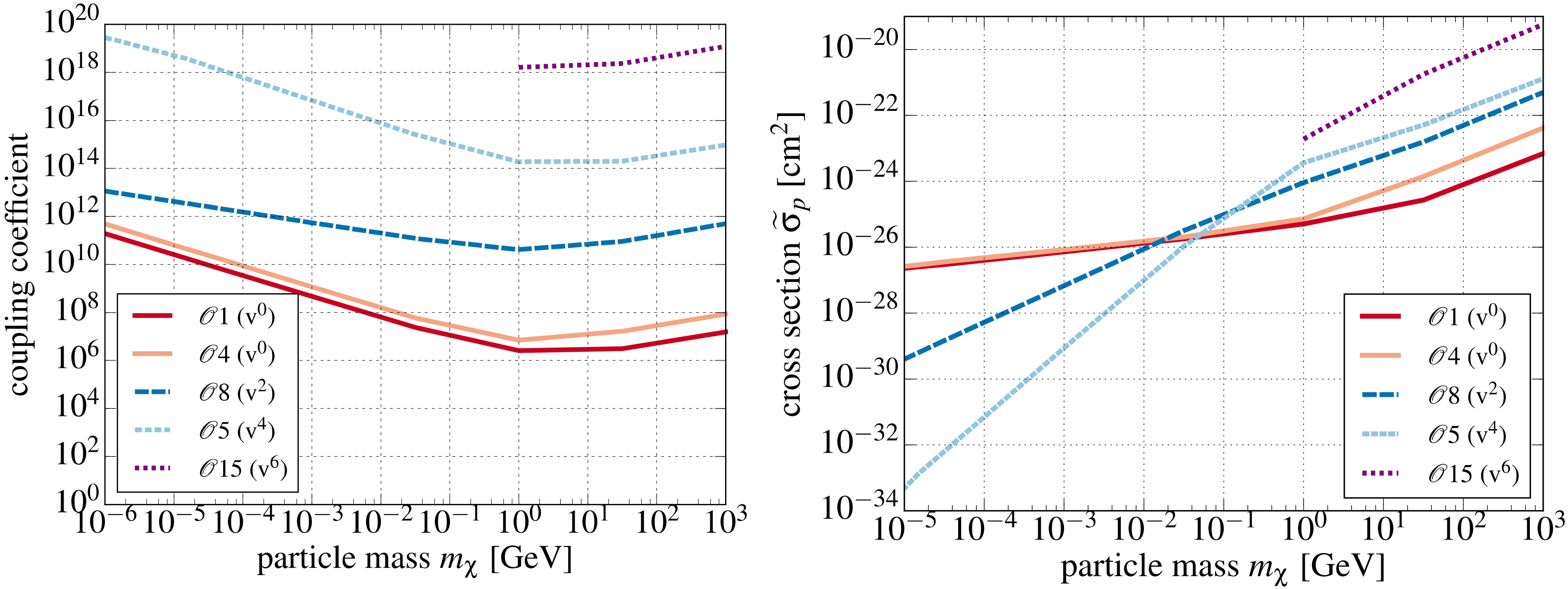


CMB observables

✧ Each operator -> cross section with a different dependence on relative particle velocity, different thermal history:



Cosmological constraint on DM-baryon EFT



Boddy and VG (2018)

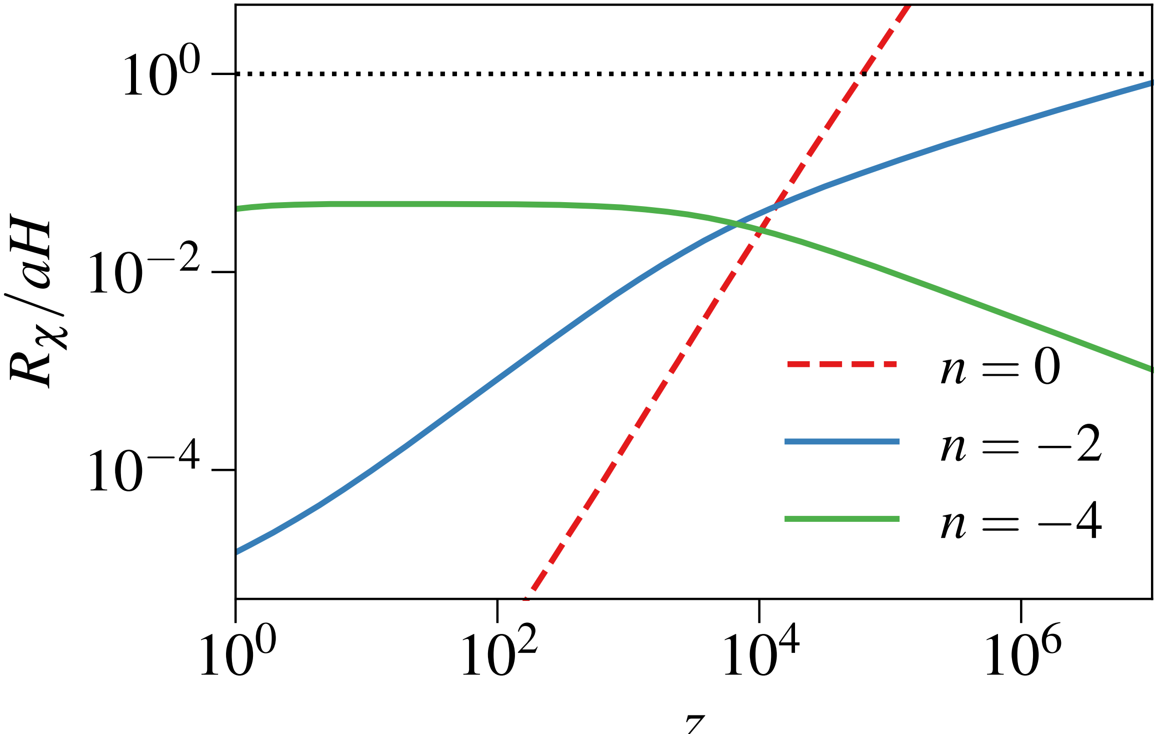
Other scattering scenarios

$\sigma \sim \sigma_0 \, v^n$

Dvorkin+ (2014); Xu+ (2018); Slatyer+ (2018)

