Positrons and Bubbles: Backgrounds for Dark Matter Indirect Searches

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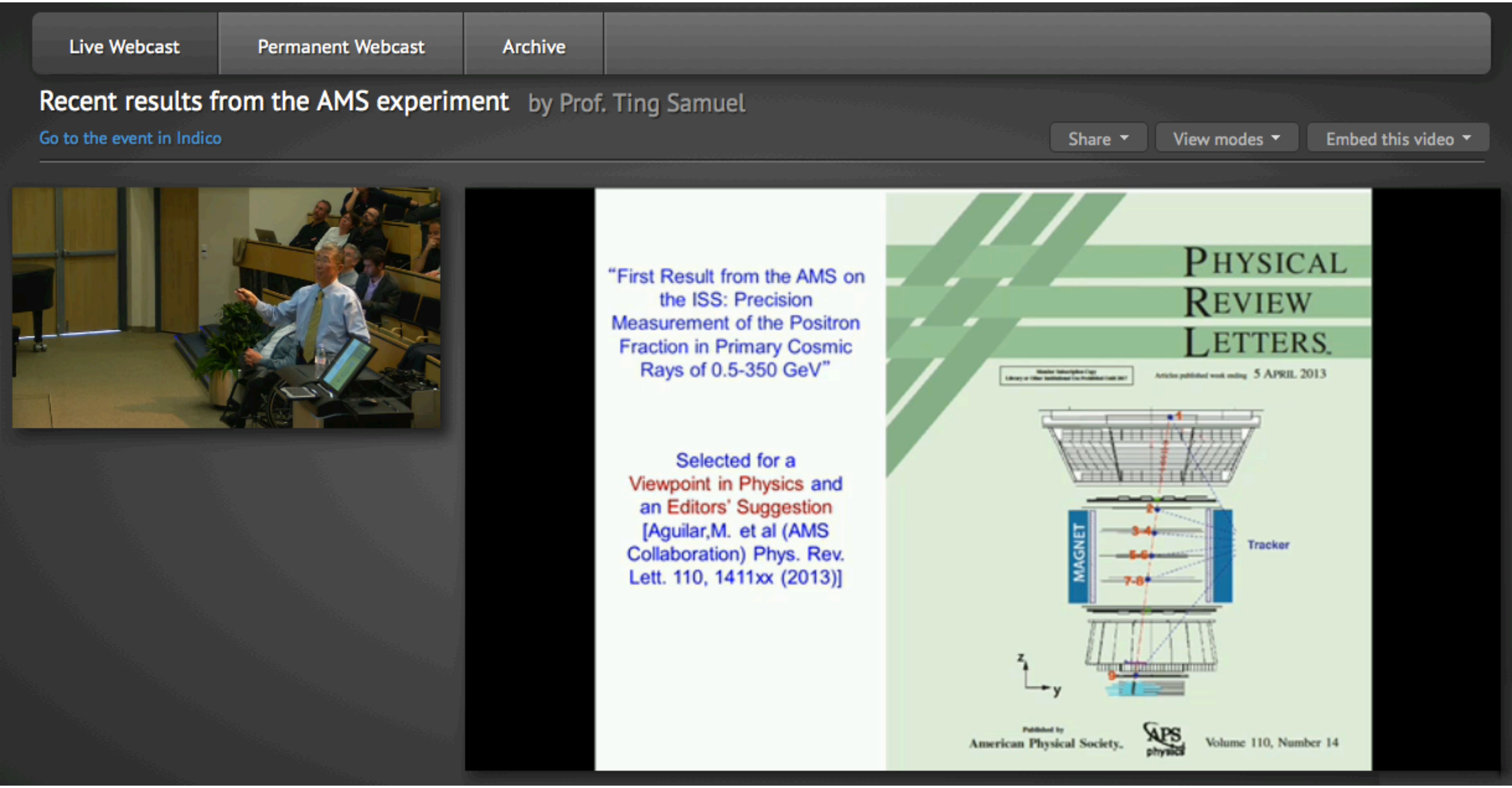
21 June 2013

SLAC NATIONAL ACCELERATOR LABORATORY

KIPAC

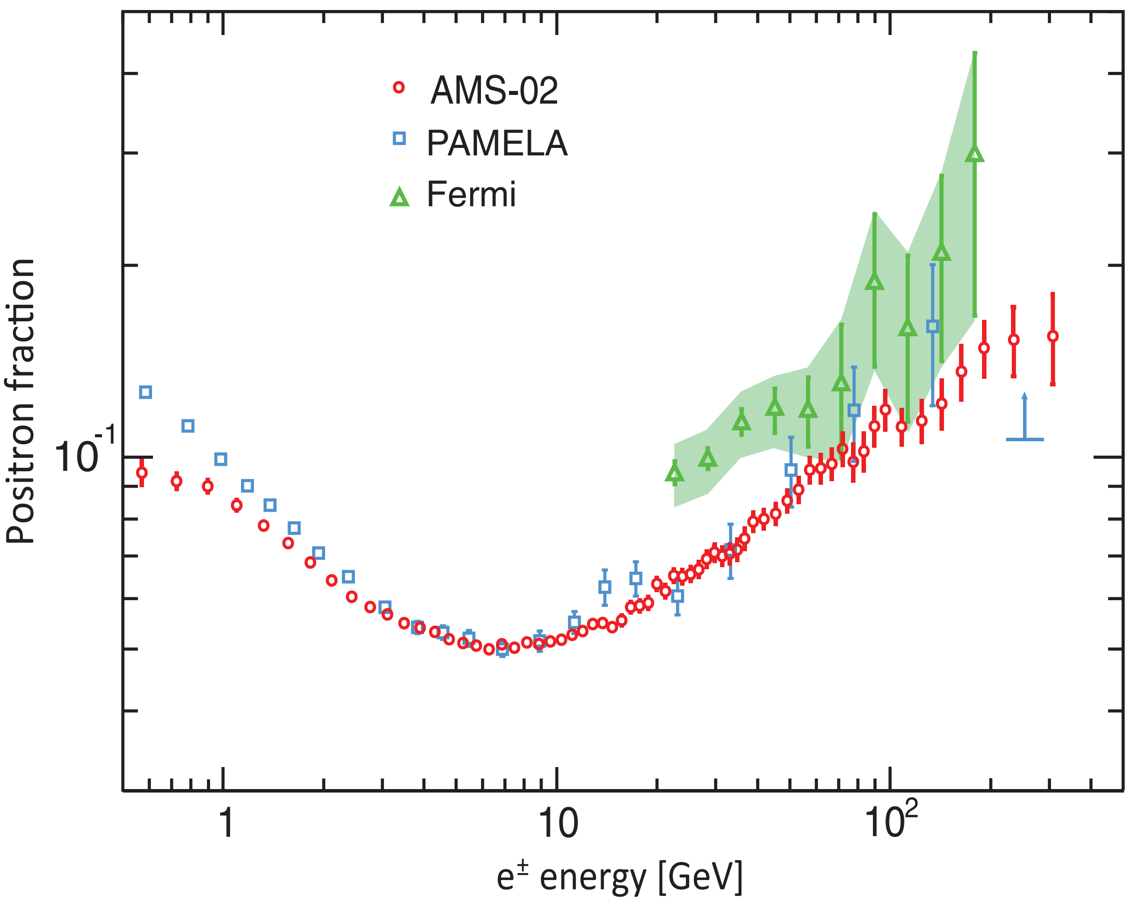
New AMS results



"There's no such thing as disappointing."

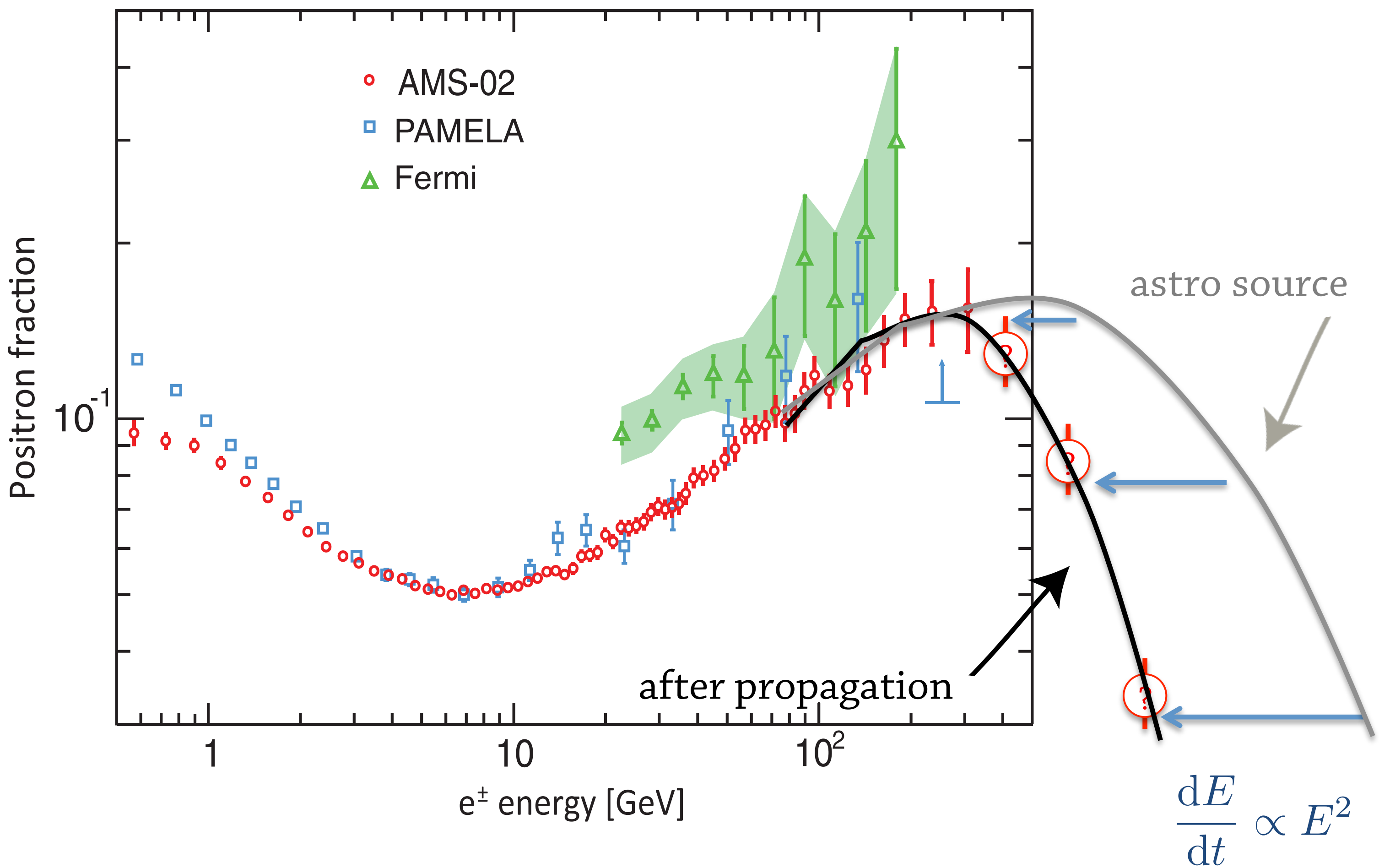
(Sam Ting)

New AMS results



"The positron fraction is turning over, so it must be dark matter."

“It’s turning over, so it must be DM.”



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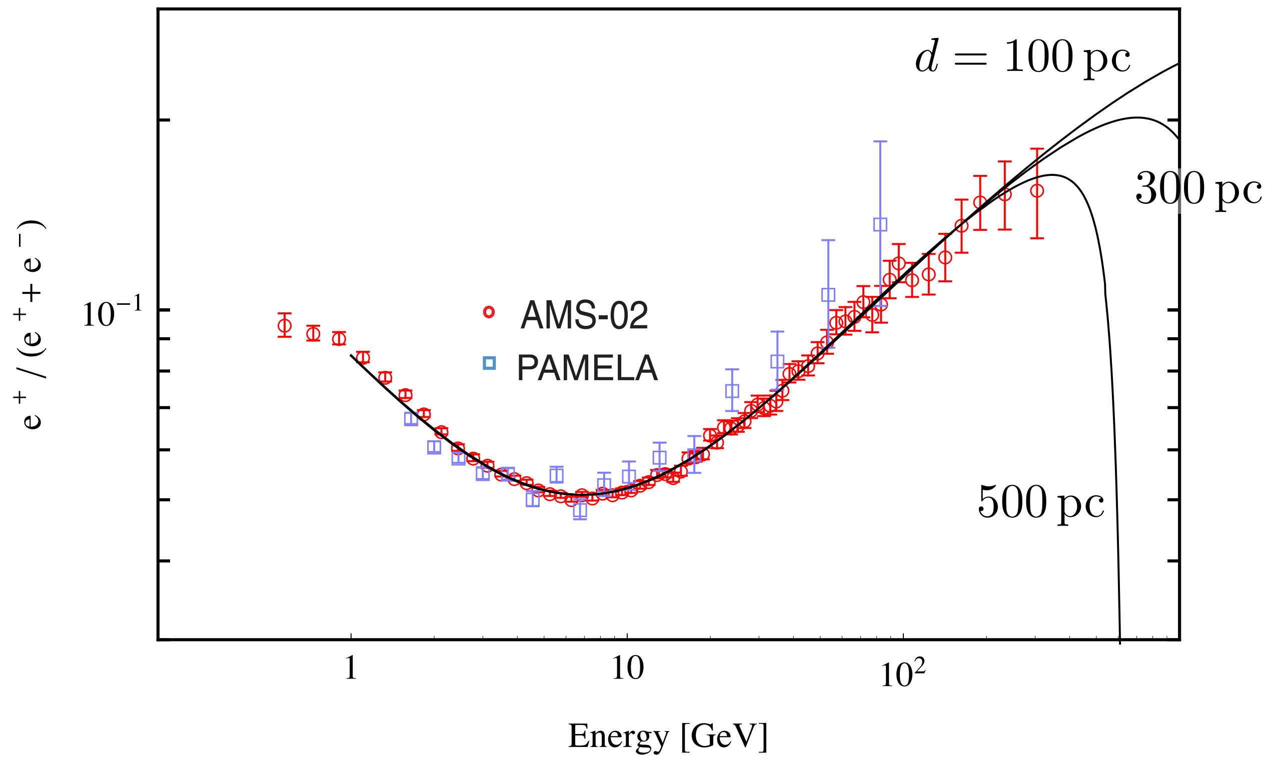
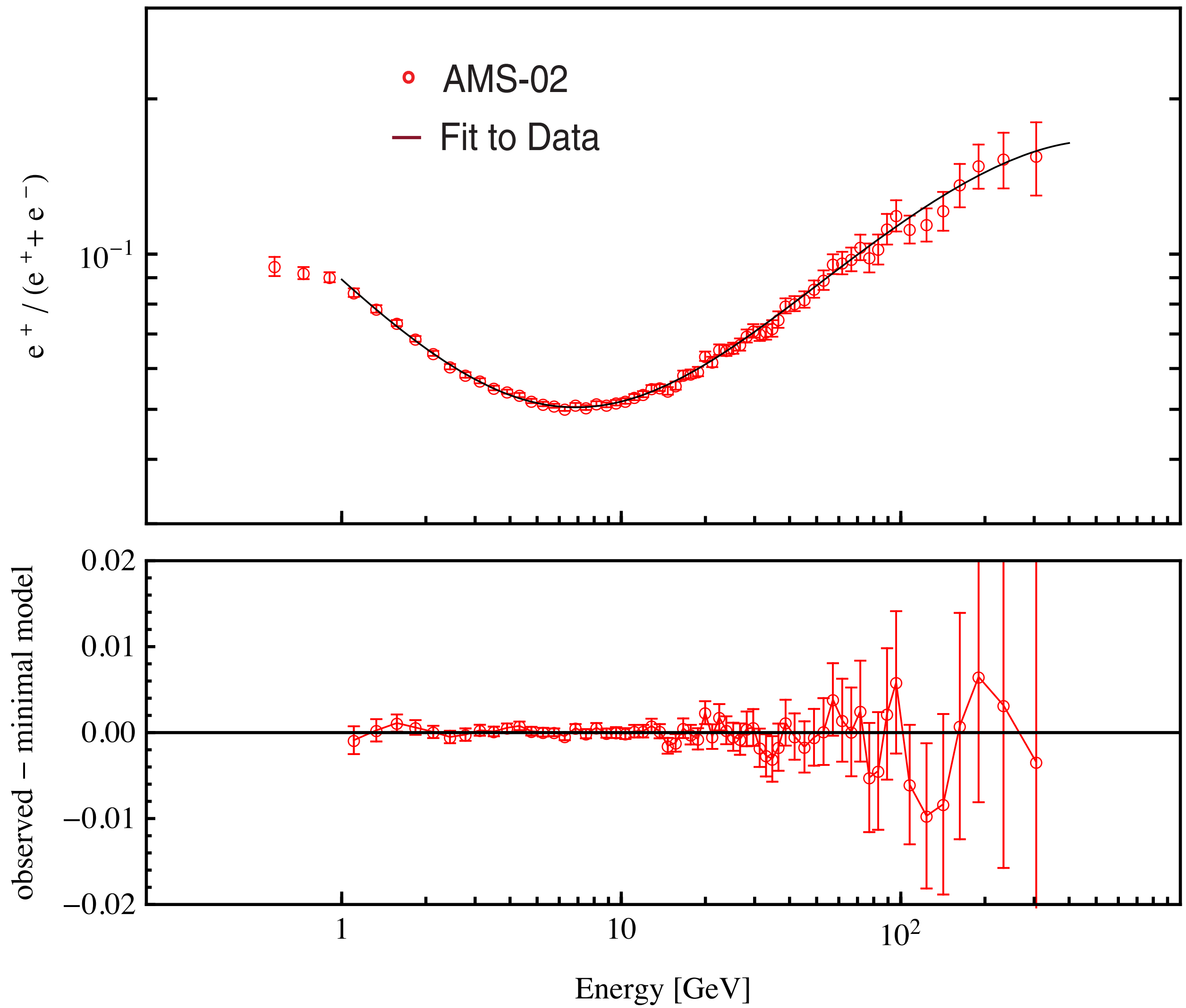


- power law spectrum with spectral index $\Gamma \sim 1.7$
- exponential cut-off at $3\,\mathrm{TeV}$
- impulsive injection $20,000 \ldots 500,000\,\mathrm{yr}$ ago

"The positron fraction has substructure, so it must be dark matter."

“There’s substructure, hence DM.”



