

N-body Halo Models and Dark Matter Direct Detection

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With: Neal Weiner

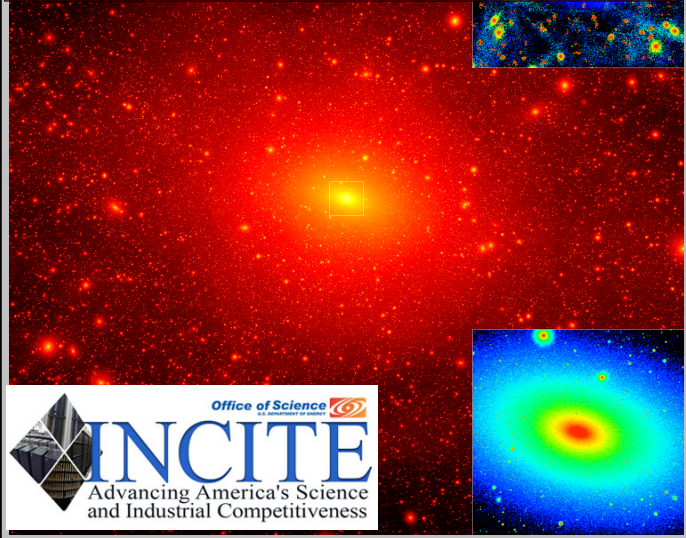
P. Madau, J. Diemand, M. Zemp, B. Moore, J. Stadel, D. Potter (Simulations: Via Lactea, GHALO)

The State-of-the-art in Cosmological DM-only N-Body Simulations

VIA LACTEA II (Diemand et al. 2008)

1.1 billion particles

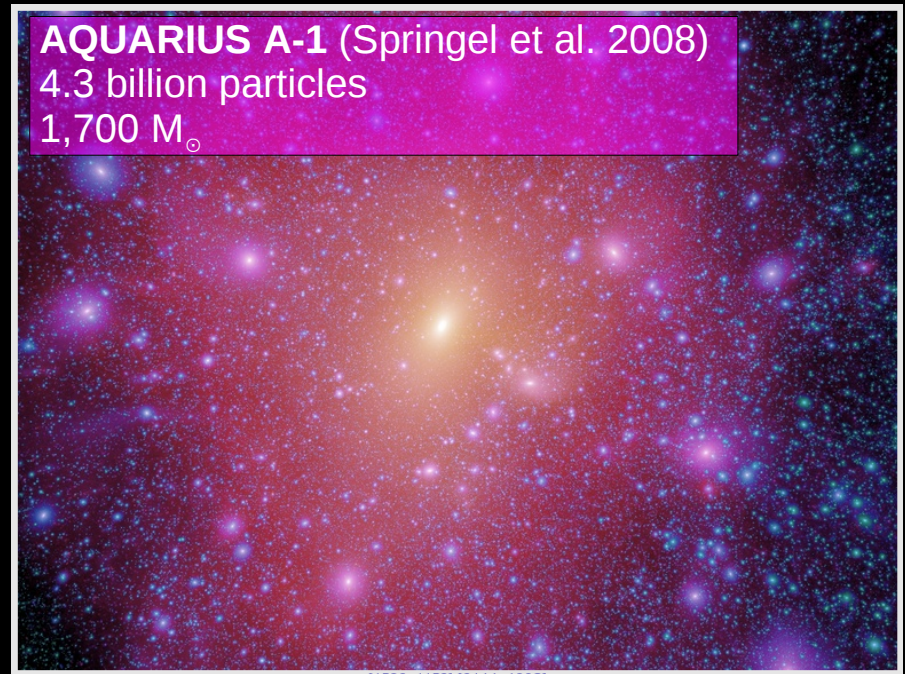
4,000 M_{\odot}



AQUARIUS A-1 (Springel et al. 2008)

4.3 billion particles

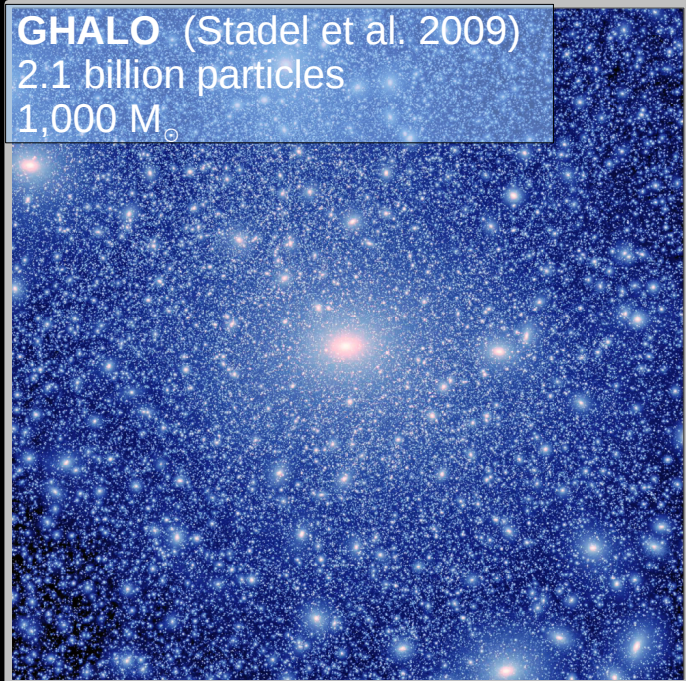
1,700 M_{\odot}



GHALO (Stadel et al. 2009)

