

Small-Scale Structure in the SDSS and LCDM

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KITP 2012: First Light/Faintest Dwarfs



In Collaboration with:
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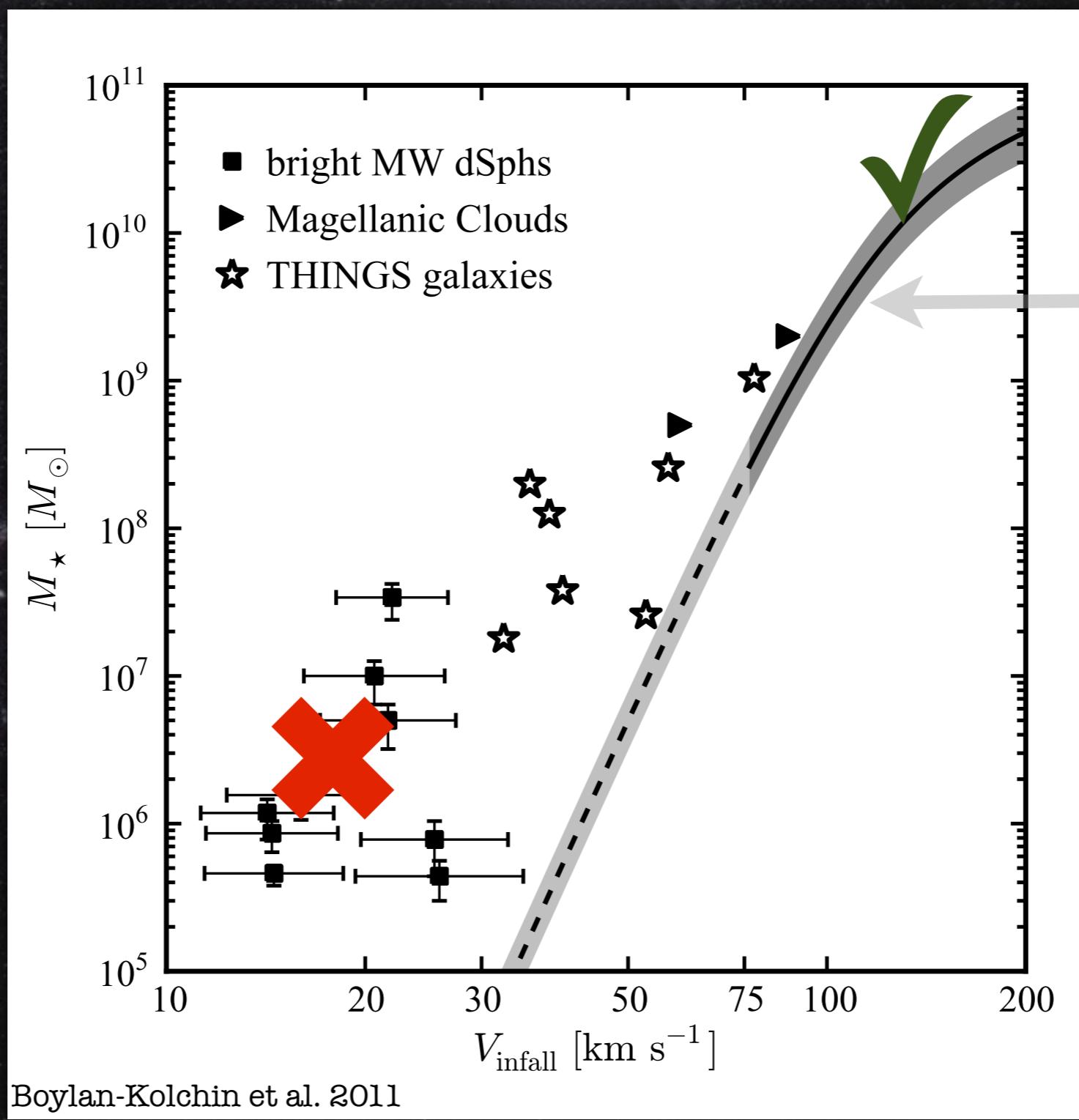
¹UCI, ²U of Sydney

ApJ, 738, 102 (2011)

Image Credit: Robert Gendler

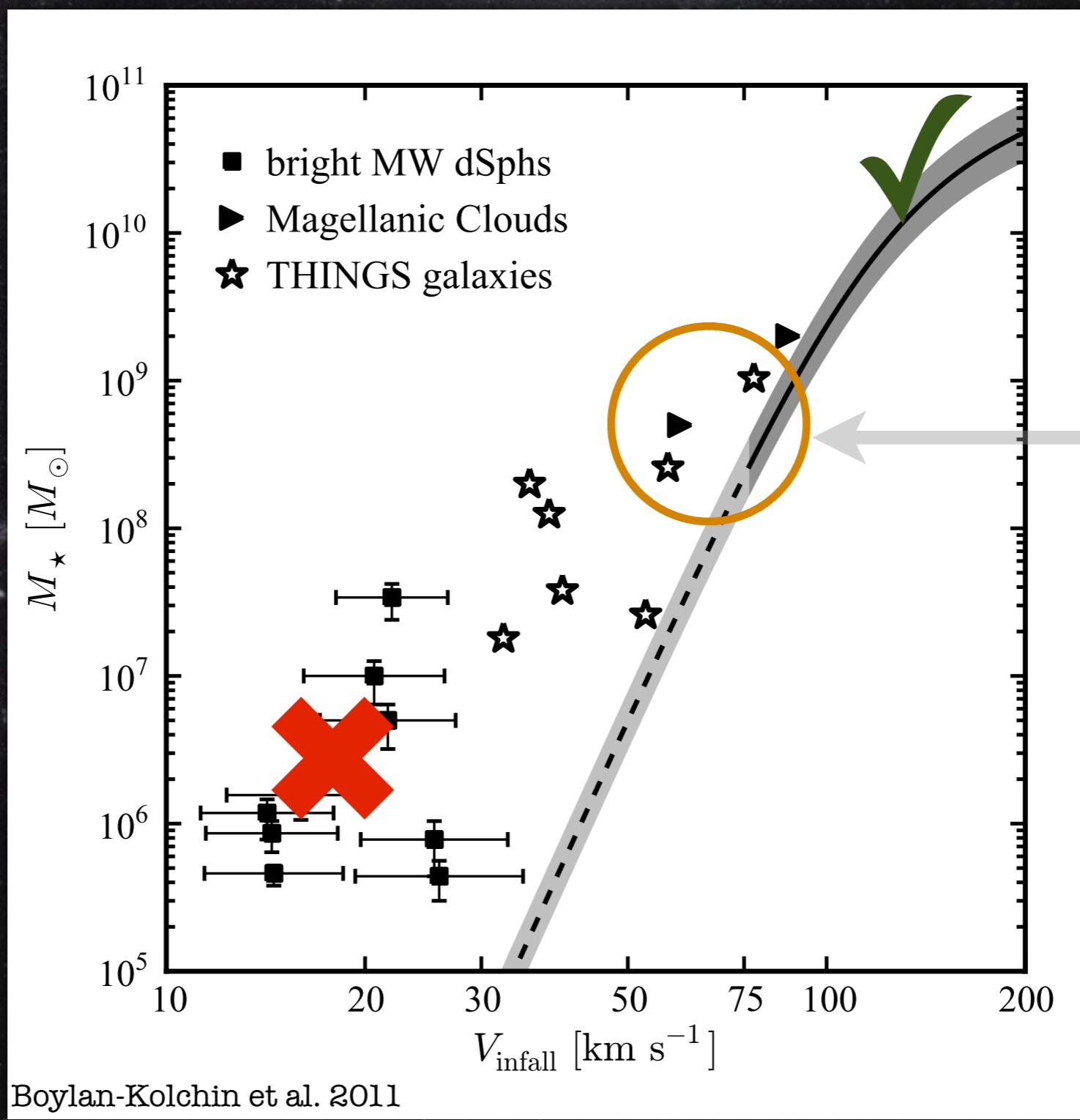
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Why LMC Scales?



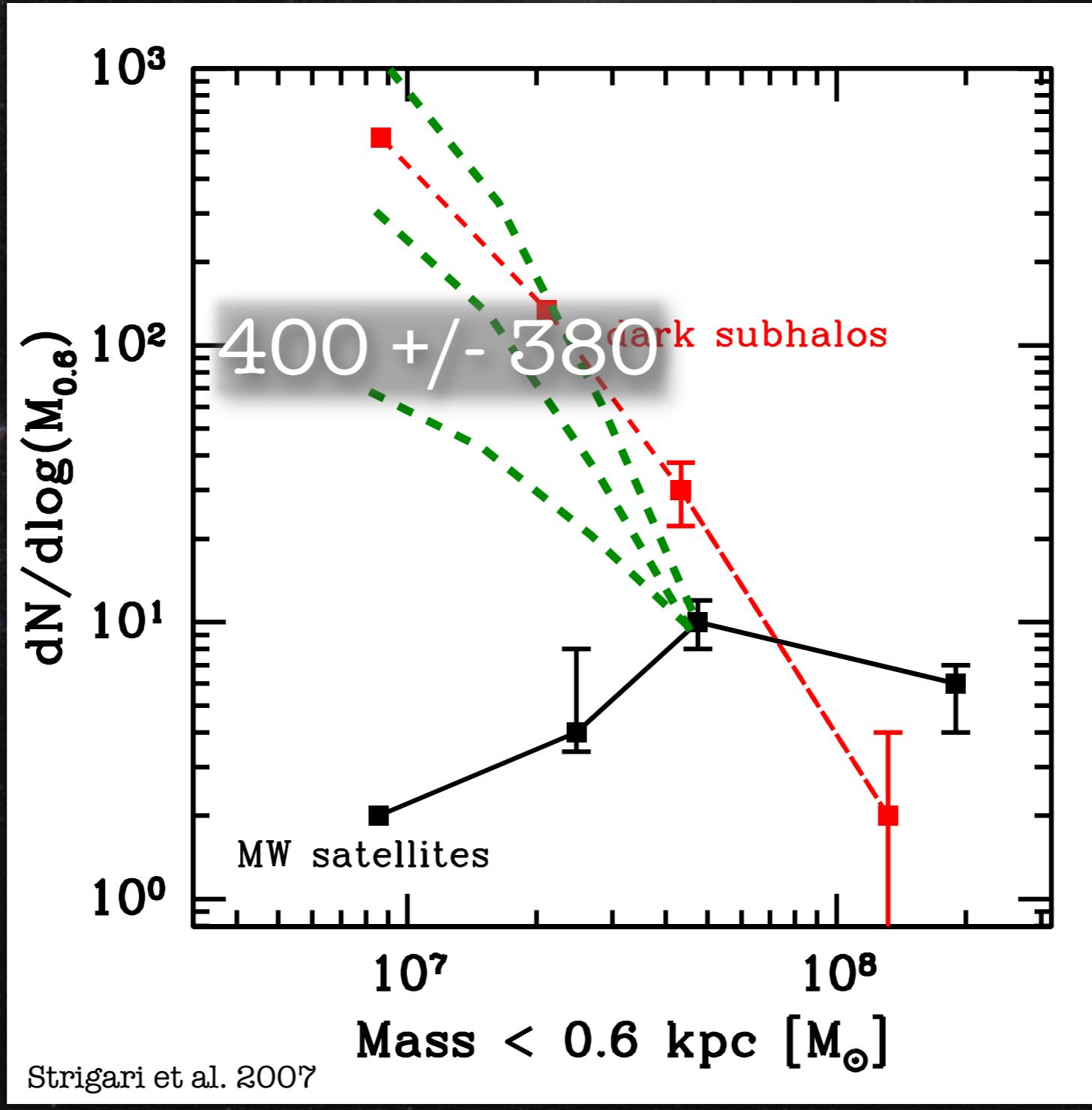
Abundance
Matching

Why LMC Scales?