Late-time: n <-2

Early-time: n≥-2



Boddy, VG, Poulin, + (coming up)

Late-time scattering: relative bulk velocity

Tseliakhovitch and Hirata (2010)

Problem: non-linear equations

$$\dot{\delta}_\chi = -\theta_\chi - \frac{\dot{h}}{2}, \qquad \dot{\delta}_b = -\theta_b - \frac{\dot{h}}{2},$$

$$\dot{\theta}_\chi = -\frac{\dot{a}}{a}\theta_\chi + c_\chi^2 k^2 \delta_\chi + R_\chi (\theta_b - \theta_\chi),$$

$$\dot{\theta}_b = -\frac{\dot{a}}{a}\theta_b + c_b^2 k^2 \delta_b + R_\gamma (\theta_\gamma - \theta_b) + \frac{\rho_\chi}{\rho_b} R_\chi (\theta_\chi - \theta_b)$$

$$R_\chi = \frac{a\, c_n\, \rho_b\, \sigma_0}{m_\chi + m_H} \left(\frac{T_b}{m_H} + \frac{T_\chi}{m_\chi} \right)^{\frac{n+1}{2}} F_{He}$$

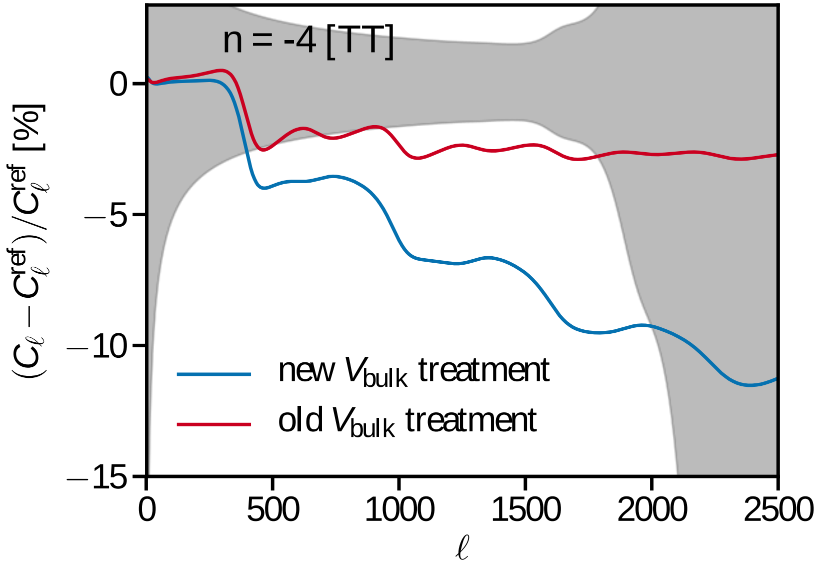
Only for Vbulk << Vthermal

Late-time scattering: relative bulk velocity

Problem: non-linear equations

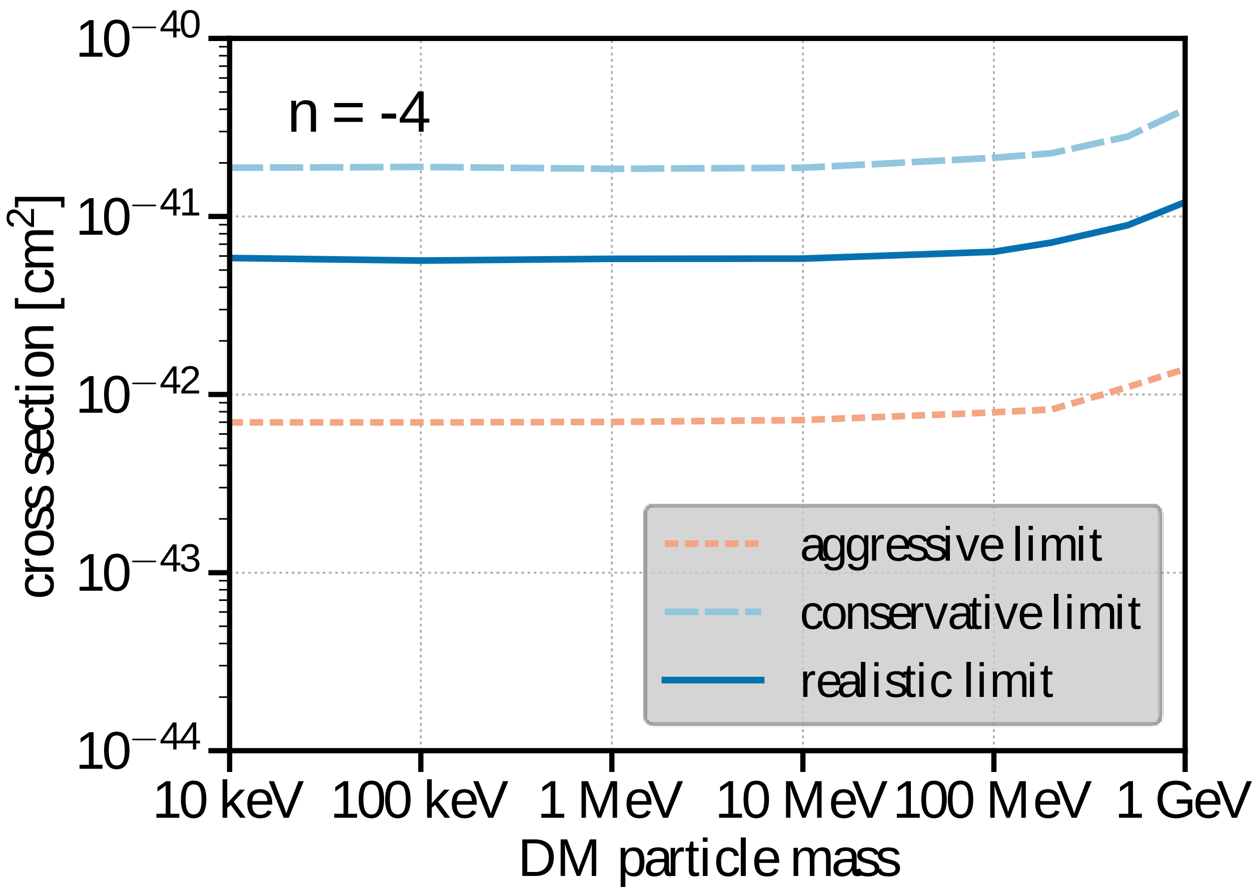
Solution: use a proxy for Vbulk

$$V_{\rm RMS}^2(k,z) \equiv \int_k^\infty \frac{dk'}{k'} \Delta_\zeta^2(k') \left[\frac{\theta_b(k',z) - \theta_\chi(k',z)}{k'}\right]^2$$



Boddy, VG, Poulin, + (coming up)

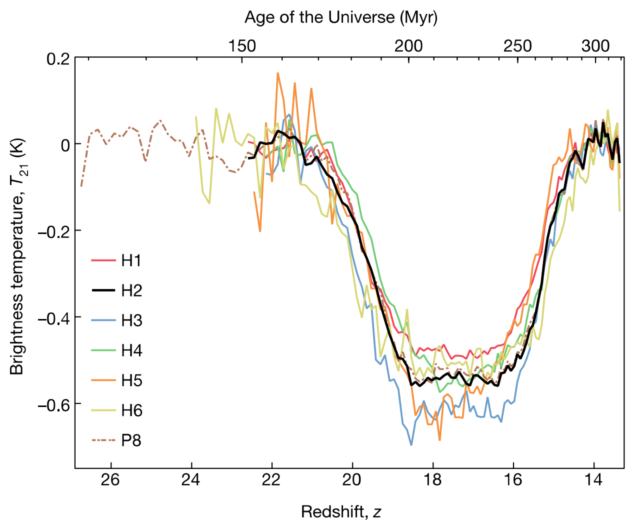
v^{-4} scattering: Planck limits



Boddy, VG, Poulin, + (coming up)

What about EDGES?

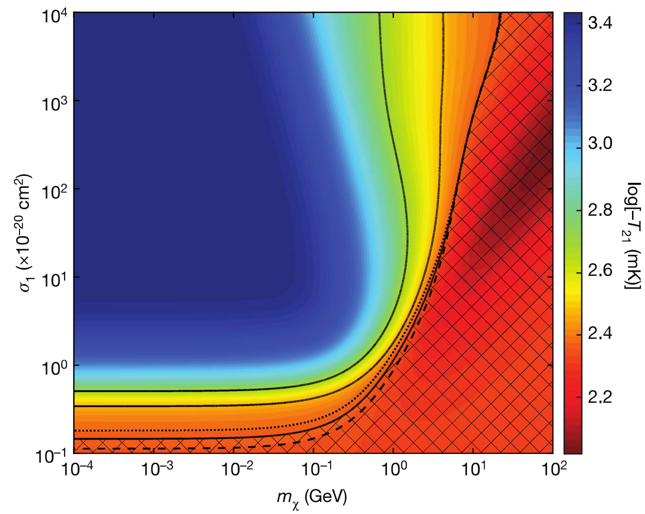
[cf. Jordan Mirocha's talk]



Bowman+ (2018)

?