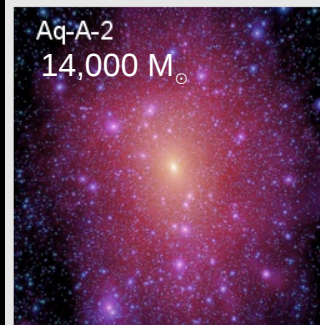
2.1 billion particles

1,000 M_{\odot}



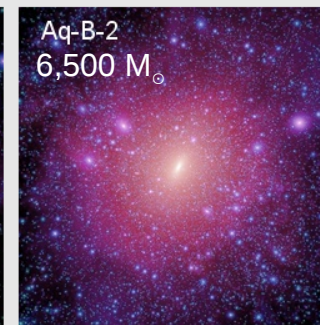
Aq-A-2

14,000 M_{\odot}



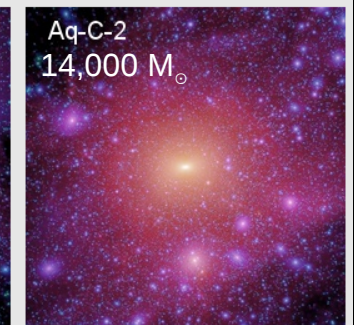
Aq-B-2

6,500 M_{\odot}



Aq-C-2

14,000 M_{\odot}



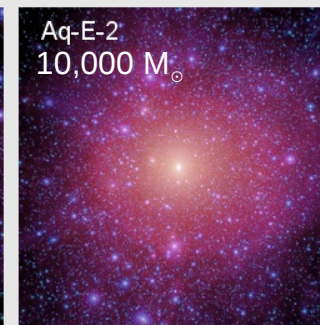
Aq-D-2

14,000 M_{\odot}



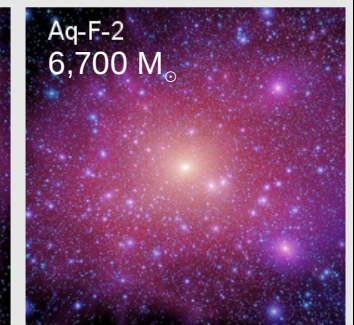
Aq-E-2

10,000 M_{\odot}

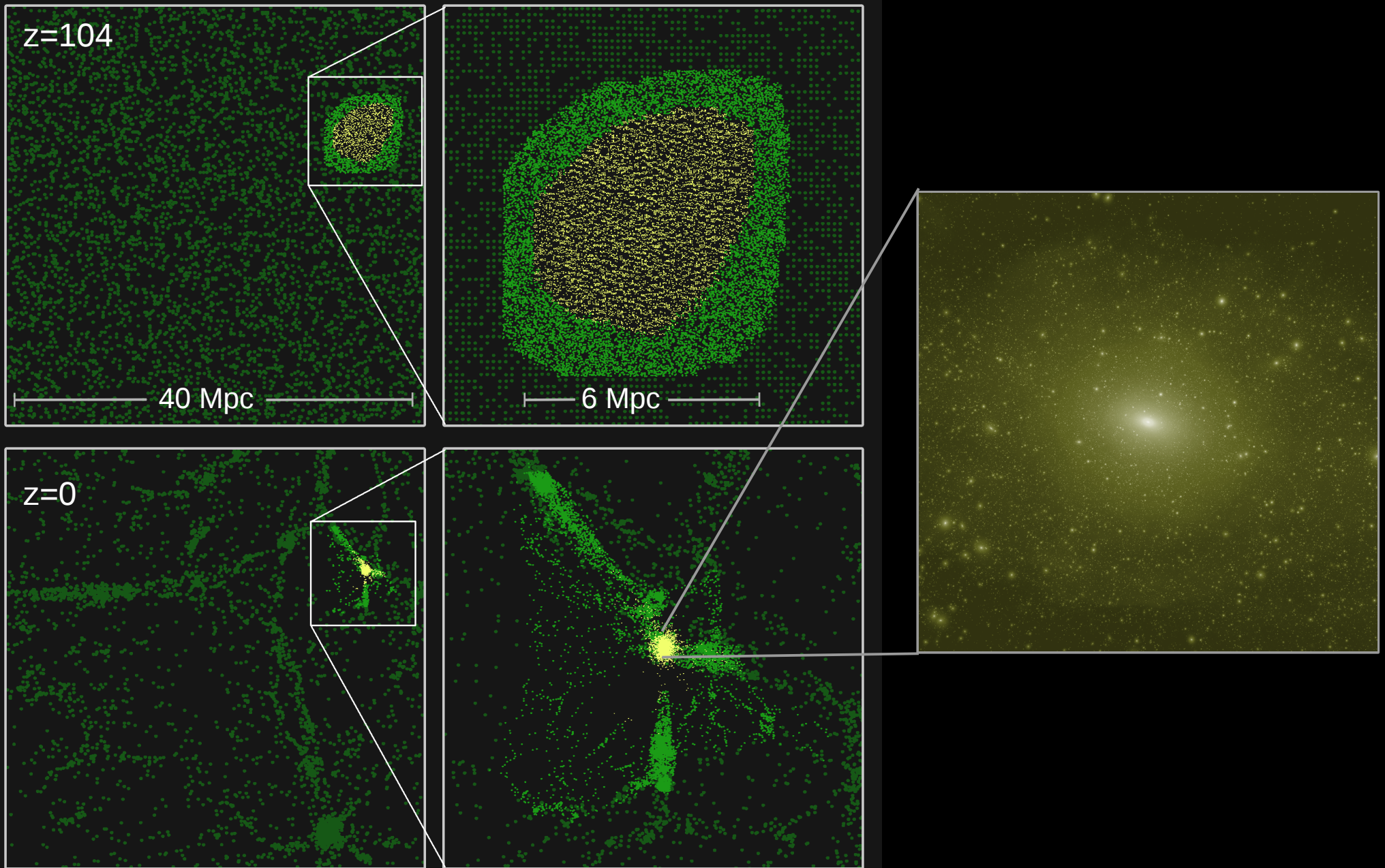


Aq-F-2

6,700 M_{\odot}

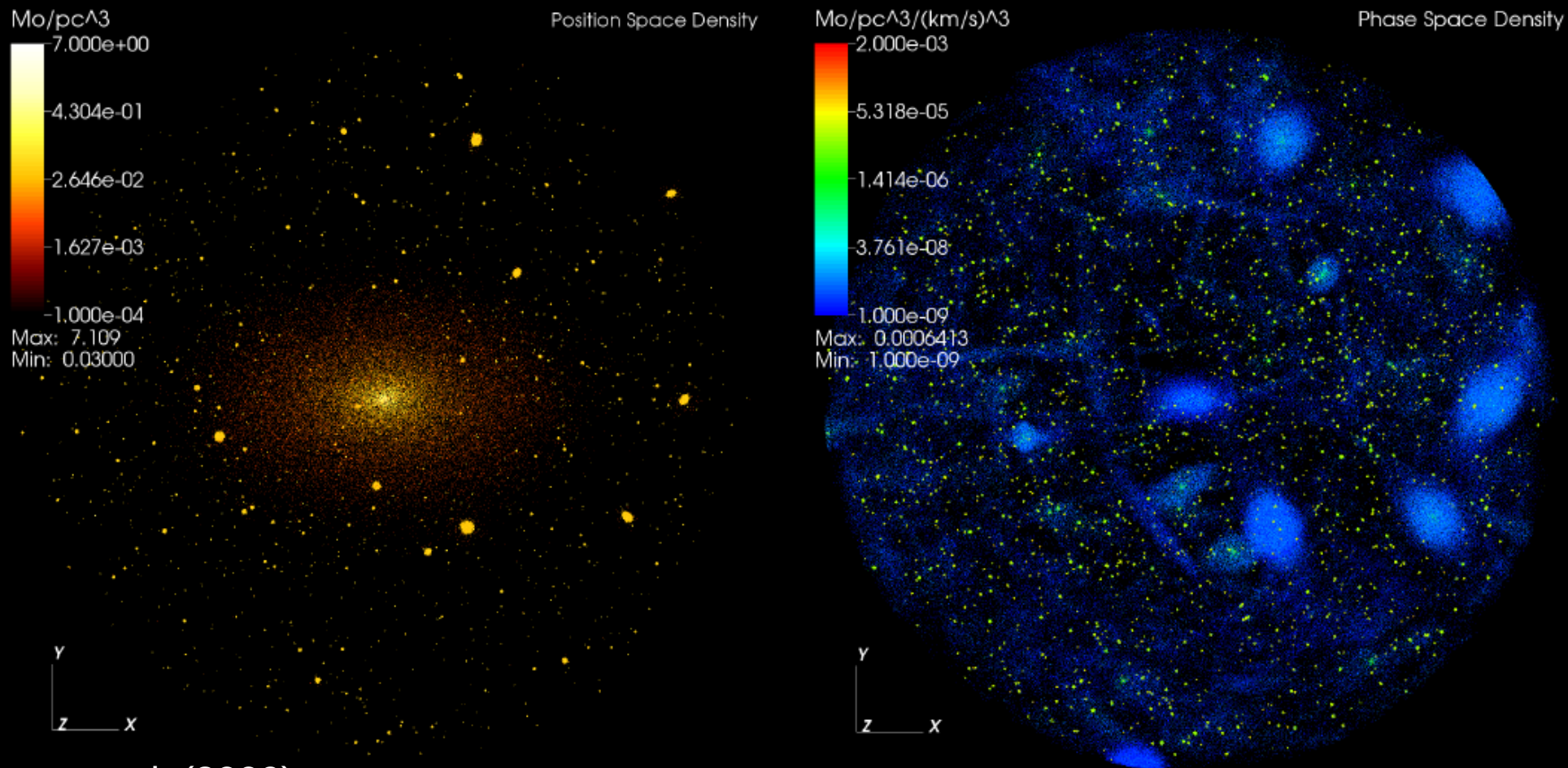


Multi-mass initial conditions



Via Lactea II – the inner 100 kpc

Whereas previous simulations were almost completely smooth in the central region, with VL-II we resolve lots of subhalos and tidal streams even down to 8 kpc.



Zemp et al. (2009)

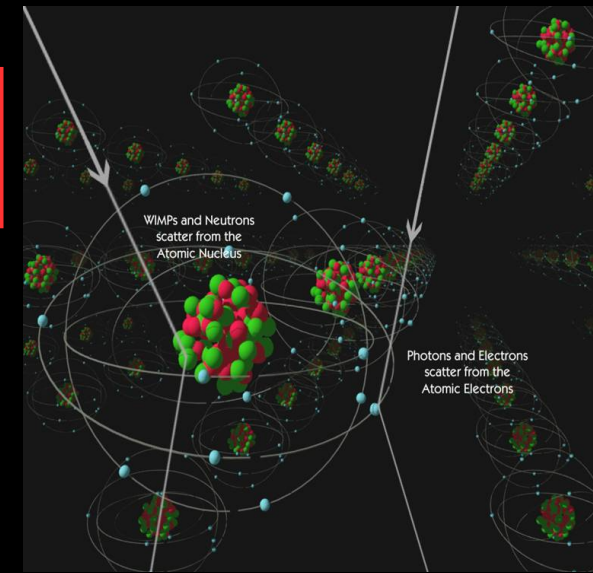
Direct Detection

The scattering event rate (events/recoil energy) is given by:

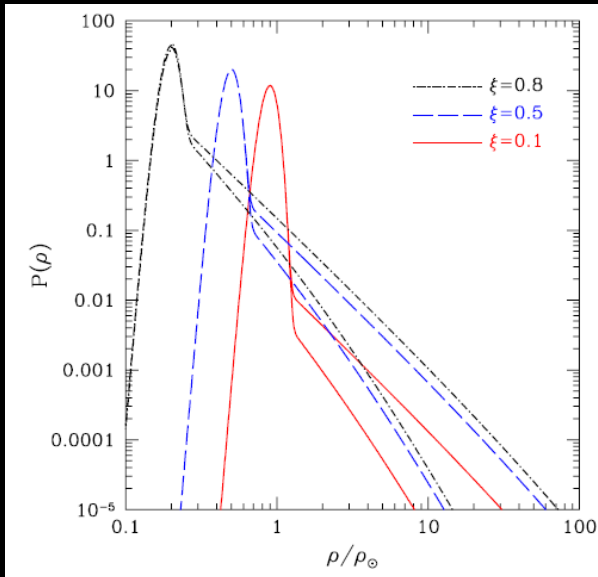
$$\frac{dR}{dE_R} = N_T M_N \frac{\rho_\chi \sigma_n}{2m_\chi \mu_{ne}^2} \frac{(f_p Z + f_n (A - Z))^2}{f_n^2} F^2[E_R] \int_{\beta_{min}}^{\infty} \frac{f(v)}{v} dv$$

This depends on the **local DM density** ρ_χ and the **velocity distribution function** $f(v)$.

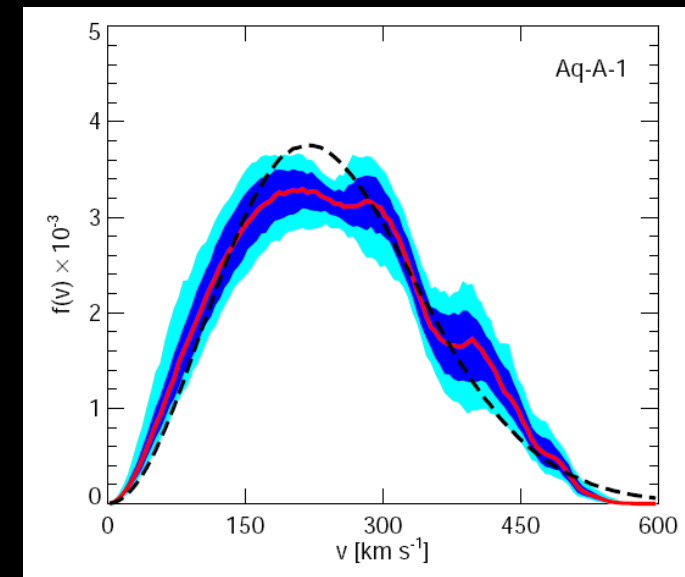
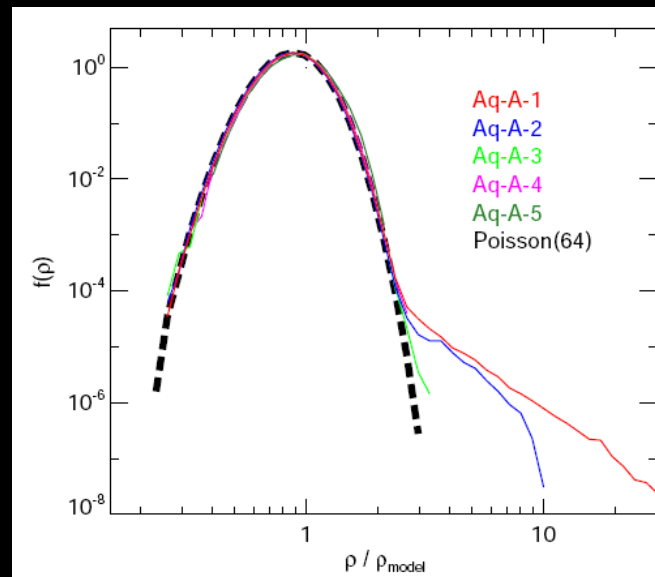
A typical assumption is $\rho_\chi = 0.3 \text{ GeV/cm}^3$ and $f(v)$ a Maxwellian with $v_p = 220 \text{ km/s}$ truncated at an escape speed of 500-600 km/s.



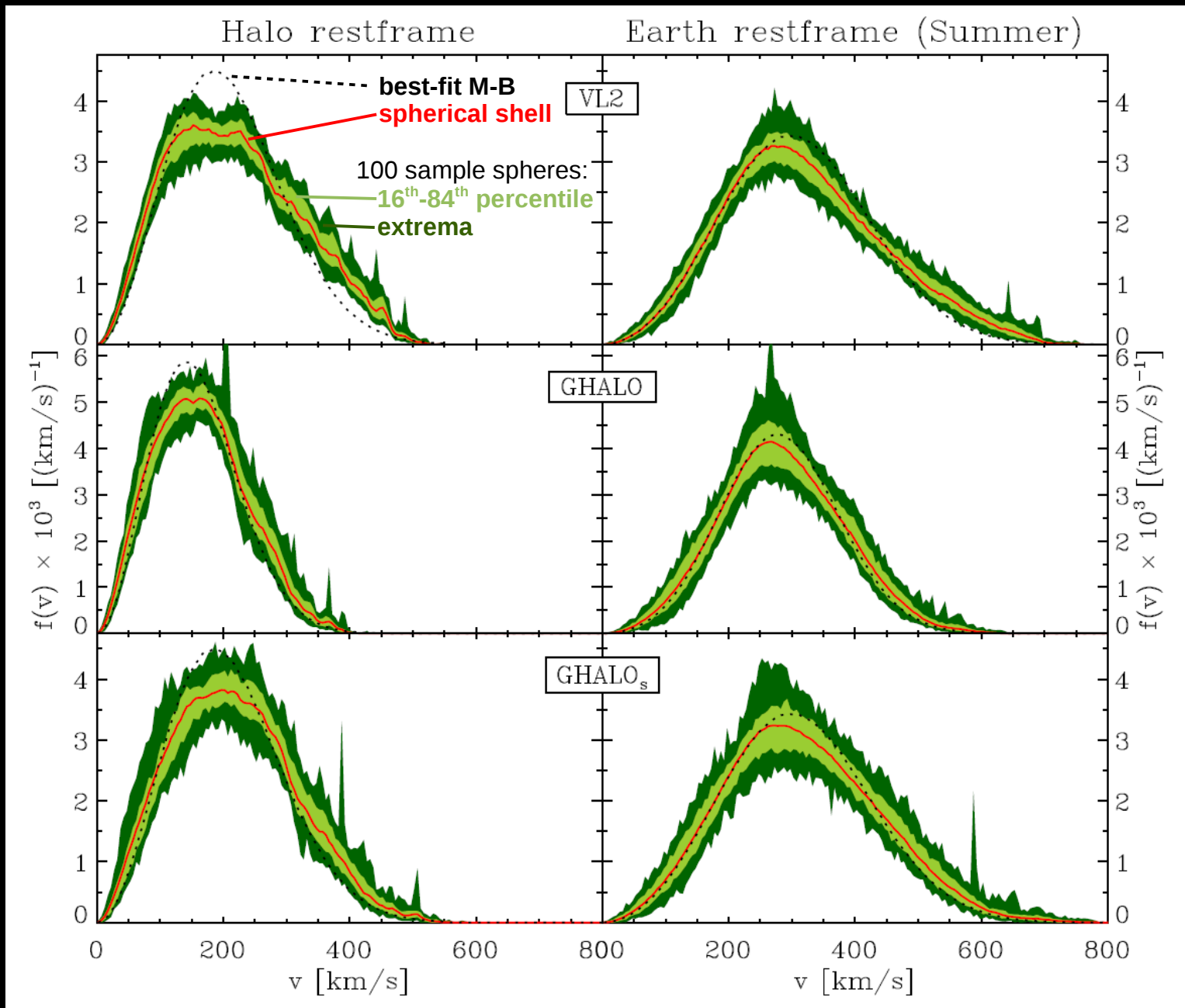
Kamionkowski & Koushiappas (2008)



