

# Status of Hot Dark Matter

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Conventional wisdom

LSND-KARMEN

Sterile neutrino (2+2 scheme)

Solar neutrino solution

## Neutrino Scenarios

Conventional wisdom today

3 light active neutrinos only

Large-mixing-angle (LMA) solar neutrino solution

Total neutrino dark matter  $\lesssim 2 \text{ eV}$  2dF Galaxy Redshift Survey

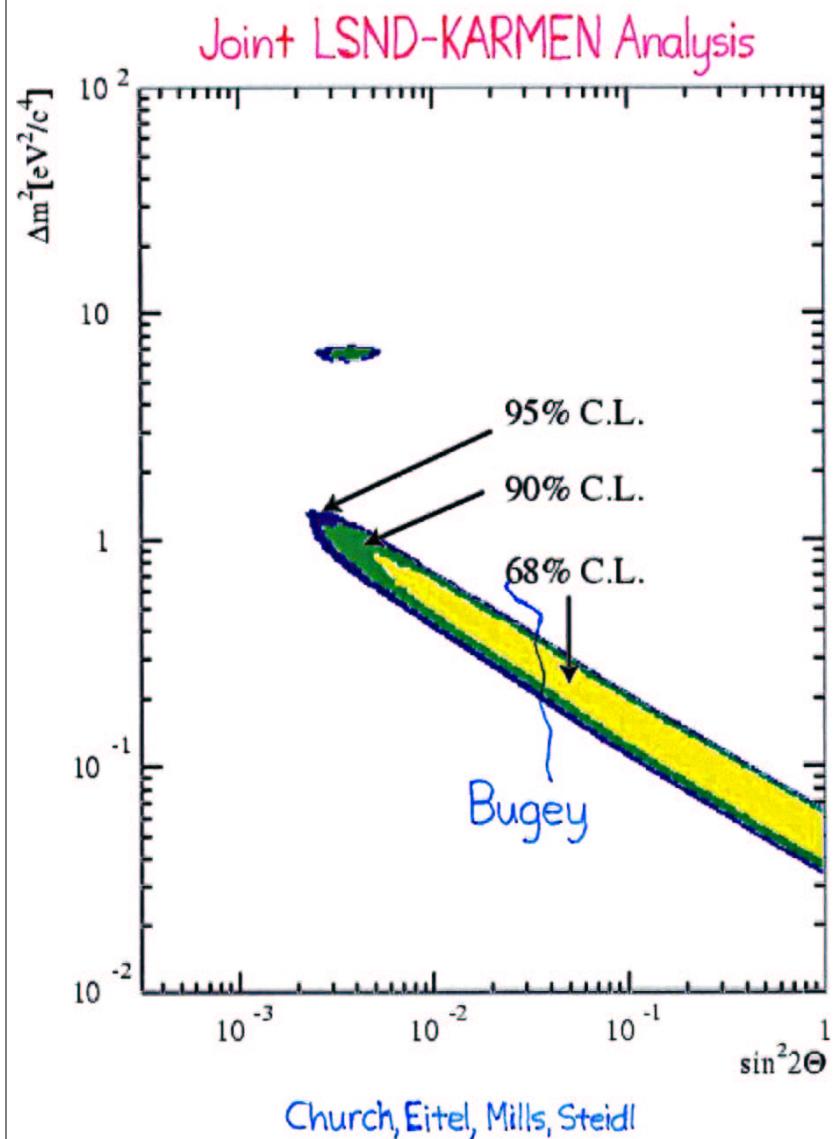
Decade ago

$3\nu$  atm.  $\overline{\nu}_e$  sol.  $\overline{\nu}_e$  (2+2)  $\overline{\nu}_s$  DOC. 1/q3  
 Also '93  $\overline{\nu}_s$   $\overline{\nu}_s$  Peltoniemi, Tommasini, Valle  
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After LSND also 3+1  $\overline{\nu}_s$  (better for WDM)