- Systematics?
- Foregrounds?
- Astrophysics?
- **Cold baryons?**

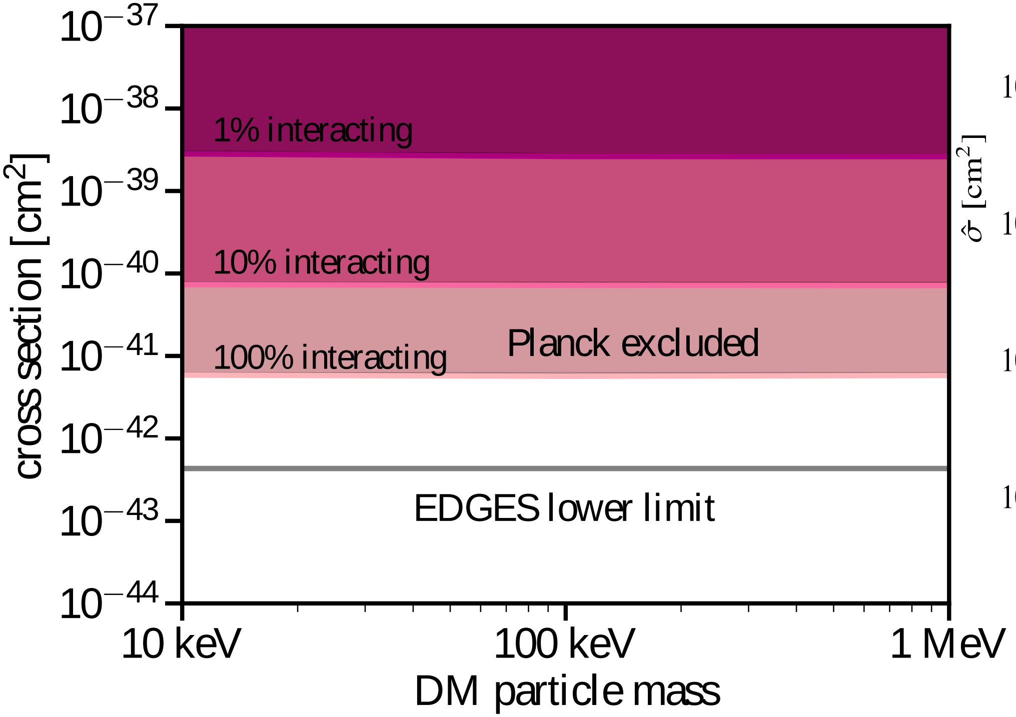


Barkana (2018)

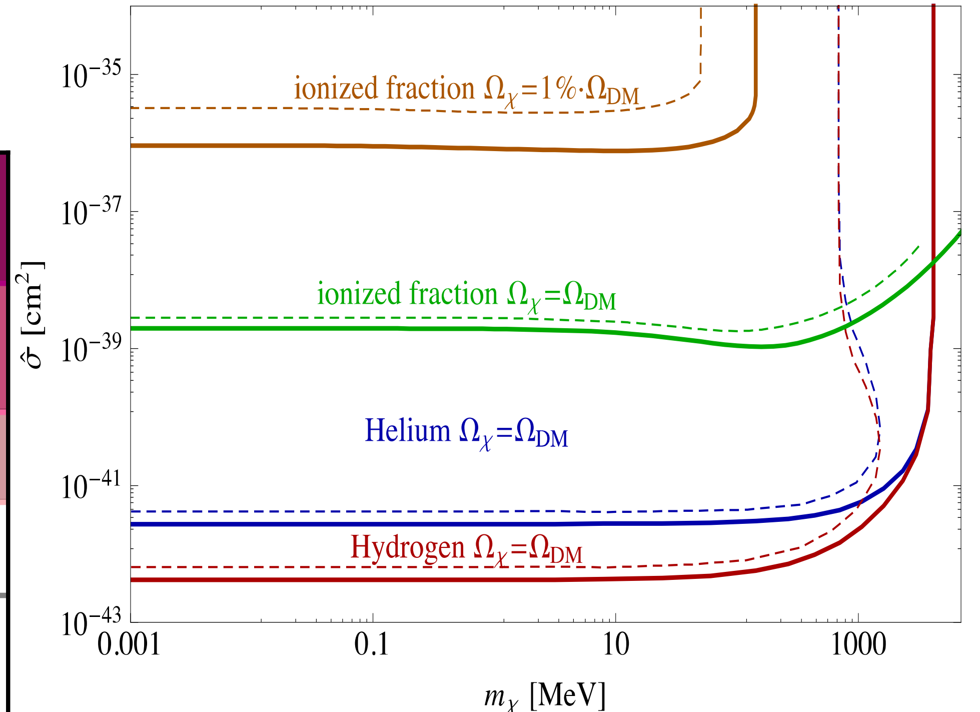
Order of business: **Is it in the sky? Is it cosmological? Is it DM?**

EDGES: v^{-4} and millicharge

From CMB limits on momentum-transfer: EDGES cannot be 1% of millicharged DM, but could be 100% with some other v^{-4} interaction.



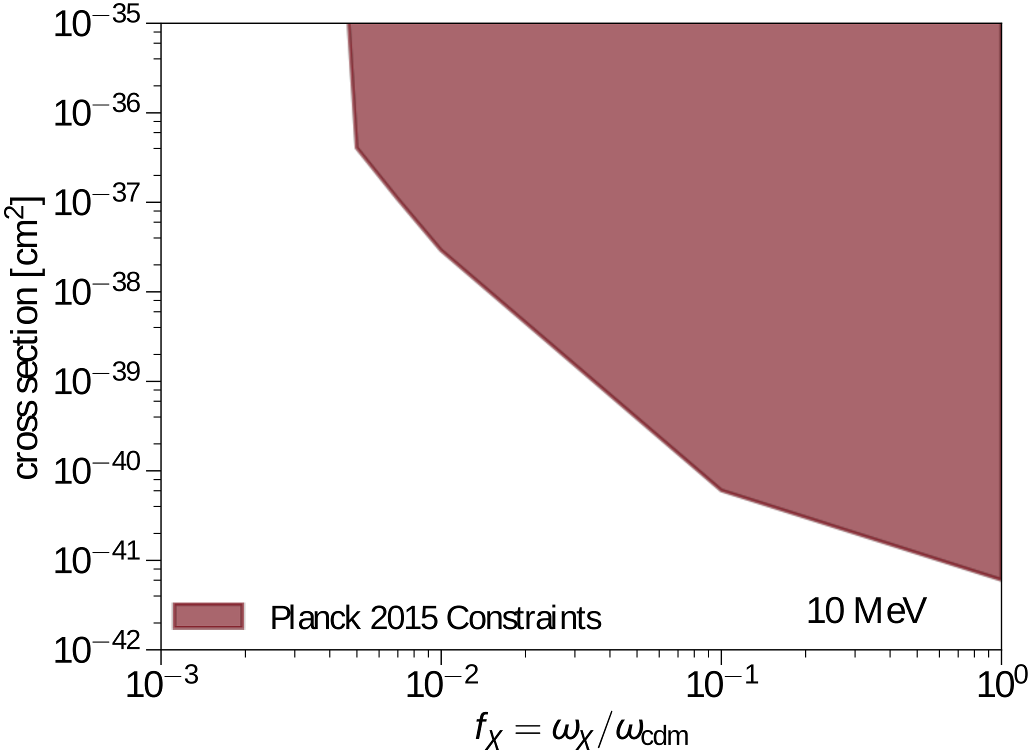
Boddy, VG, Poulin, + (coming up)