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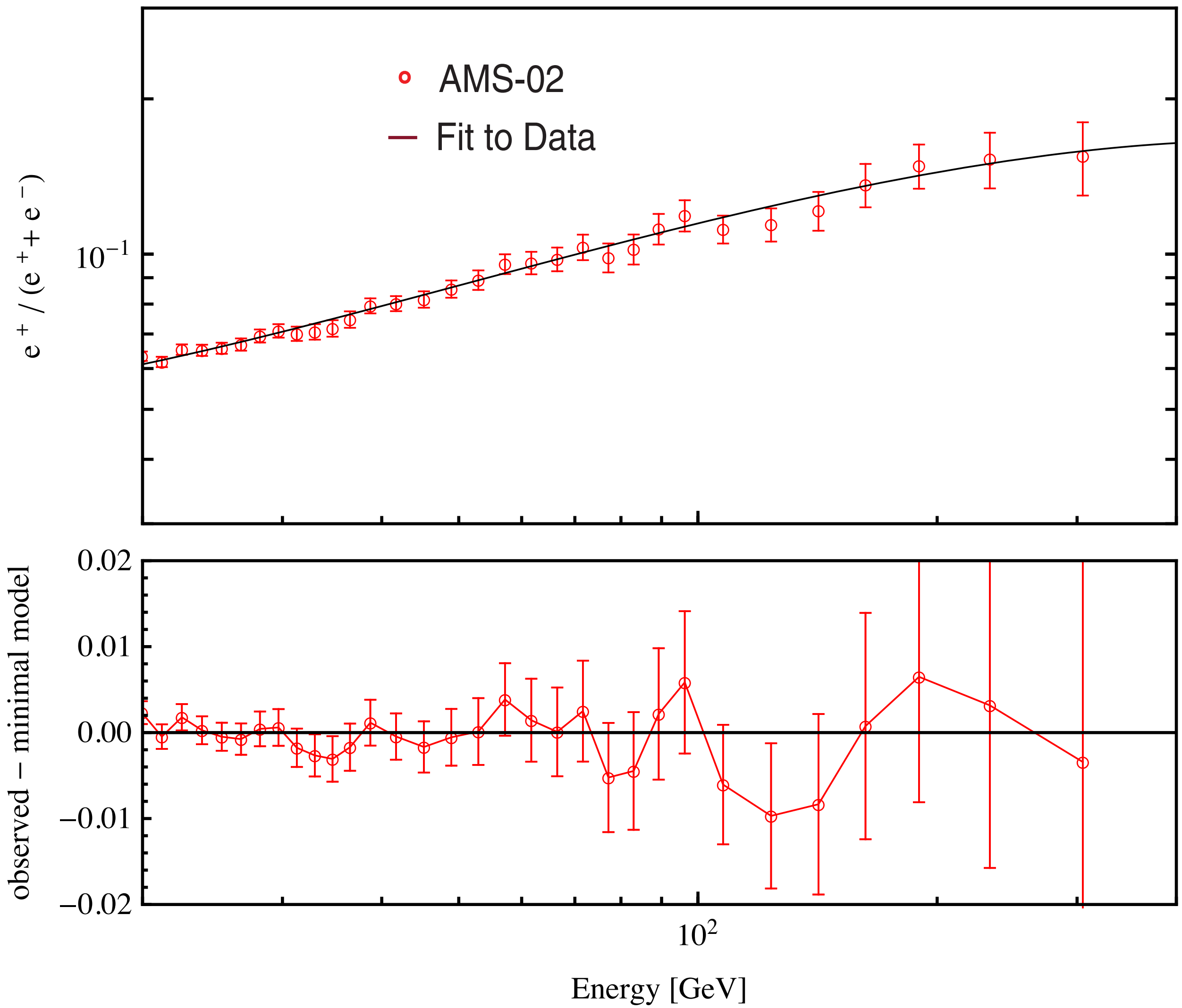

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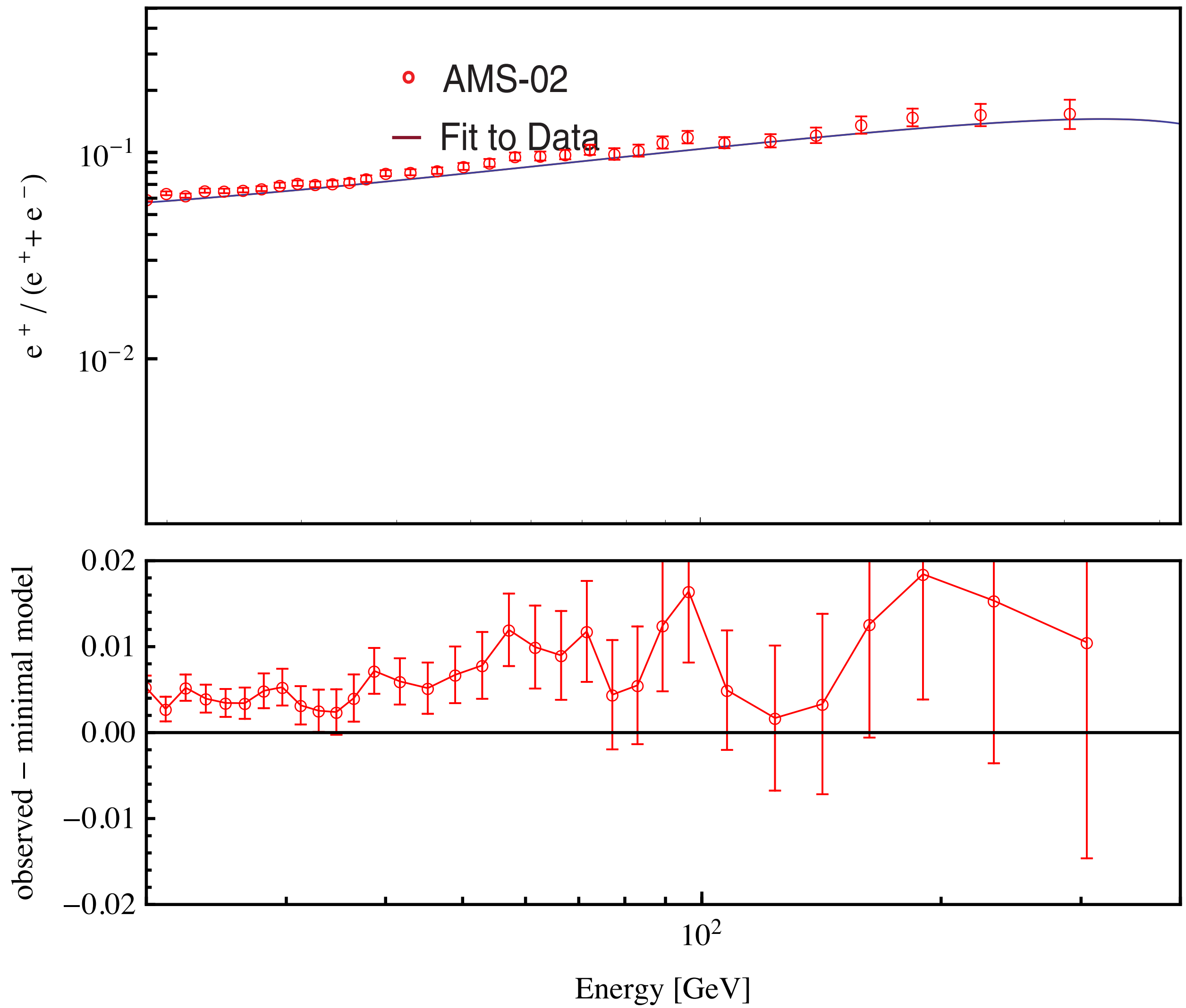

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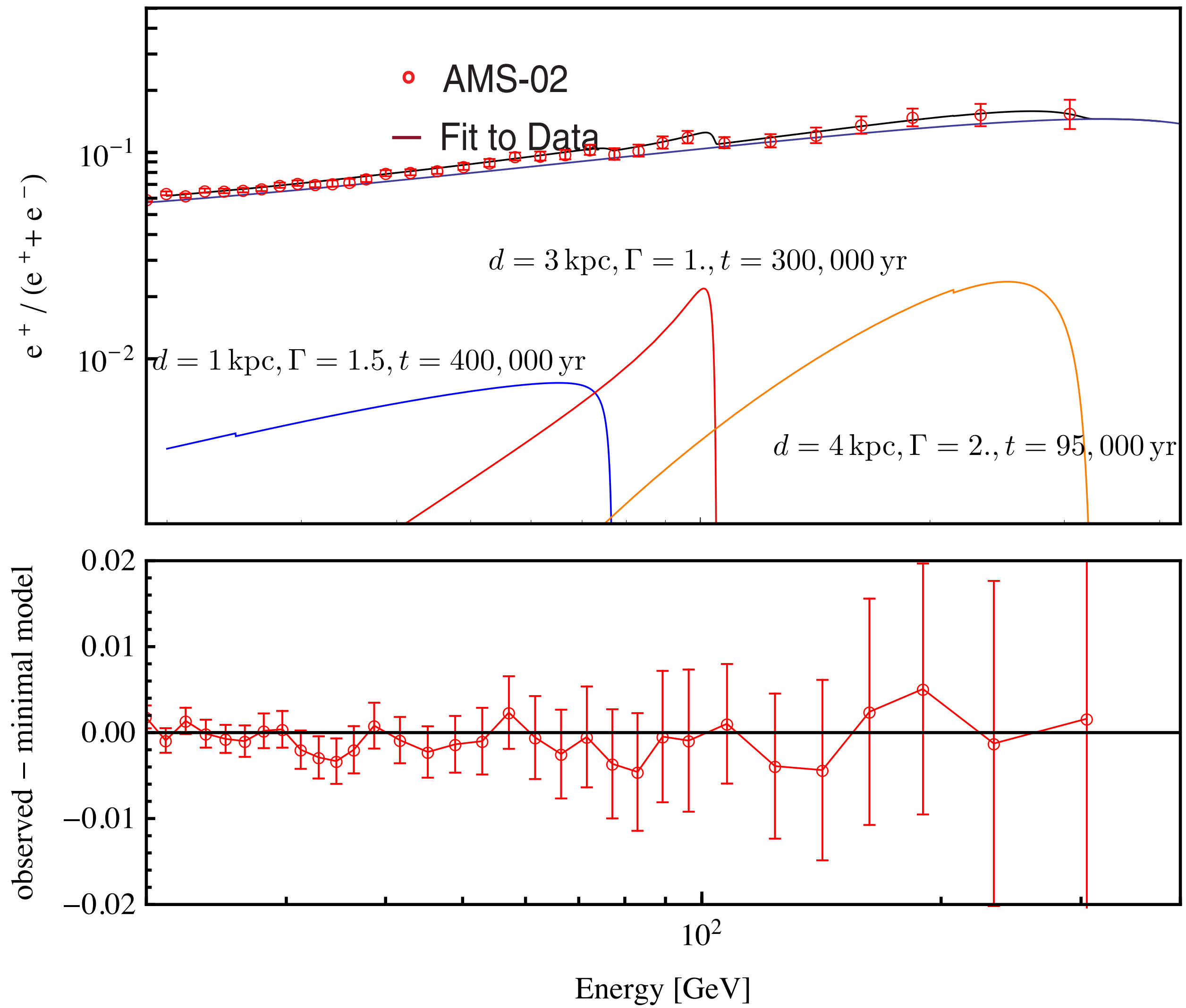
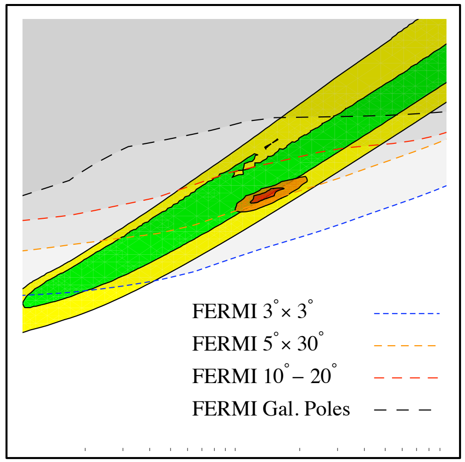
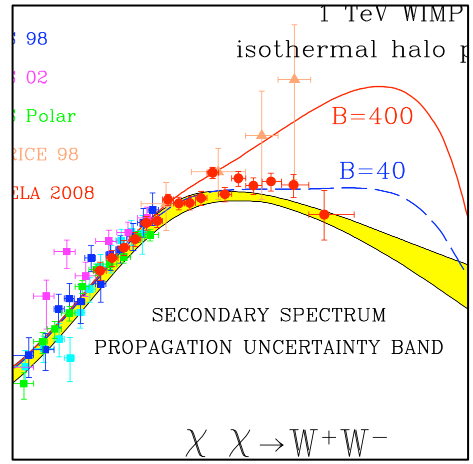


"It's either dark matter or pulsars."

Multi-messenger problem

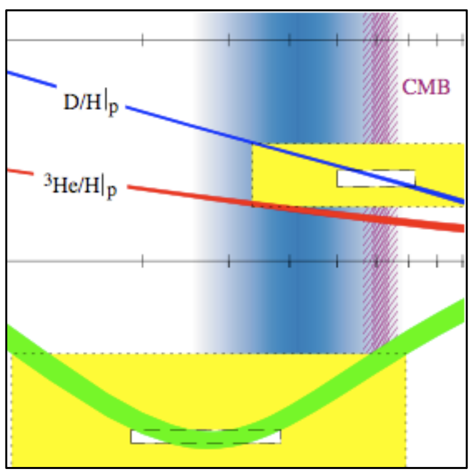
Gamma-ray constraints

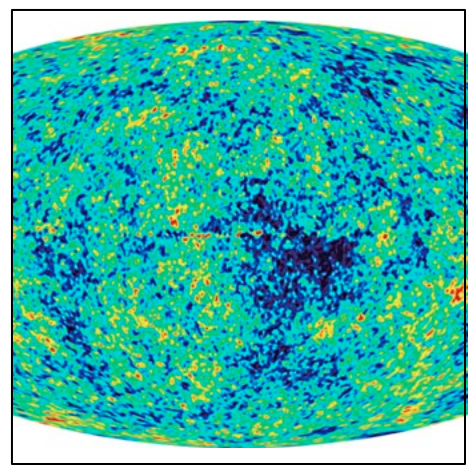




Antiproton constraints

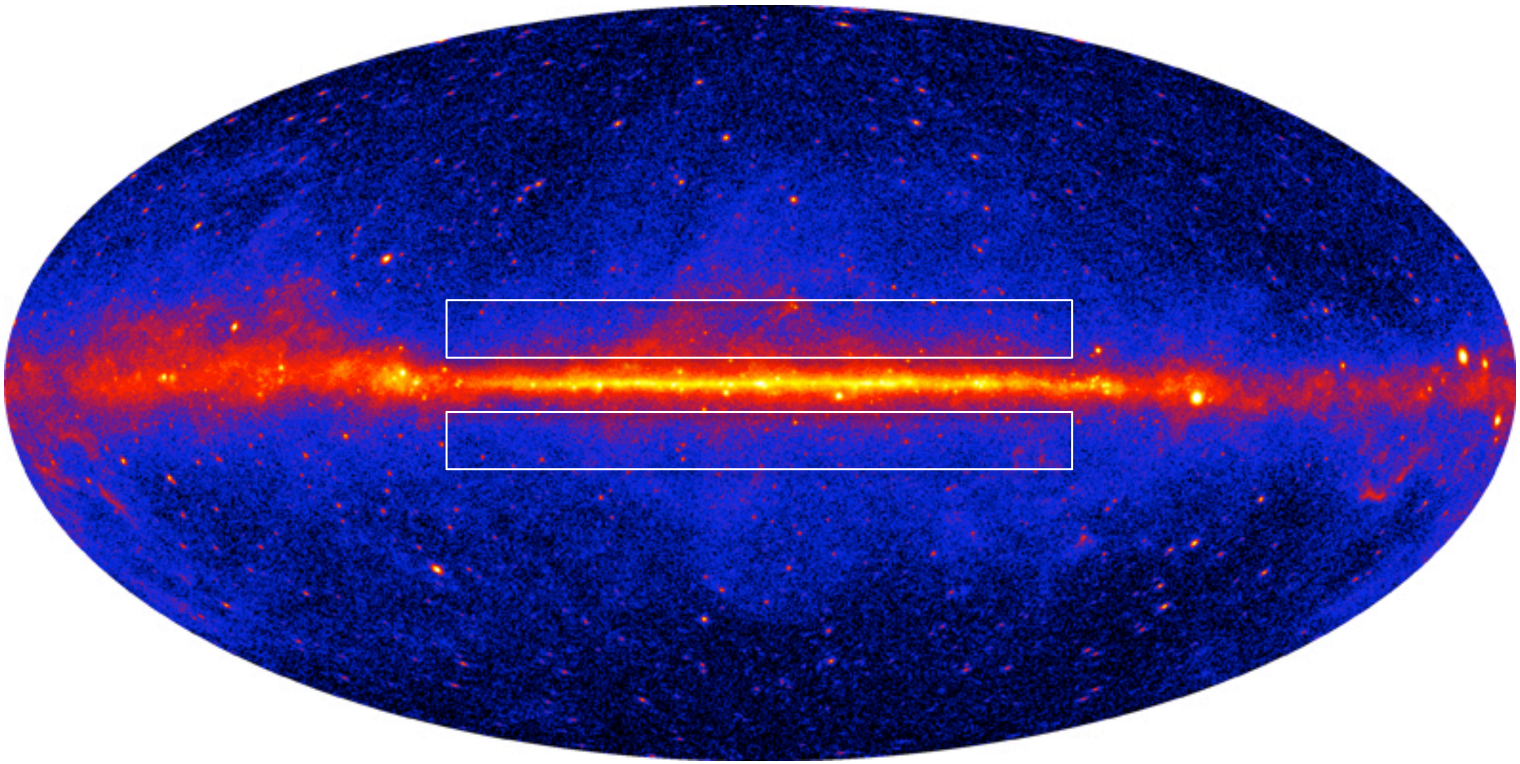
BBN constraints





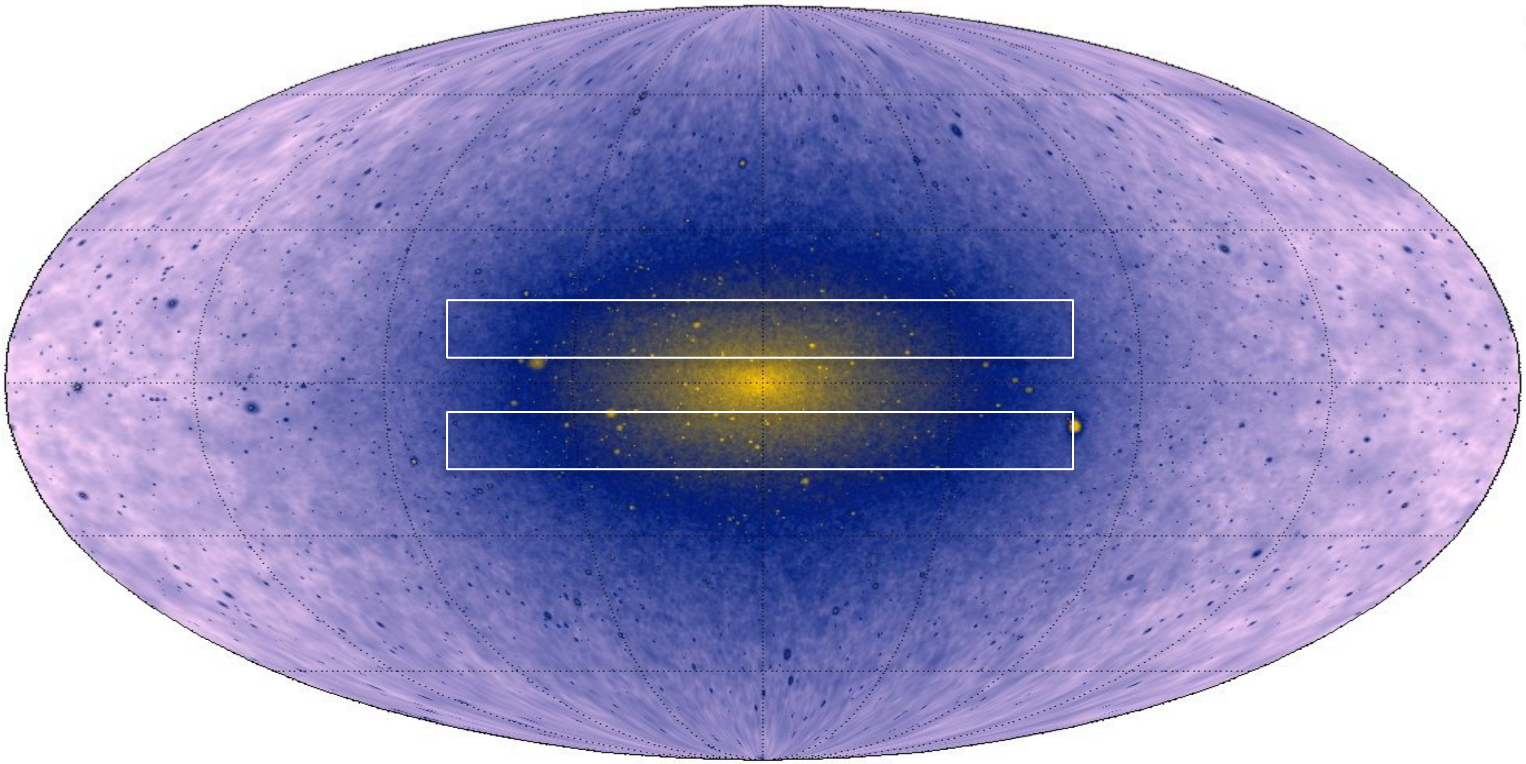
CMB constraints

Constraints from the MW halo



Fermi-LAT 3 year sky map

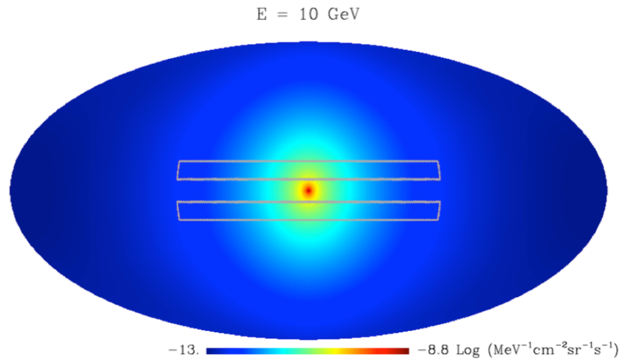
Constraints from the MW halo



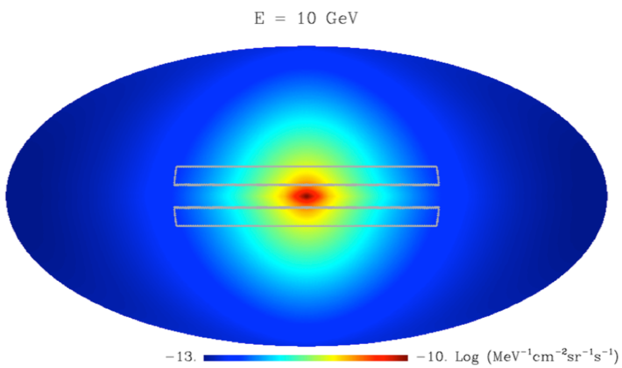
Via Lactea (Kuhlen, Diemand, Madau)