$3\nu$  HDM now a problem for Majorana neutrinos ( $\beta\beta_{0\nu}$ )

2+2 HDM enforced by LSND



### Case Against 2+2?

Negative view of LSND,  $\nu_s$ , and especially 2+2

Peres, Smirnov sum rule:  $\eta_{\odot}^s + \eta_{atm}^s = 1$  ( $\eta^s$ =fraction of  $\nu_s$ )

Gonzalez-Garcia, Maltoni, Peña-Garay

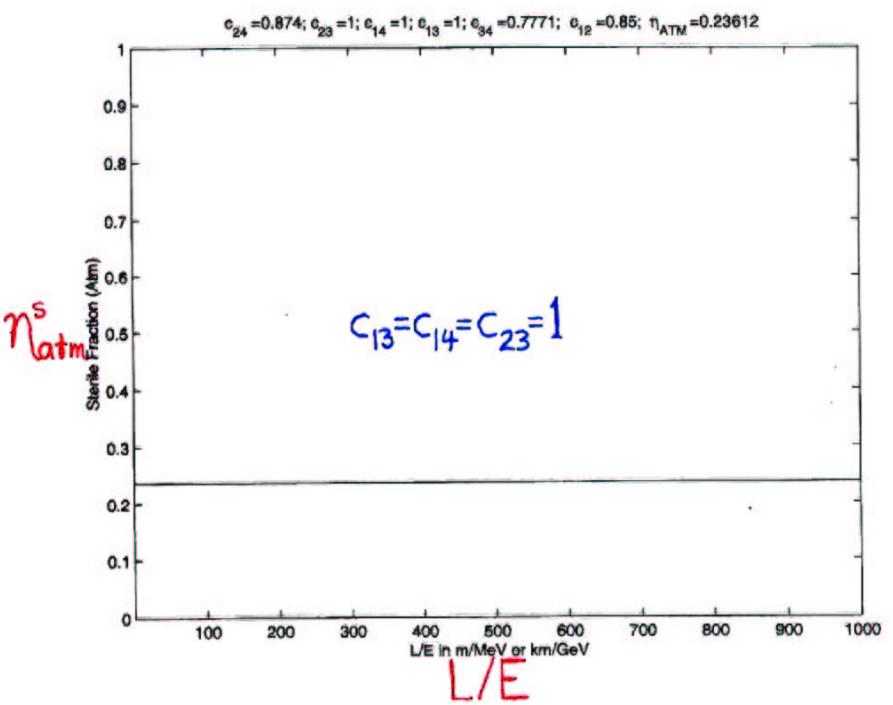
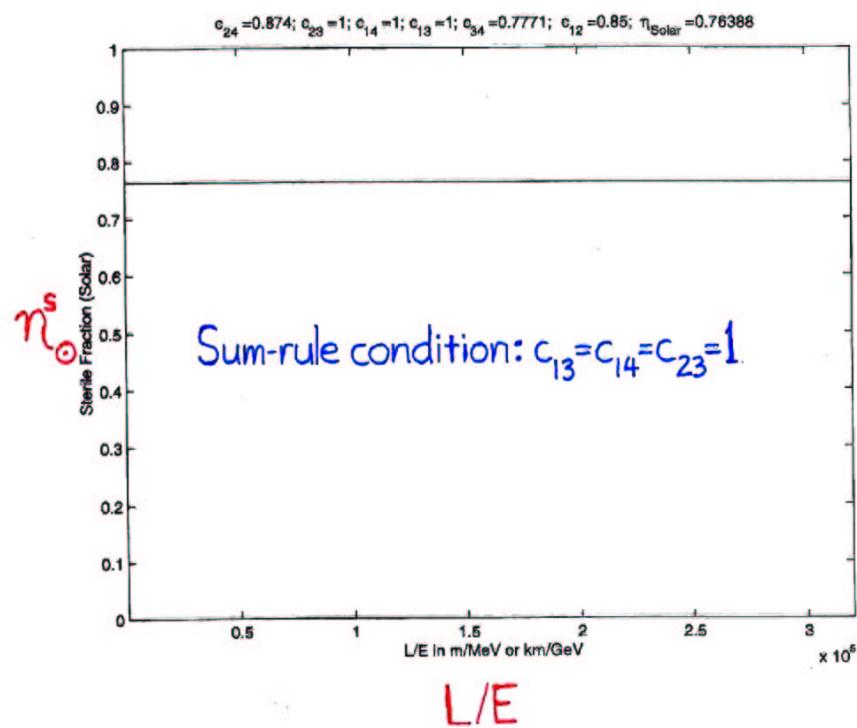
Maltoni, Schwetz, Tórtola, Valle