Barkana+ (2018)

Fractional DM with v^{-4} scattering

Take-home: EDGES with millicharged component at < 0.5% of DM is allowed.



Boddy, VG, Poulin, + (coming up)

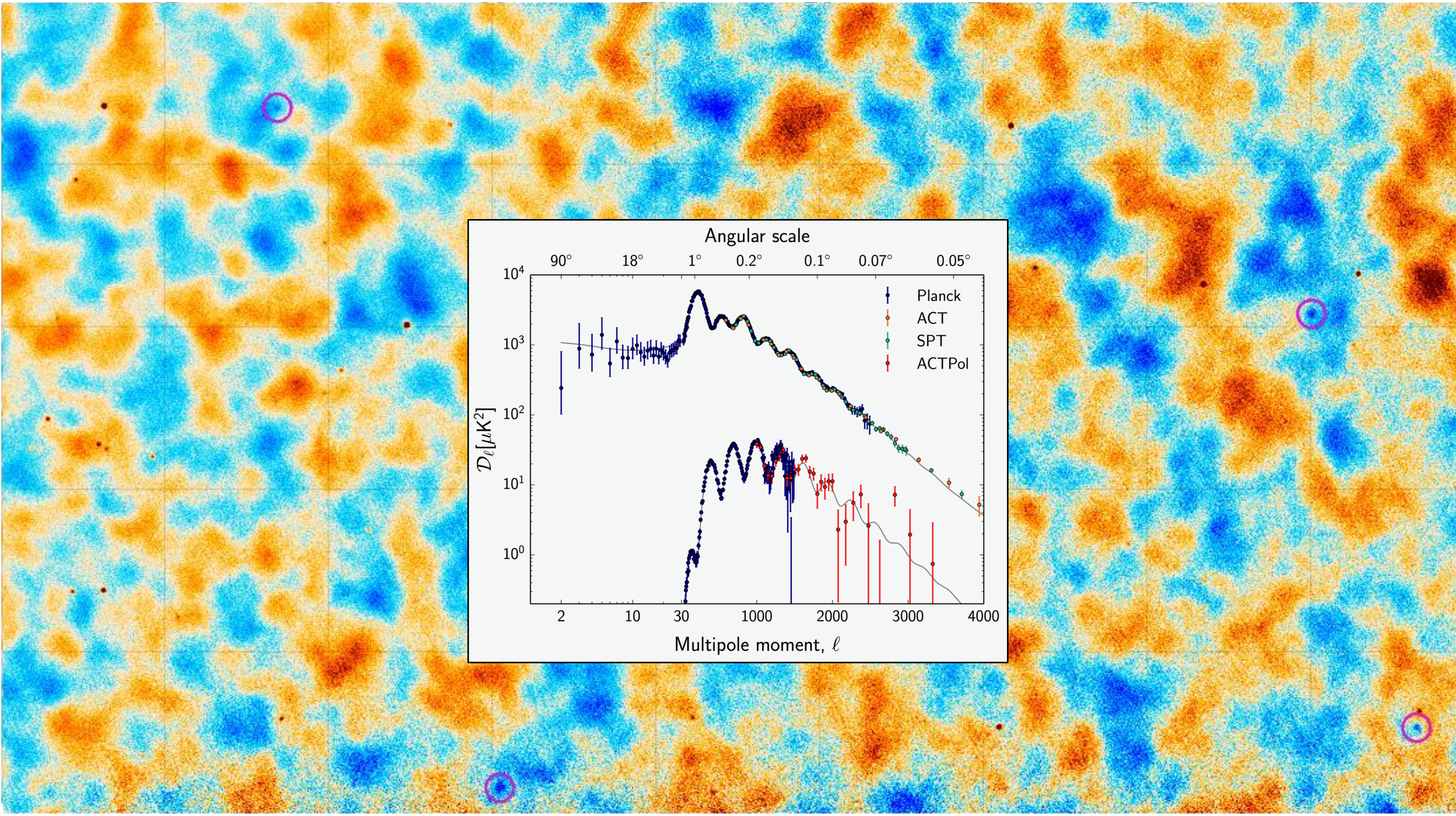
Kovetz, Poulin, VG, Boddy (coming up)

See also: Dolgov+ (2013)

What's coming?