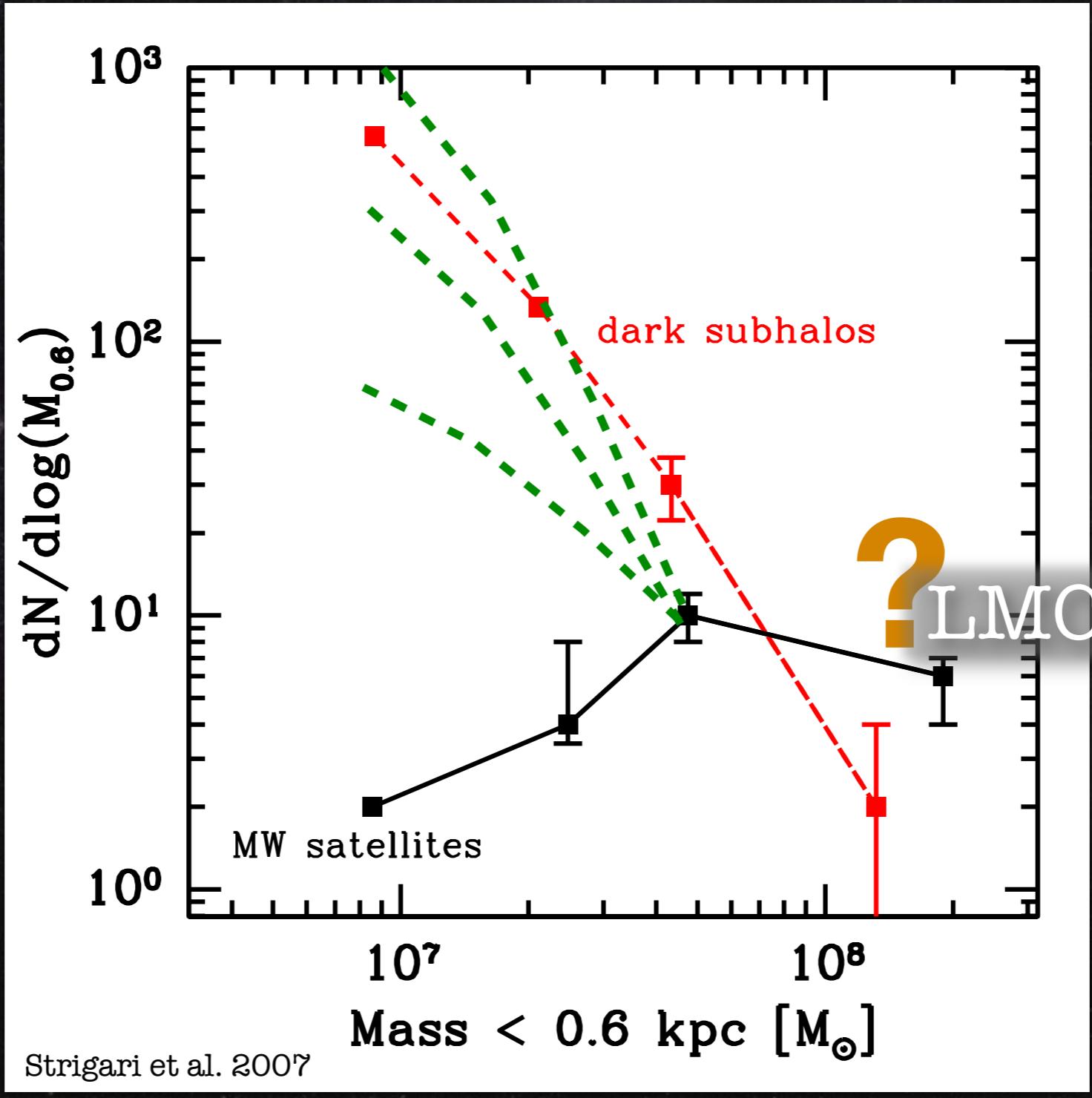


Does it
work
here?

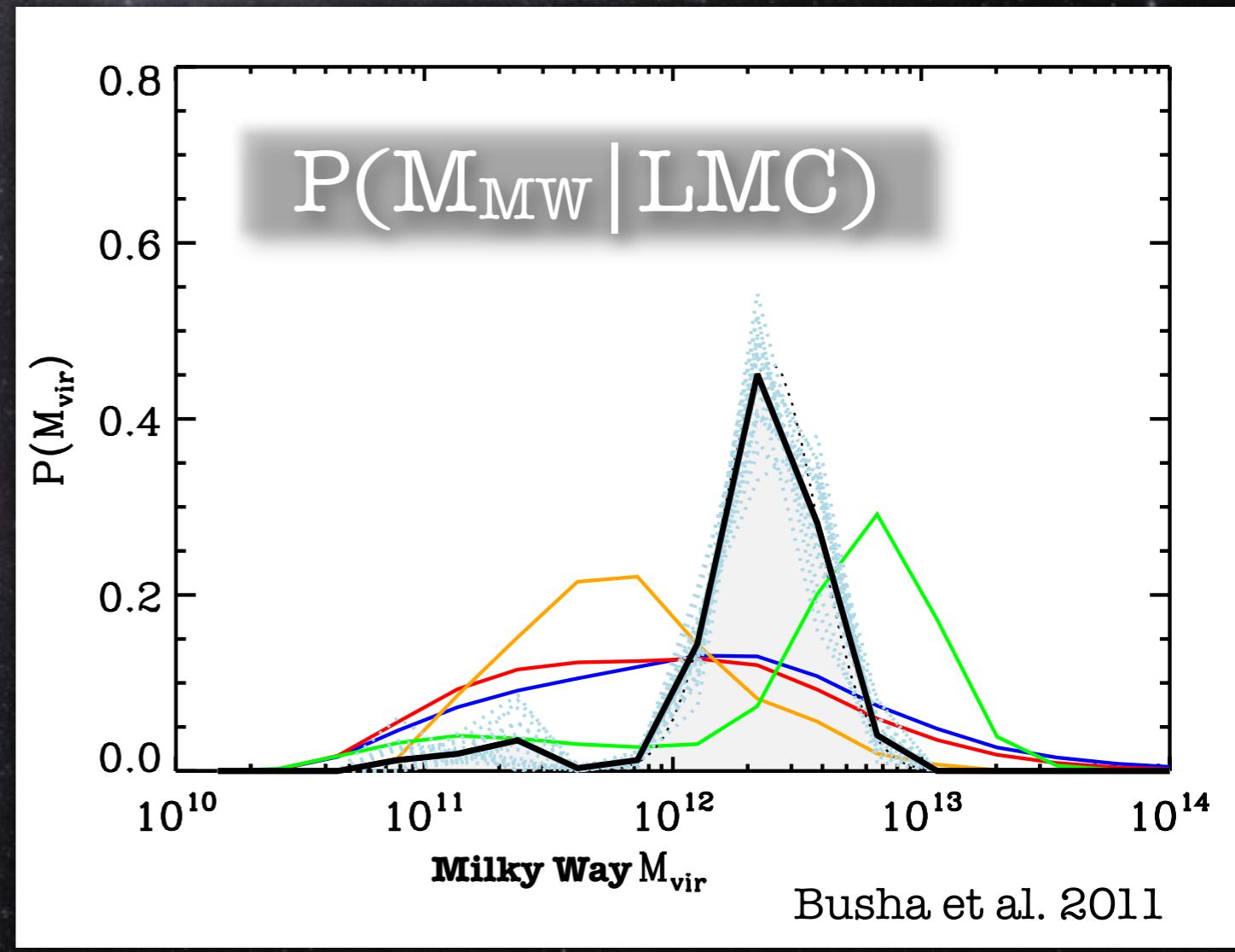
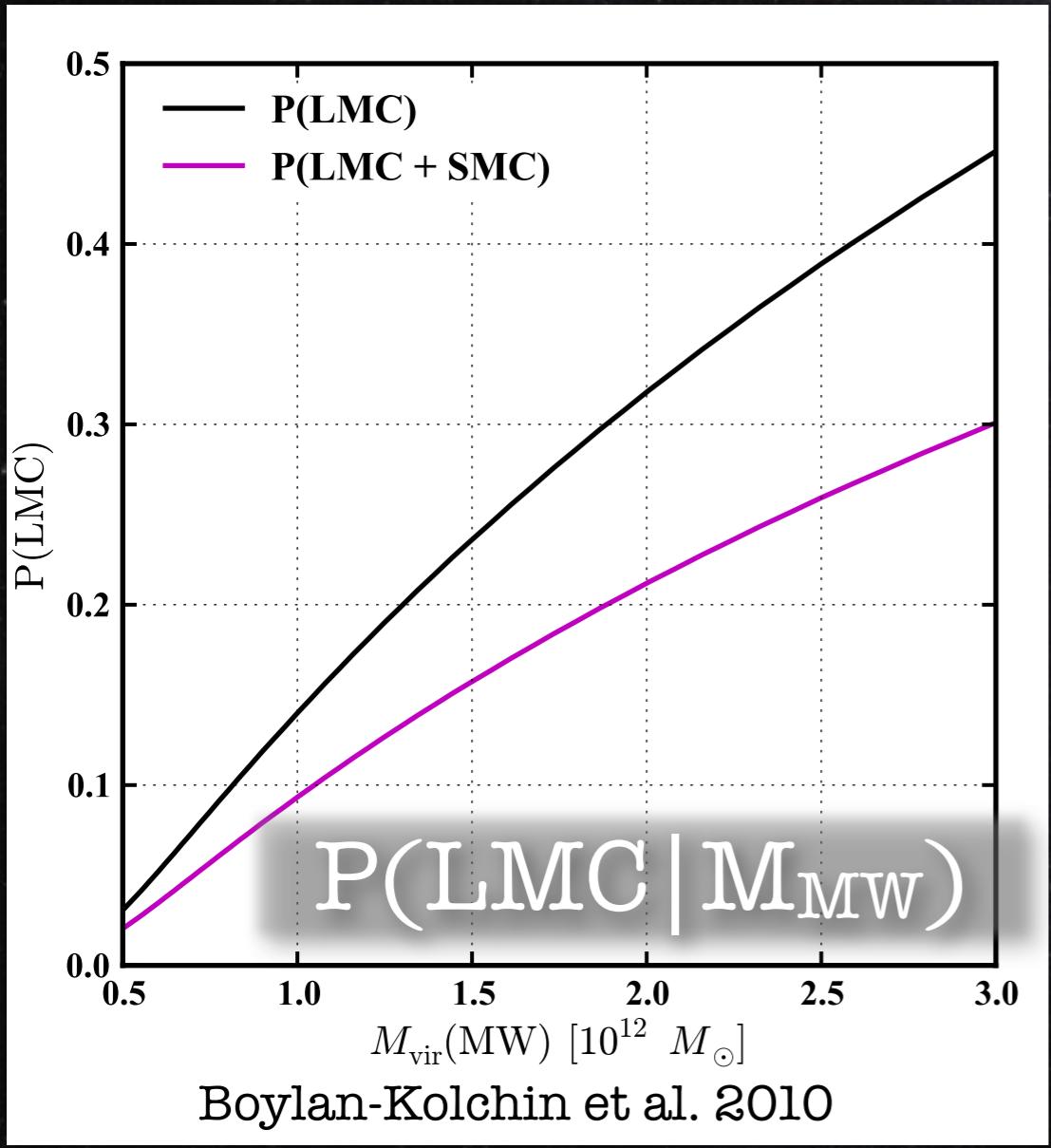
Missing Satellites Absence of Smaller Galaxies Nearby Bigger Ones Problem



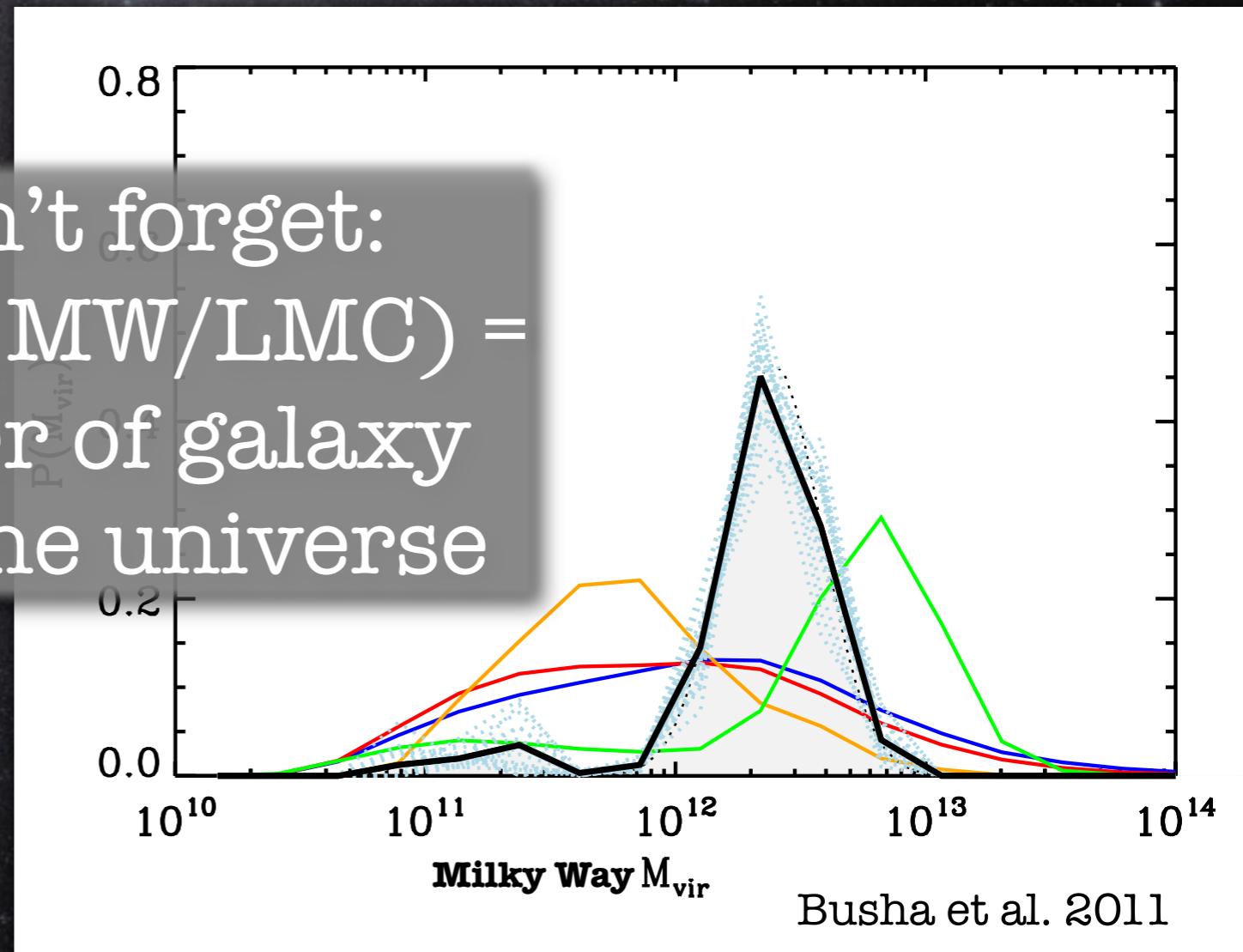
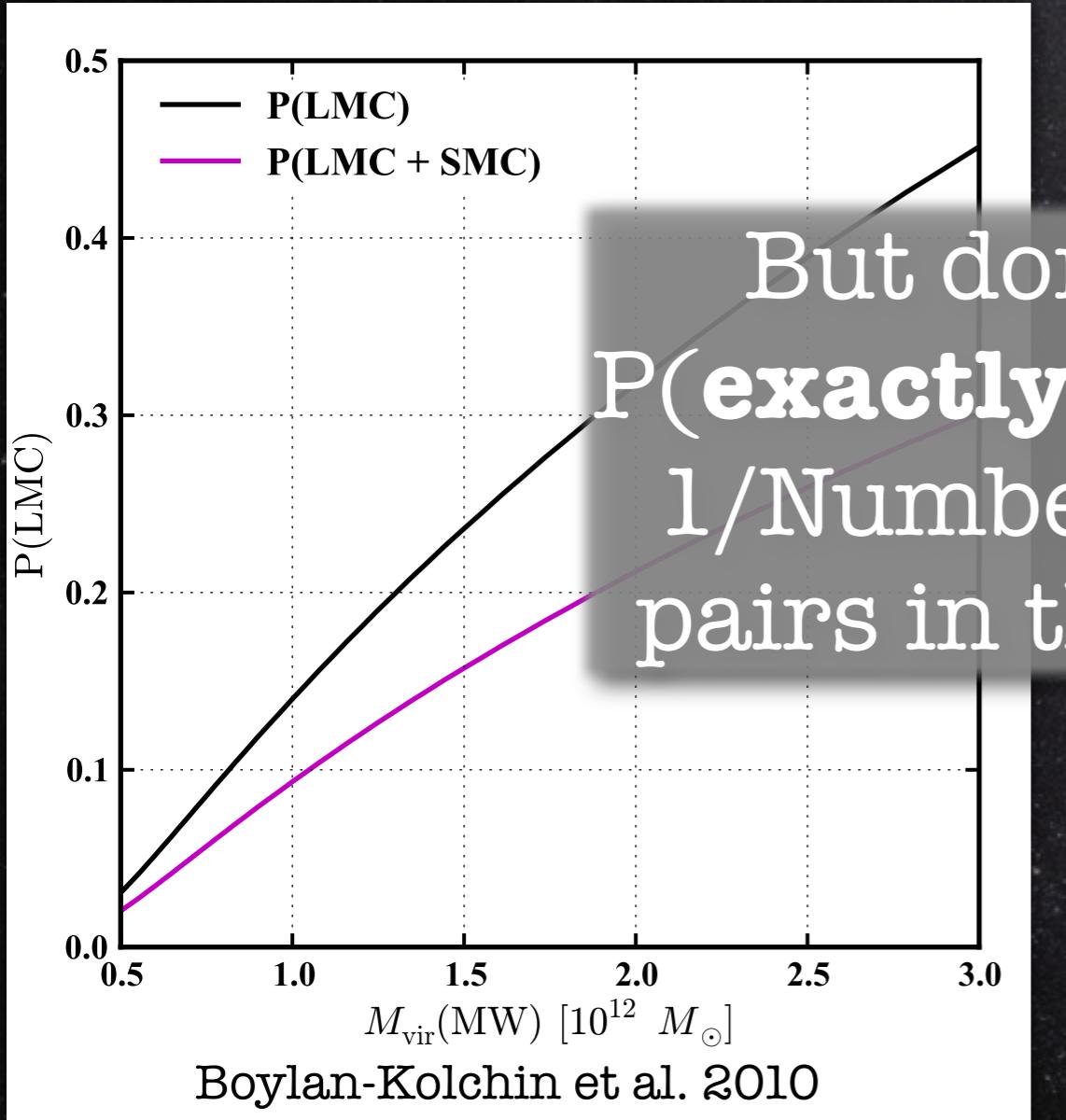
“Found” Satellites Problem?