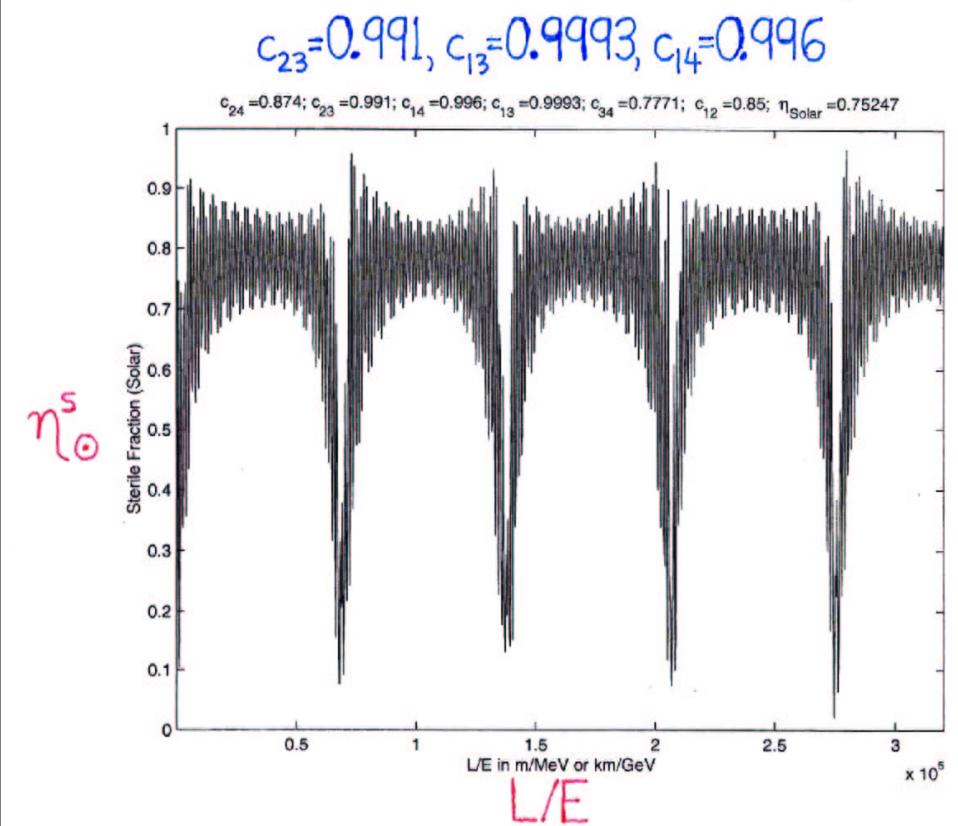
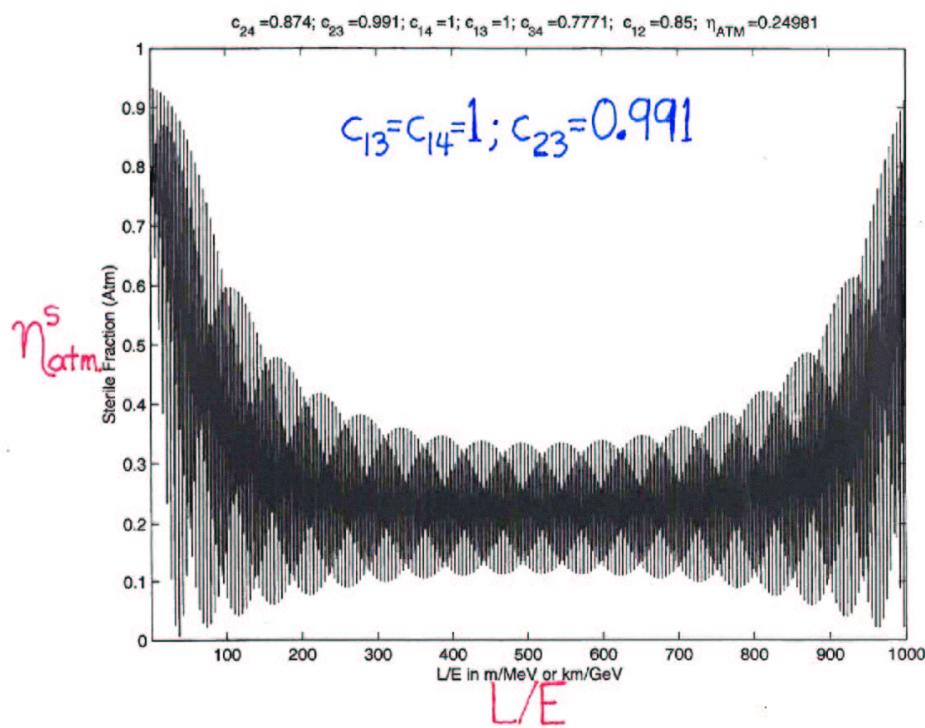
All put 4 zeros in mixing matrix; i.e., 3 angles are zero