Vogelsberger et al. (2009) Aquarius Simulations

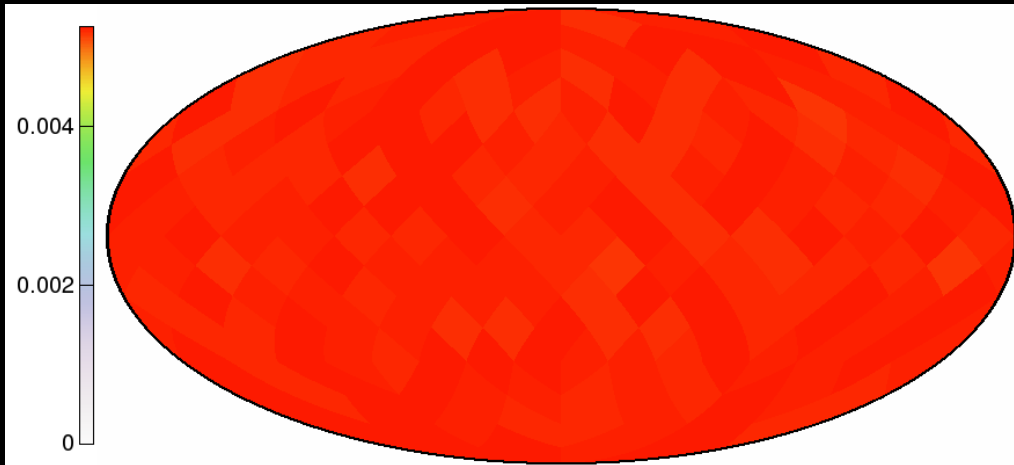


Velocity Distribution in Via Lactea/GHALO

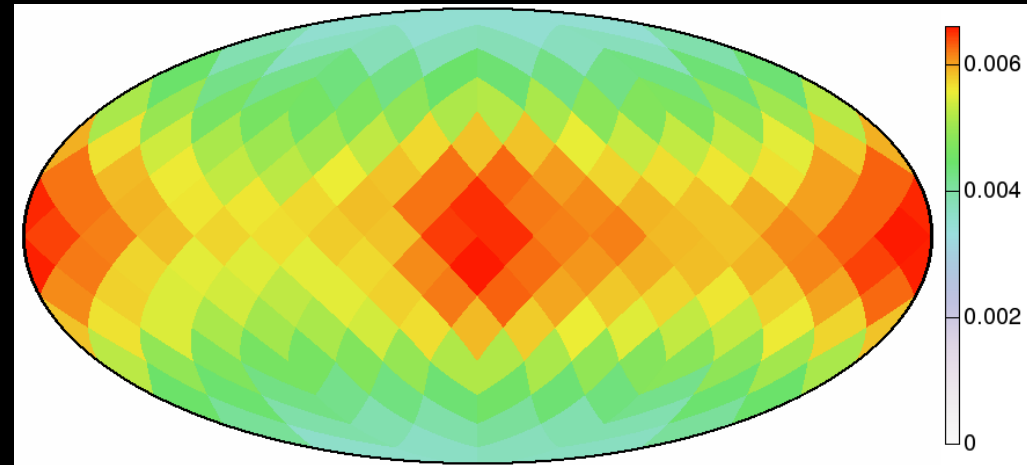


Velocity Direction in Halo Rest Frame

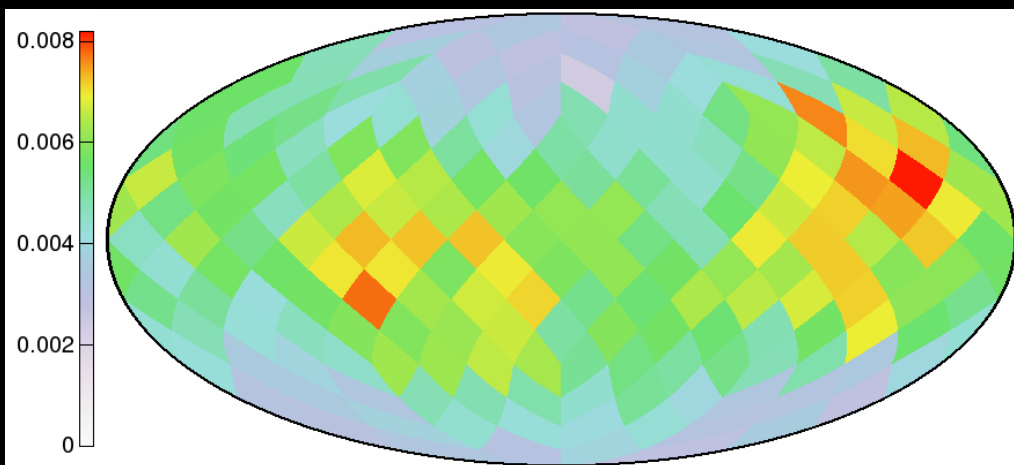
Maxwell-Boltzmann (isotropic)



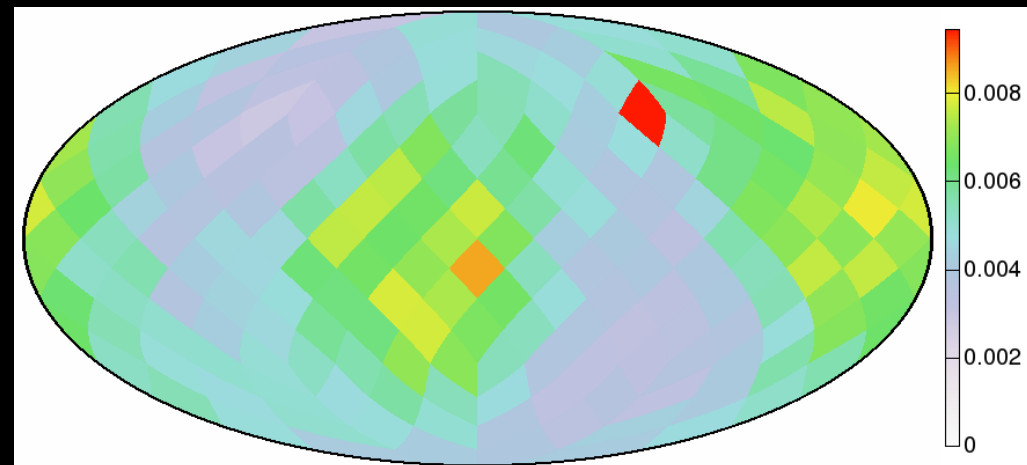
Spherical Shell ($8 \text{ kpc} < R < 9 \text{ kpc}$)



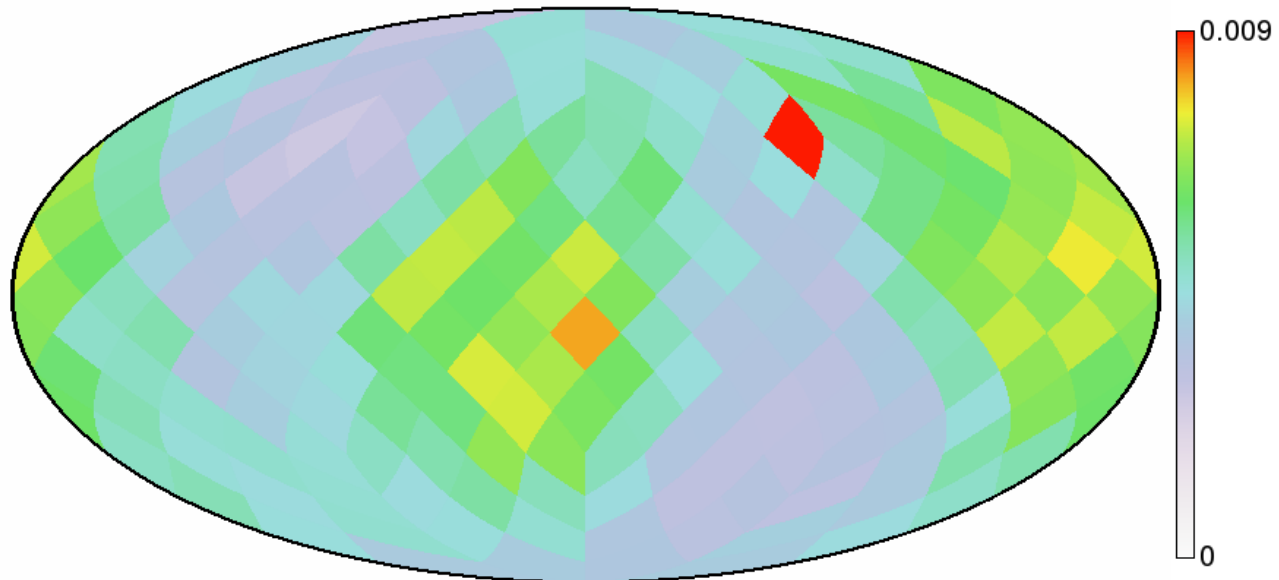
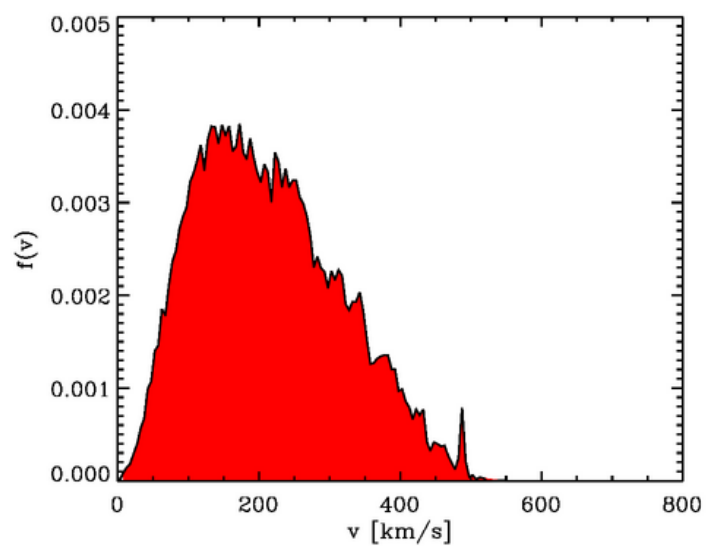
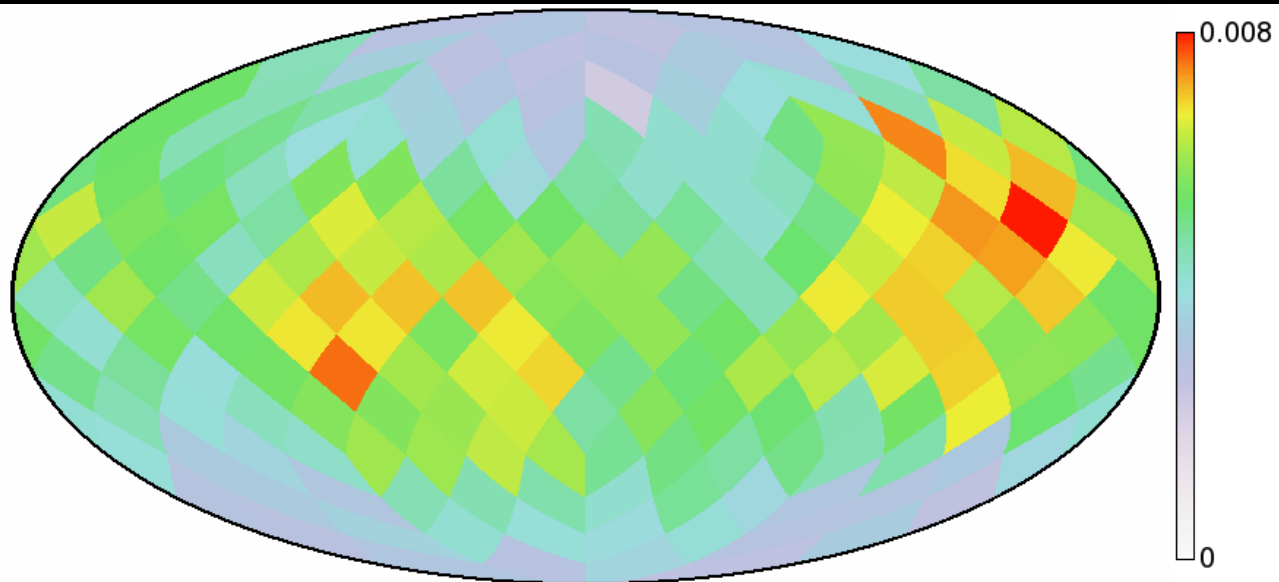
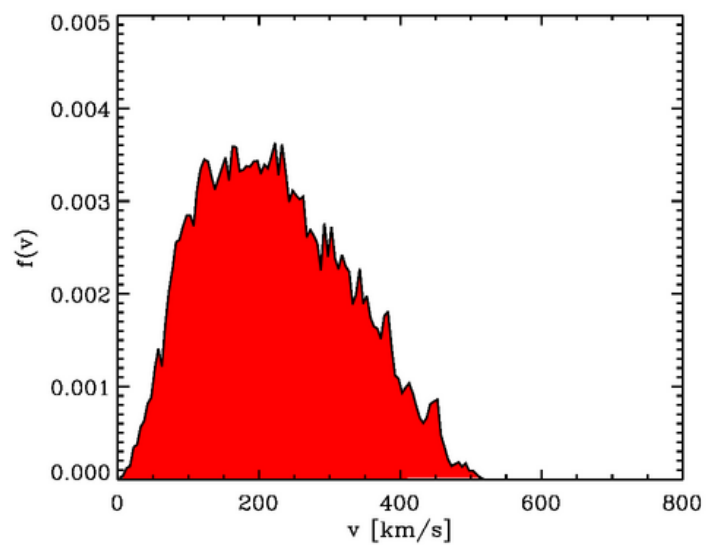
Sample Sphere #001