“Found” Satellites Problem?



“Found” Satellites Problem?



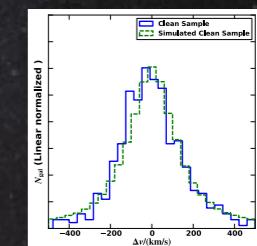
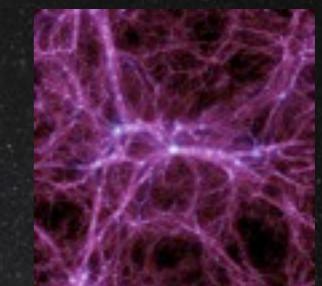
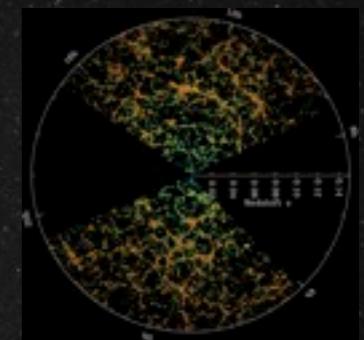
But don't forget:
 $P(\text{exactly MW/LMC}) =$
1/Number of galaxy
pairs in the universe

Two Key Questions

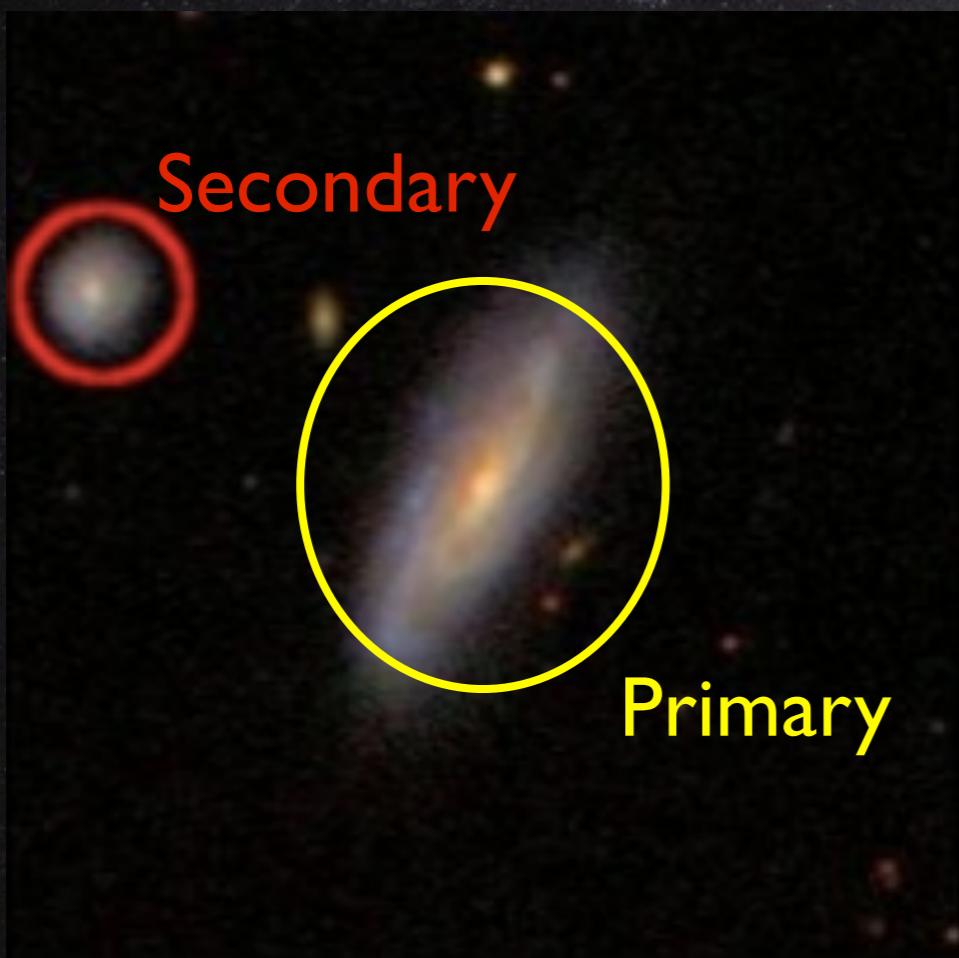
- ◆ Does LCDM reproduce simple observables of the brightest satellites of L^* galaxies?
- ◆ Is the LMC Weird? (empirically)
 - ◆ (Depends on how you ask the question!)

Approach

- ◆ SDSS (Spectroscopic) sample of $\sim 0.1 L^*$ companions of isolated $\sim L^*$ galaxies
- ◆ Identical sample from cosmological simulation
- ◆ Compare easy observables: Δv_{pair} and radial distribution
- ◆ Compare the satellites to the LMC
- ◆ (Bonus: Study the satellites themselves)

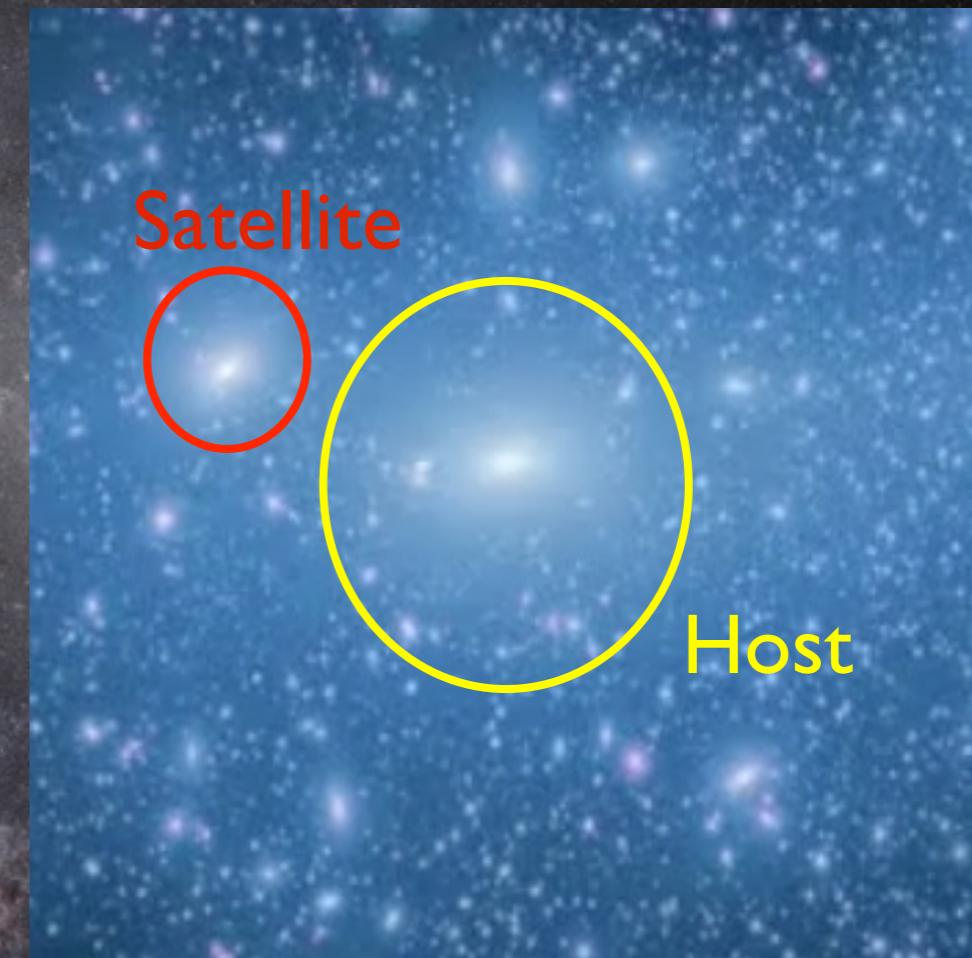


Approach



SDSS
VAGC DR7

Blanton et al. 05

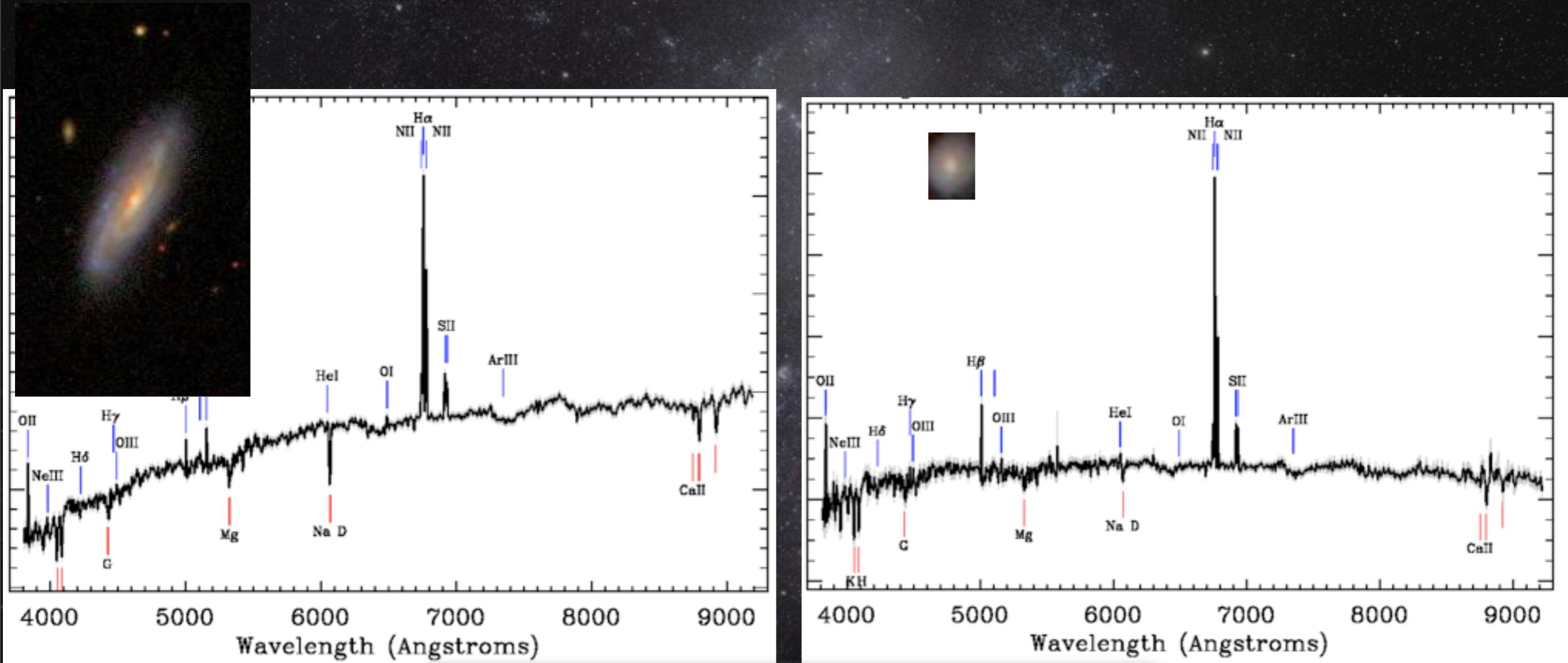


Millennium II
Simulation

Boylan-Kolchin et al. 09

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Primary and Secondary Must Have Similar Redshifts