Consequence:  $\nu_\mu$  only in atmospheric pair;  $\nu_e$  only in  $\Theta$  pair

Limits show 4 terms small, but LSND needs  $\geq 1$  non-zero

Tiny terms destroy sum rule D.O.C., Mohapatra, Yellin





## LMA solar neutrino solution

Important for preceding anti-(2+2) arguments

SNO/Super-Kamiokande  $\rightarrow$  LMA  $\rightarrow$  3  $\nu$  only

Widely believed, but is it right?

Solar neutrino flux variability

P. Sturrock, M. Weber, G. Walther, J. Scargle, M. Wheatland

## The Resonance-Spin-Flavor-Precession Solution

Invented ('88) for now-discredited solar-cycle dependence

RSFP needs neutrino magnetic moment for  $\nu_e^L \leftrightarrow \nu_i^R$  ( $i = u, \tau, s$ )

Transition moment (likely for Majorana, possible for Dirac)

Models can give  $\mu \sim 10^{-11} - 10^{-12} \mu_B$  (Standard Model,  $\mu \sim \frac{m}{eV} 10^{-19}$ )

Like MSW, resonance is a density effect:  $\Delta m^2/E = 2\beta G_F N_{\text{eff}}$

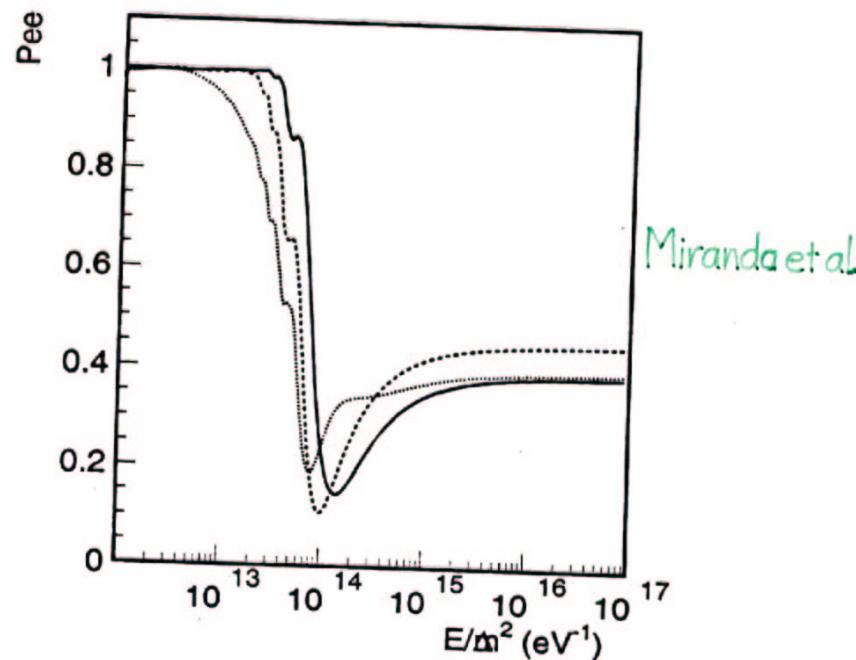
$N_{\text{eff}} = N_e - N_n$  (Majorana) or  $N_e - N_n/2$  (Dirac); here  $N_e \approx 6N_n$