Data

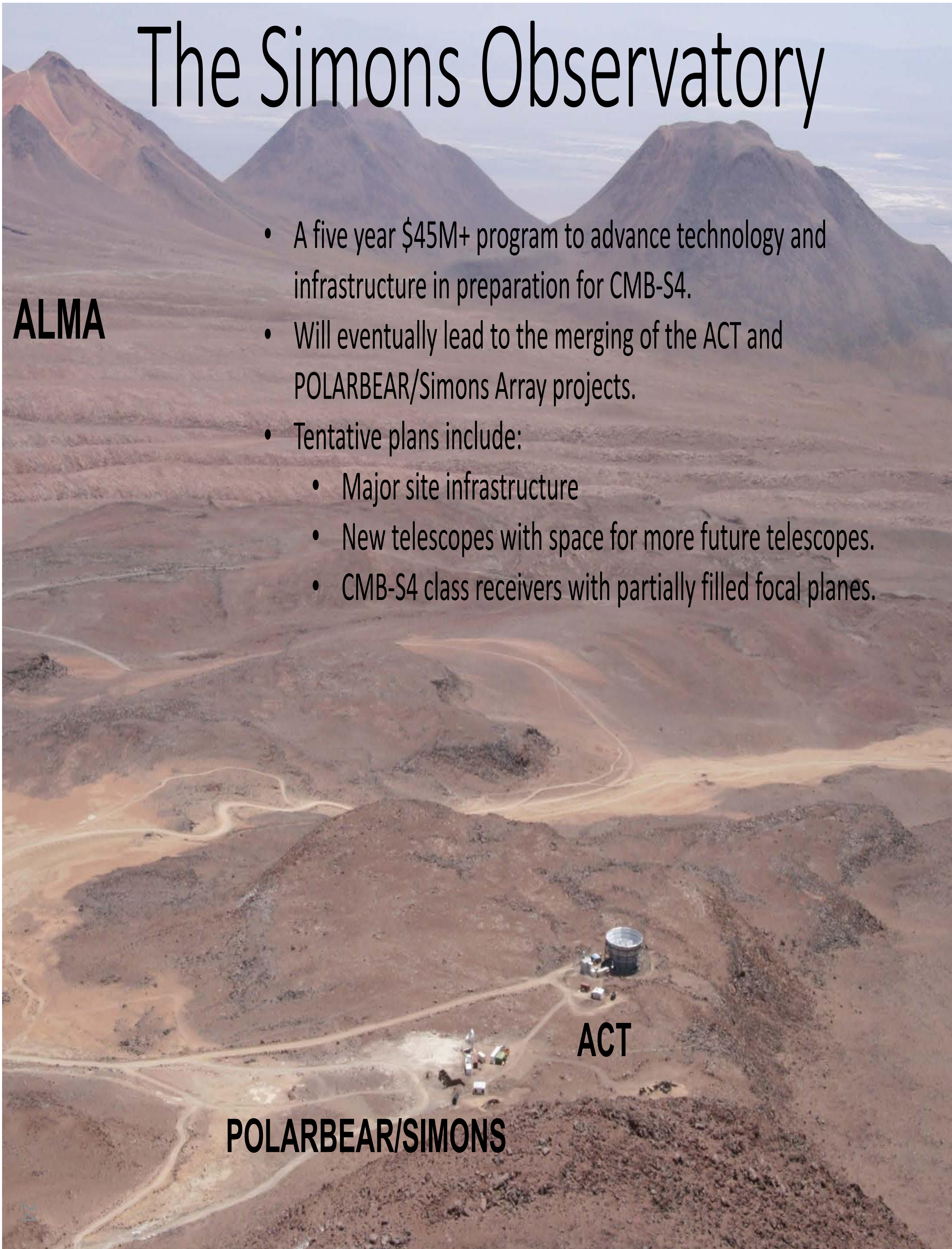
Atacama Cosmology Telescope [ACT]



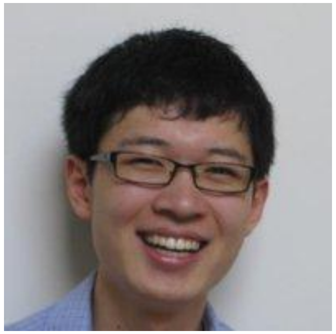
Louis et al 2016

The Simons Observatory

- A five year $45M+ program to advance technology and infrastructure in preparation for CMB-S4.
- Will eventually lead to the merging of the ACT and POLARBEAR/Simons Array projects.
- Tentative plans include:
  - Major site infrastructure
  - New telescopes with space for more future telescopes.
  - CMB-S4 class receivers with partially filled focal planes.