Need for Spectroscopic
Completeness sets the
volume limit

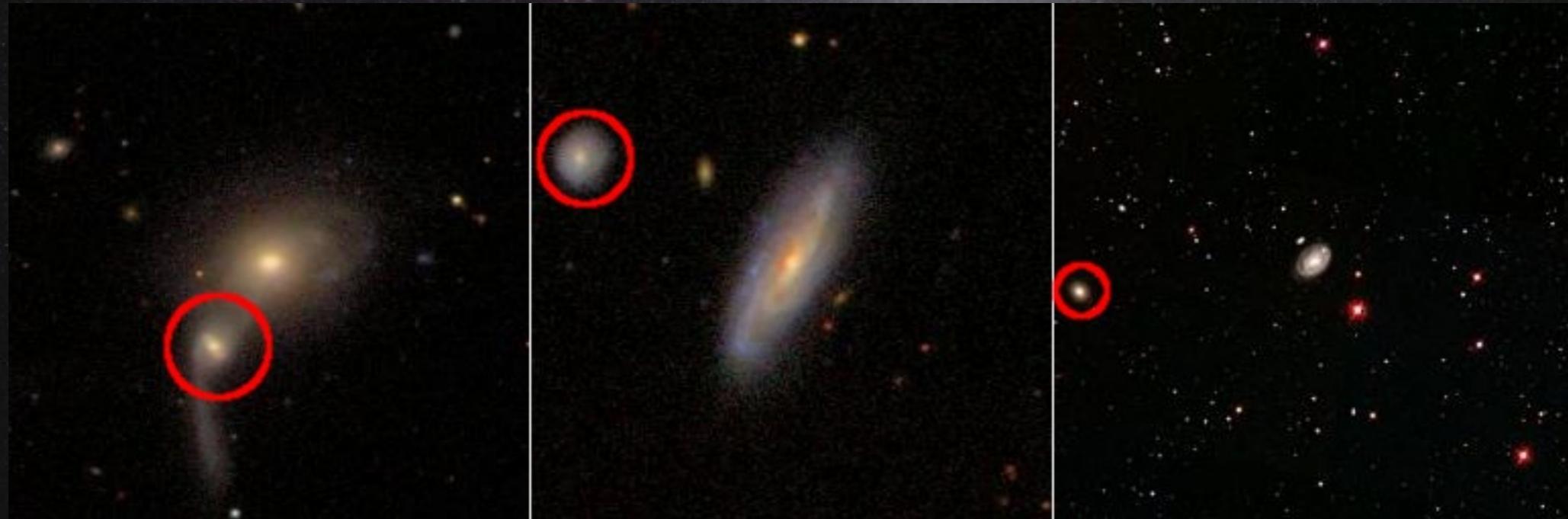
Primaries must be isolated



$250 < d_{\text{proj}} < 700 \text{ kpc/h}$

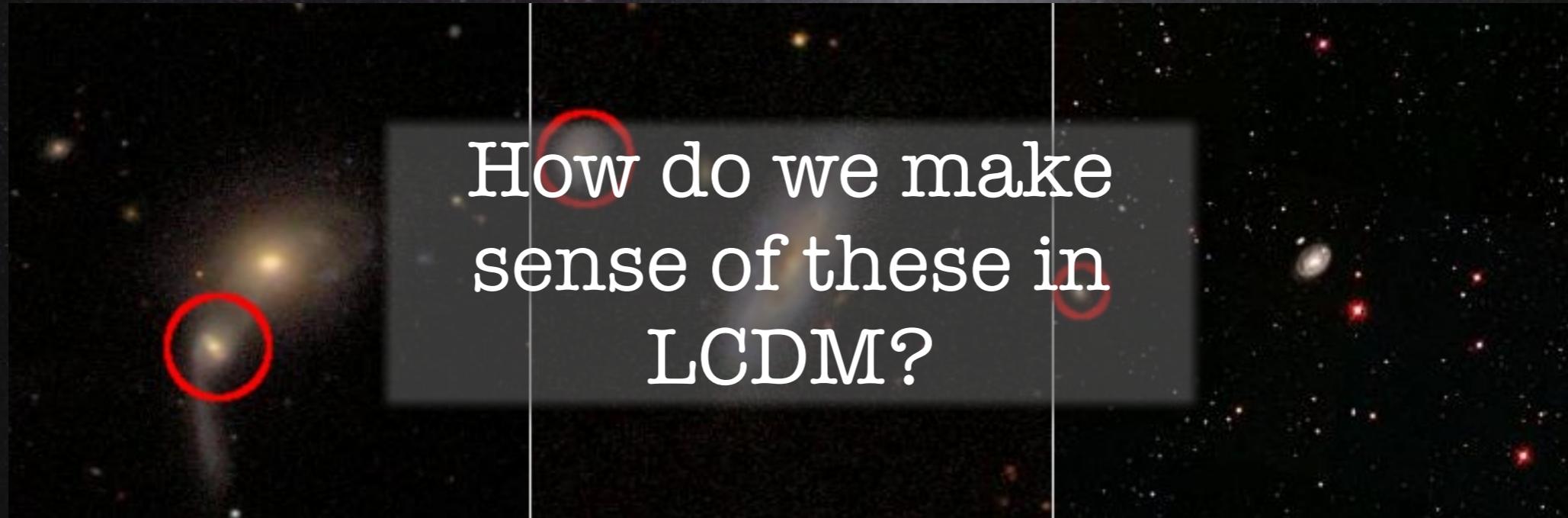
Primary/Secondary Pairs

- ◆ 1075 Primaries, 467 Secondaries
- ◆ Local ($z_{\text{median}}=.028$) due to spectroscopic limit

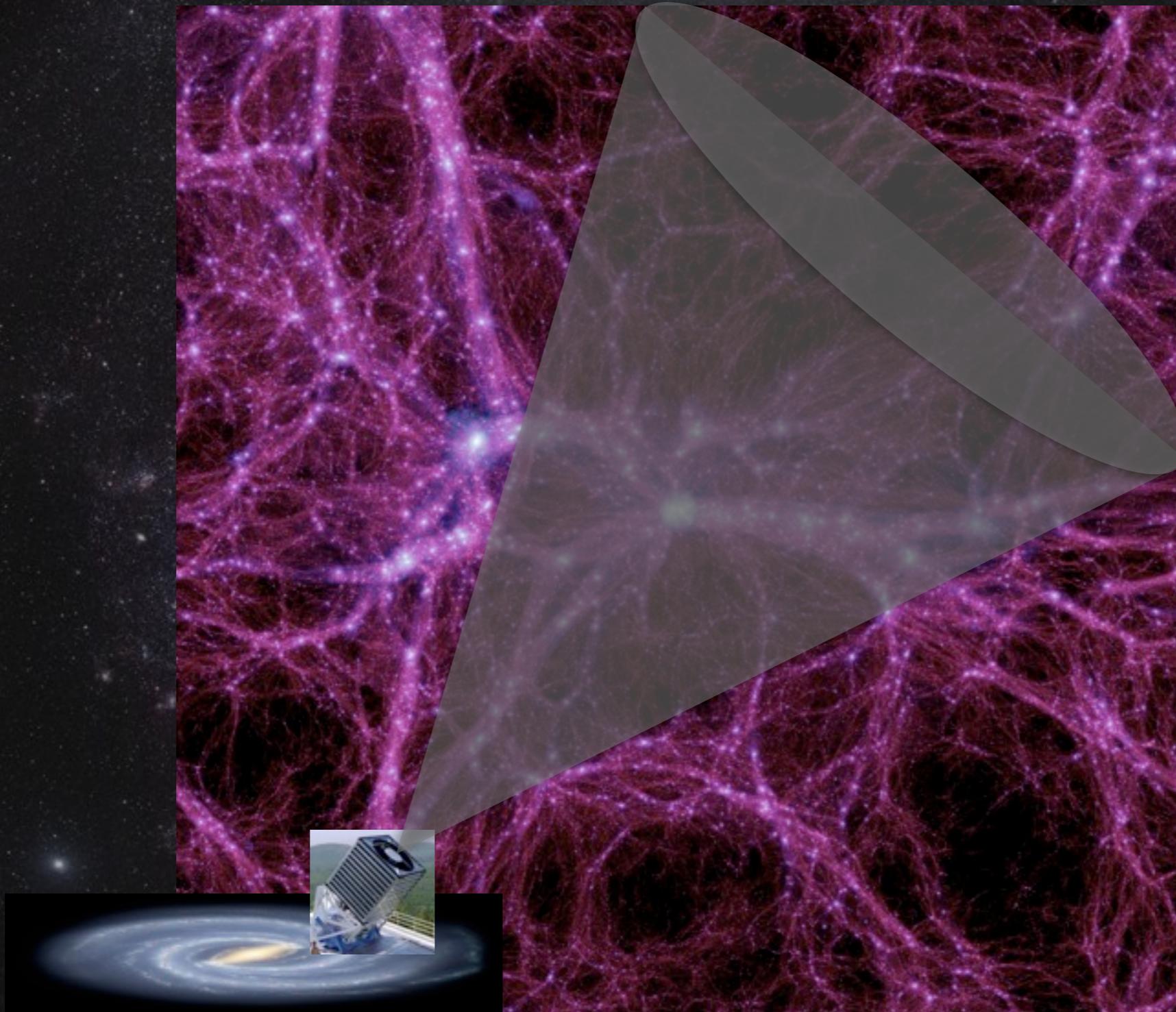


Primary/Secondary Pairs

- ◆ 1075 Primaries, 467 Secondaries
- ◆ Local ($z_{\text{median}}=.028$) due to spectroscopic limit