Degenerate morphologies

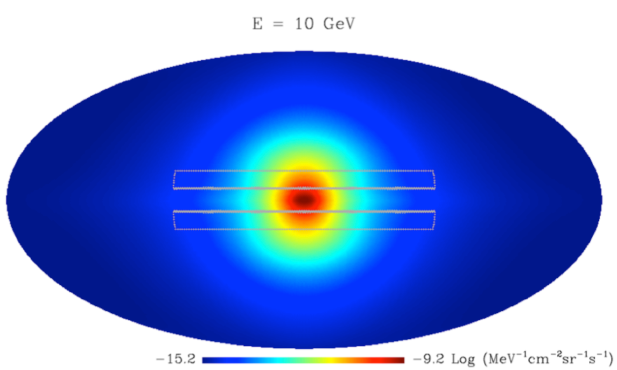
$\chi\chi \to b\bar{b}$



$\chi\chi \to \mu^+\mu^-$



astro e^-



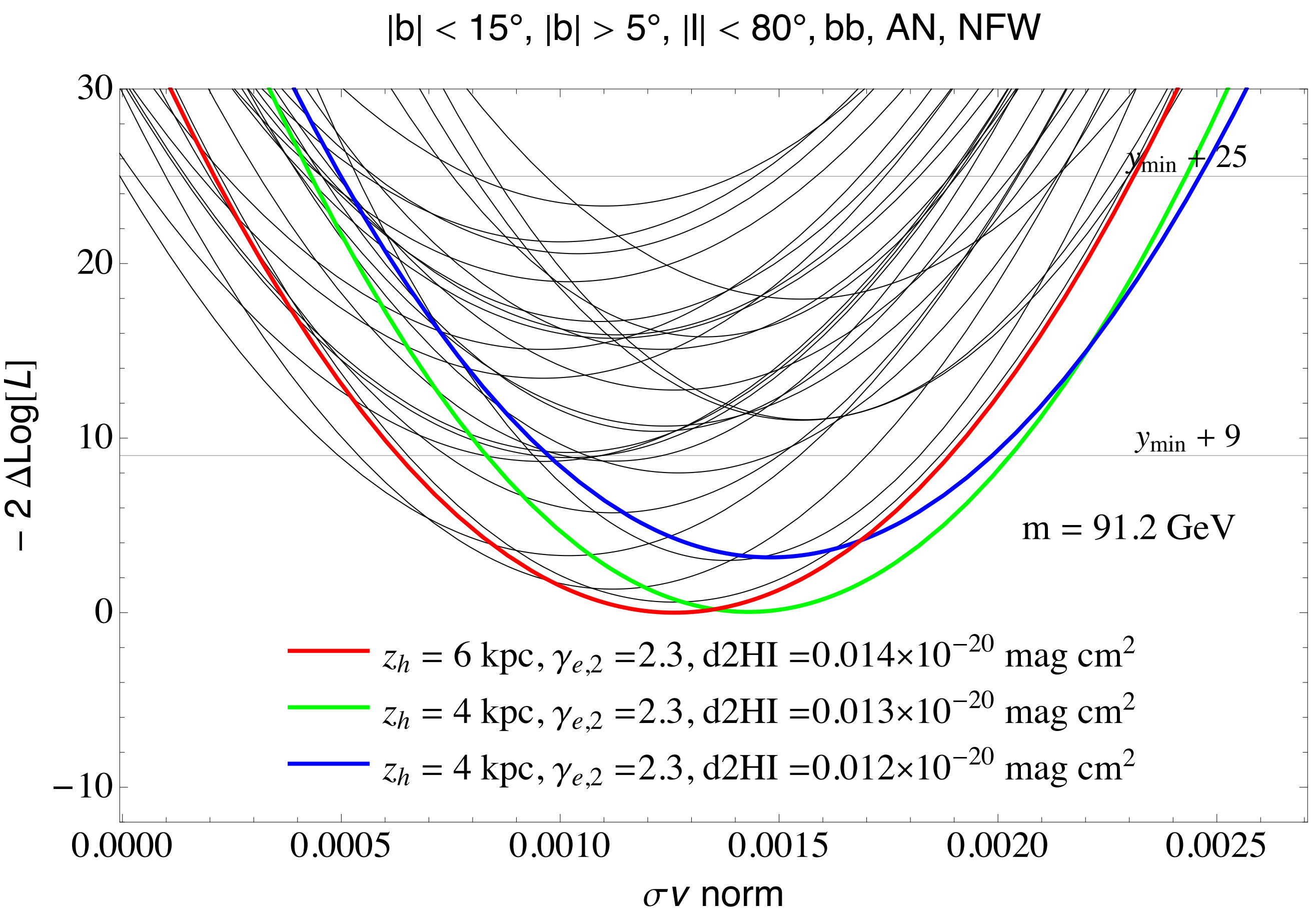
1. ultra-conservative:
   assuming no backgrounds, set bounds on annihilation x-sec.
   Papucci & Strumia JCAP 1003 (2010) 014
   Cirelli *et al.*, Nucl.Phys.B 840 (2010) 284

2. still conservative:
   adopt propagation models, vary parameters and set bound

→ marginalisation over nuisance parameters:
- CR source distribution
- electron spectral indices
- height of diffusive halo
- $X_{\rm CO}$, d2HI
- ...

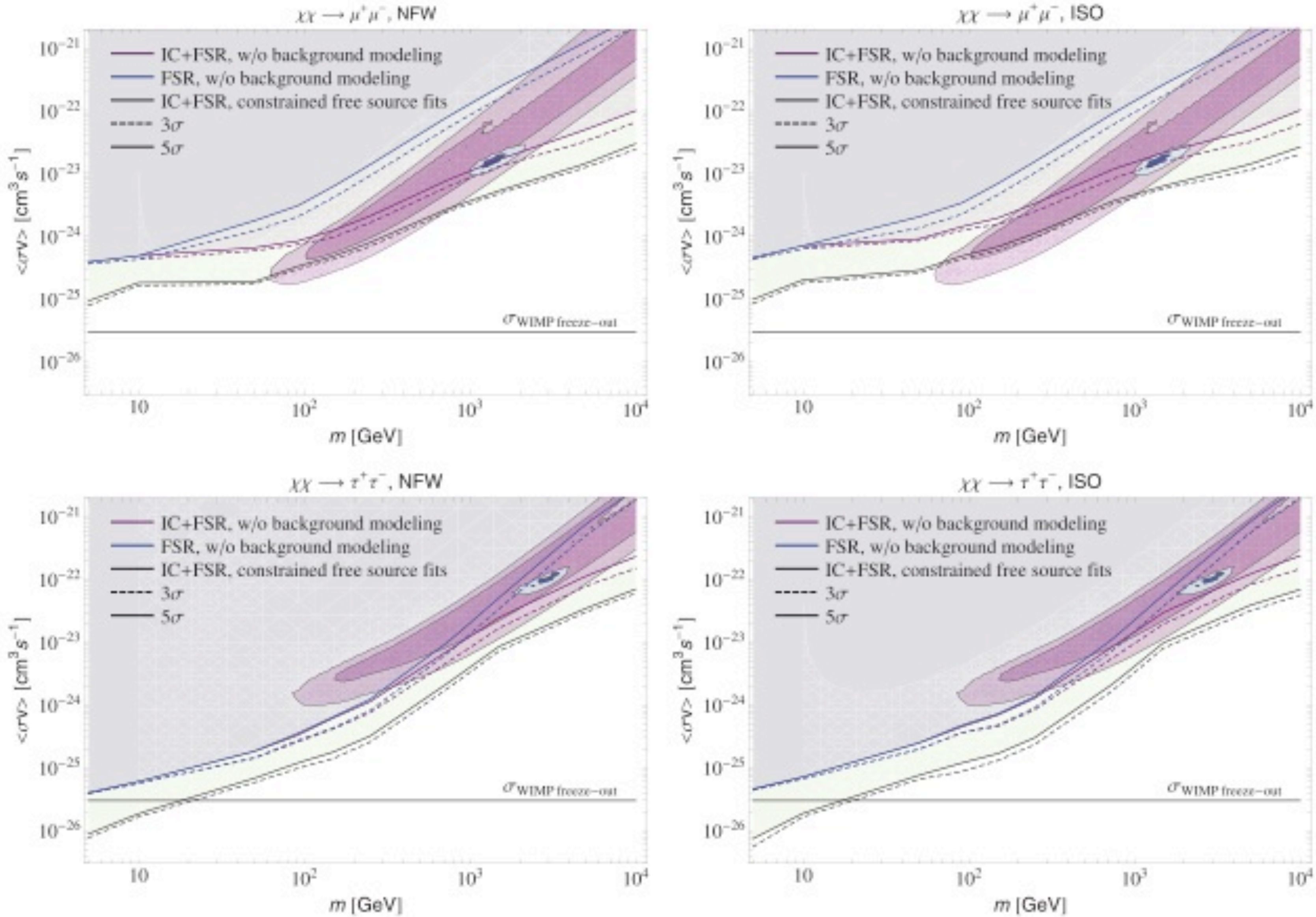
Ackermann *et al.*, ApJ **761** (2012) 91

Profile likelihood



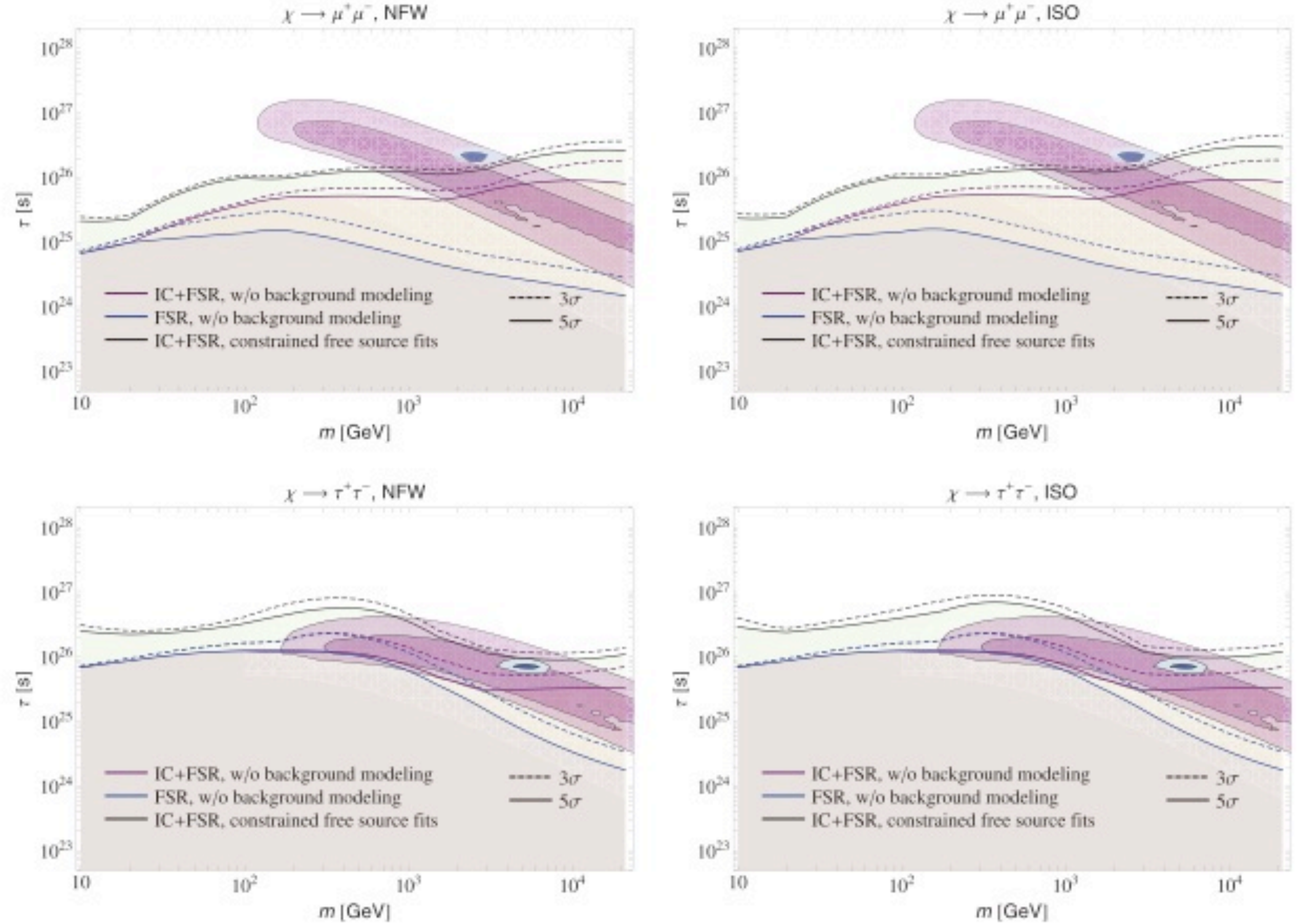
Ackermann *et al.*, ApJ **761** (2012) 91

Bounds on annihilating DM



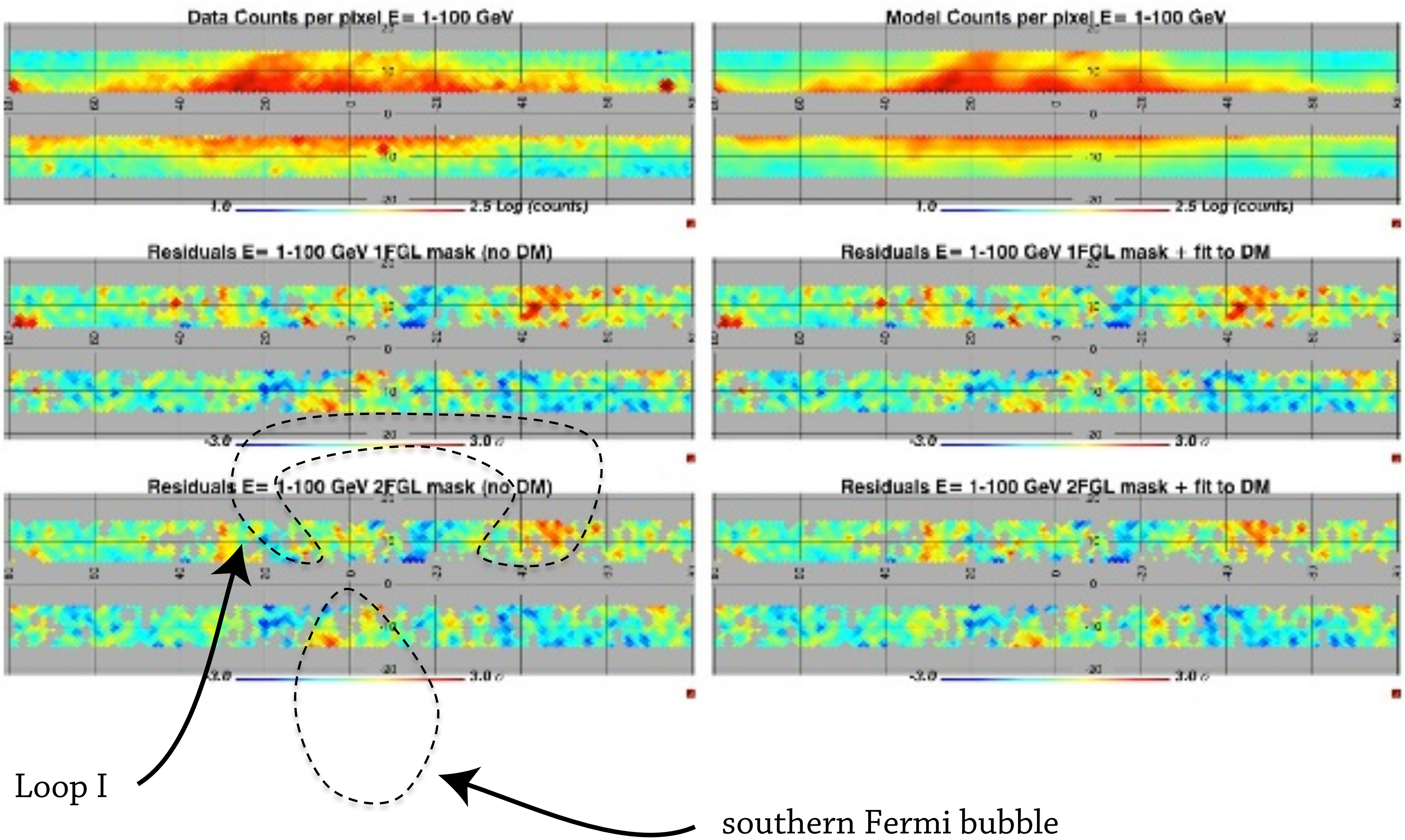
Ackermann *et al.*, ApJ **761** (2012) 91