Resonance adiabaticity:  $\frac{2E}{\Delta m^2} (2\mu B_1)^2 [N_{\text{eff}} (\frac{dN_{\text{eff}}}{dr})^{-1}] > 1$

Cf. MSW:  $N_{\text{eff}} = N_e / \cos 2\theta_\odot$  ( $\therefore$  it is close but at larger r)

MSW, RSFP resonances can overlap

$\nu_e^L$  Survival Probabilities for 3  $B_1$  Fields vs.  $E_\nu/\Delta m^2$



Choosing the dip at 0.86 MeV ( ${}^7\text{Be}$ ) makes  $\Delta m^2 \sim 10^{-8} \text{ eV}^2$

## Work of Sturrock and Collaborators

### Homestake

Are CI data compatible with a constant  $\nu_e$  flux?

Compared  $10^3/108$ -run simulated data sequences

Constant flux rejected at  $\geq 99.9\%$  confidence

Time-power spectrum analysis—what frequencies?

$12.88 \text{ yr}^{-1}$  (28.4 d) frequency dominant at 97% CL

$10.88, 11.88, 13.88, 14.88$  sidebands due to Sun's tilt adds conf.

$\therefore$  modulating field localized in latitude; seen directly at 98% CL

### SAGE and GALLEX-GNO

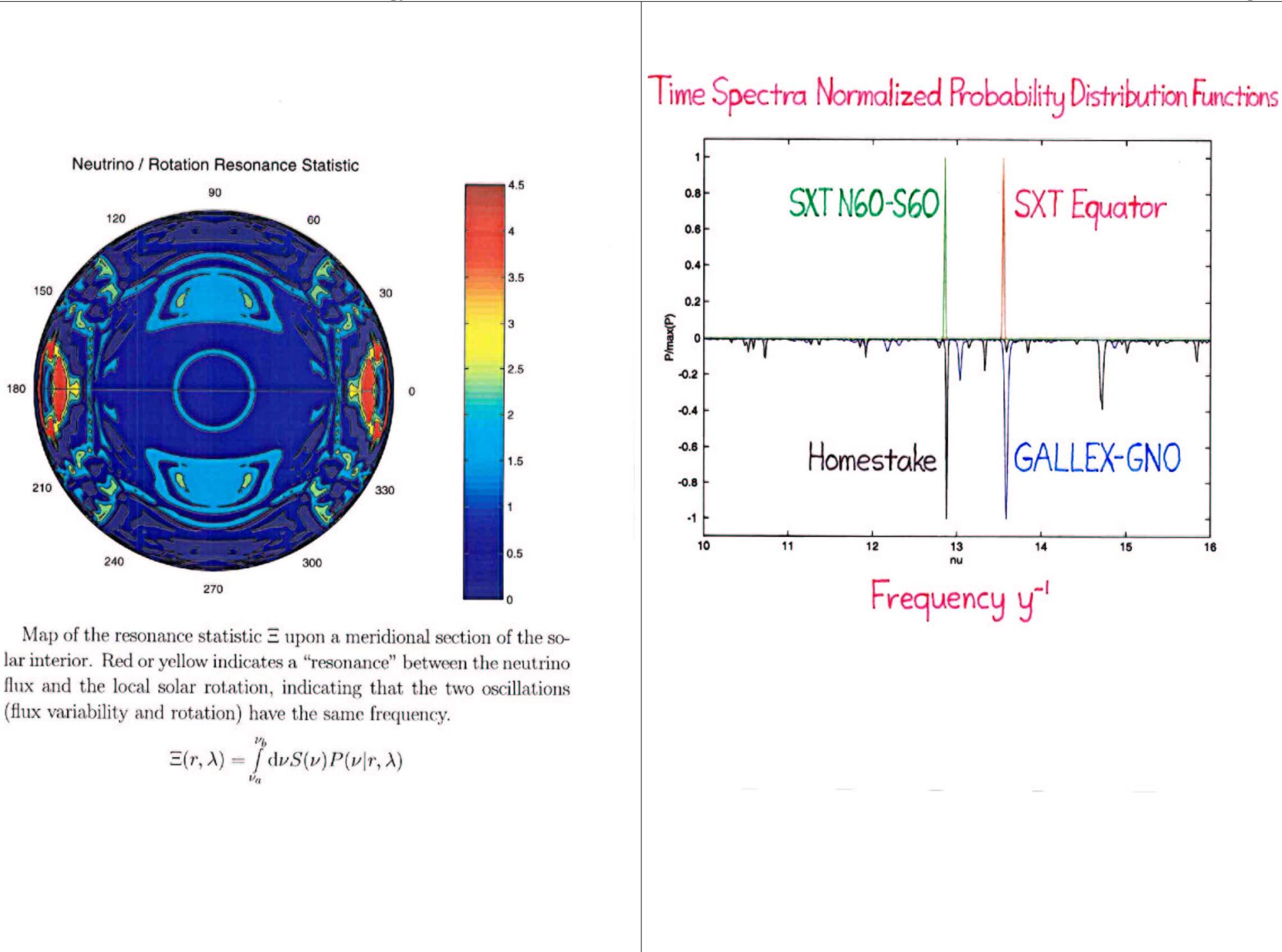
Also  $12.88$ , but  $13.59 \text{ yr}^{-1}$  (26.9 d) dominant and equatorial

Joint analysis with Homestake: stronger  $13.59$  evidence

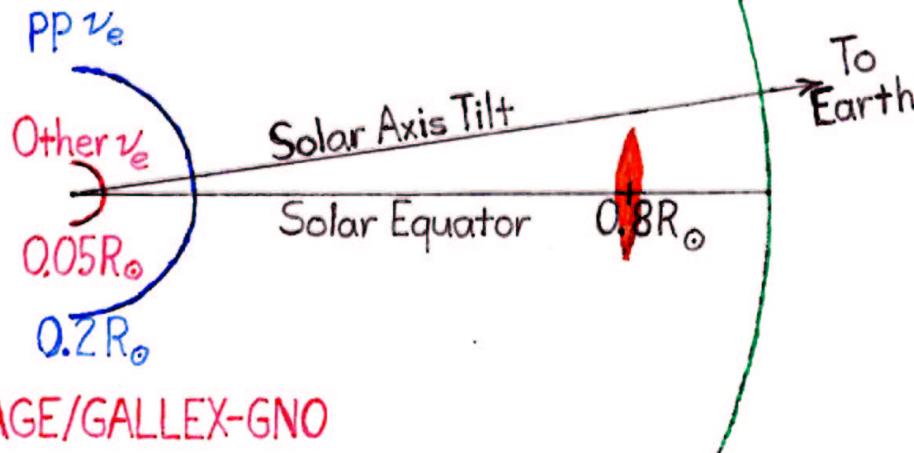
Same frequencies seen in X-rays (SXT on Yohkoh spacecraft)

$12.86 \pm 0.02 \text{ yr}^{-1}$  at high latitudes

$13.55 \pm 0.02 \text{ yr}^{-1}$  at the equator



Why Two Frequencies?