Mock Galaxy Catalogs



Abundance Matching for Luminosities



SDSS
VAGC DR7

Blanton et al. 05

Millennium II
Simulation

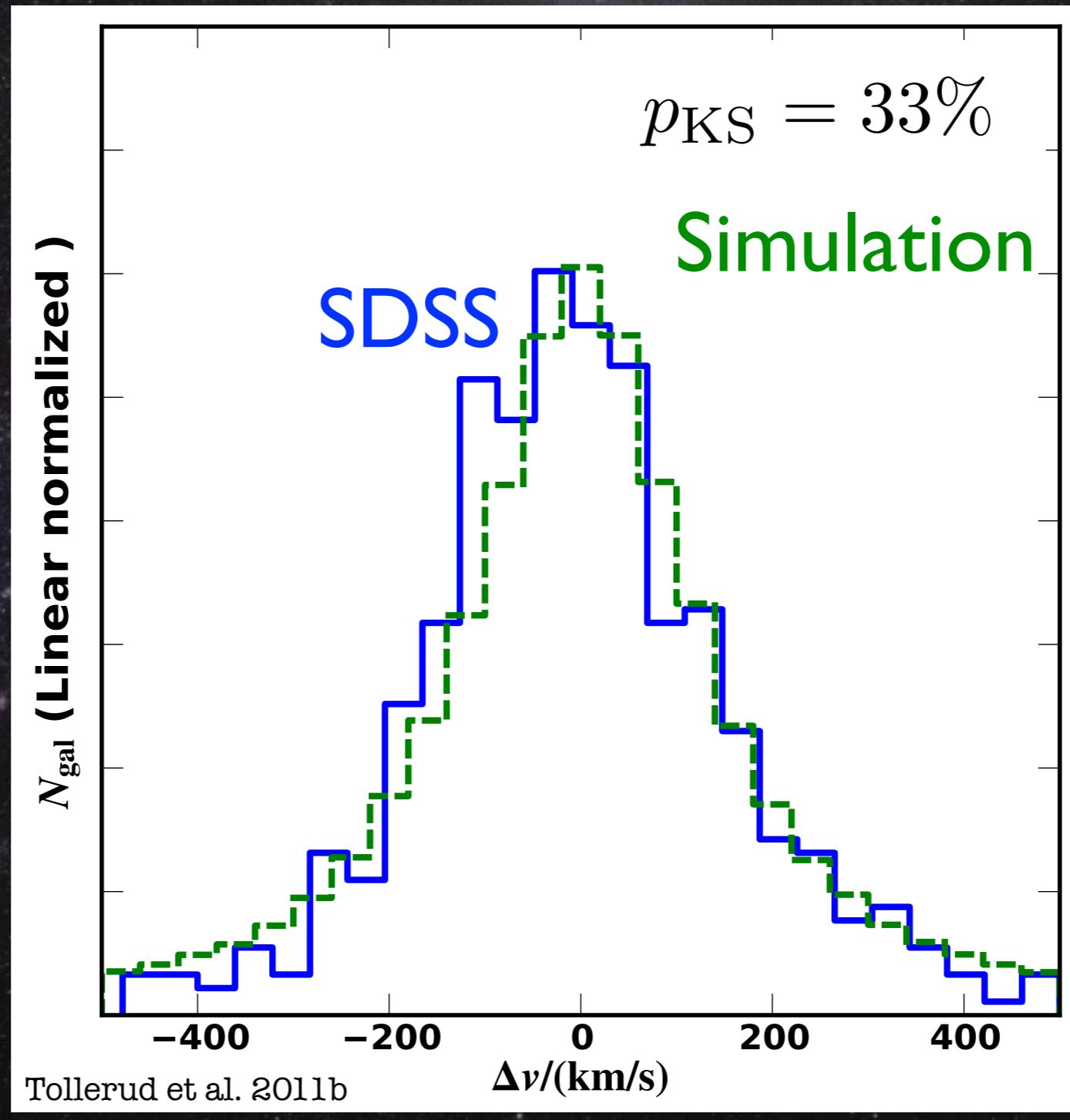
Boylan-Kolchin et al. 09

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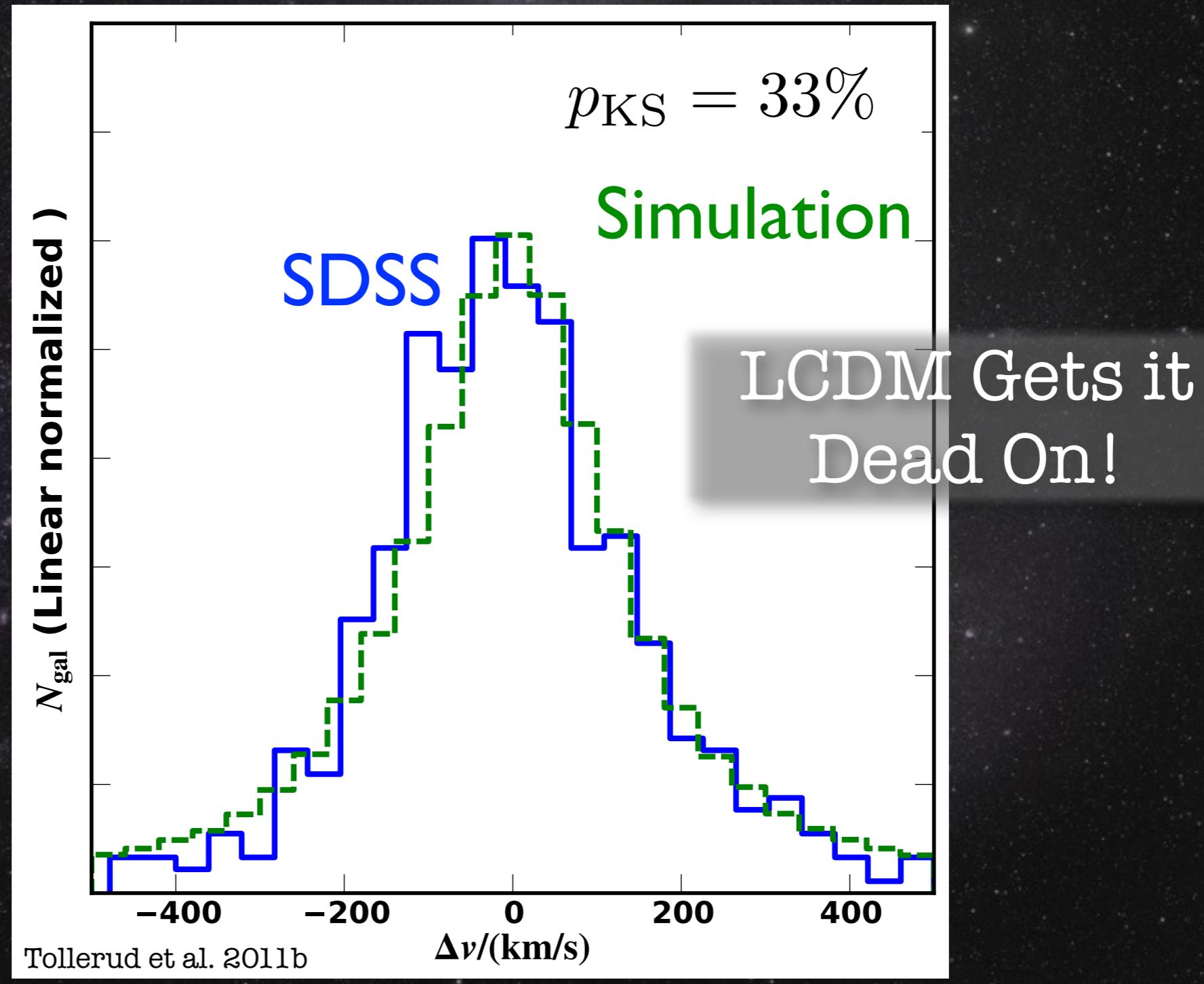
Only Three Assumptions for $\text{MS-II} \leftrightarrow \text{SDSS}$

- ♦ Gravity
- ♦ LCDM Initial Conditions/Cosmology
- ♦ Monotonic $L_{\text{gal}} \Leftrightarrow M_{\text{halo}}$