Bounds on decaying DM



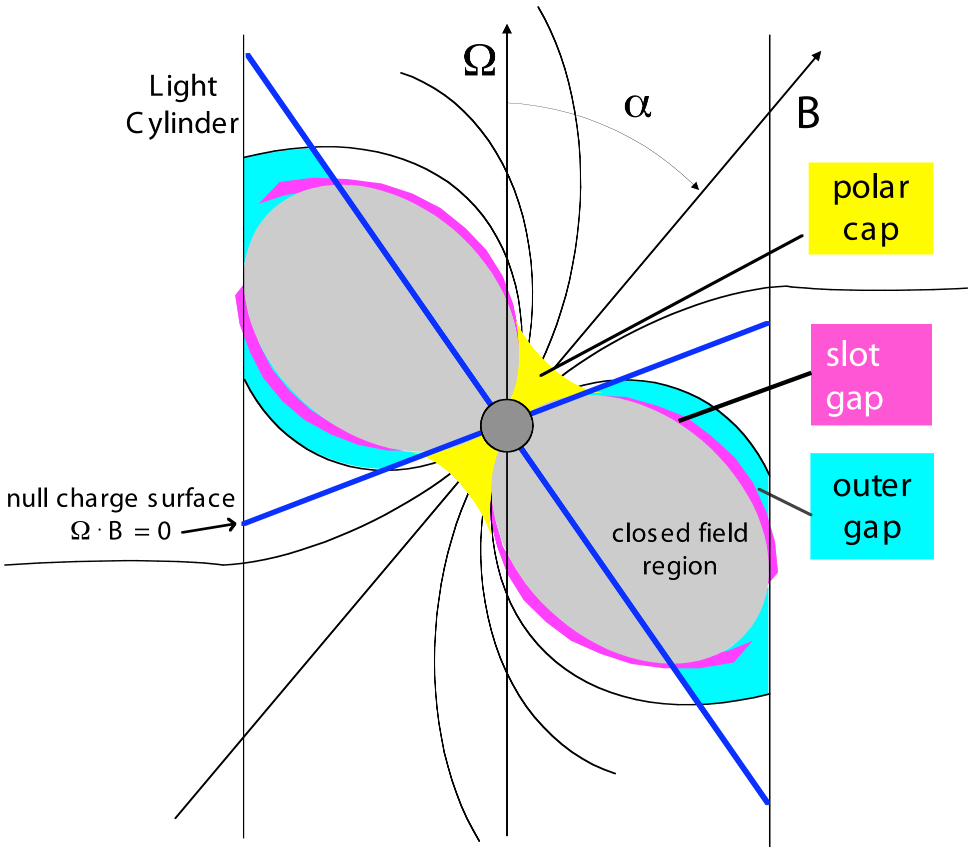
Ackermann *et al.*, ApJ **761** (2012) 91

Residuals



Ackermann *et al.*, ApJ **761** (2012) 91

“It’s either dark matter or pulsars.”



acceleration sites:

- polar cap
- outer gap
- slot gap

production of gamma-rays:

- synchrotron radiation
- inverse Compton scattering
- curvature radiation

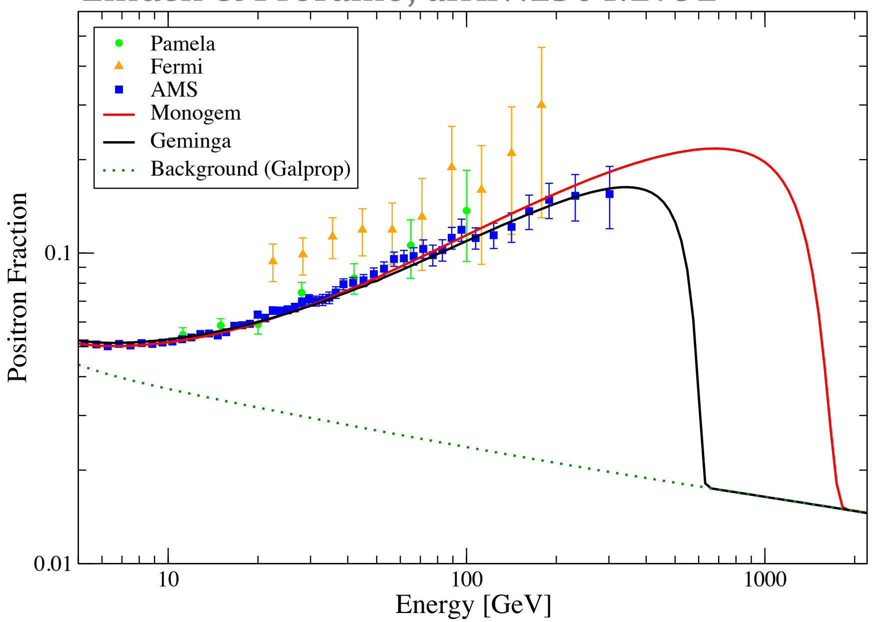
Nearby pulsars

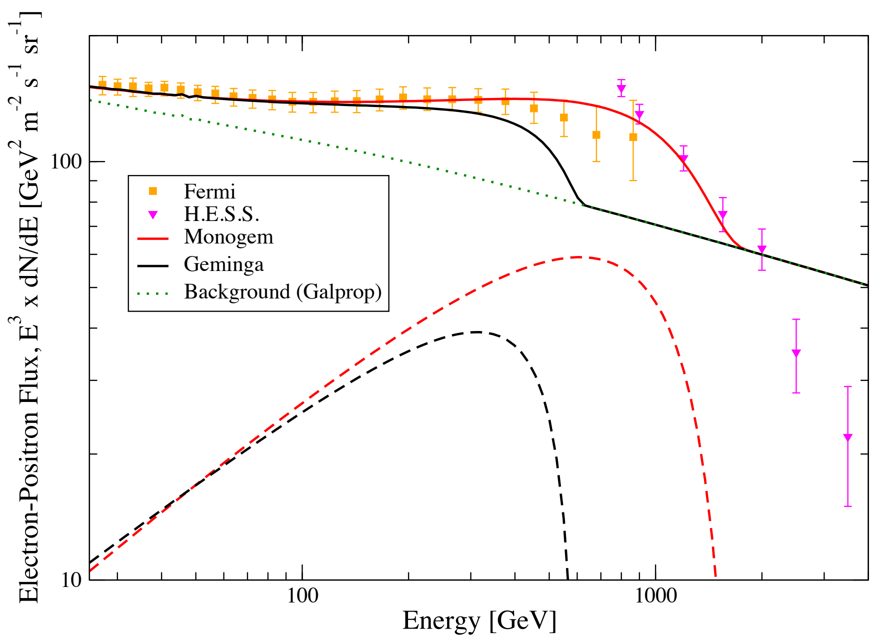
for burst-like injection from point-like source, diffusion equation can be easily solved:

$$J = \frac{\mathrm{e}^{-\vec{r}^2/\ell^2(E_0,E)}}{(\pi\ell^2(E_0,E))^{3/2}} Q(E_0) \left(\frac{E}{E_0}\right)^{-2} \quad \text{where} \quad \ell^2(E_0,E) = 4\int_{E_0}^{E} \mathrm{d}E' \frac{D(E')}{b(E')}$$

three-parameter-problem: total energy $\varepsilon_{\mathrm{tot}}$, spectra index Γ, cut-off energy E_{cut}

Linden & Profumo, arXiv:1304.1791





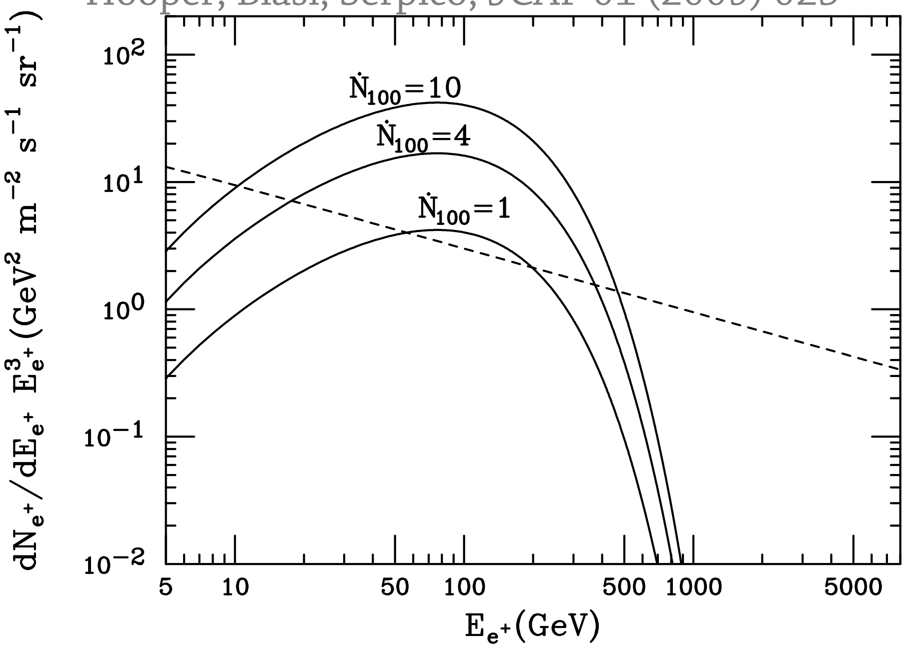
All Galactic pulsars

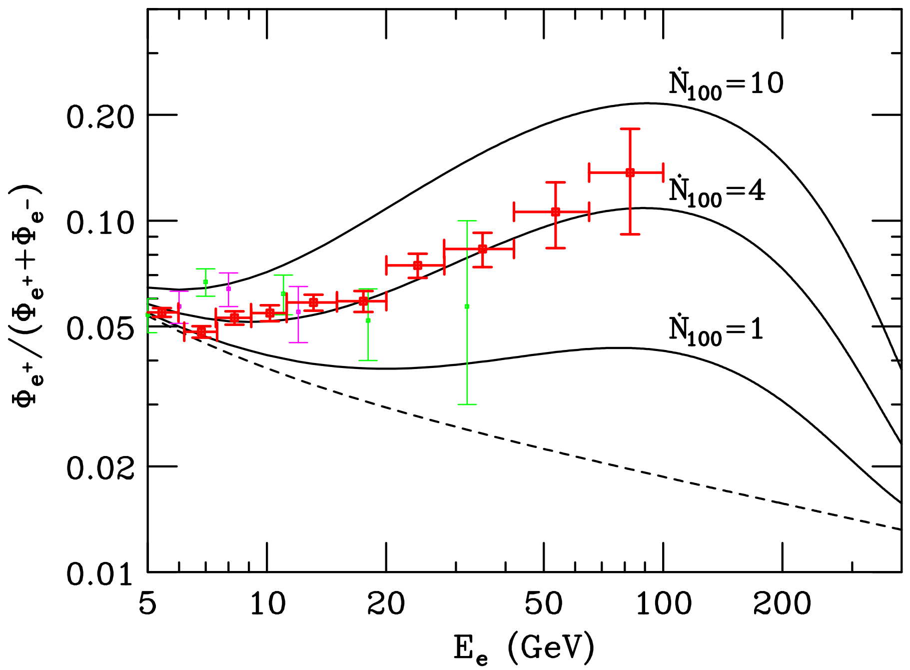
for burst-like injection from point-like source, diffusion equation can be easily solved:

$$J = \frac{\mathrm{e}^{-\vec{r}^2/\ell^2(E_0,E)}}{(\pi\ell^2(E_0,E))^{3/2}} Q(E_0) \left(\frac{E}{E_0}\right)^{-2} \quad \text{where} \quad \ell^2(E_0,E) = 4\int_{E_0}^{E} \mathrm{d}E' \frac{D(E')}{b(E')}$$

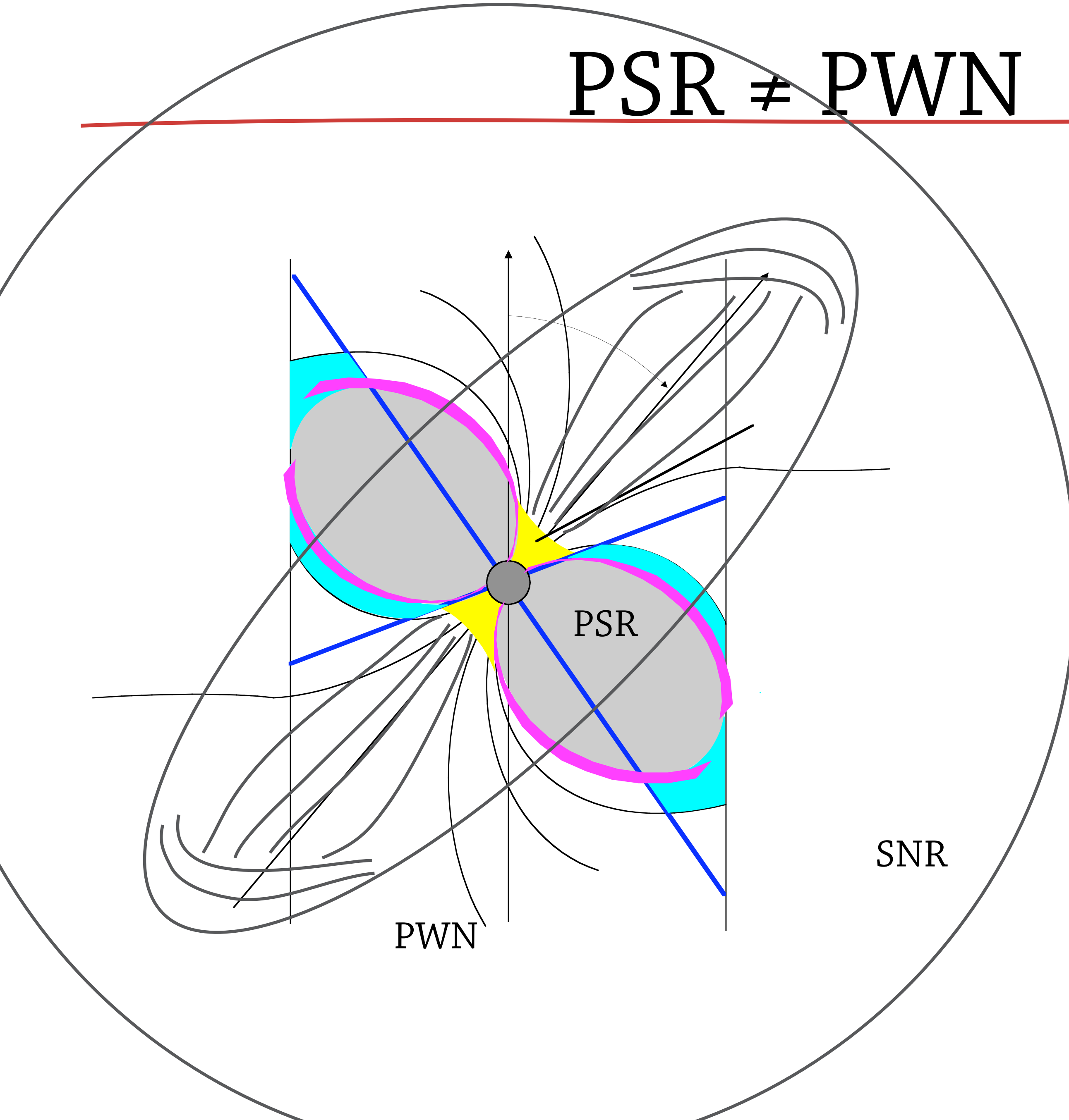
three-parameter-problem: total energy $\varepsilon_{\rm tot}$, spectra index Γ, cut-off energy $E_{\rm cut}$

Hooper, Blasi, Serpico, JCAP 01 (2009) 025





PSR ≠ PWN



freshly accelerated particles cannot escape the PWN/SNR

the PWN/SNR modifies the spectrum of $e^{\pm}$:

- cooling
- shock acceleration
- turbulent acceleration

all the parameters from pulsar observation, i.e.
$\varepsilon_{\rm tot}$, Γ, $E_{\rm cut}$
are modified

➔need statistical model and fit parameters

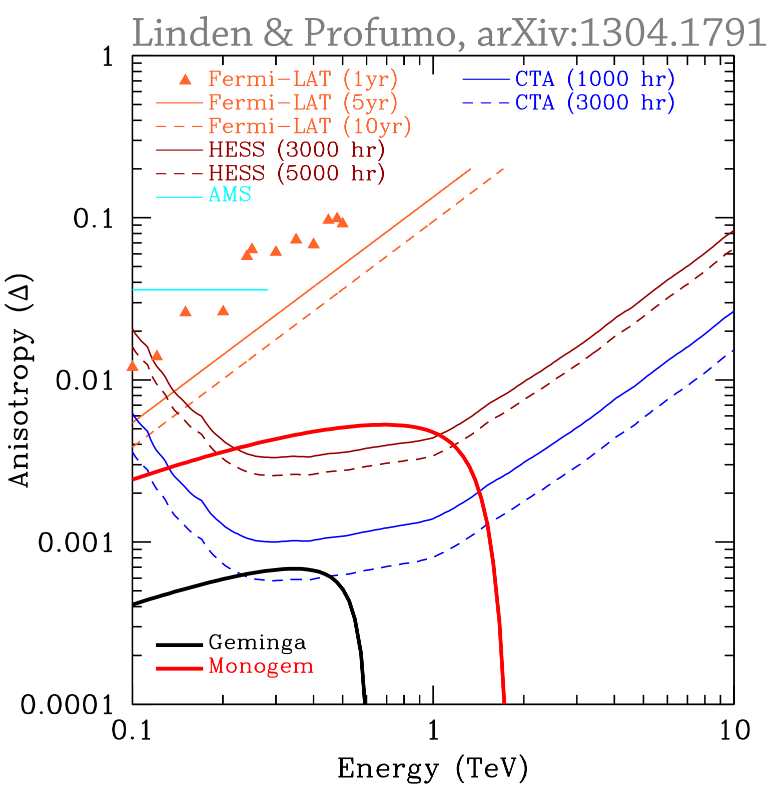
Malyshev, Cholis, Gelfand
PRD 80 (2009) 063005

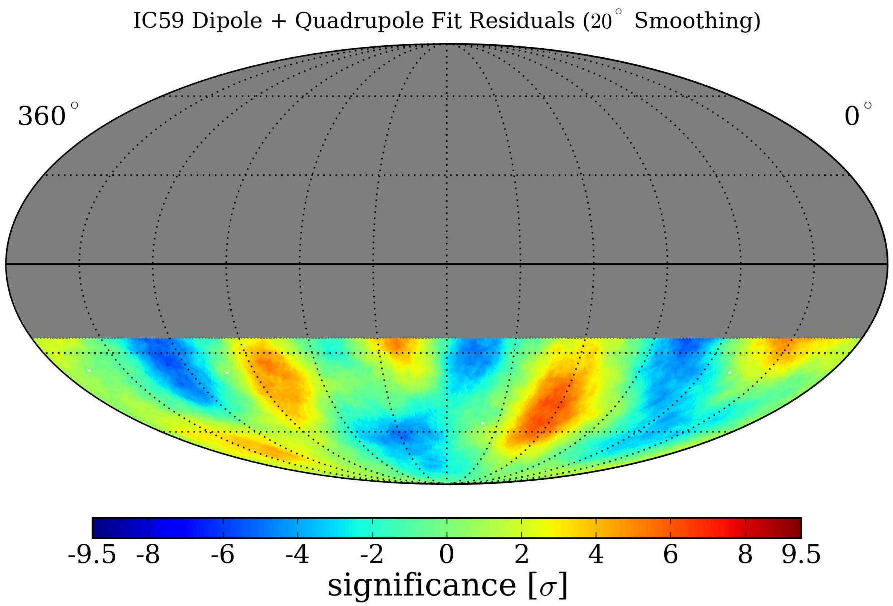
Anisotropies

can estimate anisotropy from isotropic flux:

$$\delta = \frac{3D(E)}{c} \frac{|\nabla J(E)|}{J(E)}$$

very sensitive to *local* $D(E)$





anisotropies in (hadronic) cosmic rays:

- $\delta \sim 10^{-3}$ on large (dipole) scales
- $\delta \sim 10^{-4}$ down to 10° scales

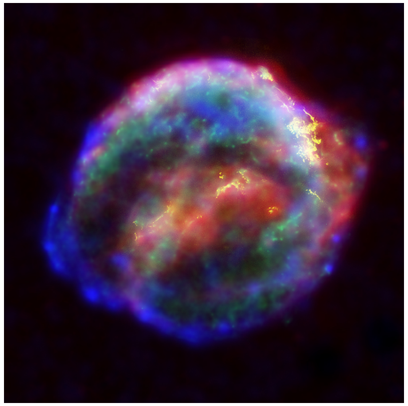
explanations modify local magnetic field

Malkov *et al.*, ApJ 721 (2010) 750

Giacinti & Sigl, PRL 109 (2012) 071101

Secondaries from the Source?