Sample Sphere #004 (containing a subhalo)

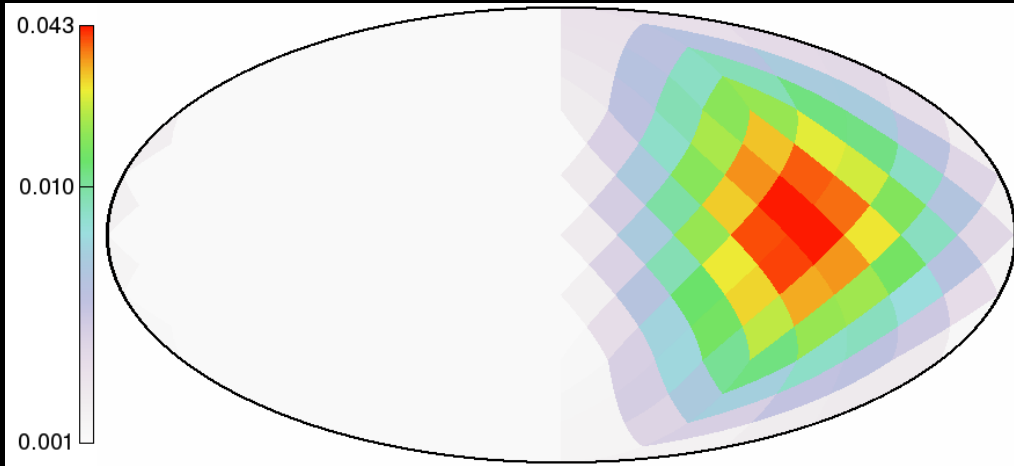


Velocity Direction in Halo Rest Frame

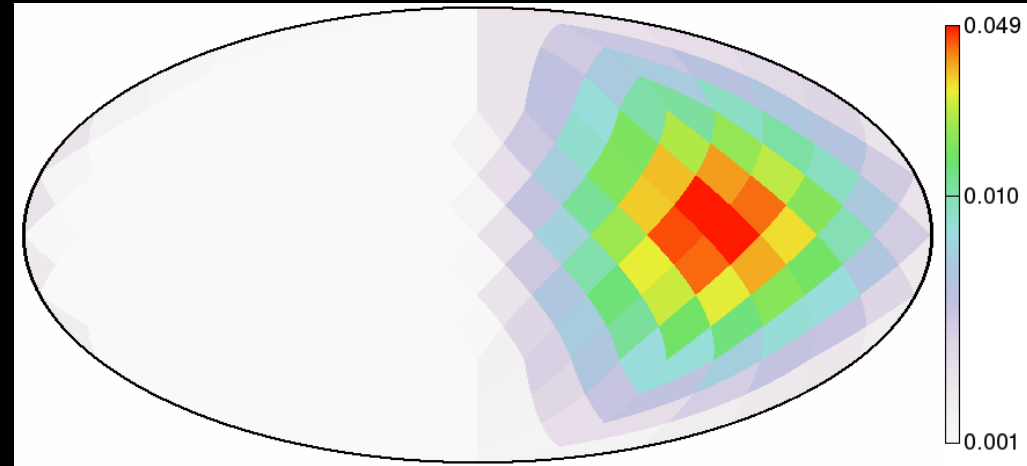


... in Earth Rest Frame $v_{\min} = 0$ km/s

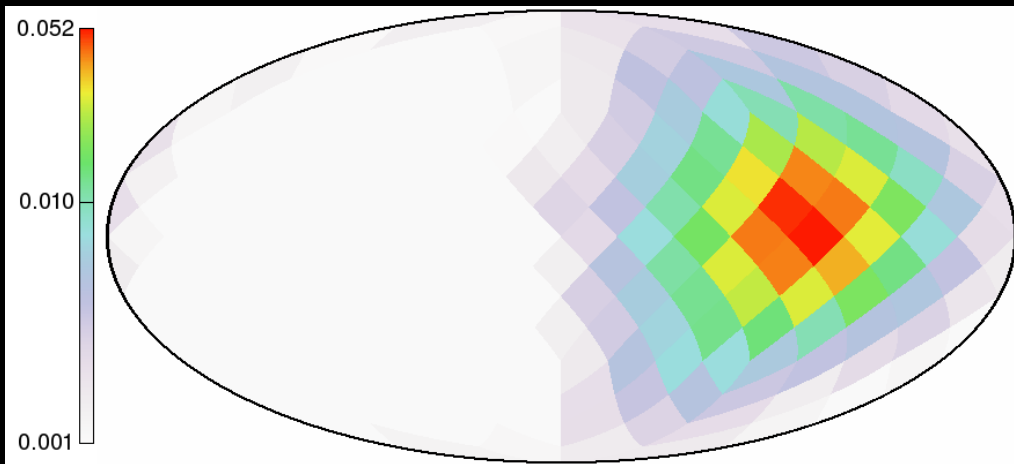
Maxwell-Boltzmann (isotropic)



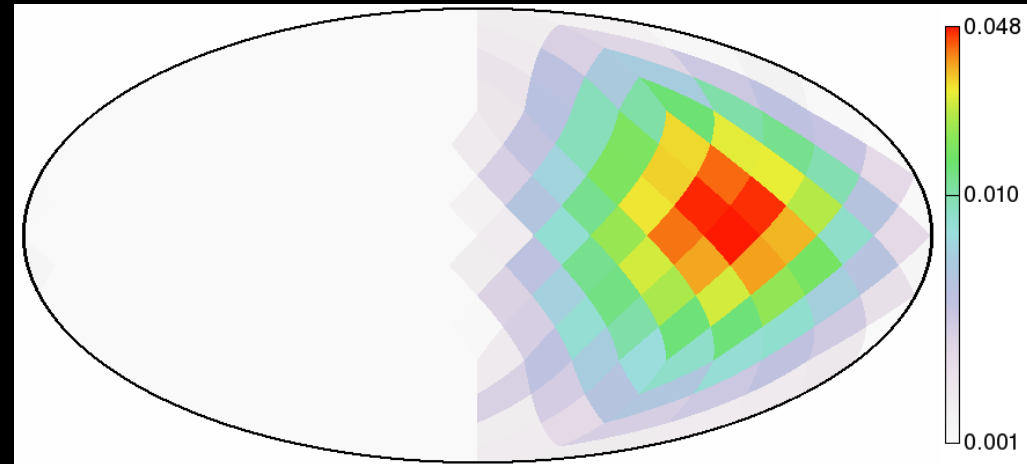
Spherical Shell (8 kpc < R < 9 kpc)



Sample Sphere #001

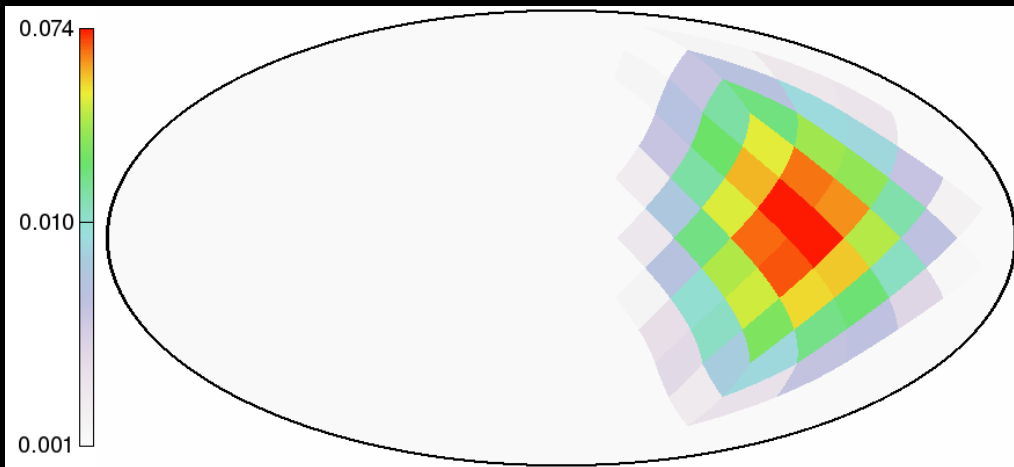


Sample Sphere #004 (containing a subhalo)

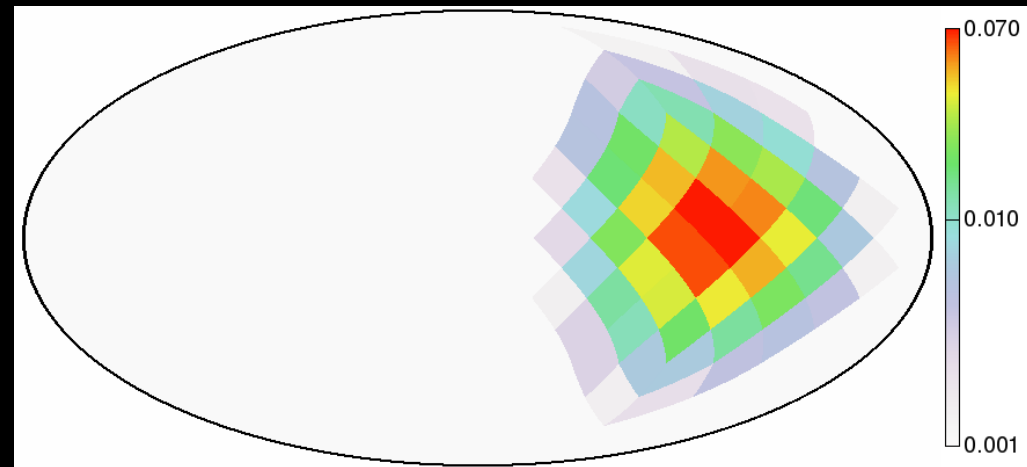


... in Earth Rest Frame $v_{\min} = 500$ km/s

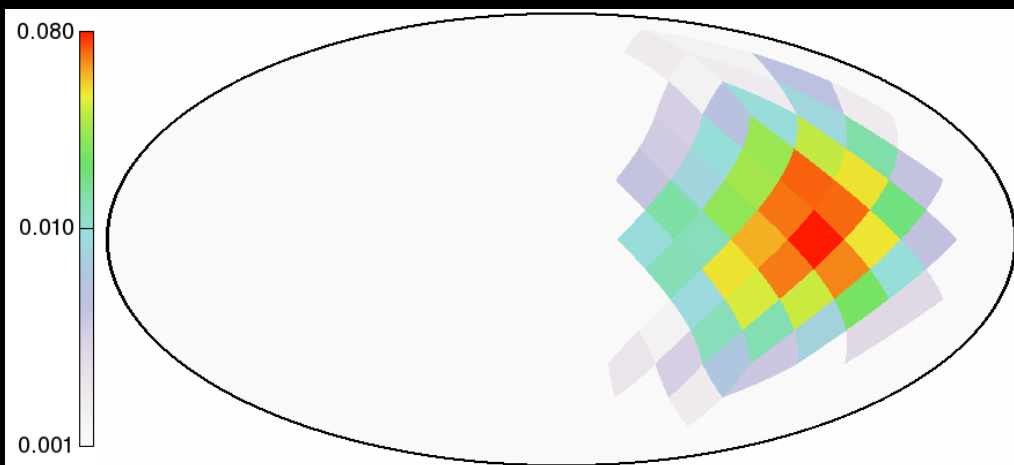
Maxwell-Boltzmann (isotropic)



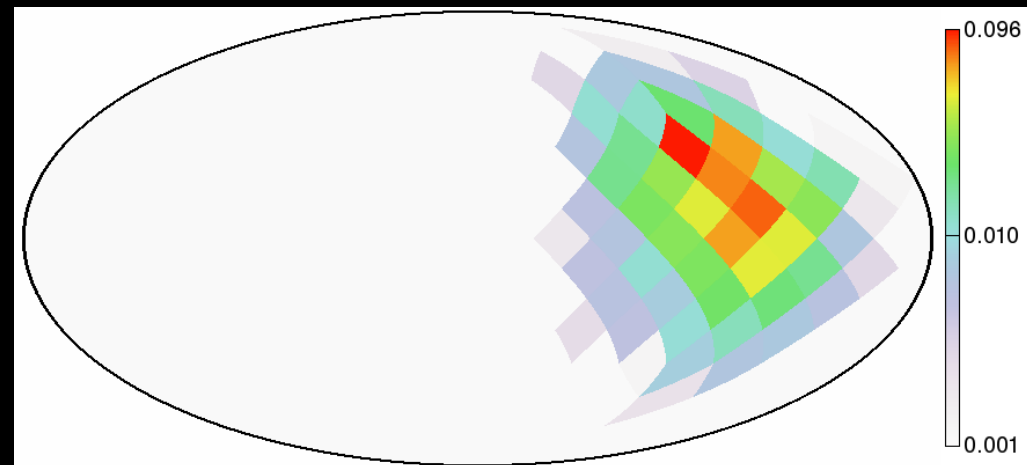
Spherical Shell (8 kpc < R < 9 kpc)