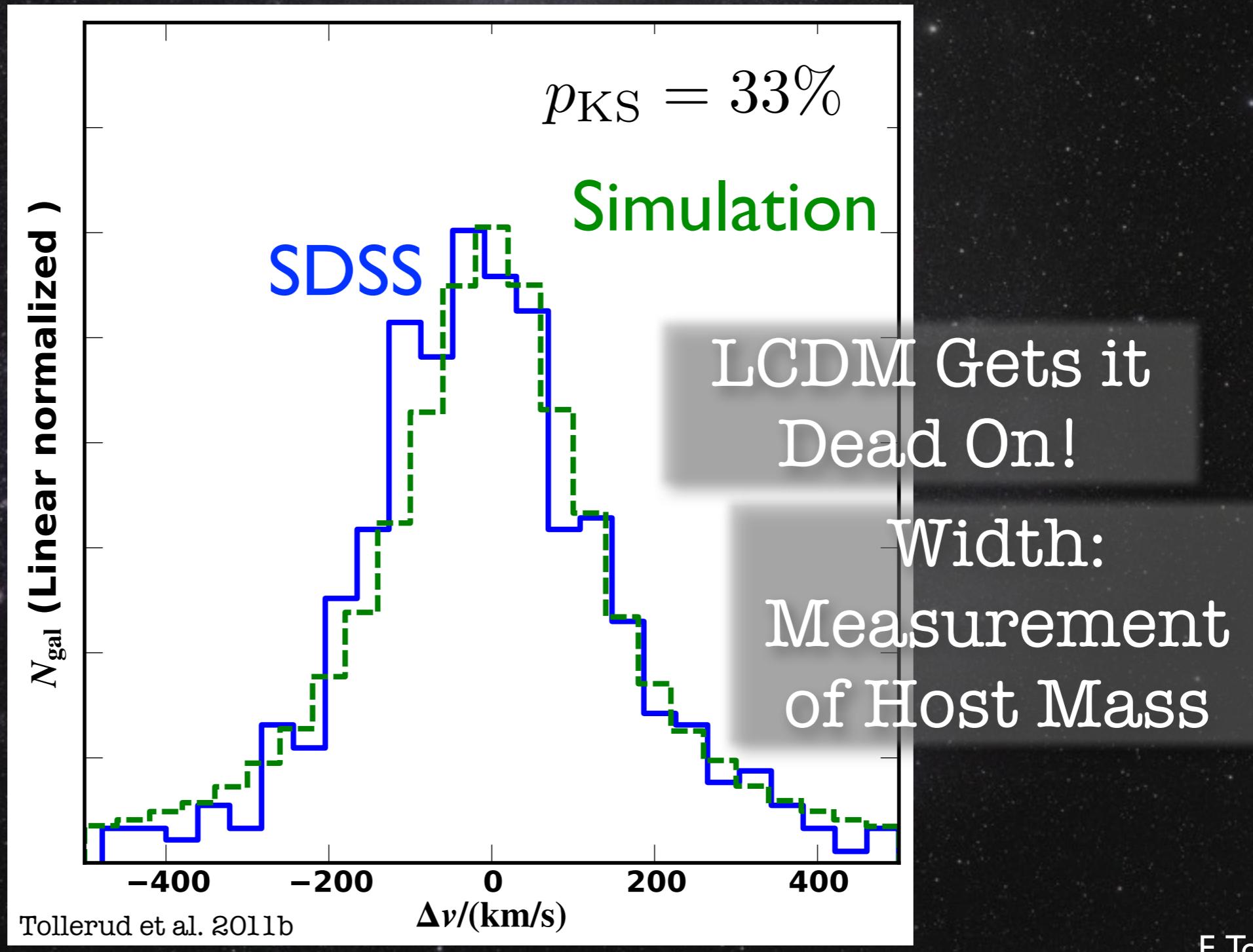
Pairwise Velocity Distribution



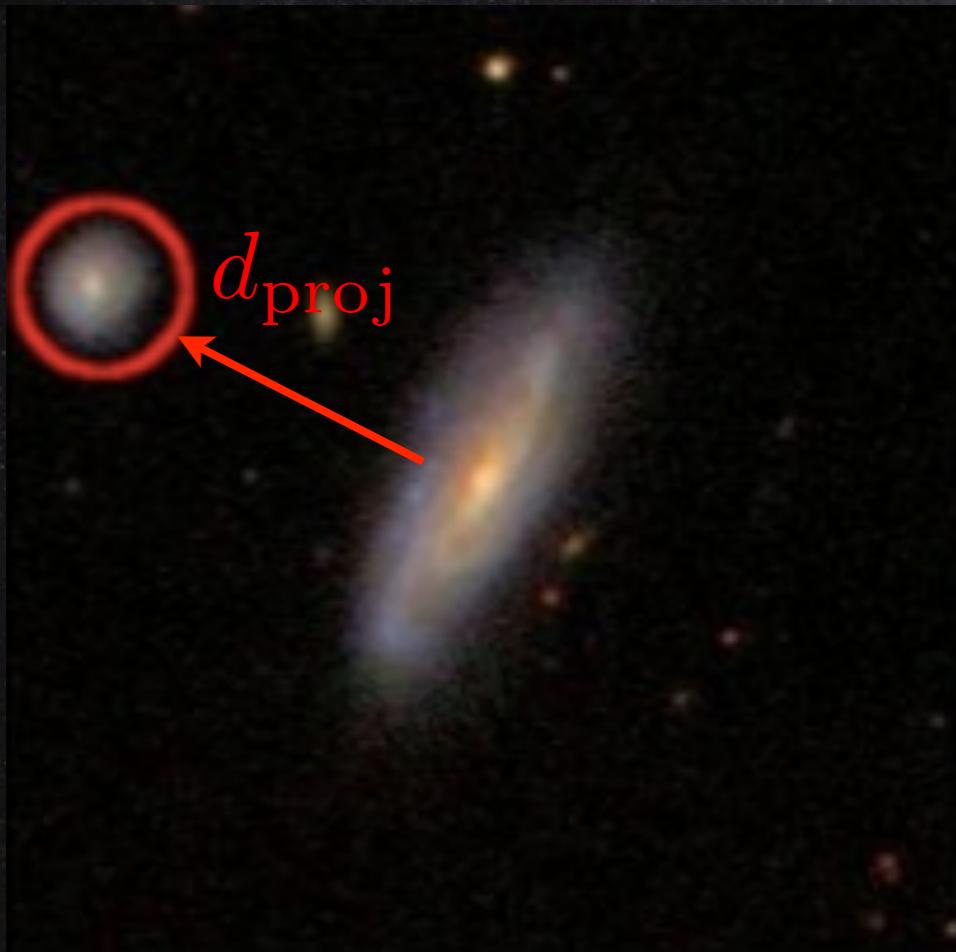
Pairwise Velocity Distribution



Pairwise Velocity Distribution



Radial Distribution

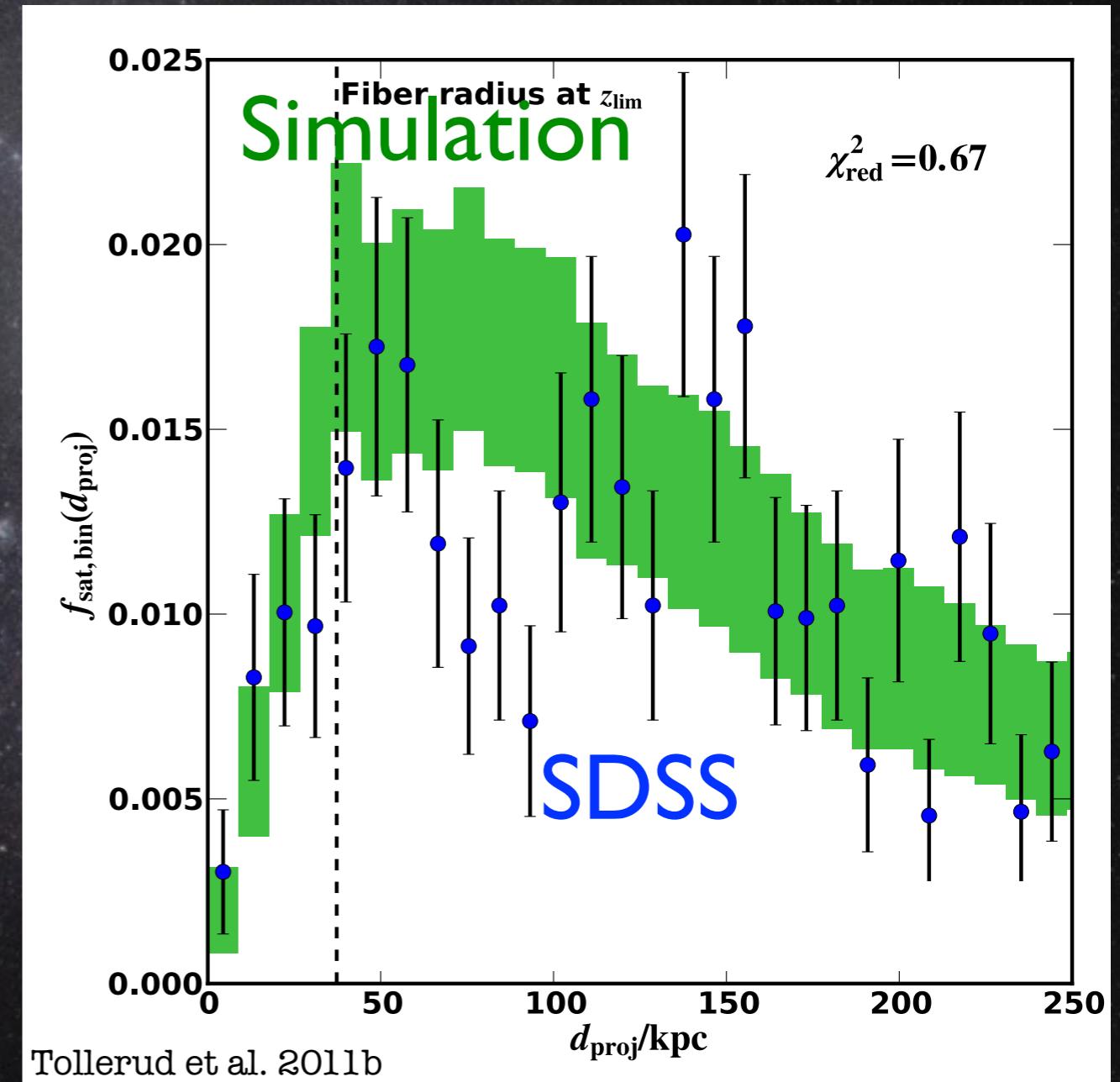
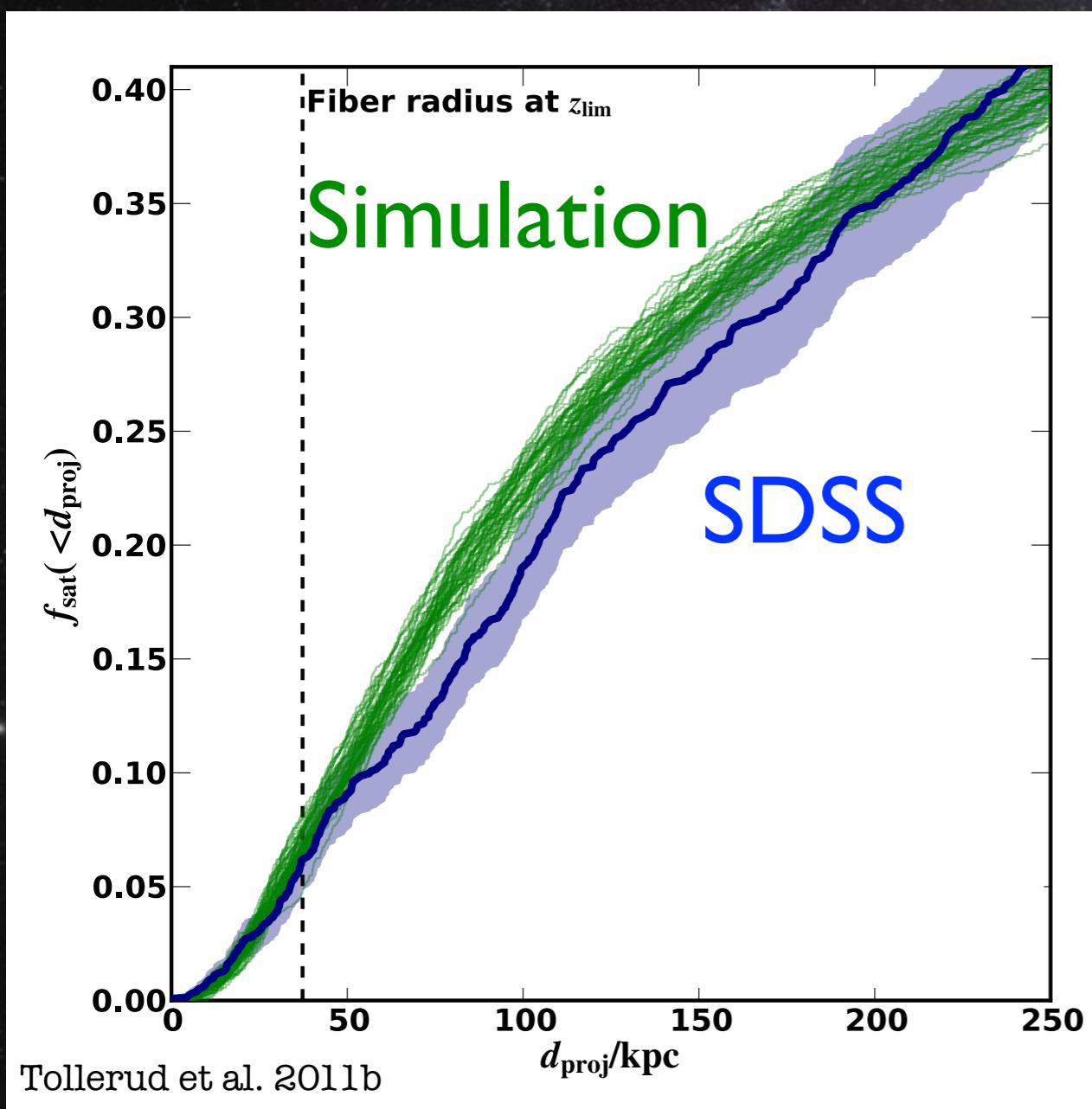


SDSS Galaxies

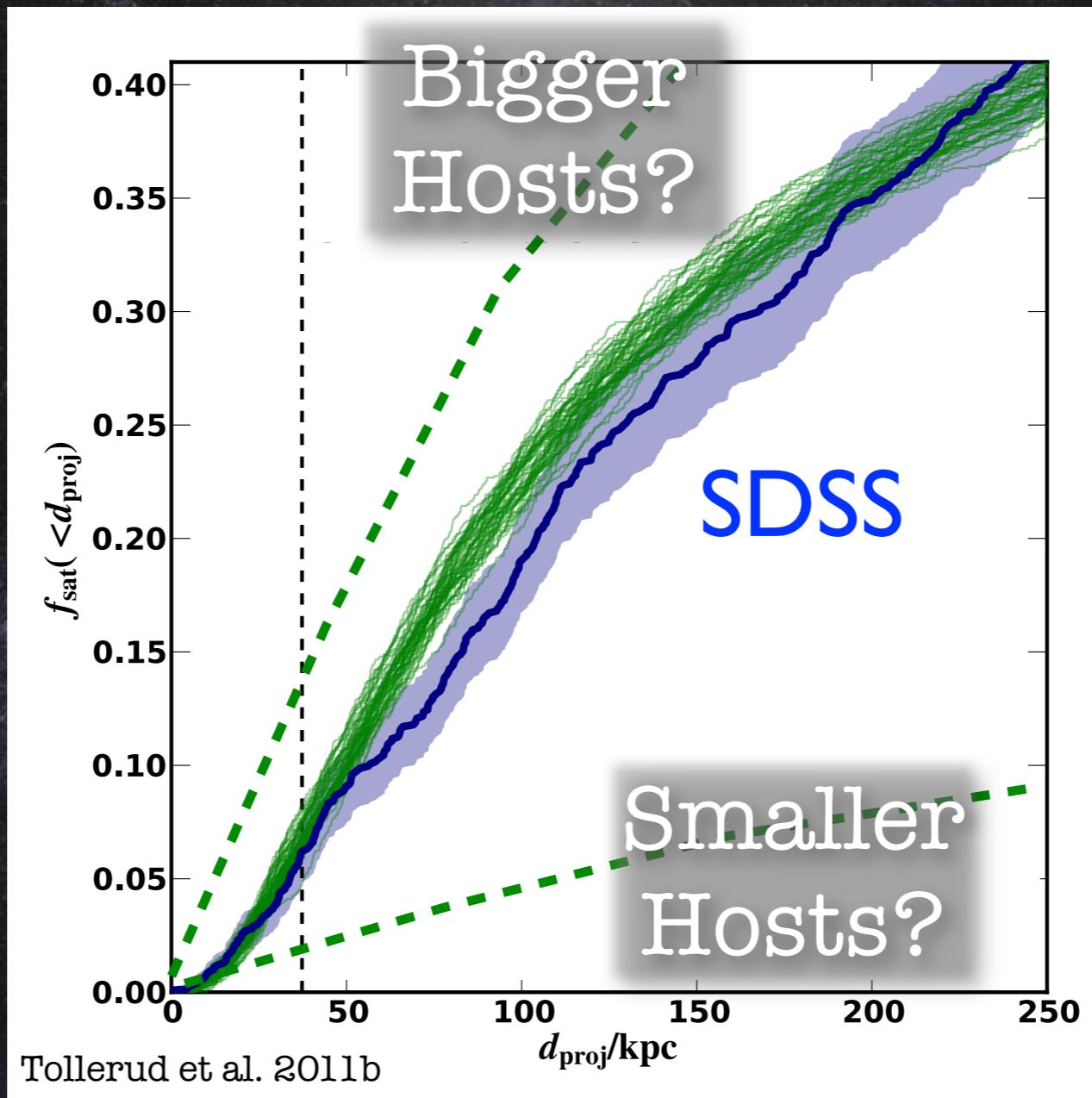


MS-II Halos

LCDM Gets The Right Radial Distribution!



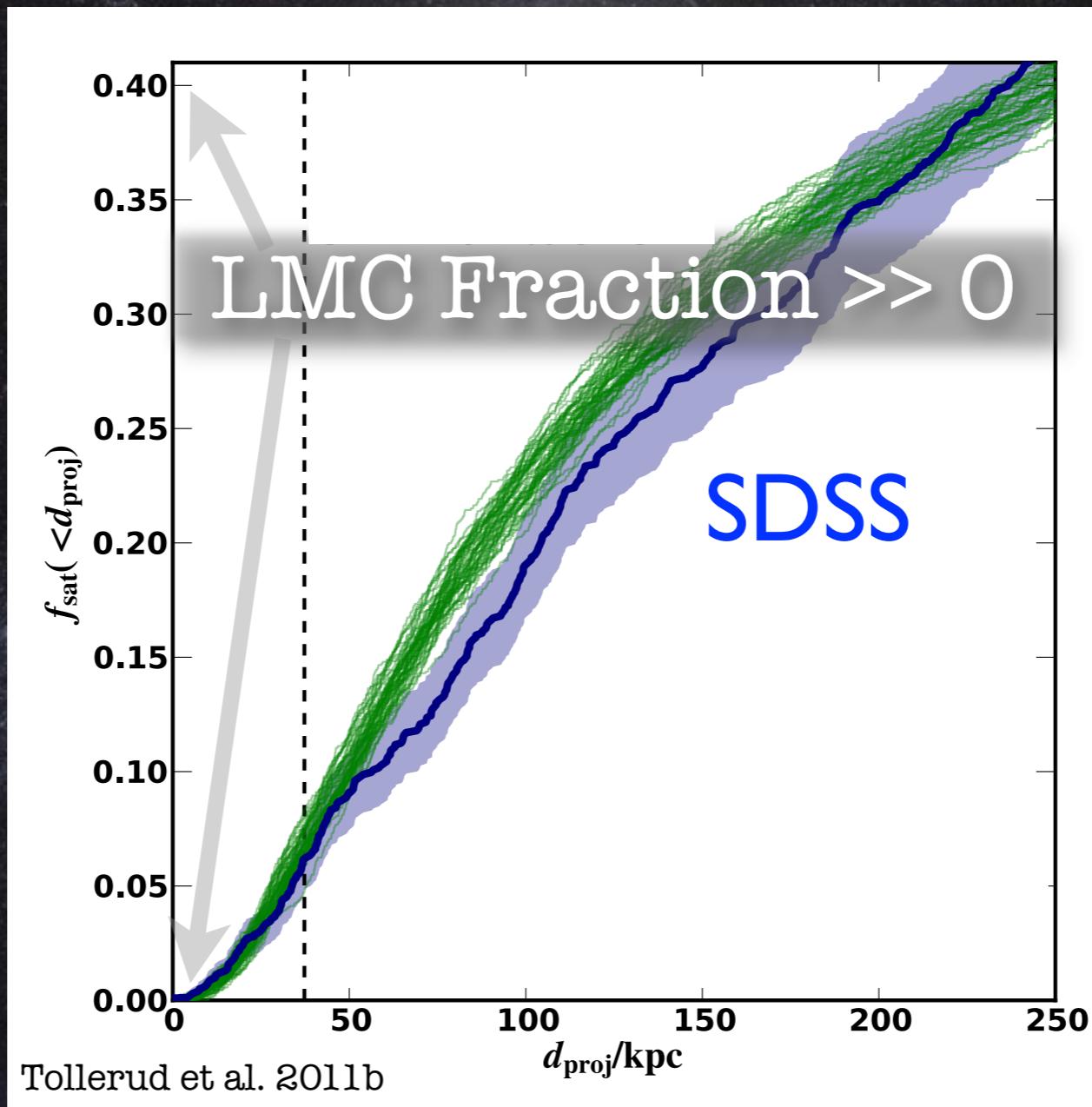
LCDM Gets The Right Radial Distribution!