Common belief: secondaries from propagation dominate since the grammage in the ISM is larger than in the source



$$\langle \tau_{\rm src} \rangle \lesssim \tau_{\rm SNR} \approx 10^{4\ldots5}\,\mathrm{yr}$$

$$n_{\rm src} \lesssim 10\,\mathrm{cm}^{-3}$$

$$\Rightarrow \lambda_{\rm src} \approx 0.2\,\mathrm{g\,cm}^{-2}$$

$$\langle \tau_{\rm ISM} \rangle \sim \tau_{\rm esc} \approx 10^{7}\,\mathrm{yr}$$

$$n_{\rm ISM} \approx 0.1\,\mathrm{cm}^{-3}$$

$$\Rightarrow \lambda_{\rm ISM} \approx \mathrm{few}\,\mathrm{g\,cm}^{-2}$$

However, the secondaries from the source can have a much harder spectrum!

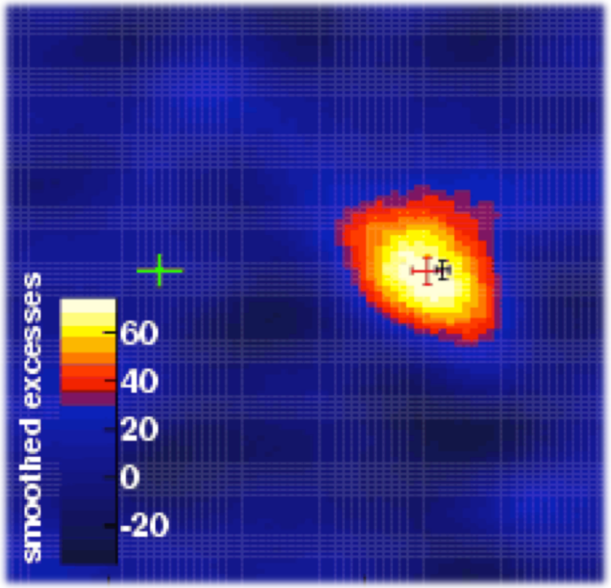
Secondary Origin of $e^{\pm}$

Rise in positron fraction could be due to secondary positrons produced during acceleration and accelerated along with primary electrons

Blasi, PRL **103** (2009) 051105

Assuming production of galactic CR in SNRs, positron fraction can be fitted

This effect is guaranteed, only its size depends on normalisation and one free parameter that needs to be fitted from observations



Cas A in γ-rays from MAGIC

DSA – Test Particle Approximation

Acceleration determined by compression ratio:

$$r = \frac{u_1}{u_2} = \frac{n_2}{n_1}, \qquad \gamma = \frac{3r}{r-1}$$

Solve transport equation,

$$u\frac{\partial f}{\partial x} = D\frac{\partial^2 f}{\partial x^2} + \frac{1}{3}\frac{\mathrm{d}u}{\mathrm{d}x}p\frac{\partial f}{\partial p}$$

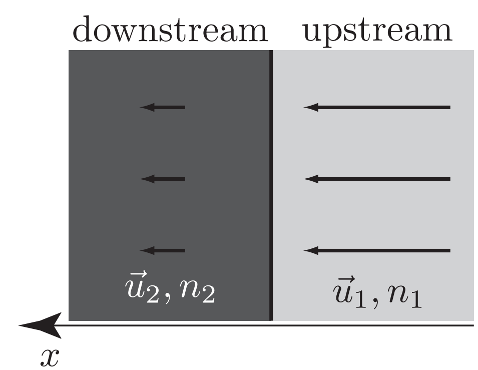
$$f \xrightarrow{x\to-\infty} f_{\rm inj}(p), \qquad \left|\lim_{x\to\infty} f\right| \ll \infty$$

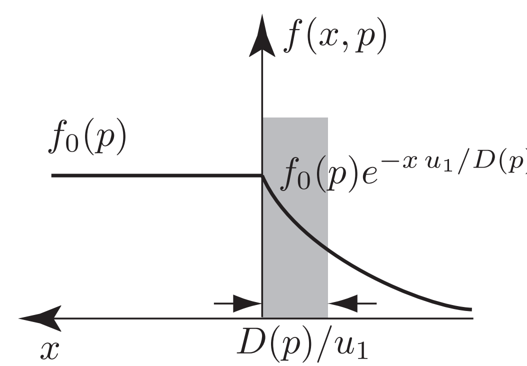
Solution for $x < 0$:

$$f = f_{\rm inj}(p) + (f^0(p) - f_{\rm inj}(p))e^{-x\,u_1/D(p)}$$

where

$$f^0(p) = \gamma\int_0^p \frac{\mathrm{d}p'}{p'}\left(\frac{p'}{p}\right)^\gamma f_{\rm inj}(p') + Cp^{-\gamma}$$





As long as $f_{\rm inj}(p)$ is softer than $p^{-\gamma}$, at high energies:

$$f(x,p) \sim p^{-\gamma}$$

DSA with Secondaries

- Secondaries get produced with primary spectrum:

$$q_{e^\pm} \propto f_{\rm CR} \propto p^{-\gamma}$$

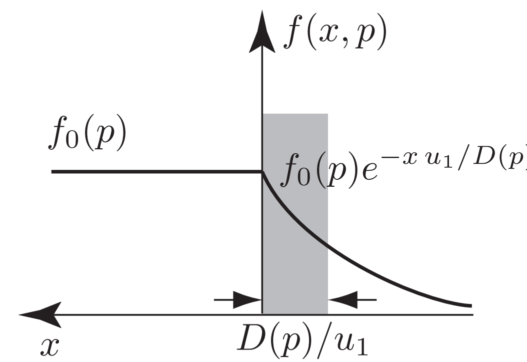
- Only particles with $|x| \lesssim D(p)/u$ can be accelerated

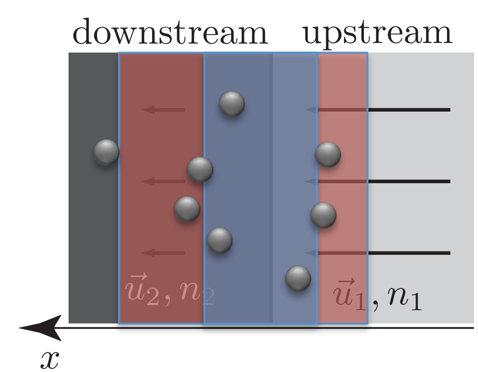
- Bohm diffusion: $D(p) \propto p$

- Fraction of secondaries that go into acceleration $\propto p$

- Equilibrium spectrum

$$n_{e^\pm} \propto q_{e^\pm} \left(1 + \frac{p}{p_0}\right) \propto p^{-\gamma} + p^{-\gamma+1}$$





$\mathbf{p_2} > \mathbf{p_1}$

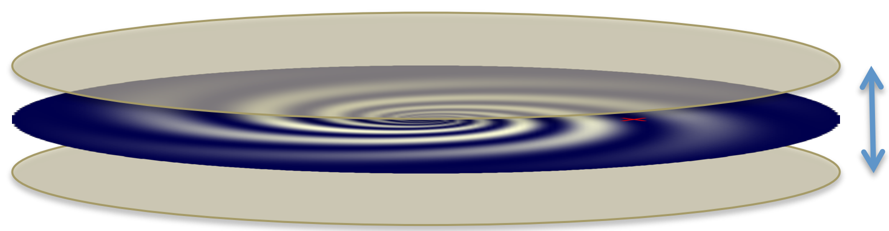
Rising positron fraction at source

Diffusion of GCRs

Transport equation:

$$\frac{\mathrm{d}n(\vec{r},t)}{\mathrm{d}t} = \underbrace{\nabla(D\nabla n(\vec{r},t))}_{\text{diffusion}} - \underbrace{\frac{\partial}{\partial E}(b(E)n(r,t))}_{\text{energy losses}} + \underbrace{q(\vec{r},t)}_{\text{injection}}$$

Boundary conditions:

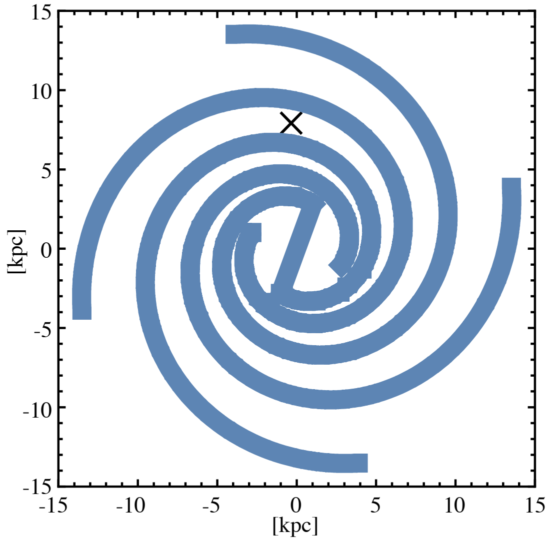


$2h$

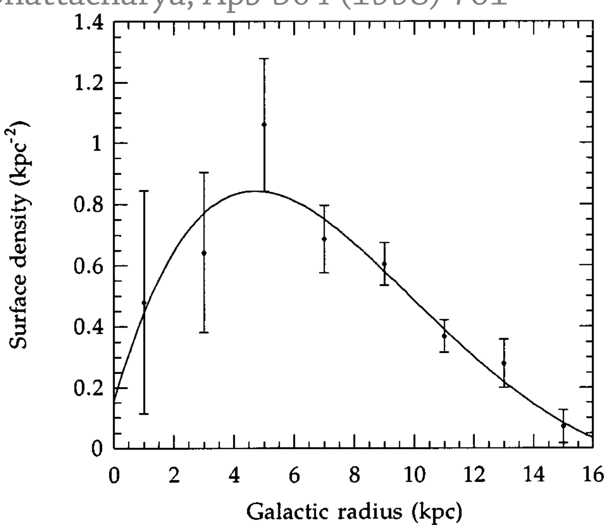
Green's function:

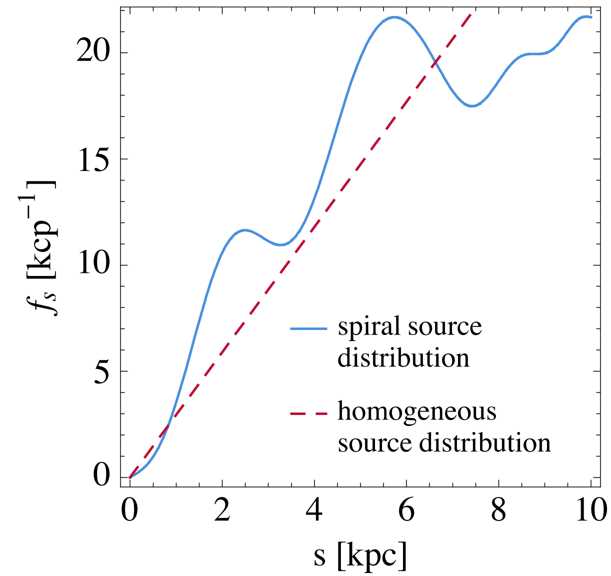
describes flux from one discrete, burst-like source

Statistical Distribution of Sources



Case, Bhattacharya, ApJ 504 (1998) 761



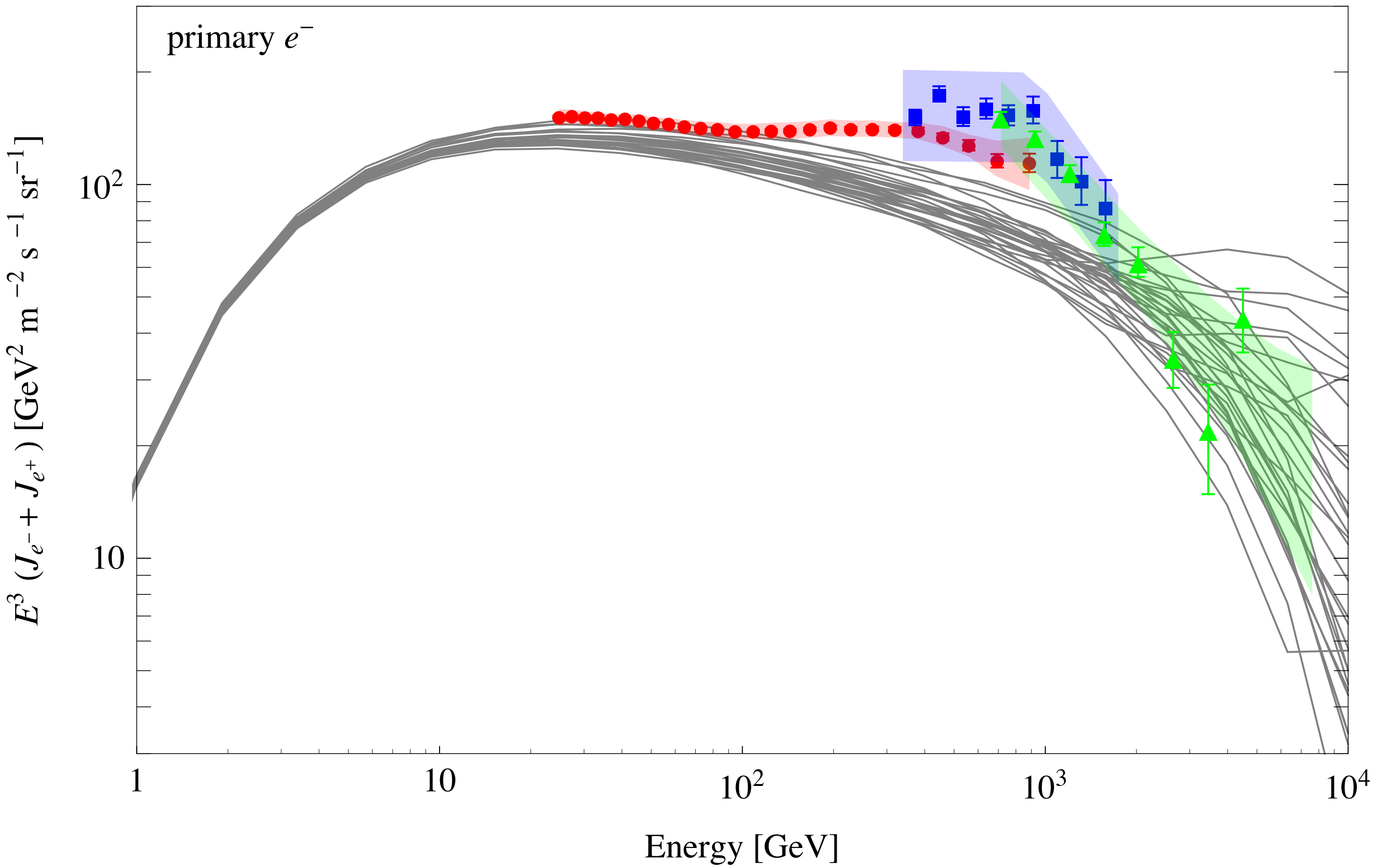


probability density for distances, $f_s(s)$

probability density for ages, $f_t(t) = \text{const.}$

Ahlers, Mertsch, Sarkar, PRD **80** (2009) 123017

The Total $(e^+ + e^-)$ Flux



Ahlers, Mertsch, Sarkar, PRD **80** (2009) 123017
Mertsch & Sarkar, *in preparation*

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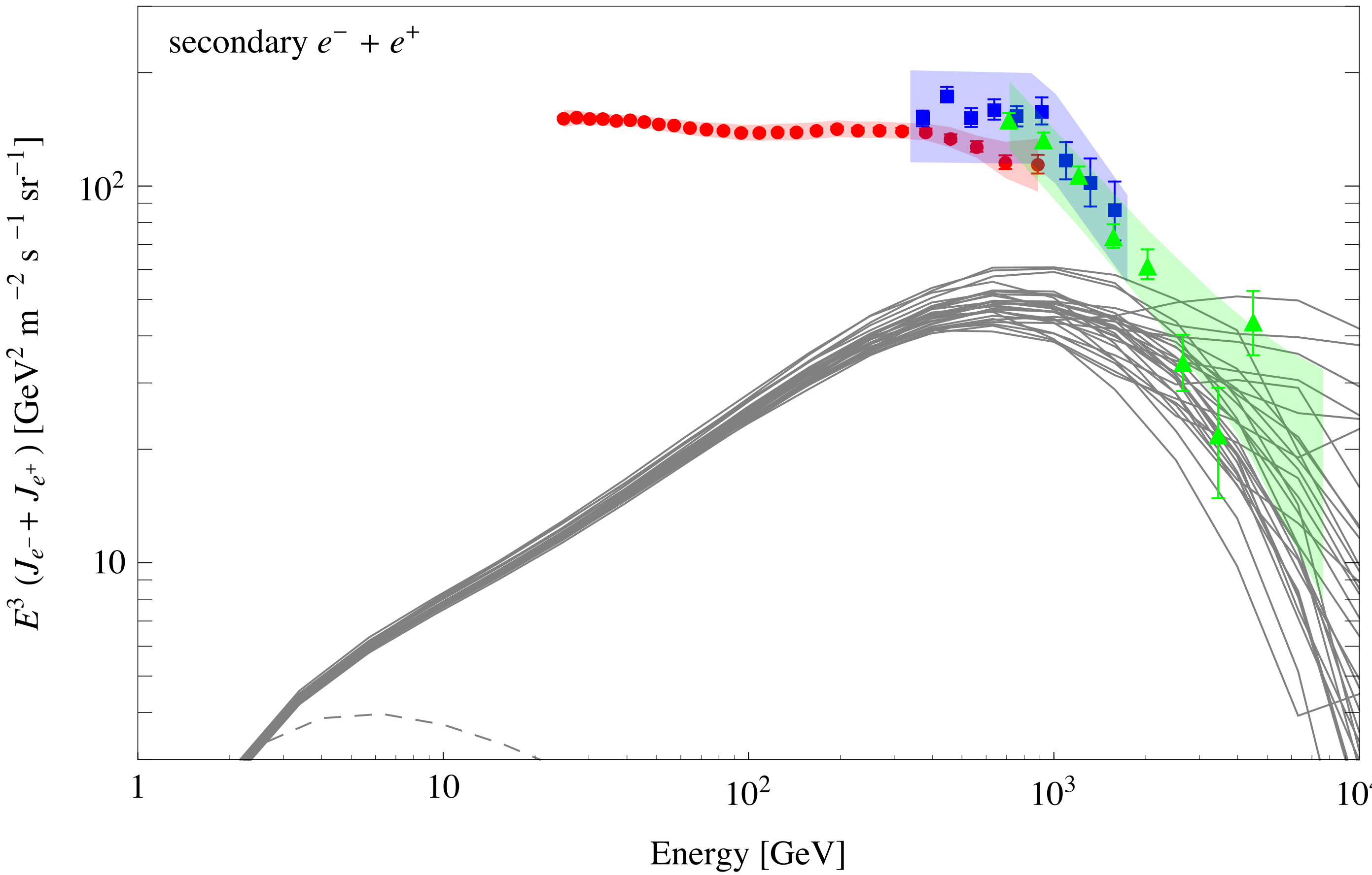