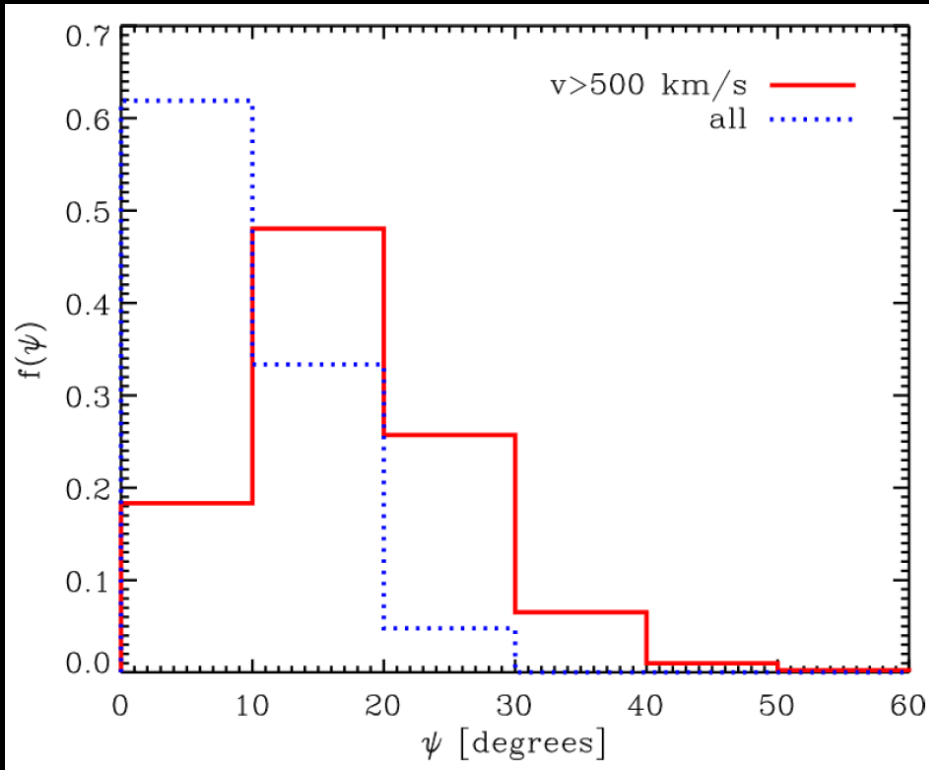
Sample Sphere #001



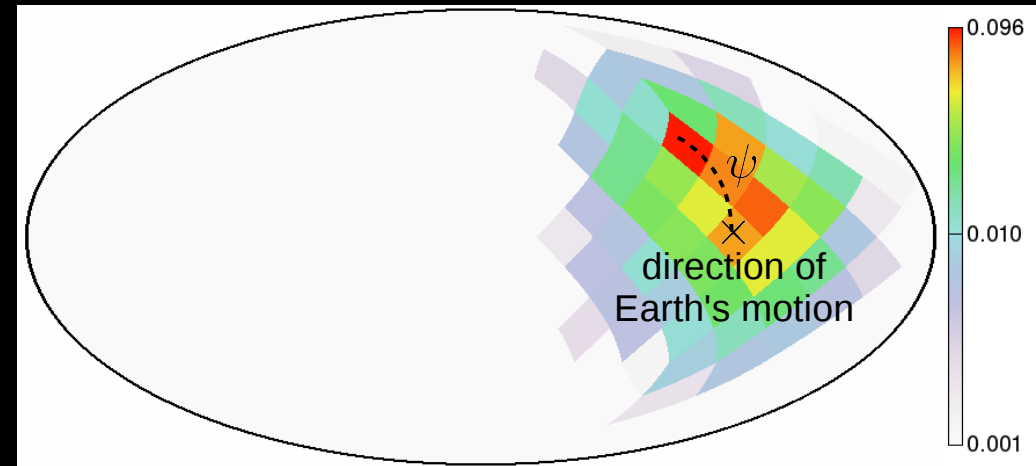
Sample Sphere #004 (containing a subhalo)



... in Earth Rest Frame $v_{\min} = 500$ km/s



Sample Sphere #004 (containing a subhalo)



At $v_{\min} = 500$ km/s the hotspot is more than 10° away from the direction of Earth's motion in $\sim 80\%$ of all cases!

Focus on High Velocity Structure

➤ Directionally sensitive experiments

- often require high recoil energies to determine direction and to distinguish head from tail.

➤ Inelastic Dark Matter

- higher v_{\min} for a given E_R
- lower $E_R \rightarrow$ *higher* v_{\min}

$$\beta_{\min} = \sqrt{\frac{1}{2m_N E_R} \left(\frac{m_N E_R}{\mu} + \delta \right)}$$

- “heavy inelastic”: scattering of Iodine, $m_\chi \gtrsim 100$ GeV, $\delta \gtrsim 100$ keV

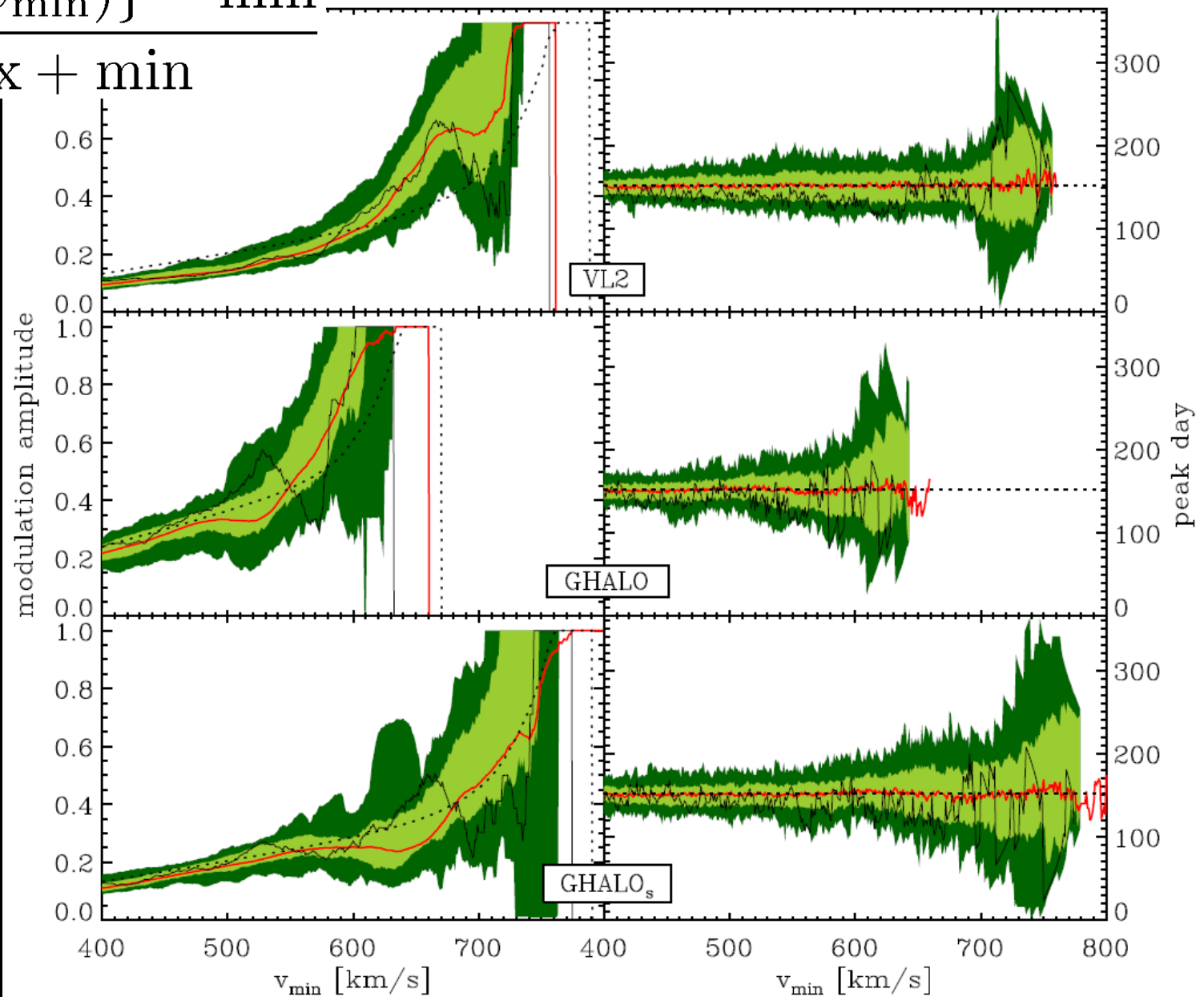
➤ Light Dark Matter

- relies on “channeling” effect
- light particle ($m_\chi \lesssim 10$ GeV) requires higher speed to get a given E_R

Annual Modulation and Peak Day

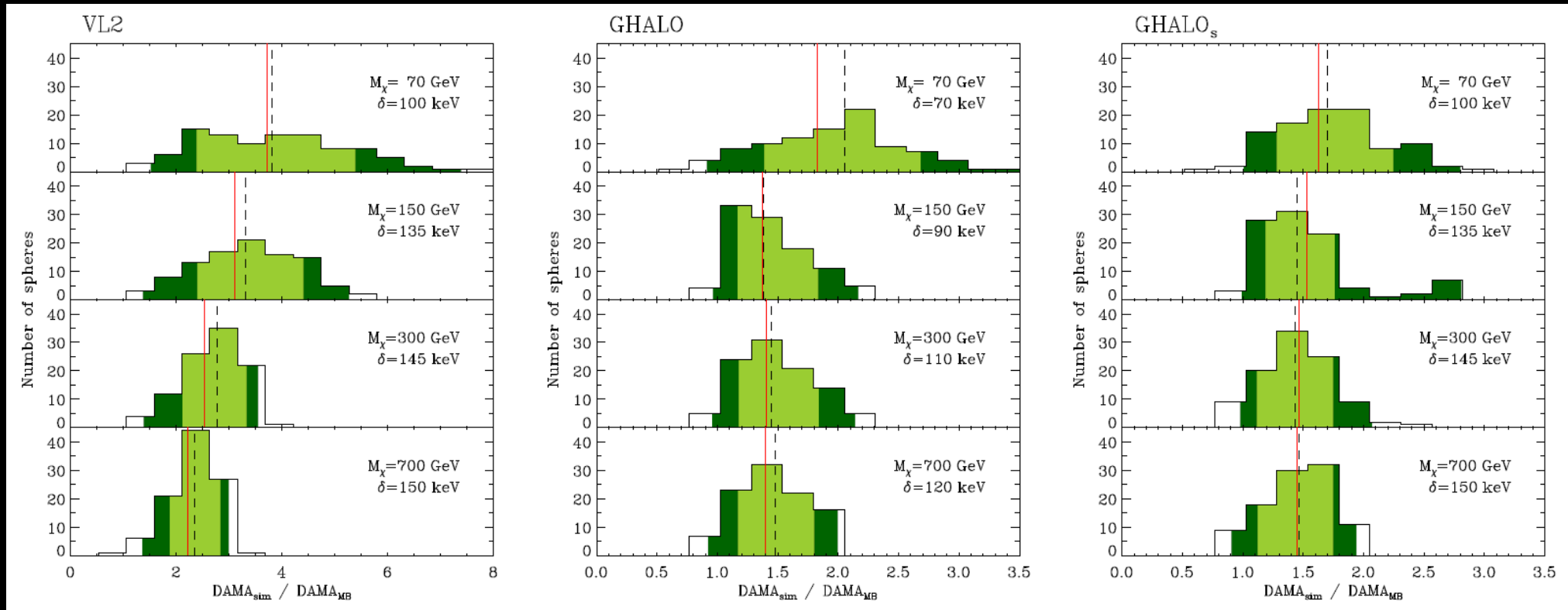
$$\frac{\max\{g(v_{\min})\} - \min}{\max + \min}$$

modulation amplitude



DAMA Modulation Amplitude

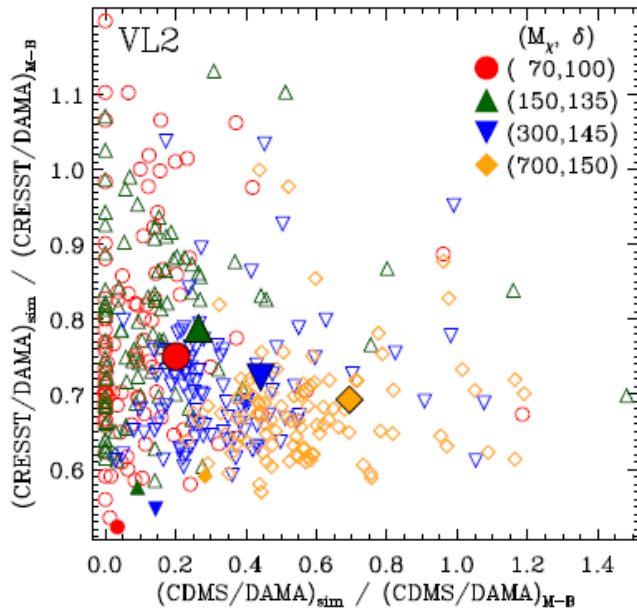
DAMA reports a modulation signal in the 2 – 6 keVee range.
We assume a quenching factor of $q=0.08$, corresponding to $E_R = 25 - 75$ keV.