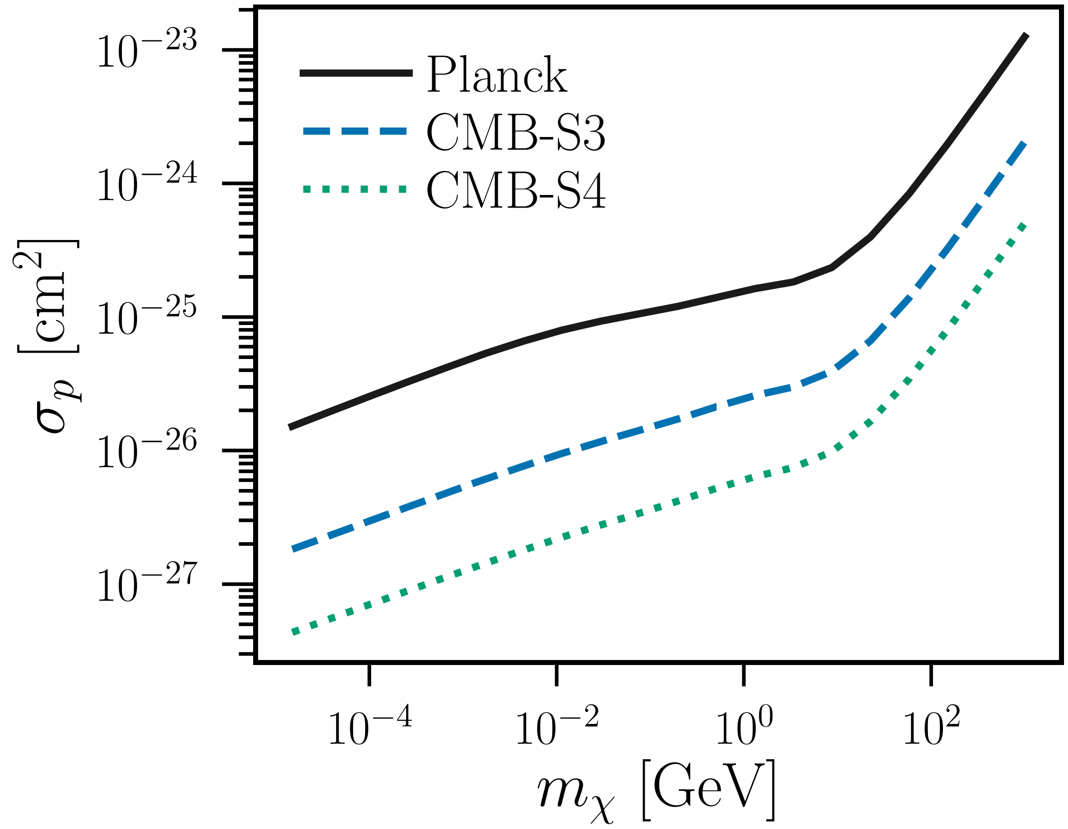


Forecasts



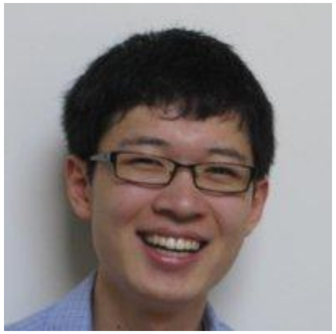
Zack Li (Princeton)

Gain 2 orders of magnitude with CMB-Stage 4.



[Li, VG, +, in prep]

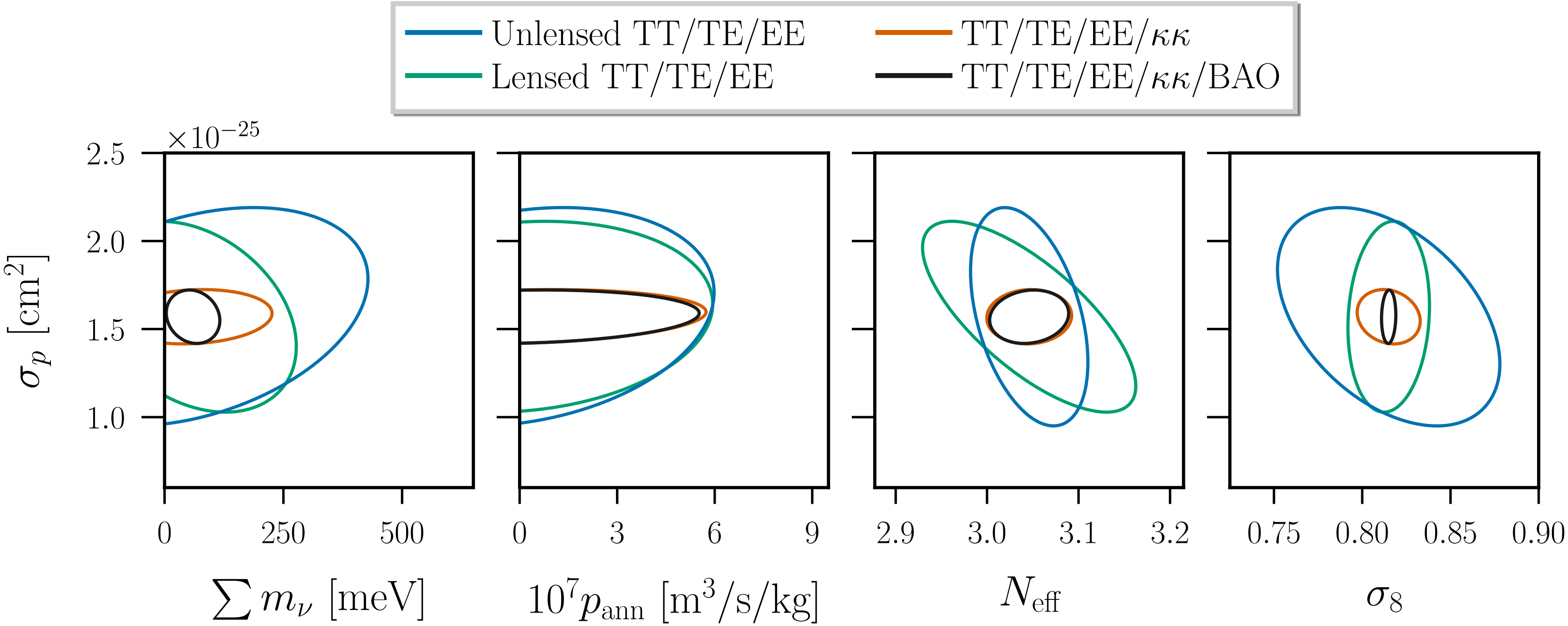
Distinguishability?



Zack Li (Princeton)

DM-baryon scattering does **NOT** look like neutrino mass, DM annihilations, Neff, nor LCDM parameters.