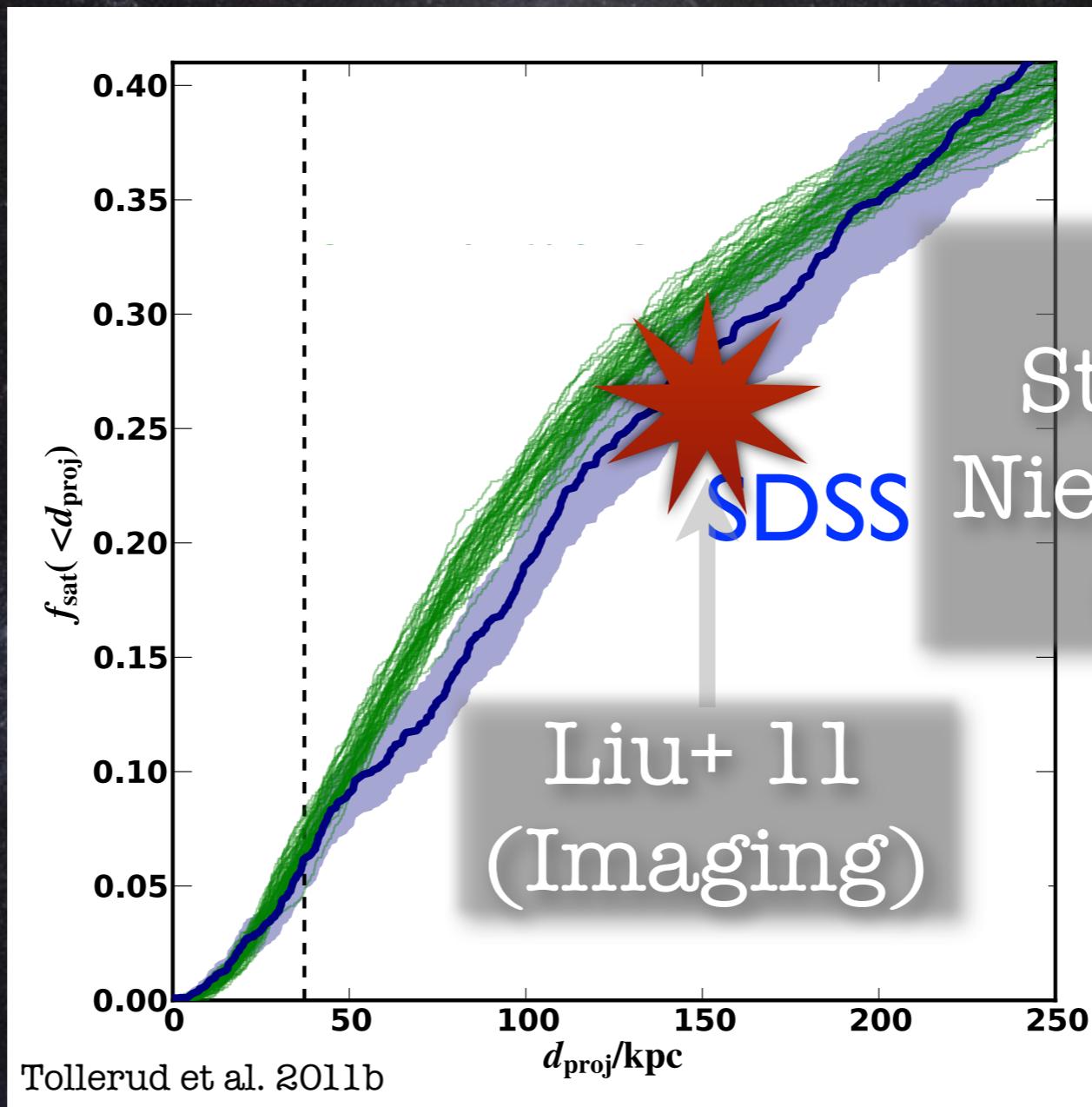
Tollerud et al. 2011b

E. Tollerud

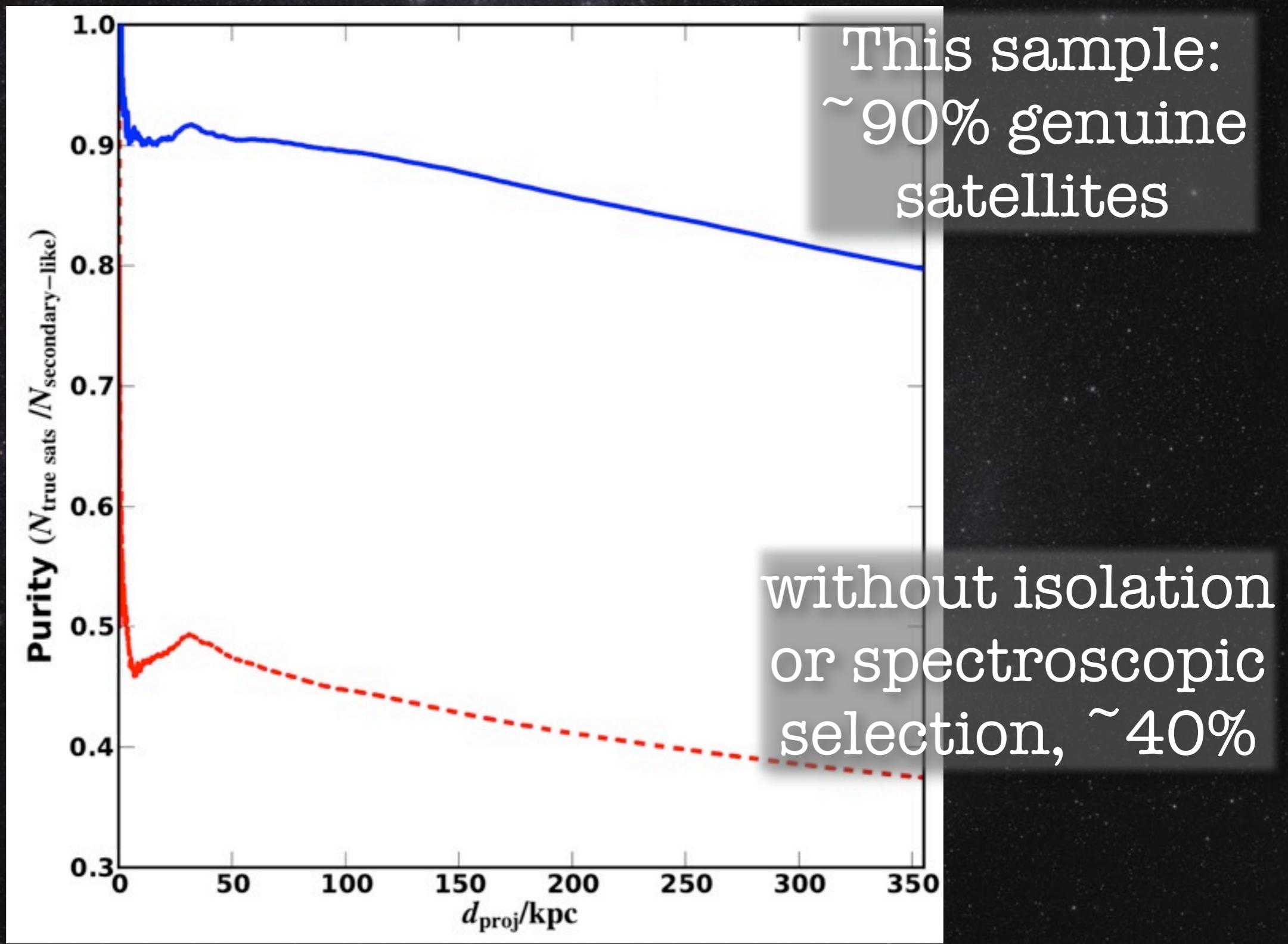
Empirically, MW/LMC Isn't Too Rare



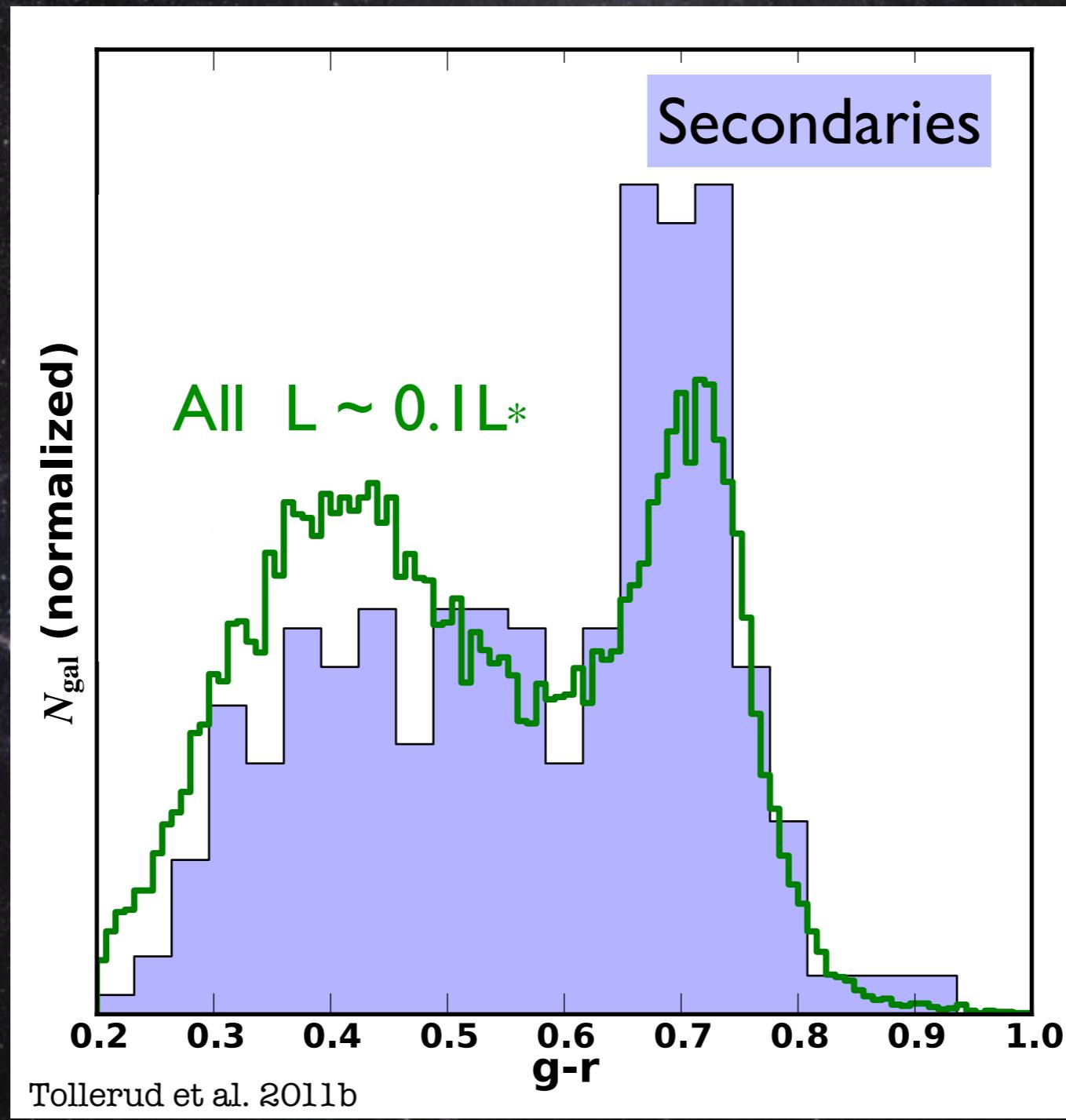
Imaging Can Do This Statistically



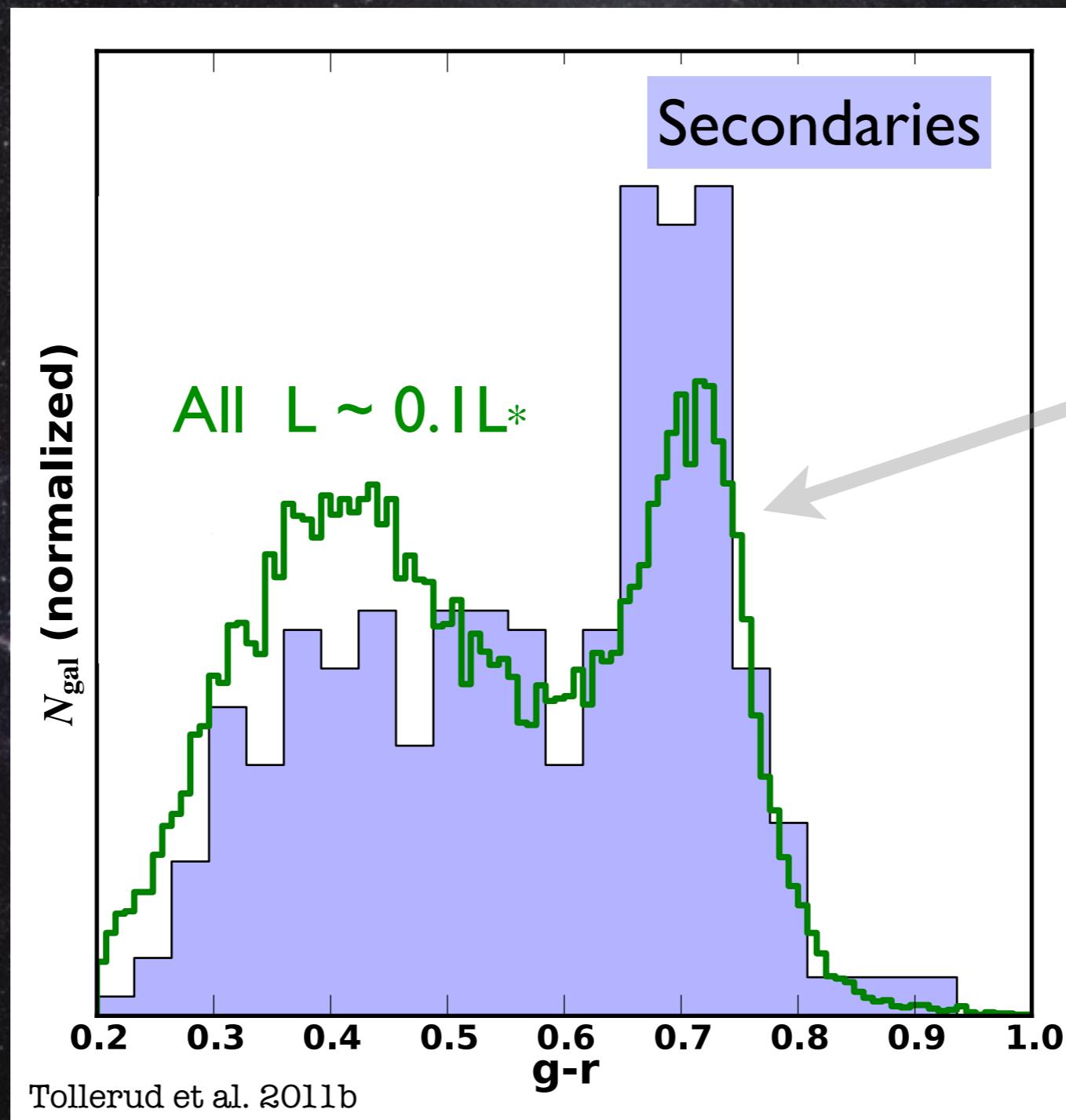
LCDM \rightarrow SDSS



Satellite Colors