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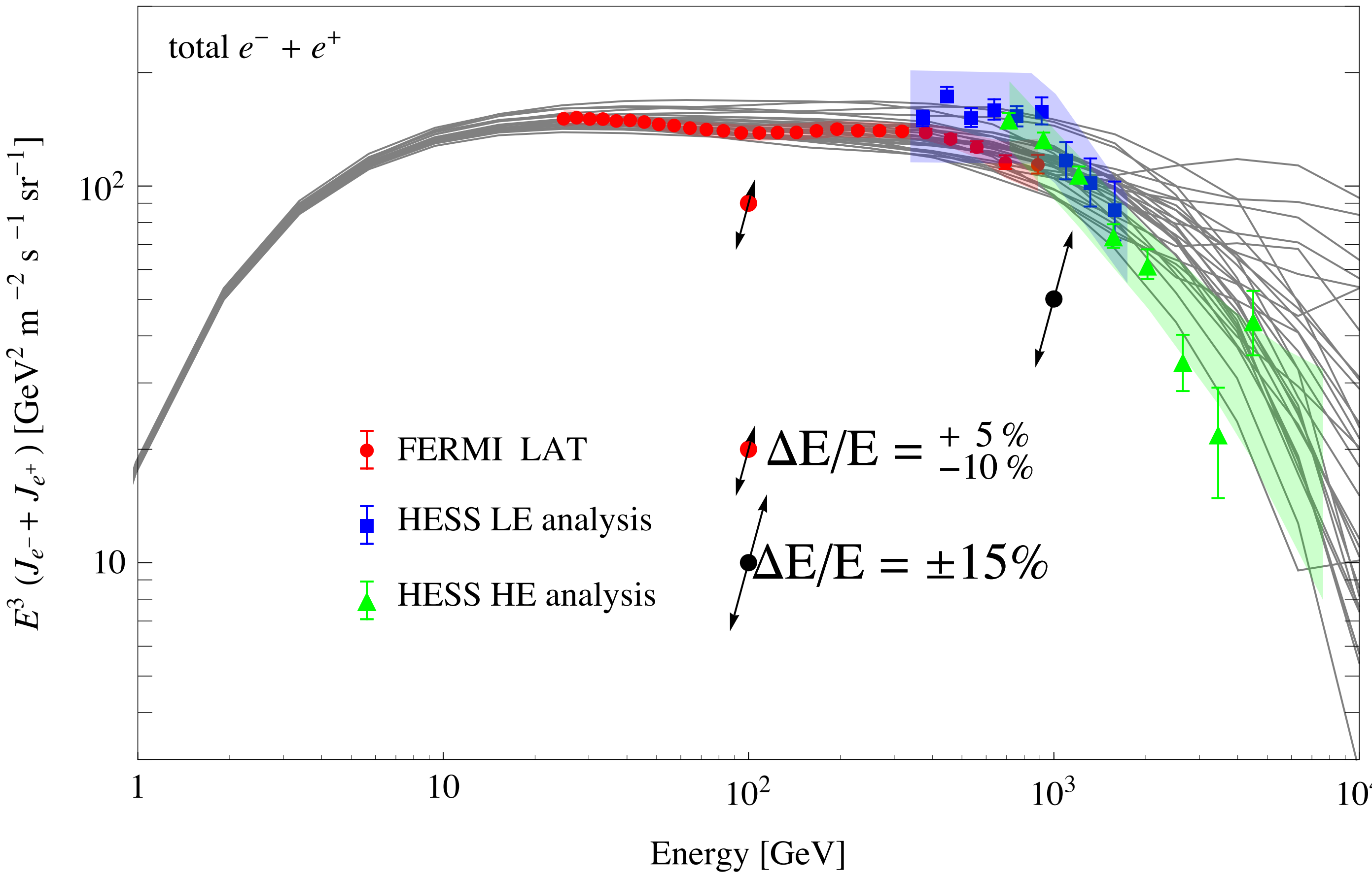

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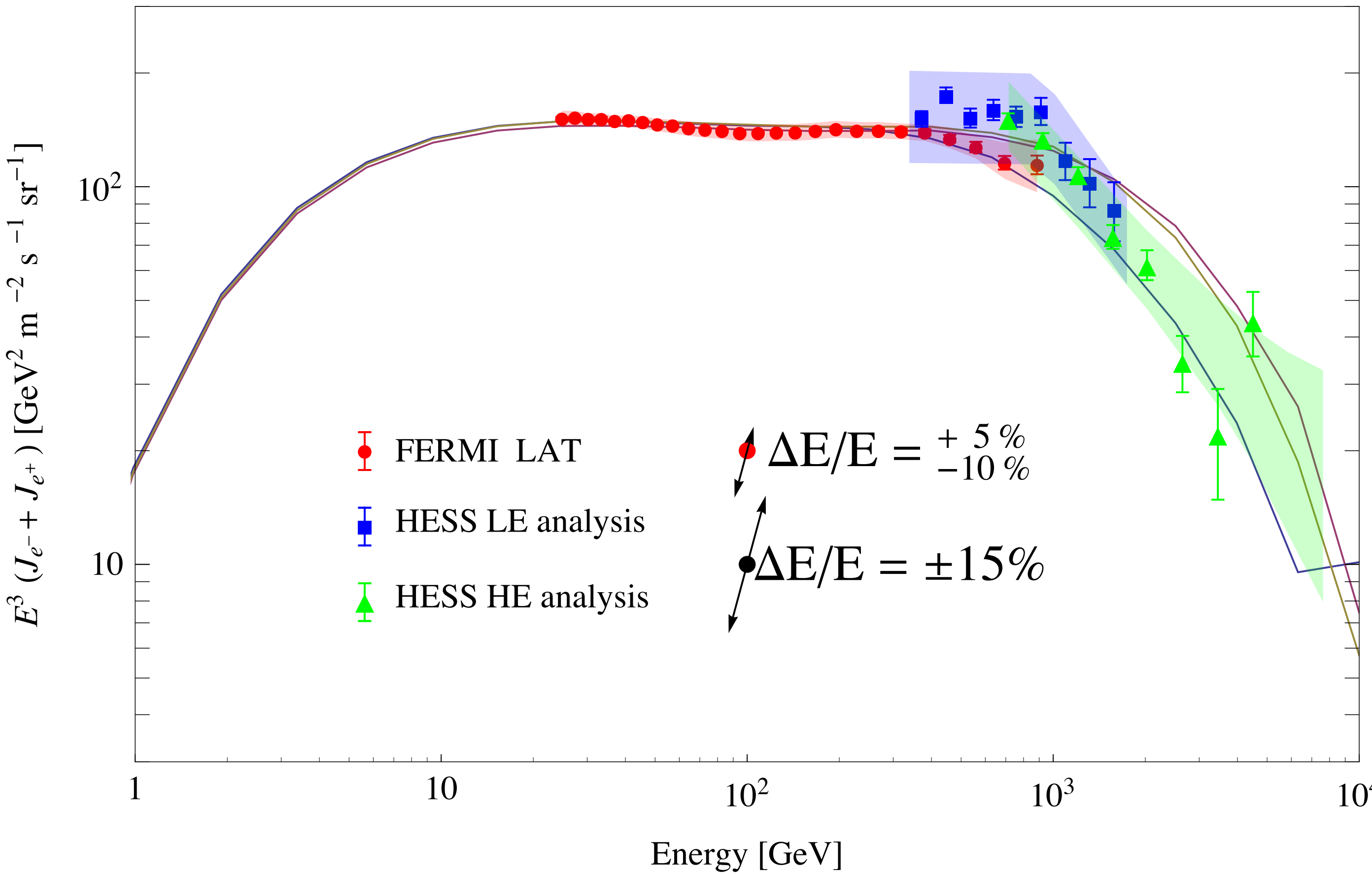
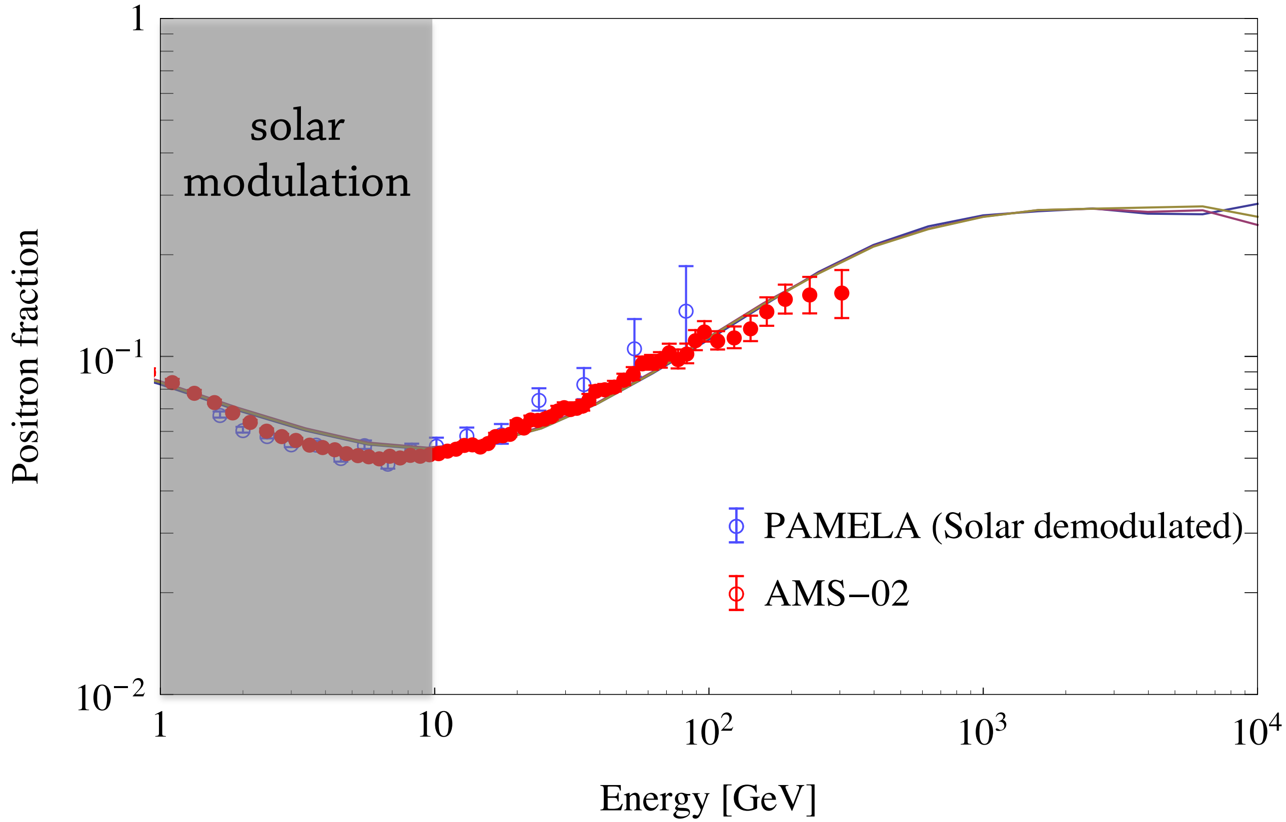


Ahlers, Mertsch, Sarkar, PRD **80** (2009) 123017
Mertsch & Sarkar, *in preparation*

The Positron Fraction



Ahlers, Mertsch, Sarkar, PRD **80** (2009) 123017
Mertsch & Sarkar, *in preparation*

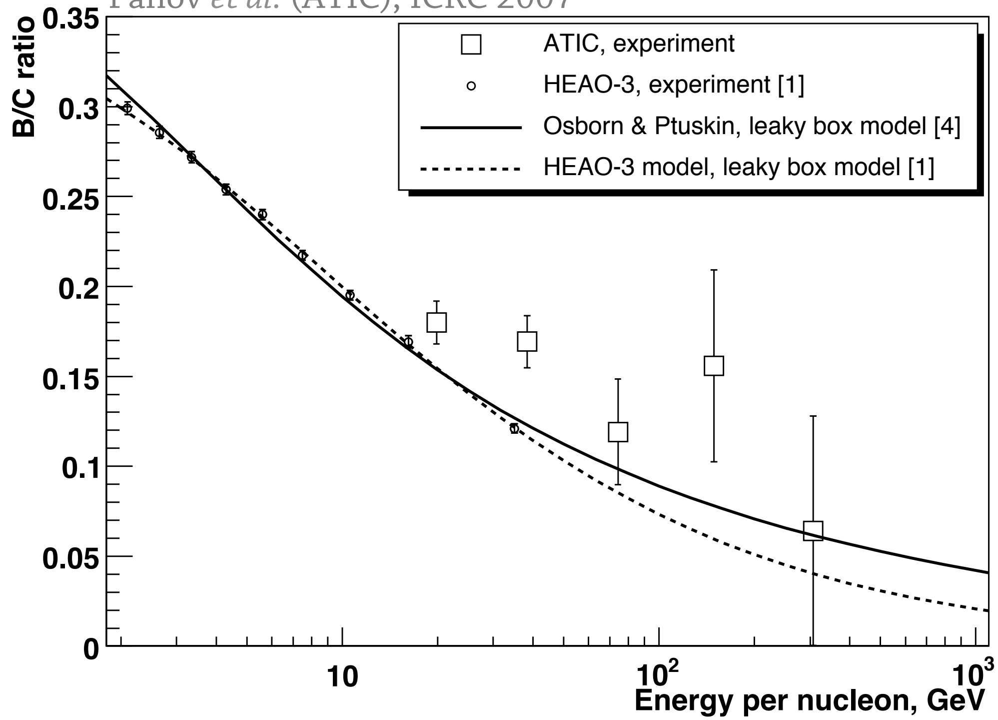
Nuclear Secondary-to-Primary Ratios

rise in...	nuclei
DM	✗
Pulsars	✗

DM and pulsars do not produce nuclei!

Nuclear secondary-to-primary ratios used for testing and calibrating propagation models

Panov *et al.* (ATIC), ICRC 2007



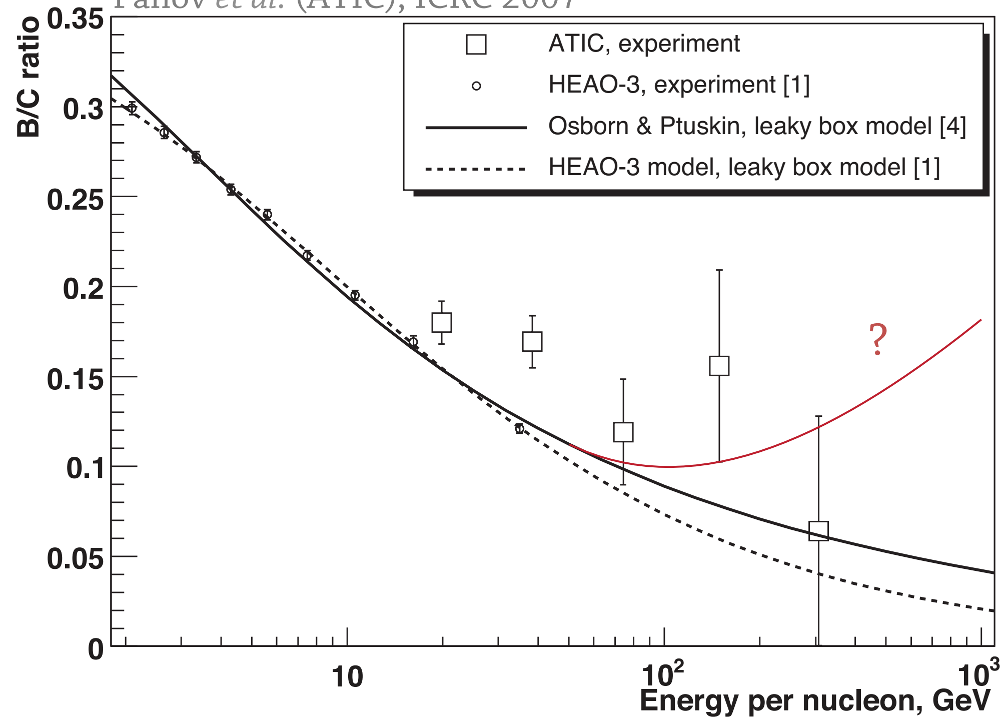
Nuclear Secondary-to-Primary Ratios

rise in...	nuclei
DM	✗
Pulsars	✗
Acceleration of Secondaries	✓

This would be a clear indication for acceleration of secondaries!