Kuhlen, Weiner et al. (arXiv/0912.2358)

For iDM, the DAMA modulation is larger by factor of a few in the simulations than in the best-fitting Maxwell-Boltzmann model.

CRESST and CDMS vs. DAMA

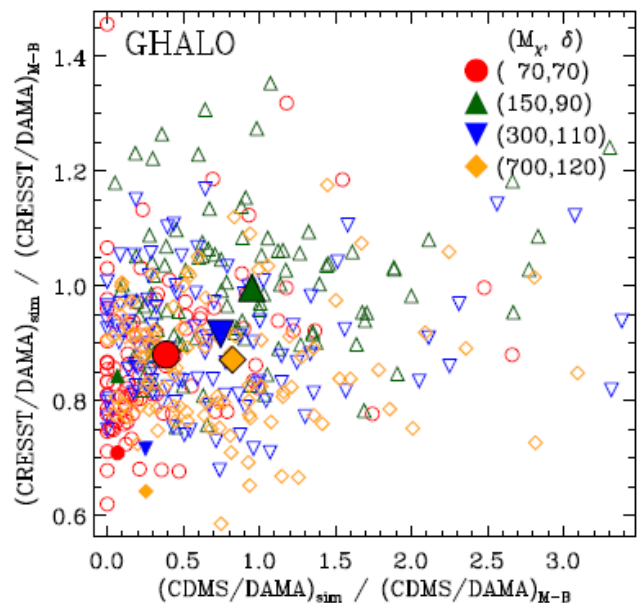


Calculate signal at CDMS and CRESST normalized by the DAMA modulation amplitude:

$$\left(\frac{\text{CDMS}}{[\text{DAMA}]}\right)_{\text{sim}} \quad \left(\frac{\text{CRESST}}{[\text{DAMA}]}\right)_{\text{sim}}$$

Look at variation of this ratio in the simulation to the best-fit Maxwellian case:

$$\frac{\left(\frac{\text{CDMS}}{[\text{DAMA}]}\right)_{\text{sim}}}{\left(\frac{\text{CDMS}}{[\text{DAMA}]}\right)_{\text{MB}}} \quad \text{vs.} \quad \frac{\left(\frac{\text{CRESST}}{[\text{DAMA}]}\right)_{\text{sim}}}{\left(\frac{\text{CRESST}}{[\text{DAMA}]}\right)_{\text{MB}}}$$



When this Ratio-of-Ratios is <1, the departure from M-B weakens CDMS and CRESST limits relative to DAMA.

DAMA: Iodine, A=127

E = 25 – 75 keV

CDMS: Germanium, A=73

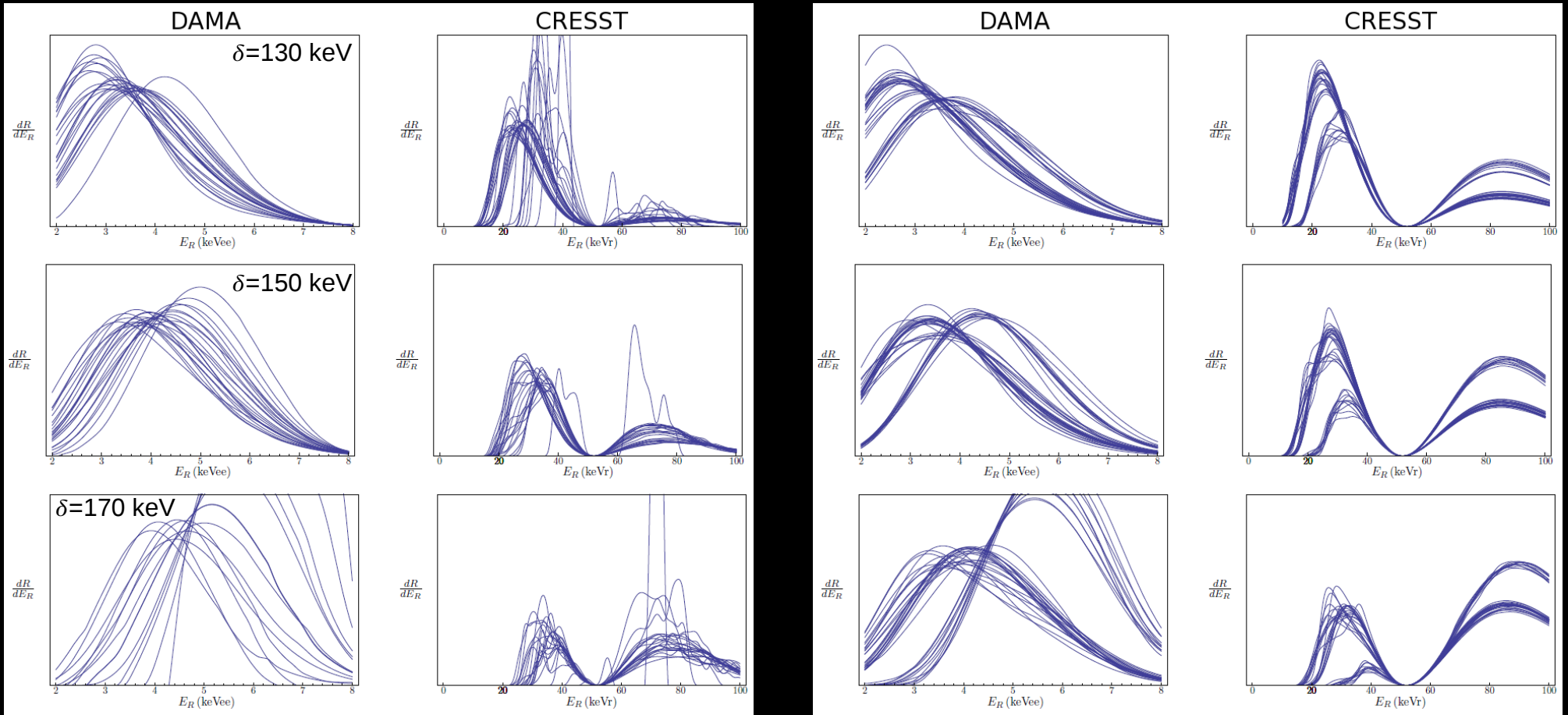
E = 10 – 100 keV

CRESST: Tungsten, A=184

Effect on iDM Recoil Spectrum

$M_\chi = 100 \text{ GeV}$

$M_\chi = 700 \text{ GeV}$



Velocity substructure can substantially change the shape and the peak location of the recoil spectrum.

Effect on iDM Recoil Spectrum

$M_\chi = 100 \text{ GeV}$

S. YELLIN

PHYSICAL REVIEW D 66, 032005 (2002)

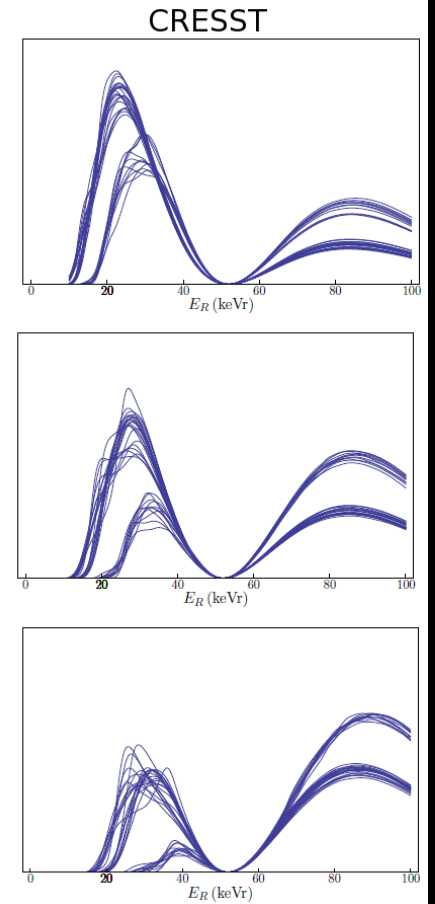
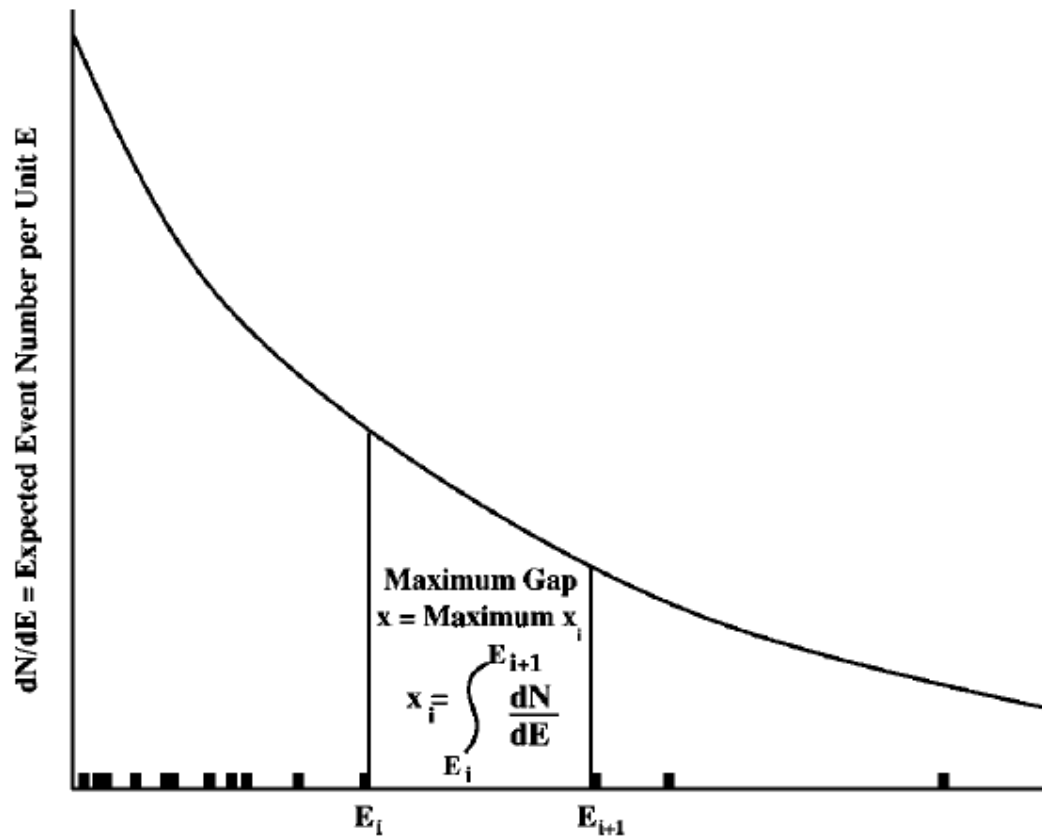
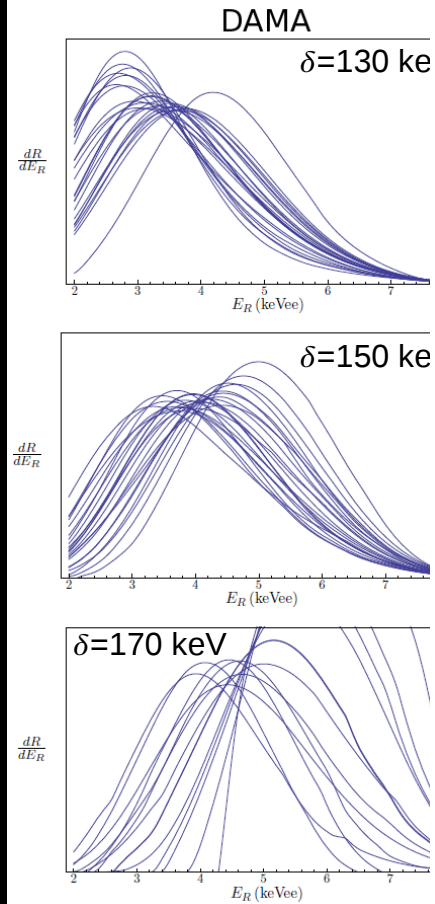