### SAGE/GALLEX-GNO

Mainly pp  $\nu_e$  ( $^7\text{Be}$  suppressed) produced at  $\sim 0.2 R_\odot$

Most  $\nu_e$  modulated by equatorial field rotation (13.6 g)

### Homestake

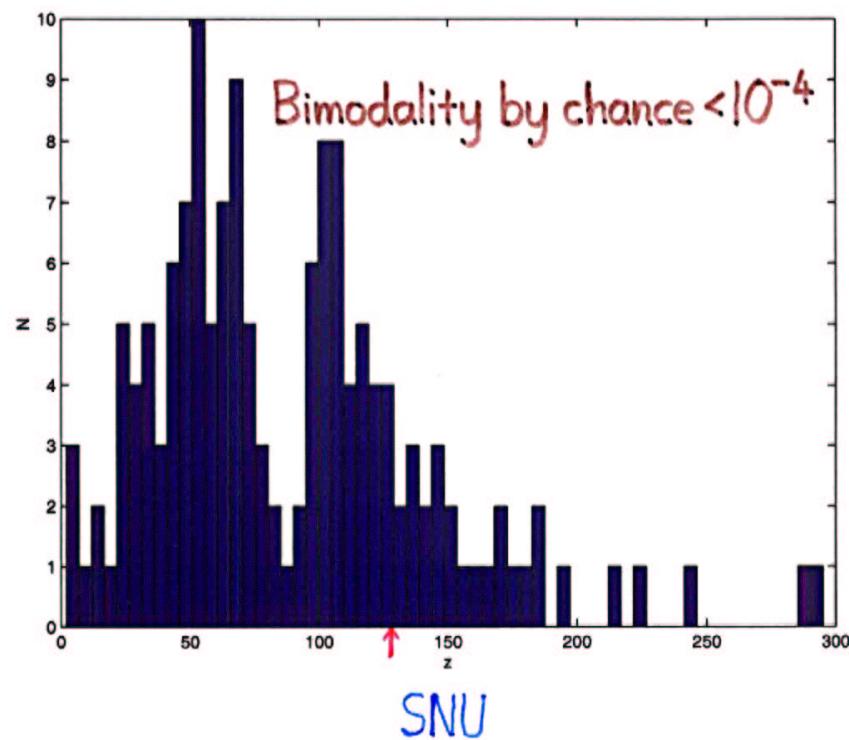
$\nu_e$  made near Sun's center at  $\sim 0.05 R_\odot$