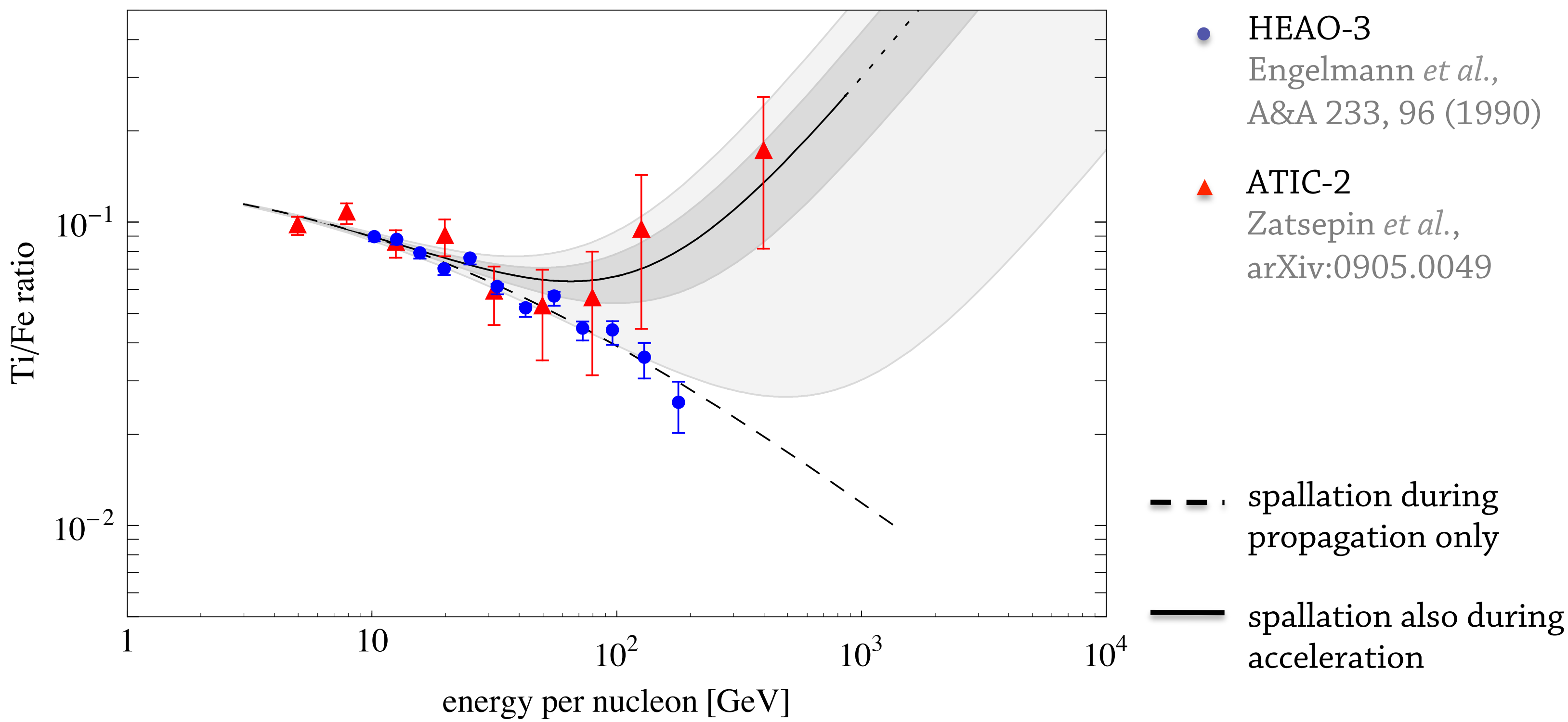
If nuclei are accelerated in the same sources as electrons and positrons, nuclear ratios *must* rise eventually

Panov *et al.* (ATIC), ICRC 2007



Titanium-to-Iron Ratio

PM and Sarkar, PRL **103** (2009) 081104

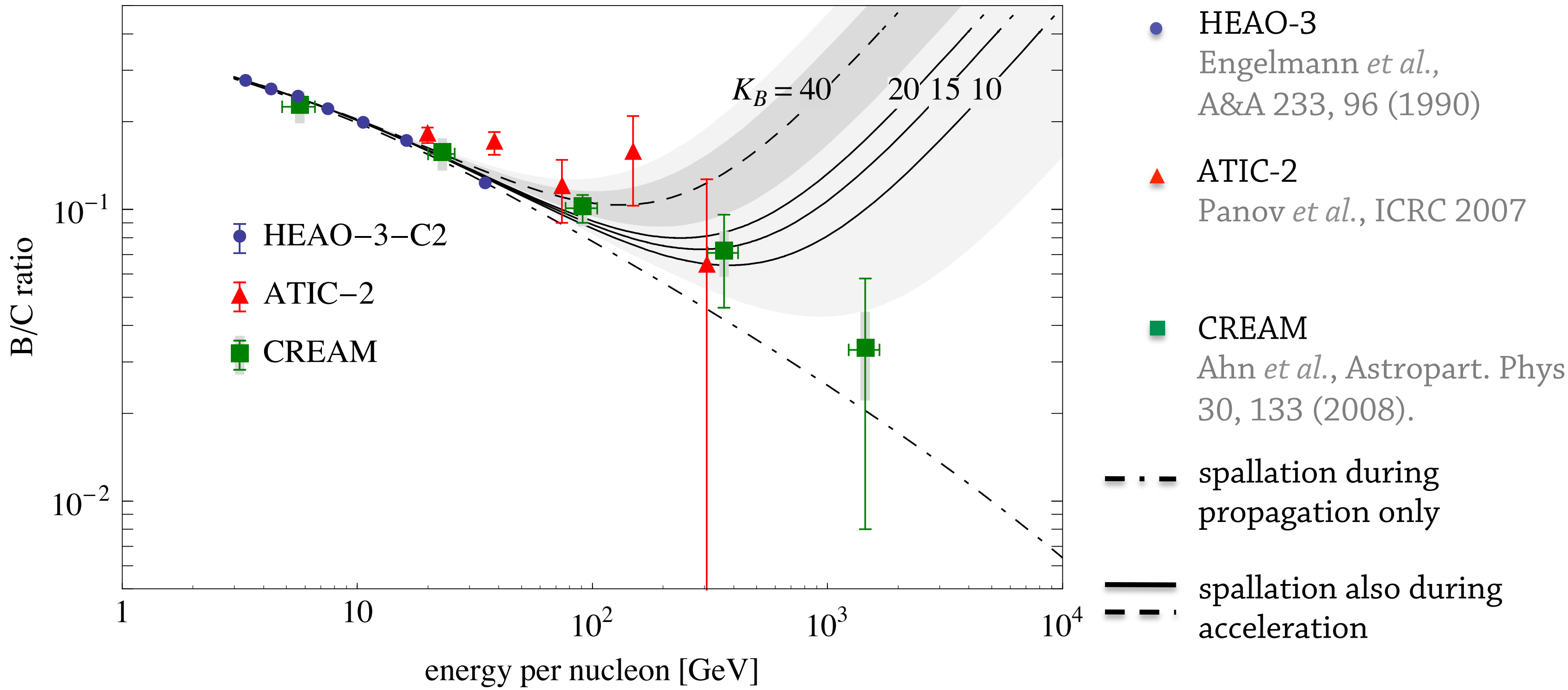


Titanium-to-iron ratio used as calibration point for diffusion coefficient:

$$K_{\mathrm{B}} \simeq 40$$

Boron-to-Carbon Ratio

PM and Sarkar, PRL **103** (2009) 081104; Ahlers *et al.*, PRD **80** (2009) 123017

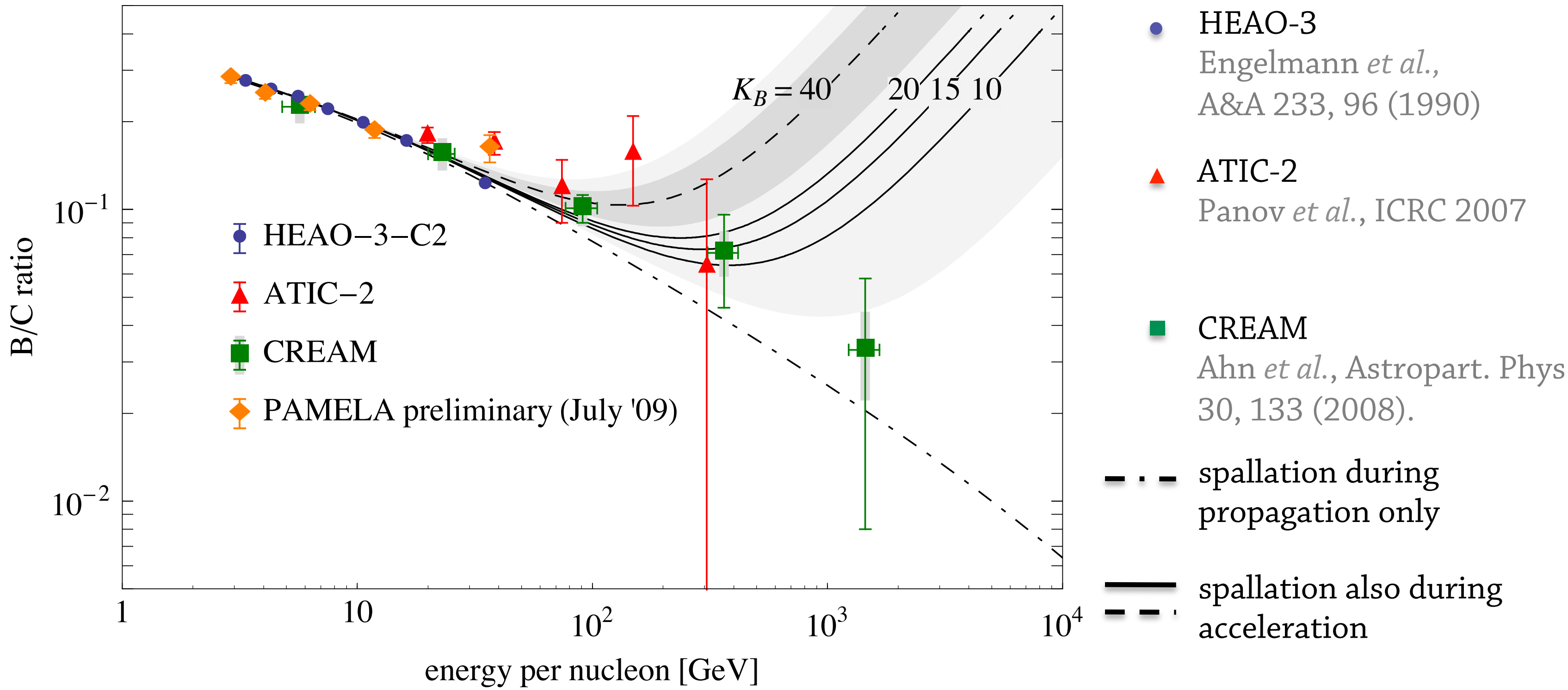


PAMELA is currently measuring B/C with unprecedented accuracy

A rise would rule out the DM and pulsar explanation of the PAMELA e^+/e^- excess.

Boron-to-Carbon Ratio

PM and Sarkar, PRL **103** (2009) 081104; Ahlers *et al.*, PRD **80** (2009) 123017

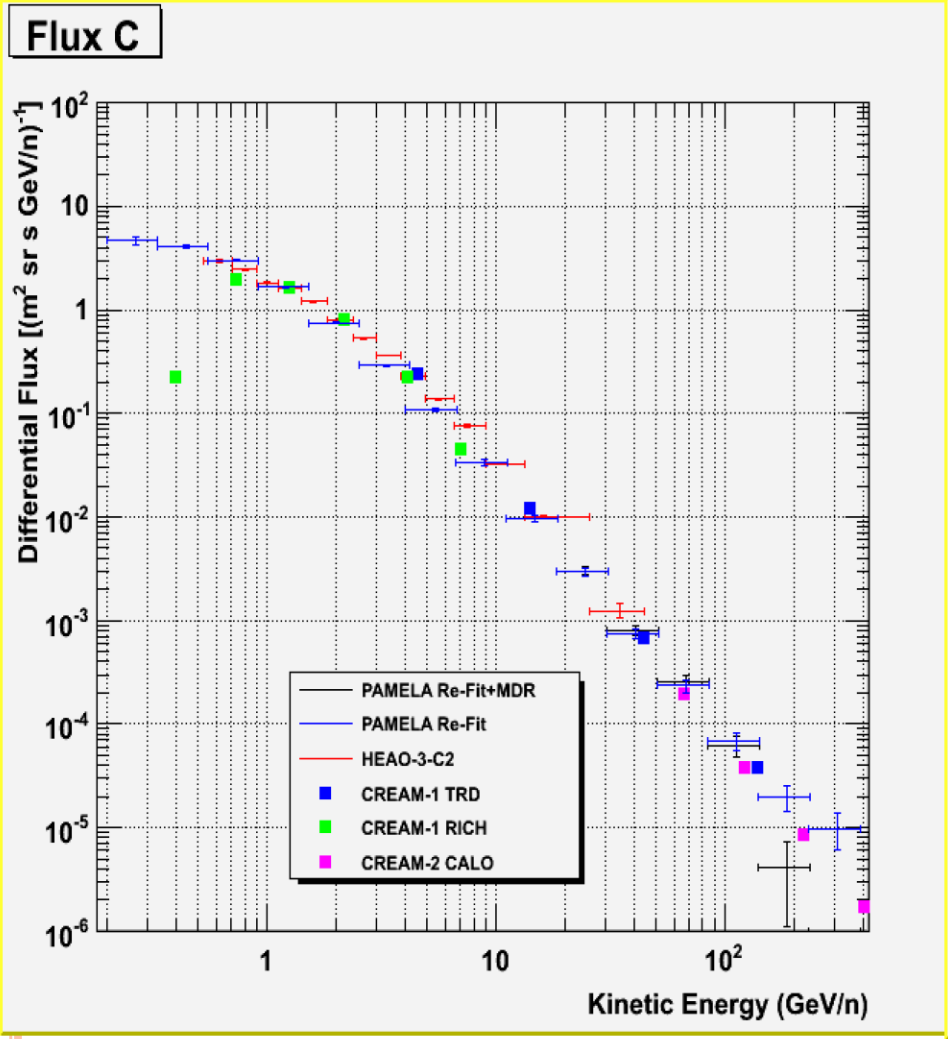


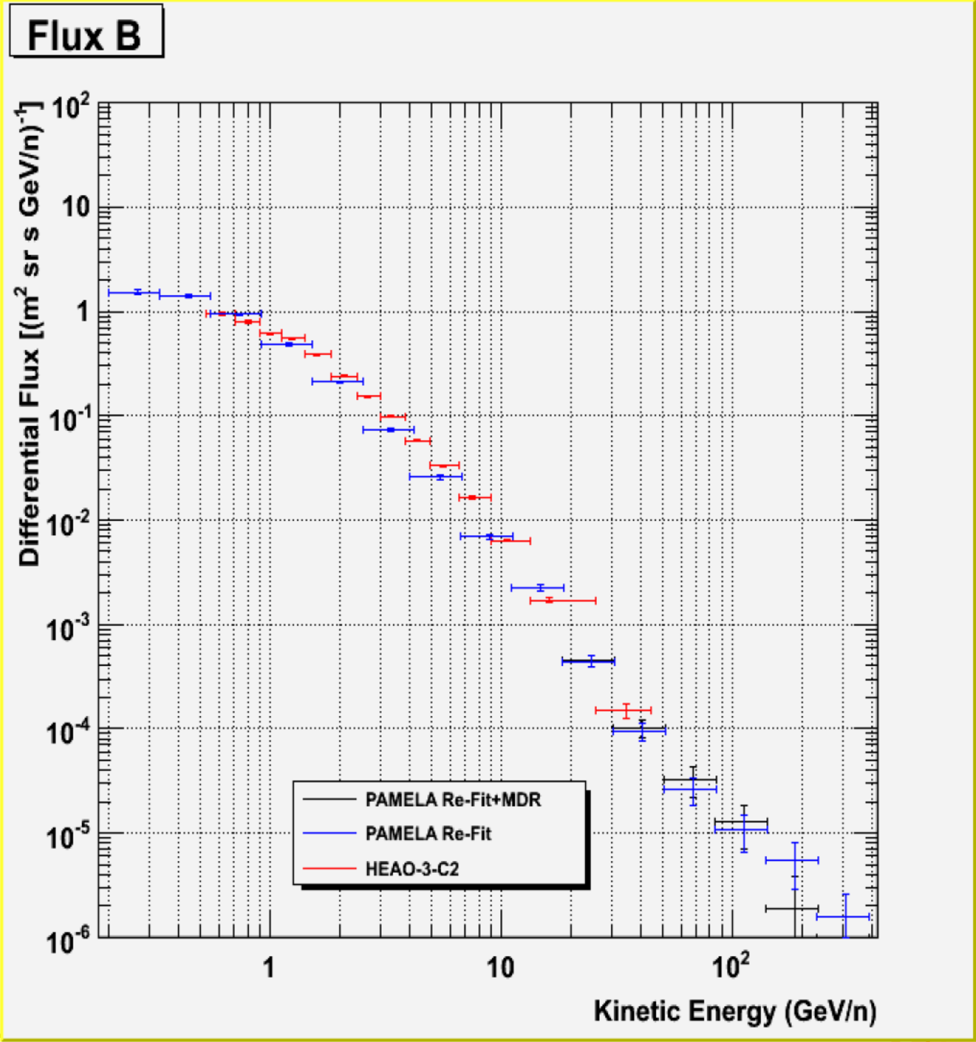
PAMELA is currently measuring B/C with unprecedented accuracy

A rise would rule out the DM and pulsar explanation of the PAMELA e^+/e^- excess.

Boron and Carbon flux – in progress

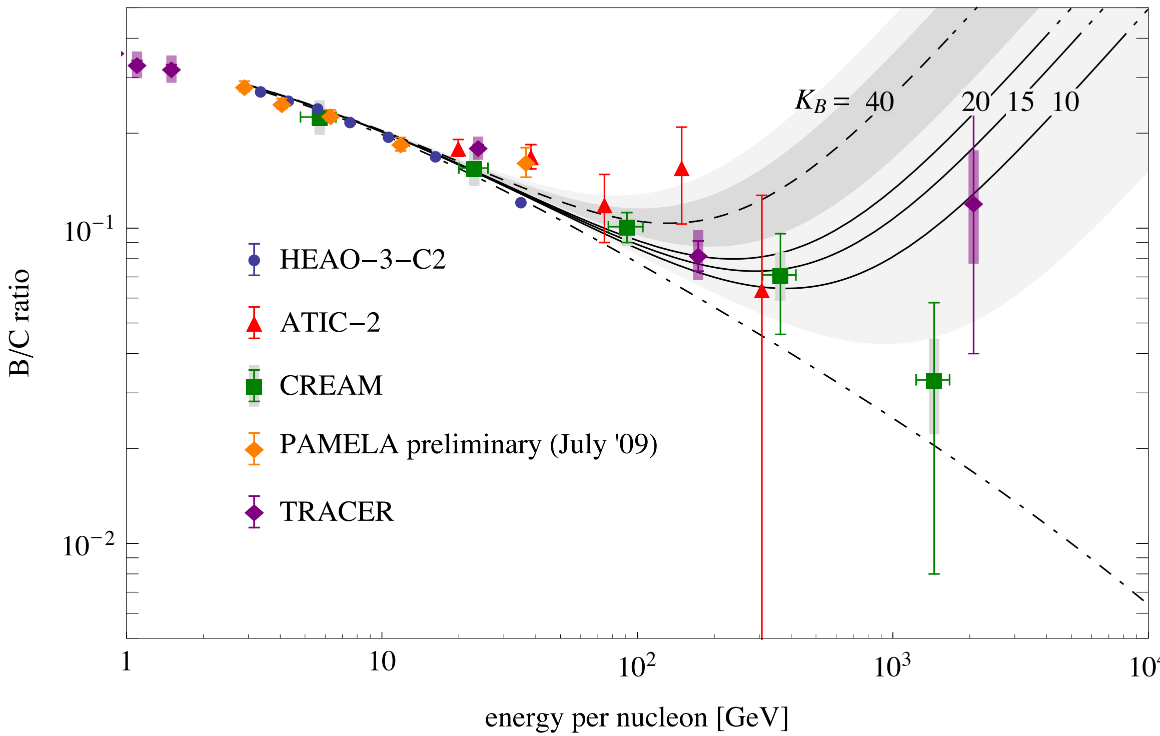
R. Sparvoli, 6th Patras Workshop on Axions, WIMPs and WISPs, 5-9 July 2010





Boron-to-Carbon Ratio

PM and Sarkar, PRL **103** (2009) 081104; Ahlers *et al.*, PRD **80** (2009) 123017



HEAO-3
Engelmann *et al.*, A&A 233, 96 (1990)

ATIC-2
Panov *et al.*, ICRC 2007

CREAM
Ahn *et al.*, Astropart. Phys. 30, 133 (2008).

spallation during propagation only

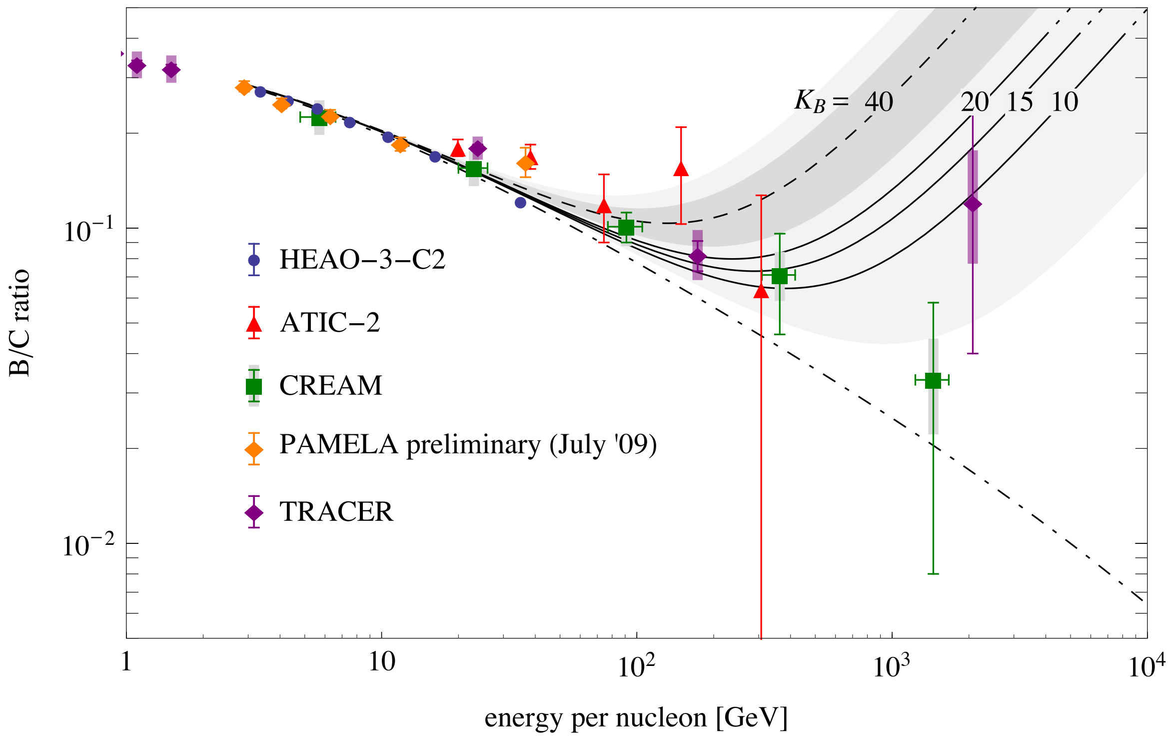
spallation also during acceleration

PAMELA is currently measuring B/C with unprecedented accuracy

A rise would rule out the DM and pulsar explanation of the PAMELA e^+/e^- excess.