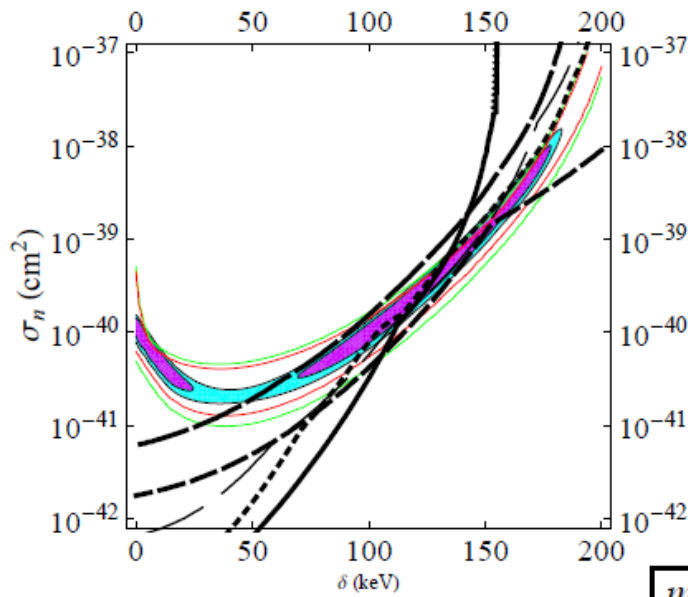
FIG. 1. Illustration of the maximum gap method. The horizontal axis is some parameter “ E ” measured for each event. The smooth curve is the signal expected for the proposed cross section, including any known background. The events from signal, known background, and unknown background are the small rectangles along the horizontal axis. The integral of the signal between two events is “ x_i .”

Velocity subst

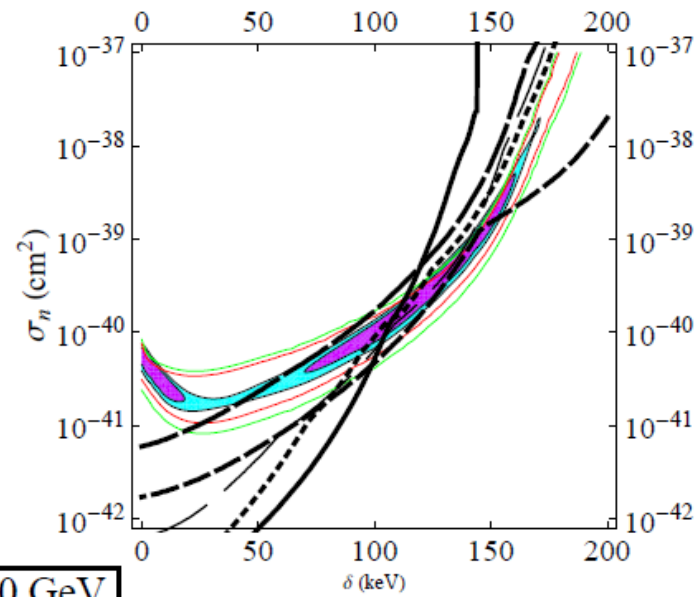
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Inelastic DM and Via Lactea/GHALO

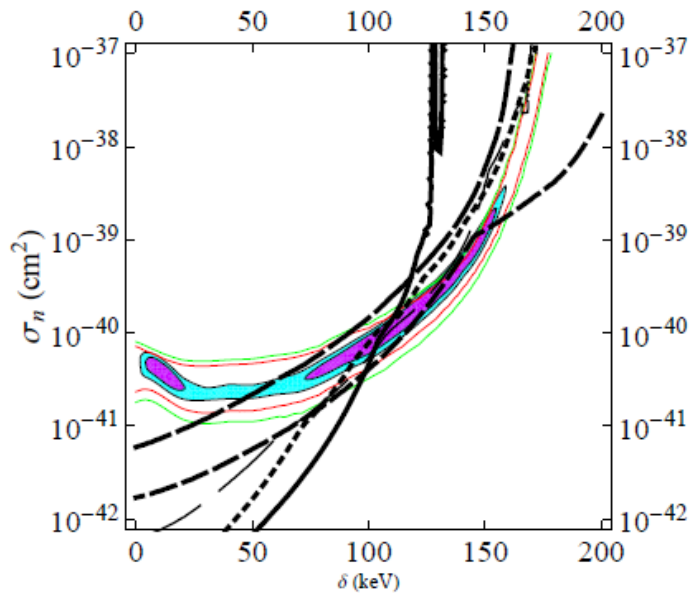
M-B ($v_0 = 220$ km/s, $v_{\text{esc}} = 550$ km/s)



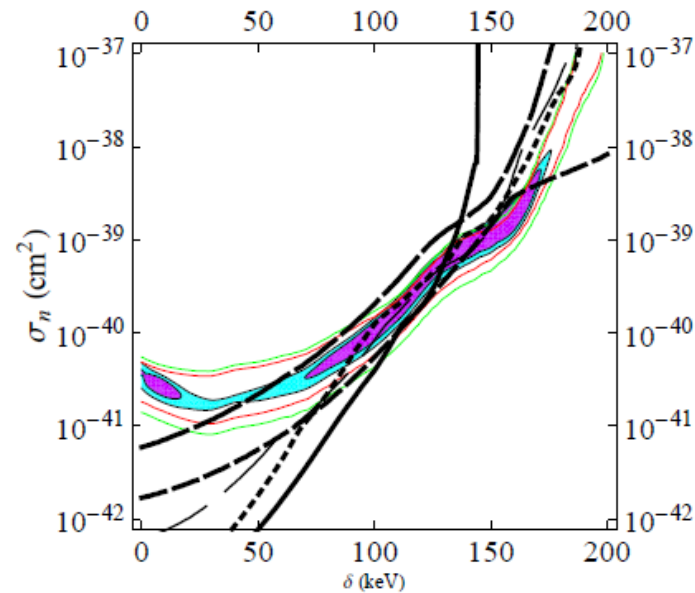
VL2 Shell ($8 \text{ kpc} < r < 9 \text{ kpc}$)



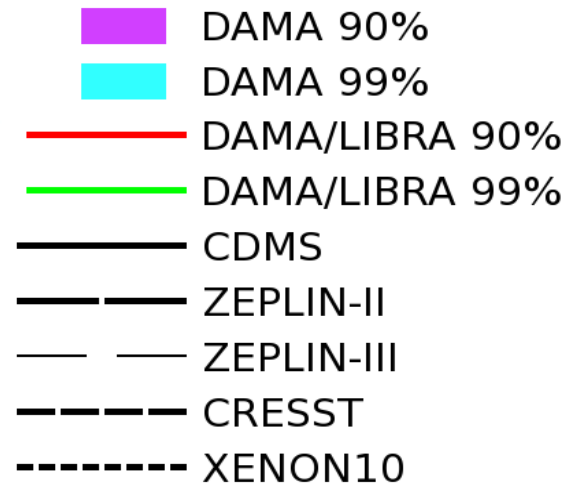
VL2 sample sphere



GHALO sample sphere

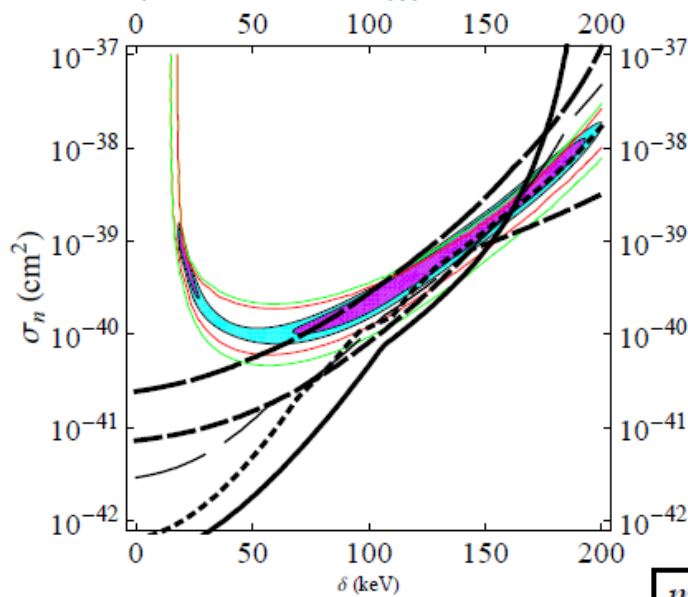


$m_\chi = 150$ GeV

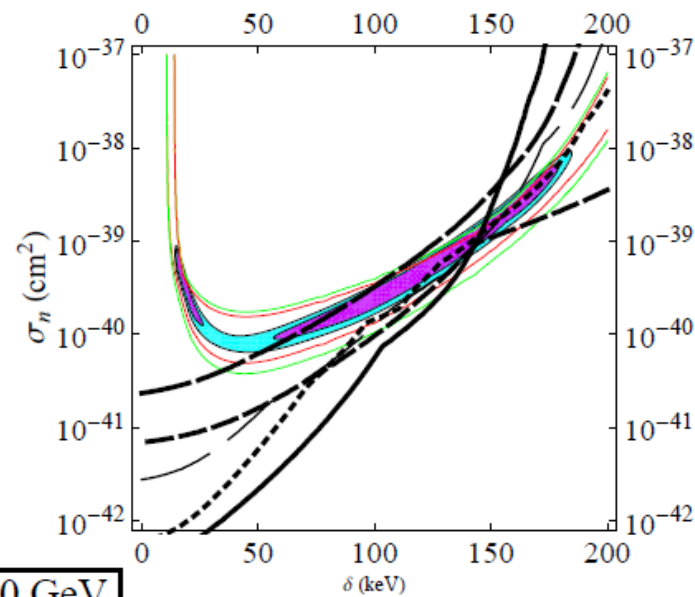


Inelastic DM and Via Lactea/GHALO

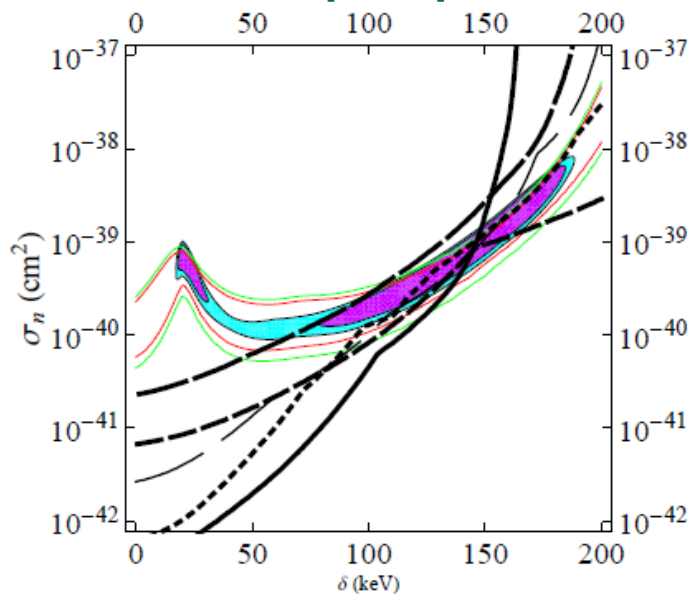
M-B ($v_0 = 220$ km/s, $v_{\text{esc}} = 550$ km/s)



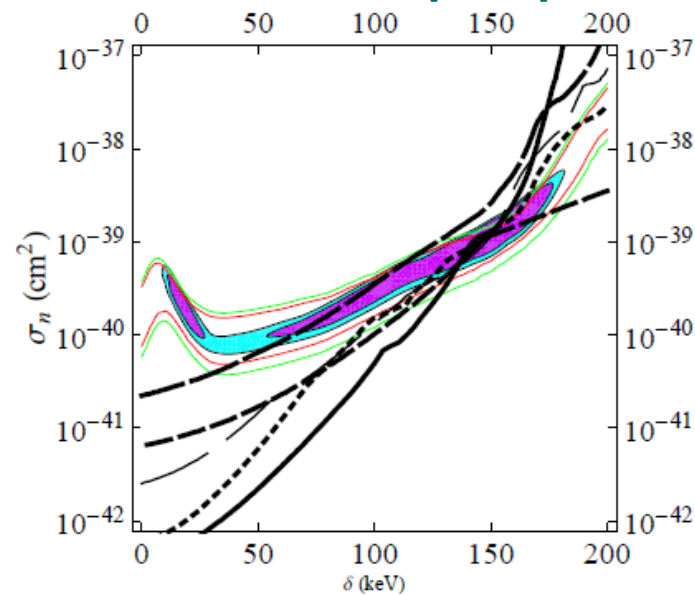
VL2 Shell ($8 \text{ kpc} < r < 9 \text{ kpc}$)