7° axis tilt makes most  $\nu_e$  miss equatorial field

Higher latitude field modulates most  $\nu_e$  as it rotates

Get mainly  $12.9 \text{ y}^{-1}$  rate

### GALLEX-GNO+SAGE Event Distribution

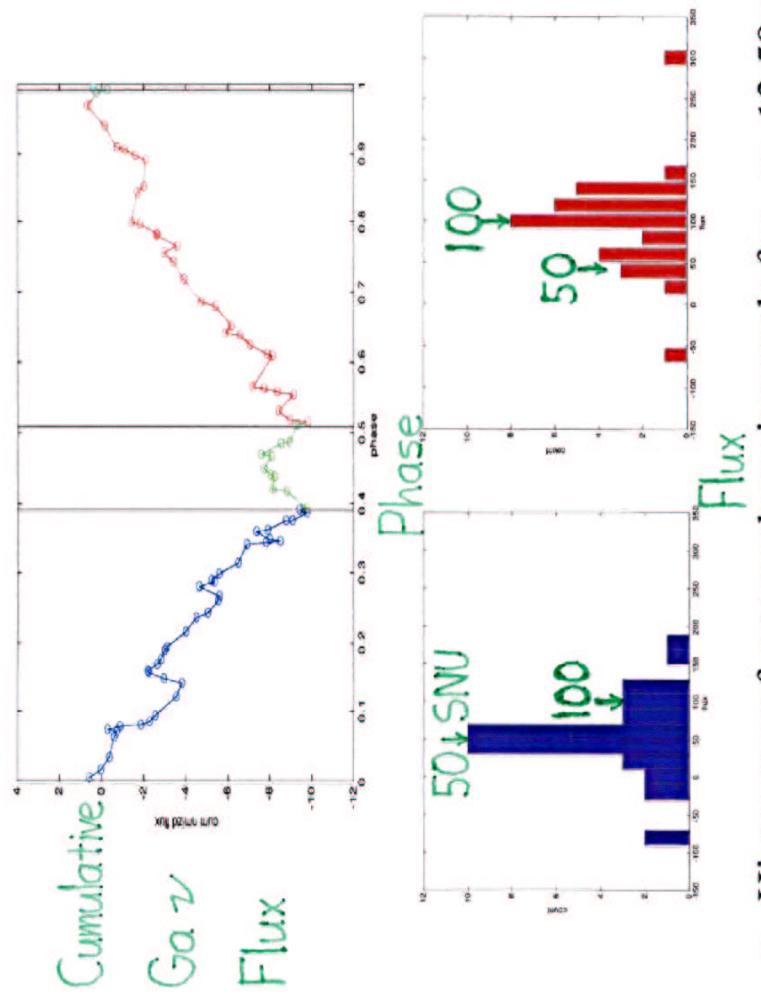


SNU

For no  $\nu$  oscillations, expect  $128^{+9}_{-7}$  SNU

Lower peak is half that; upper peak is 0.8-0.9

## Ga SOLAR NEUTRINO FLUX VARIABILITY



• Histograms for two phase intervals for  $\nu = 13.59 \text{ y}^{-1}$

### Understanding a Bimodal $\nu$ Flux Distribution

If GALLEX-GNO is correct, can the model give two peaks?

Main effect: convection-zone field modulates pp  $\nu$  at  $13.6 \text{ y}^{-1}$

Low  $B_1$  half cycle (no RSFP;  $\Delta m^2 \sim 10^{-8} \text{ eV}^2$  MSW only)

Little effect for small mixing angle ( $\sin 2\theta \sim 0.1$ )

$^7\text{Be}$ ,  $^8\text{B}$ , CNO neutrinos reduced by radiative zone  $B_1$

Overall flux reduced somewhat

High  $B_1$  half cycle

RSFP cuts into pp distribution below Ga threshold

Overall flux reduced by  $\sim 1/2$

### Location of the Resonance

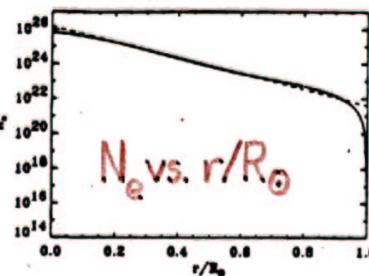
SOHO/MDI helioseismology convoluted with GALLEX-GNO data

SOHO/MDI( $\nu, \lambda$ ) matching Ga( $\nu$ )

Near equator at  $r=0.8R_{\odot}$

Locating  $\nu=13.6\text{ yr}^{-1}$  determines  $\frac{\Delta m^2}{E}$

$$\frac{\Delta m^2}{E} = 2\sqrt{2}G_F(N_e - N_n) = 1 \times 10^{-14}\text{ eV}$$



Recall  $\nu_e$  survival probability for RSFP fit

Exactly the  $\Delta m^2/E$  needed

Exponential ( $N_e - N_n$ ) vs.  $r$  could have given very different  $\frac{\Delta m^2}{E}$

RSFP resonance for  $12.9\text{ yr}^{-1}$  must be at  $r=0.8R_{\odot}$  also

Higher latitudes, however

Either radiative-zone field or latitudinal wave

Recall  $12.9, 13.6\text{ yr}^{-1}$  frequencies seen out to corona