Boron-to-Carbon Ratio

PM and Sarkar, PRL **103** (2009) 081104; Ahlers *et al.*, PRD **80** (2009) 123017



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AMS-02

~~PAMELA~~ is currently measuring B/C with unprecedented accuracy

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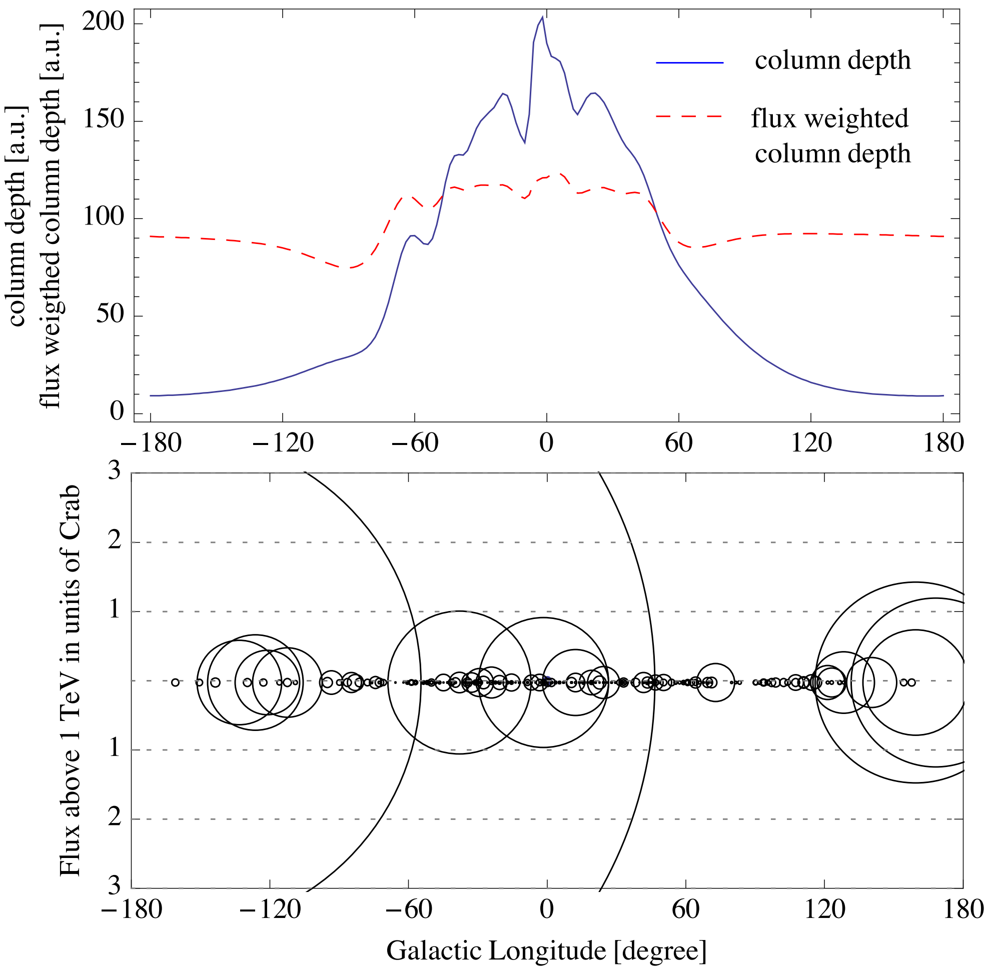
Hint at Hadronic SNRs



Maximum column depth around galactic centre

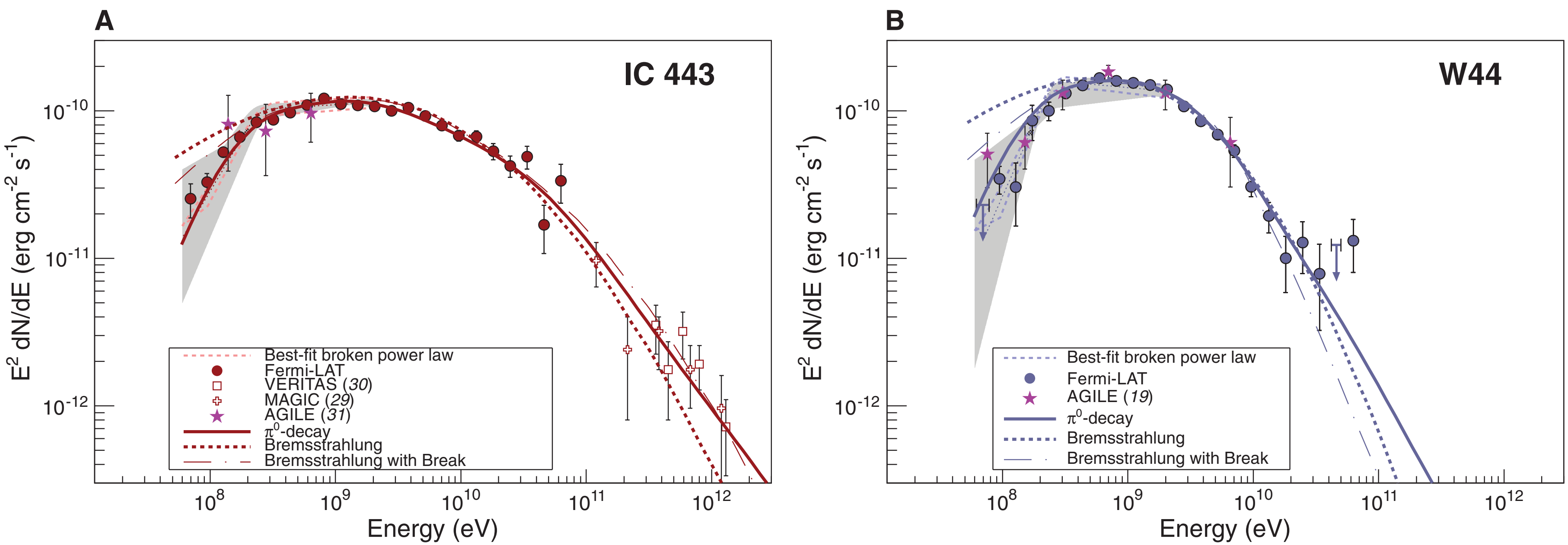
However brightness only smaller by 30% in rest of sky

on average:

- 3 sources brighter than Crab
- 7 sources brighter than 50% Crab

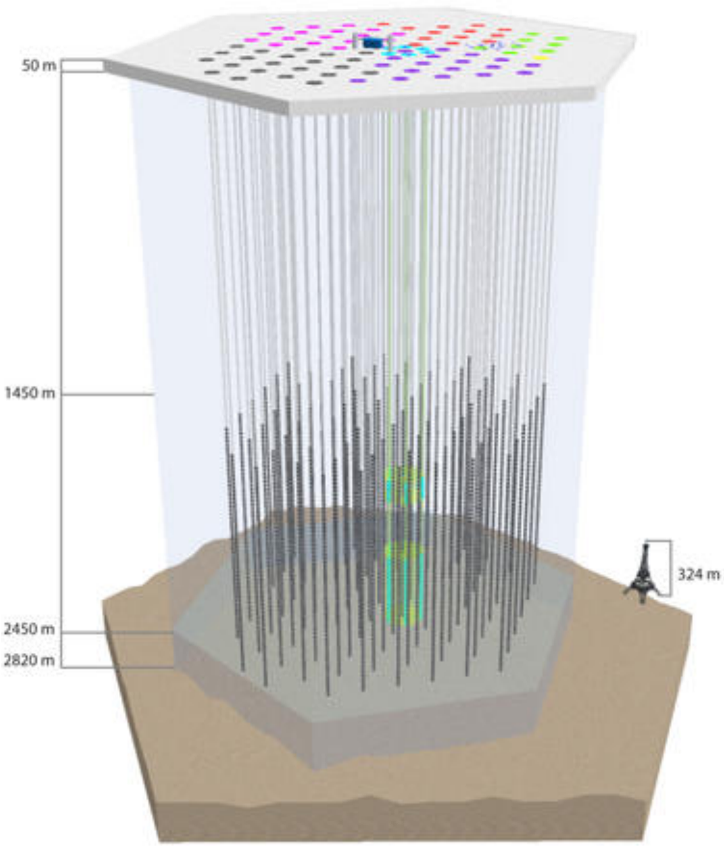
We have seen hadronic accelerators!

Ackermann *et al.*, Science 339 (2013) 807



- longstanding issue: hadronic or leptonic?
- low-energy analysis down to $100\ \mathrm{MeV}$
- π^0-bump observed, bremsstrahlung disfavoured

Prospects for IceCube



Flux from SNR at 2 kpc with $\Gamma = 2.4$ and above normalisation:

$$F_{\nu_\mu}(> 3\ \mathrm{TeV}) \simeq 7 \times 10^{-13} \mathrm{cm}^{-2}\,\mathrm{s}^{-1}$$

To be compared with IC40 + IC59 point source limit (90% CL upper limit on muon neutrino flux for energies between 3 TeV and 3 PeV):

$$F_{\nu_\mu} \lesssim 1.3 \times 10^{-12} \mathrm{cm}^{-2}\mathrm{s}^{-1}$$

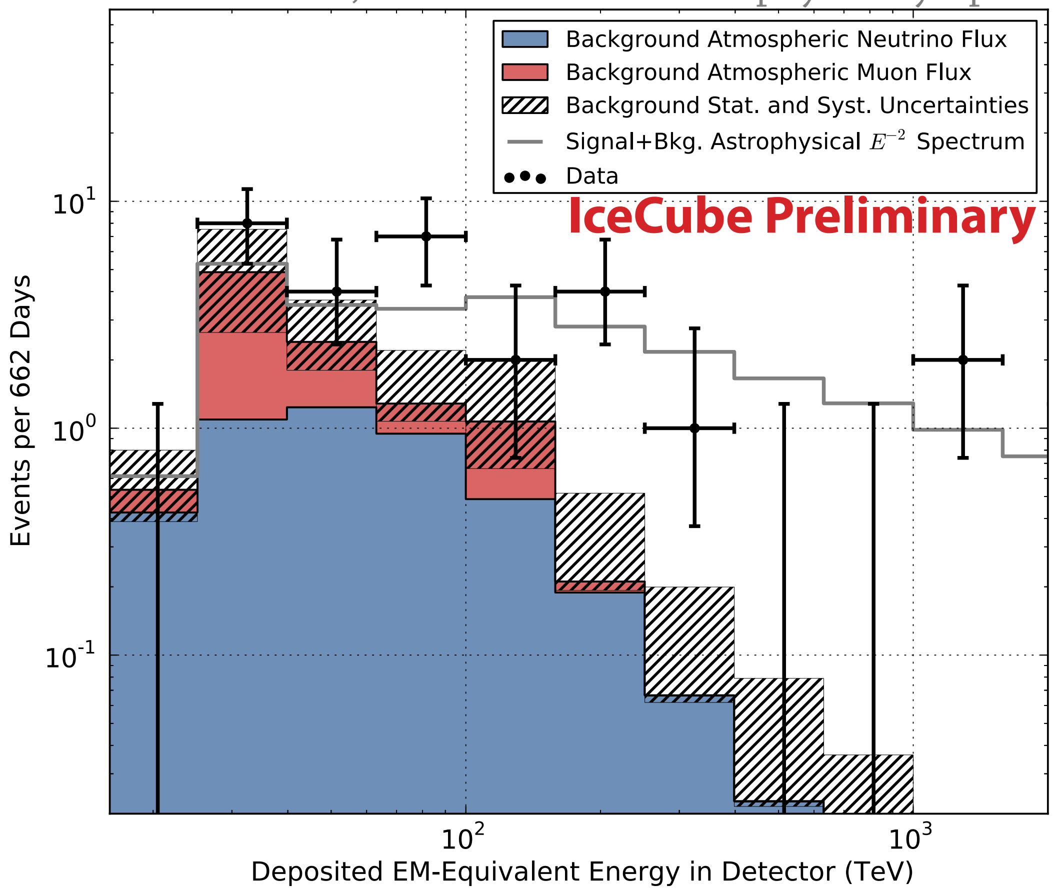
However, E^{-2} point source with

$$F_{\nu_\mu} \simeq 7.2 \times 10^{-12} \mathrm{cm}^{-2}\,\mathrm{s}^{-1}$$

can be detected in full IceCube (80 strings) with 5 σ significance in 3 years .

28 neutrinos

N. Whitehorn, IceCube Particle Astrophysics Symposium

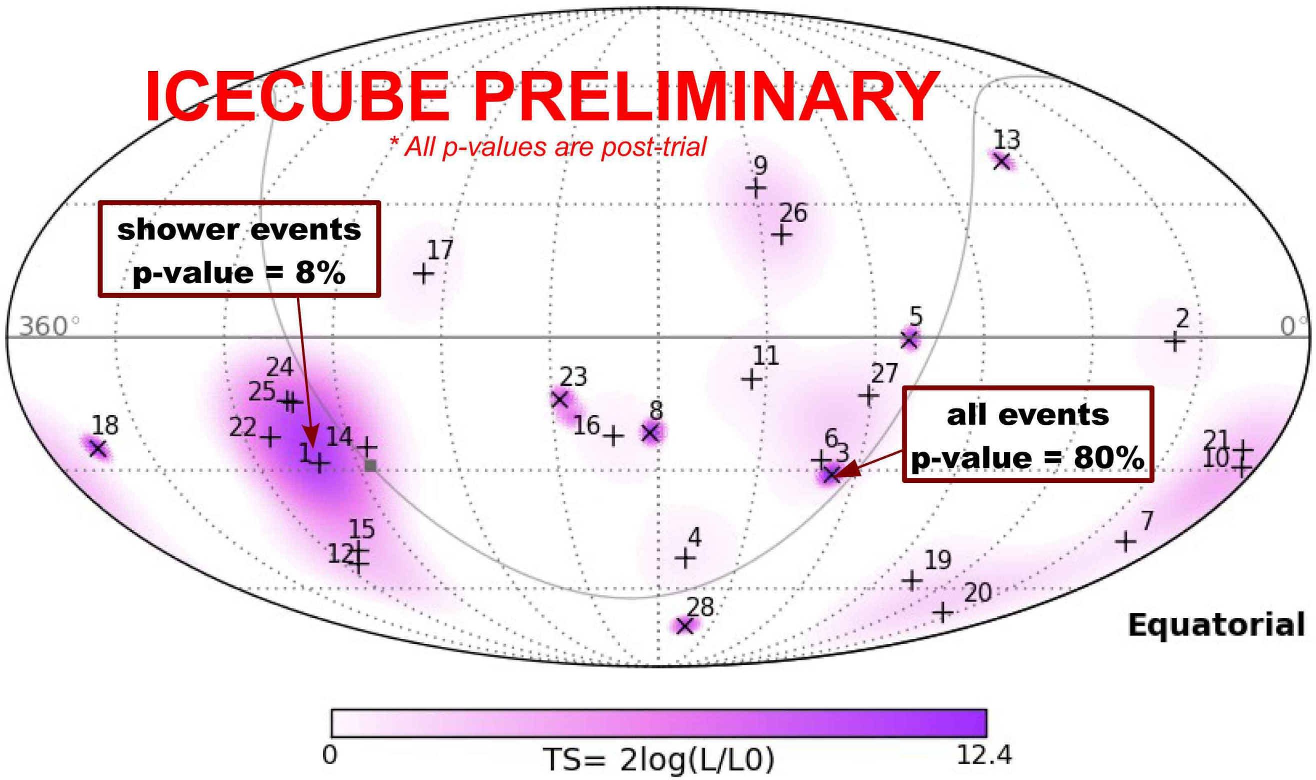


IceCube search for TeV-PeV νs:

- 2 yr of data (IC79 & IC86)
- look for contained events

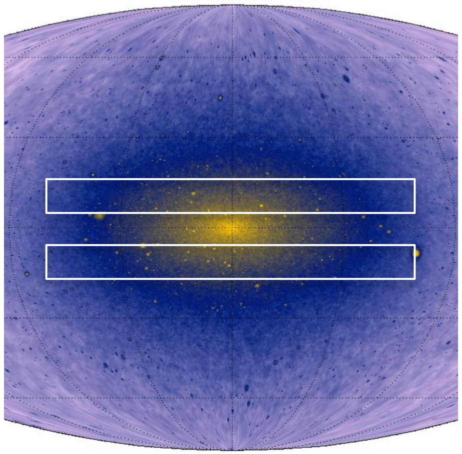
- 28 events (7 with visible μ, 21 without) on background of ~11
- 4.3σ excess over standard atmospheric backgrounds
- hard, $E^{-2.2}$ flux with cut-off at few PeV
- most likely not a *single* point source

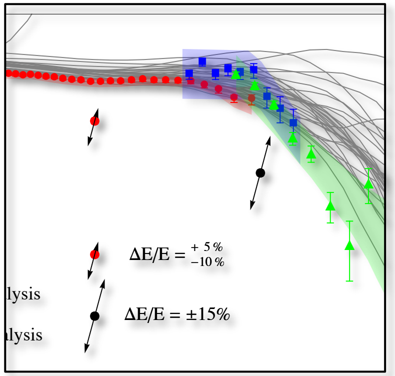
28 neutrinos



- "diffuse" flux: $E^2 \phi_{\nu_\mu} \sim 10^{-8}\,\mathrm{GeV\,cm^{-2}\,s^{-1}\,sr^{-1}}$
- total integrated flux: $F_{\nu_\mu}(> 3\,\mathrm{TeV}) \sim 4 \times 10^{-11}\,\mathrm{cm^{-2}\,s^{-1}}$
- e.g. 1 src @ $d \sim 260\,\mathrm{pc}$ or 2 srcs @ $d \sim 370\,\mathrm{pc}$ etc.

Conclusion





Gamma-ray limits from the halo put very tight constraints on DM explanation of positron excess

Acceleration of secondary e^+ in SNRs could explain PAMELA and Fermi-LAT excess