VL2 sample sphere



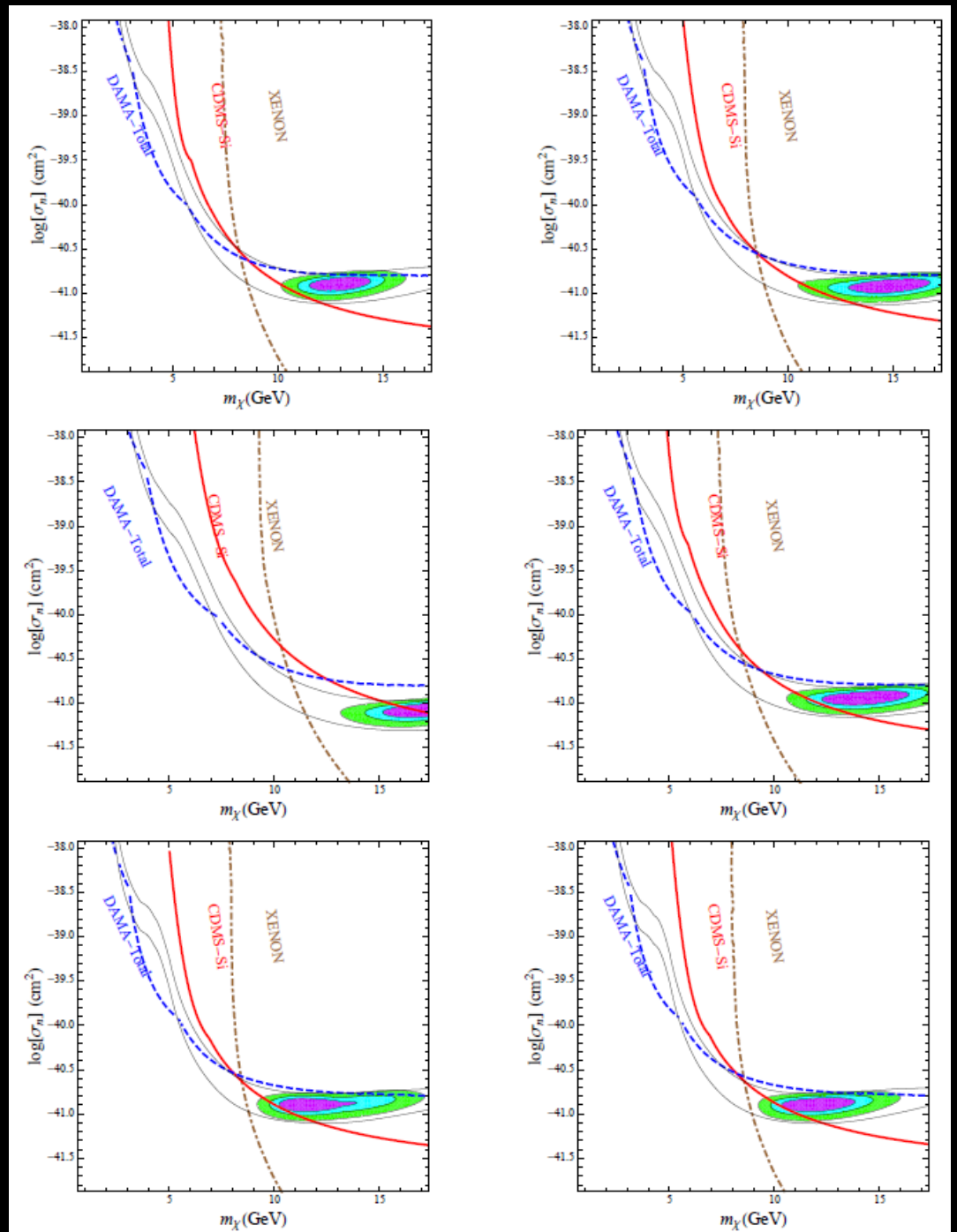
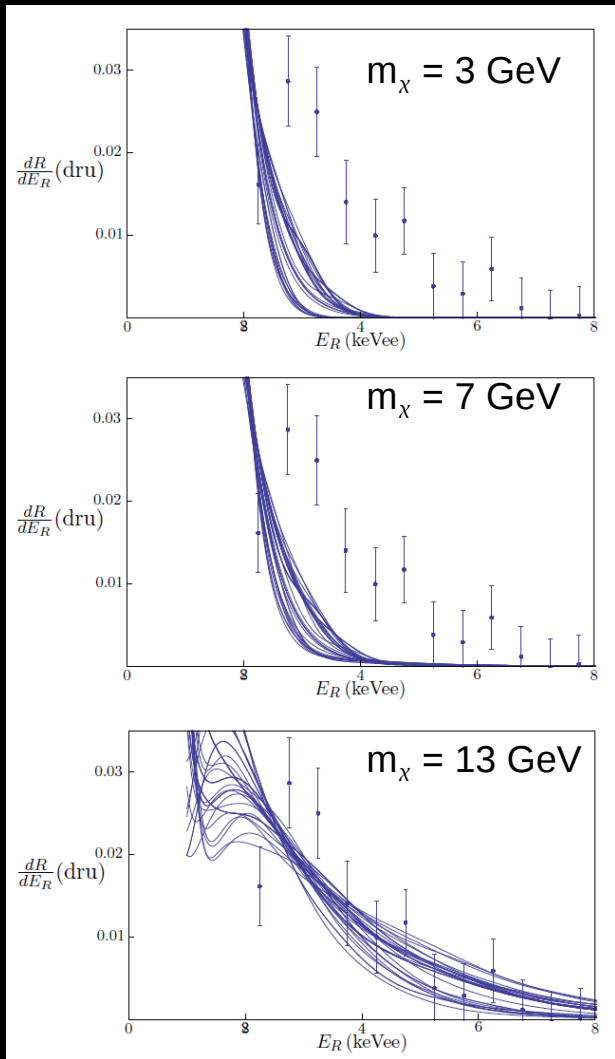
GHALO sample sphere



$m_\chi = 700$ GeV

- DAMA 90%
- DAMA 99%
- DAMA/LIBRA 90%
- DAMA/LIBRA 99%
- CDMS
- ZEPLIN-II
- ZEPLIN-III
- CRESST
- XENON10

Effects on Light DM Constraints



The simulations improve agreement, but still tension with CDMS and XENON.

Tables of $g(v_{\min}) = \int_{v_{\min}}^{\infty} \frac{f(v)}{v} dv$ are available for download, at:

Dark Matter Direct Detection with Non-Maxwellian Velocity Structure - Mozilla Firefox

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http://astro.berkeley.edu/~mqk/dmdd/

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Dark Matter Direct Detection...

We provide tables for the VL2, GHALO, and GHALO_s simulations. (See the paper and references therein for details on the simulations.) Additionally we provide tables for a case where we scaled the velocities of the particles to give $v_0=220$ km/s, where v_0 is the peak of halo rest frame $f(v)$ (i.e. the most probable speed). The velocity scaling factor is $f_v = 1.1875$ for VL2 and GHALO_s and 1.5556 for GHALO. An escape velocity cutoff of $v_{\text{esc}}=550$ km/s was applied. Note that we *did not use* these scaled distributions in our analysis, because we found that a lot of the interesting structure was pushed to beyond the escape velocity.

For each simulation, we provide one table for the spherical shell average (all particles with $8 \text{ kpc} < r < 9 \text{ kpc}$), a tarball containing tables for all 100 sample spheres, and a smaller tarball containing 10 DAMA-favorable sample spheres (i.e. spheres lying in the lower left corner of Fig.8).

VL2

<p>original (everything_VL2.tar.gz)</p> <p>$V_{\max} = 201.3$ km/s $V_{\text{circ}}(R_{\odot}) = 158.1$ km/s $V_0 = 185.3$ km/s</p> <p>spherical shell sample: gym_VL2_shell.txt Nvmin_VL2_shell.txt</p> <p>all spheres: gym_VL2_spheres.tar.gz Nvmin_VL2_spheres.tar.gz</p> <p>favorable spheres: gym_VL2_fav.tar.gz Nvmin_VL2_fav.tar.gz</p>	<p>scaled to $v_0=220$ km/s (everything_VL2_220.tar.gz)</p> <p>$V_{\max} = 201.3$ km/s $V_{\text{circ}}(R_{\odot}) = 158.1$ km/s $V_0 = 220.0$ km/s</p> <p>spherical shell sample: gym_VL2_220_shell.txt Nvmin_VL2_220_shell.txt</p> <p>all spheres: gym_VL2_220_spheres.tar.gz Nvmin_VL2_220_spheres.tar.gz</p> <p>favorable spheres: gym_VL2_220_fav.tar.gz Nvmin_VL2_220_fav.tar.gz</p>
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GHALO

<p>original (everything_ghalo.tar.gz)</p> <p>$V_{\max} = 151.5$ km/s $V_{\text{circ}}(R_{\odot}) = 121.7$ km/s $V_0 = 141.4$ km/s</p> <p>spherical shell sample: gym_ghalo_shell.txt Nvmin_ghalo_shell.txt</p>	<p>scaled to $v_0=220$ km/s (everything_ghalo_220.tar.gz)</p> <p>$V_{\max} = 151.5$ km/s $V_{\text{circ}}(R_{\odot}) = 121.7$ km/s $V_0 = 220.0$ km/s</p> <p>spherical shell sample: gym_ghalo_220_shell.txt Nvmin_ghalo_220_shell.txt</p>
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Done zotero

Tables of $g(v_{\min}) = \int_{v_{\min}}^{\infty} \frac{f(v)}{v} dv$ are available for download, at:

Mozilla Firefox

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http://astro.berkeley.edu/~mqk/dmdd/data/gvmin_VL2_shell.txt

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http://astro.be...n_VL2_shell.txt

```
# Via Lactea II - spherical shell sample - g(v_min)
#
# v_min      g00      g01      g02      g03      g04      g05      g06      g07      g08      g09      g10      g11
# [km/s]    [(km/s)^-1]
#
0.0  4.095e-03  4.073e-03  4.014e-03  3.933e-03  3.859e-03  3.807e-03  3.791e-03  3.810e-03  3.862e-03  3.933e-03  4.012e-03  4.072e-03
1.0  4.095e-03  4.072e-03  4.014e-03  3.932e-03  3.859e-03  3.807e-03  3.791e-03  3.810e-03  3.862e-03  3.933e-03  4.012e-03  4.072e-03
2.0  4.095e-03  4.072e-03  4.013e-03  3.932e-03  3.859e-03  3.807e-03  3.791e-03  3.810e-03  3.862e-03  3.933e-03  4.012e-03  4.071e-03
3.0  4.095e-03  4.074e-03  4.015e-03  3.935e-03  3.861e-03  3.808e-03  3.791e-03  3.809e-03  3.860e-03  3.930e-03  4.009e-03  4.070e-03
4.0  4.094e-03  4.071e-03  4.010e-03  3.929e-03  3.856e-03  3.806e-03  3.791e-03  3.811e-03  3.864e-03  3.935e-03  4.014e-03  4.072e-03
5.0  4.094e-03  4.070e-03  4.010e-03  3.929e-03  3.856e-03  3.806e-03  3.790e-03  3.810e-03  3.863e-03  3.935e-03  4.014e-03  4.072e-03
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```

Done

zotero

Conclusions

- **Numerical simulations with $O(10^9)$ particles** are beginning to resolve the phase-space structure even at $8 \text{ kpc} \ll R_{\text{vir}}$.
- The inner halo is smoother than the outskirts, but there is **residual substructure, both in density and velocity**. Probably not converged...
- The velocity modulus distribution function **$f(v)$ shows marked departures from the standard Maxwell-Boltzmann** assumption.
- Both **global (bumps, wiggles) and local (spikes)** velocity substructure.
- This is **especially important at high v_{min}** , so for directional detection, inelastic DM, light DM (for example).
- Possible effects of velocity substructure:
 - **change the direction** of scattering particles
 - increase the DAMA **modulation amplitude**
 - **change the recoil spectrum** (with implications for “maximum gap” technique)
- Drawback: must marginalize of halo model uncertainty ?
[make use of my $g(v_{\text{min}})$ tables...]
- Advantage: In the future maybe we can learn something about the Galaxy halo