[Li, VG, +, in prep]



What's coming?

Analysis

Work in progress

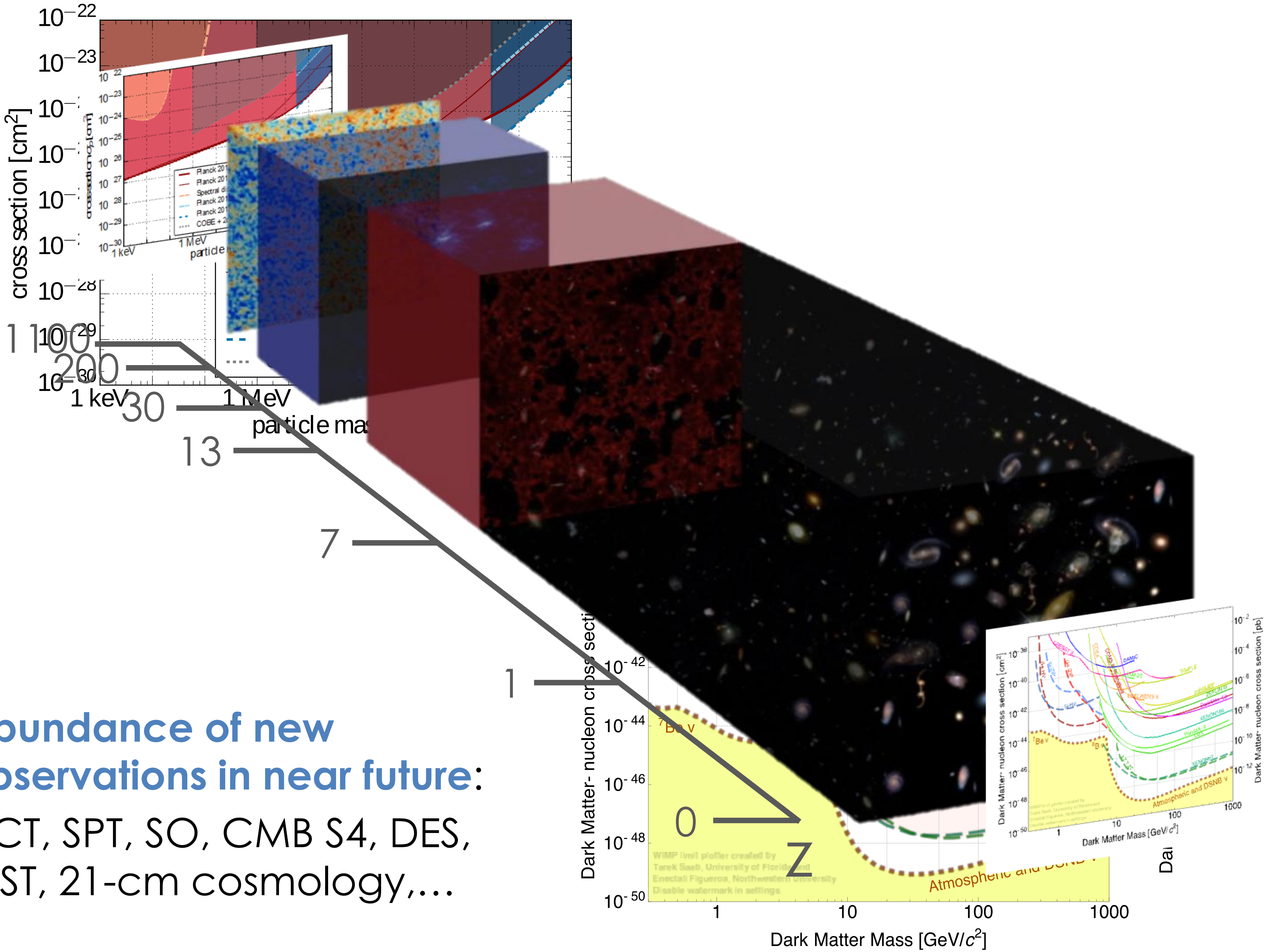
(with K. Boddy, Z. Li, M. Madhavacheril, the ACTPol collaboration)

- Cross-correlation with large-scale structure.
- Scattering with electrons (better sensitivity to lower mass).
- Specific well-motivated models.
- **Ultimate goal: combine analyses of experimental and observational data, find and confirm the signal, robustly test DM physics.**

Yes to putting it all on the same plot!

[c.f. talks by Rouven and Jocelyn]

Future

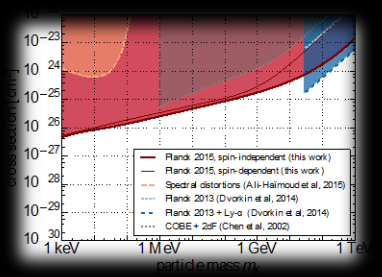


Abundance of new observations in near future:

ACT, SPT, SO, CMB S4, DES, LSST, 21-cm cosmology,…

Timeslices by E. Kovetz

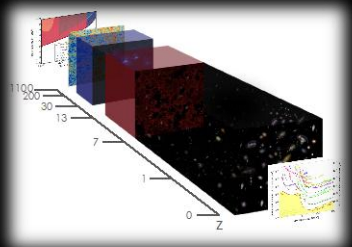
Summary



✓ CMB and cosmology probe vast parameter space (sub-GeV mass and large cross sections).

✓ Abundance of new data on the horizon:

CMB, galaxy surveys, 21-cm experiments, direct detection, LHC, fixed targets, +



✓ Synthesizing information is important to guide searches and will be essential in discovery era.