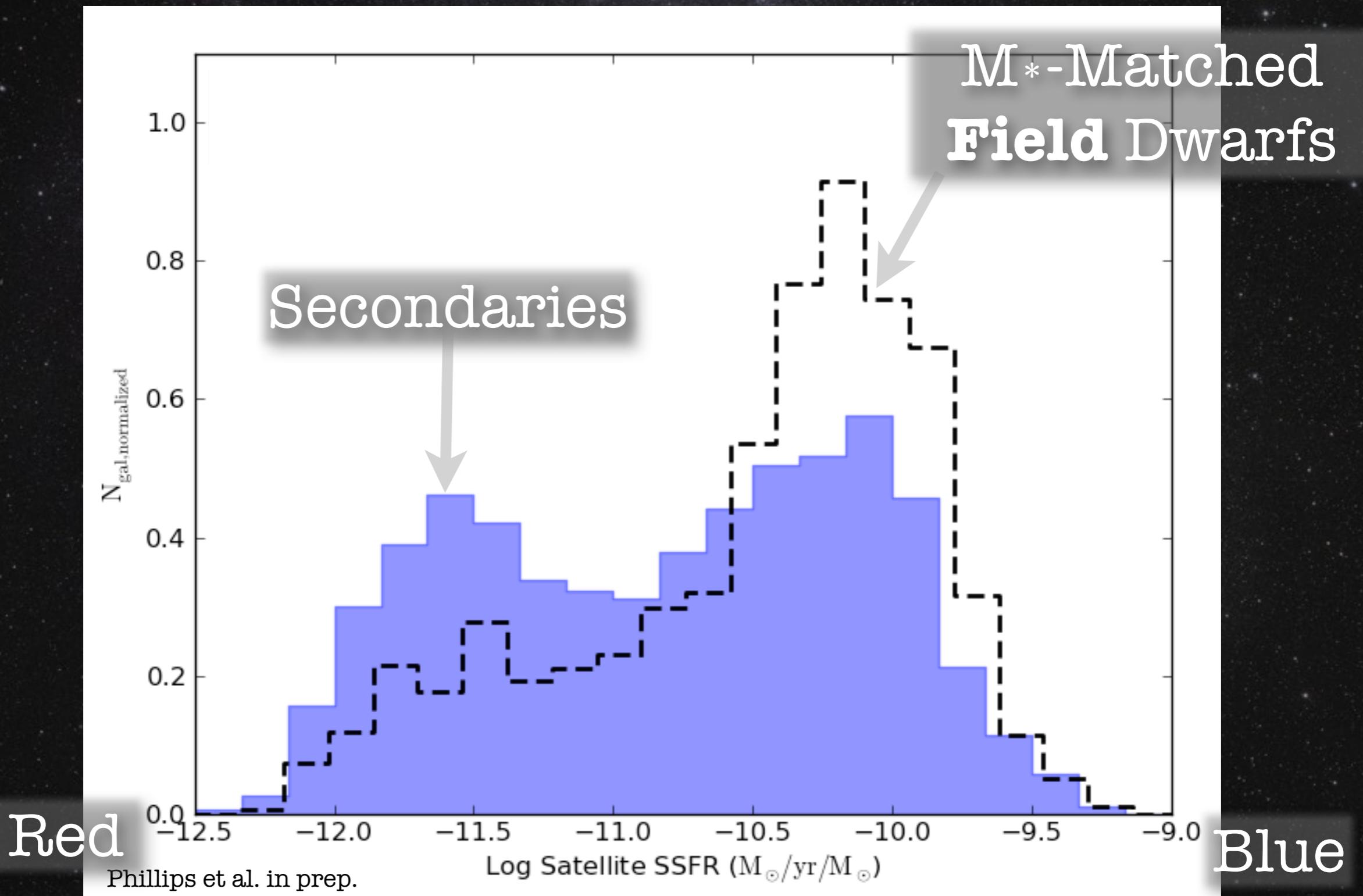


Satellite Colors

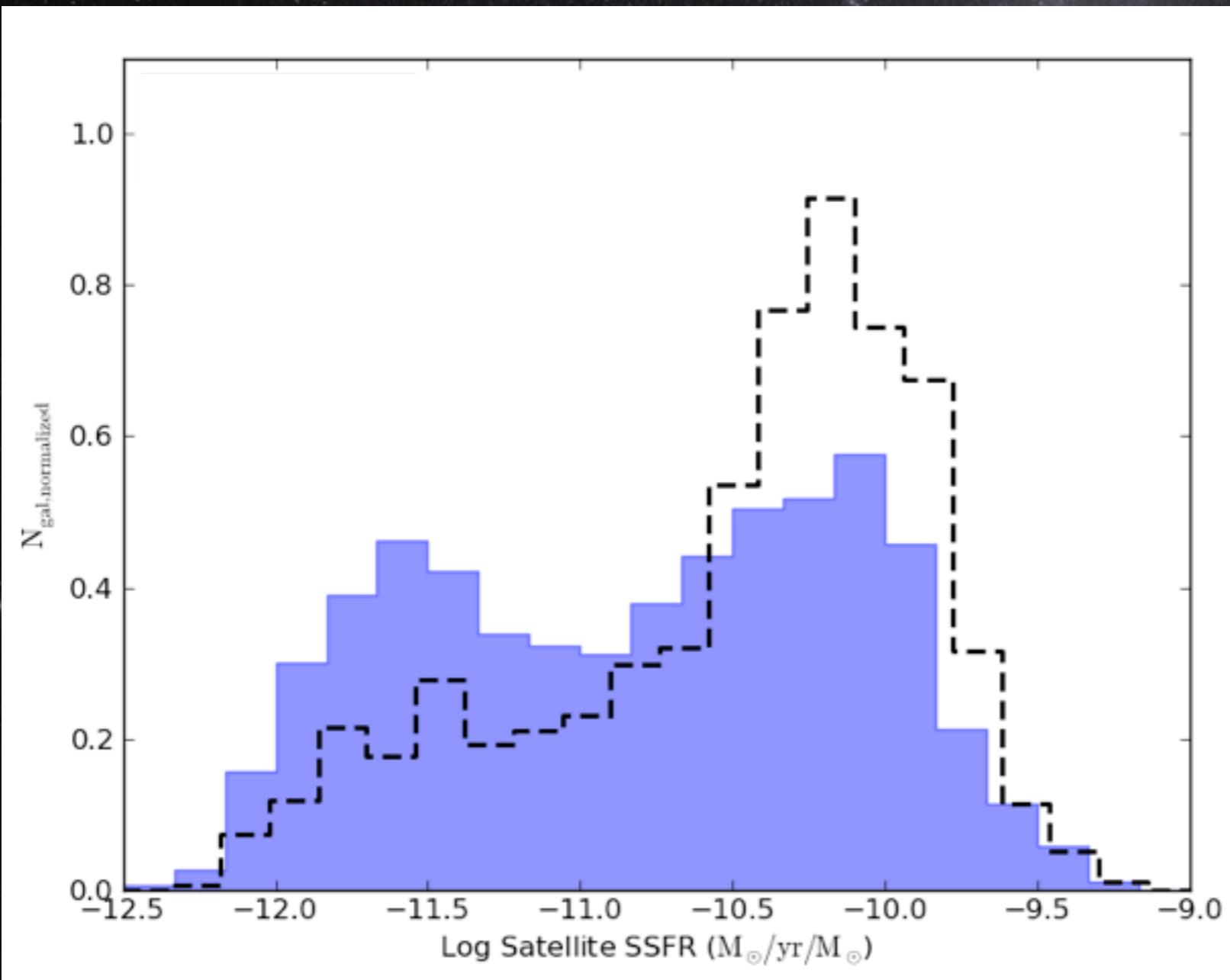


Satellites of
Isolated L^*
Galaxies
Are Red

Satellite SFRs



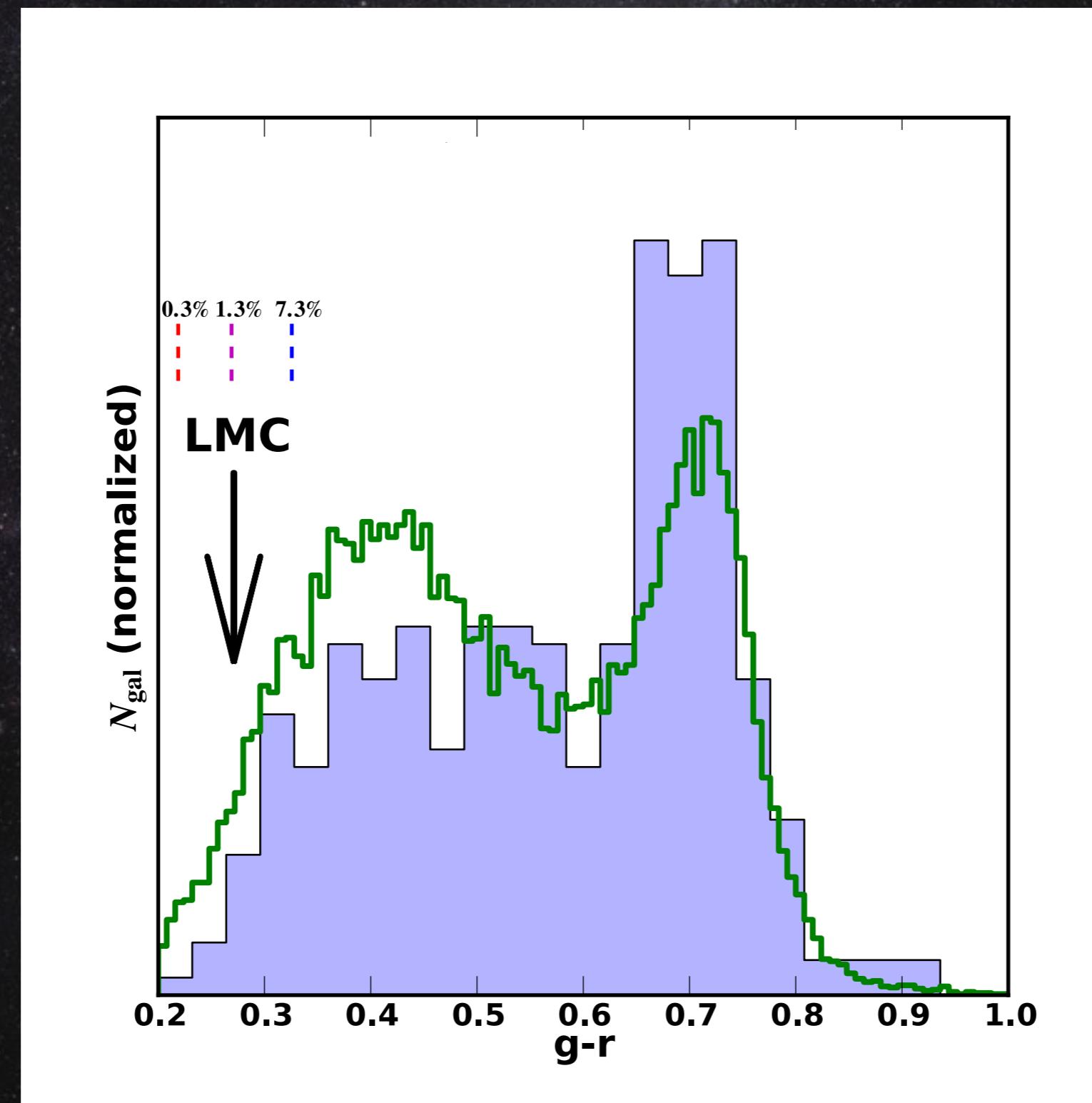
Why Are Secondaries Quenched?



- ◆ Not Harassment
- ◆ Ram Pressure Stripping?
- ◆ Strangulation?

LMC is Very Blue

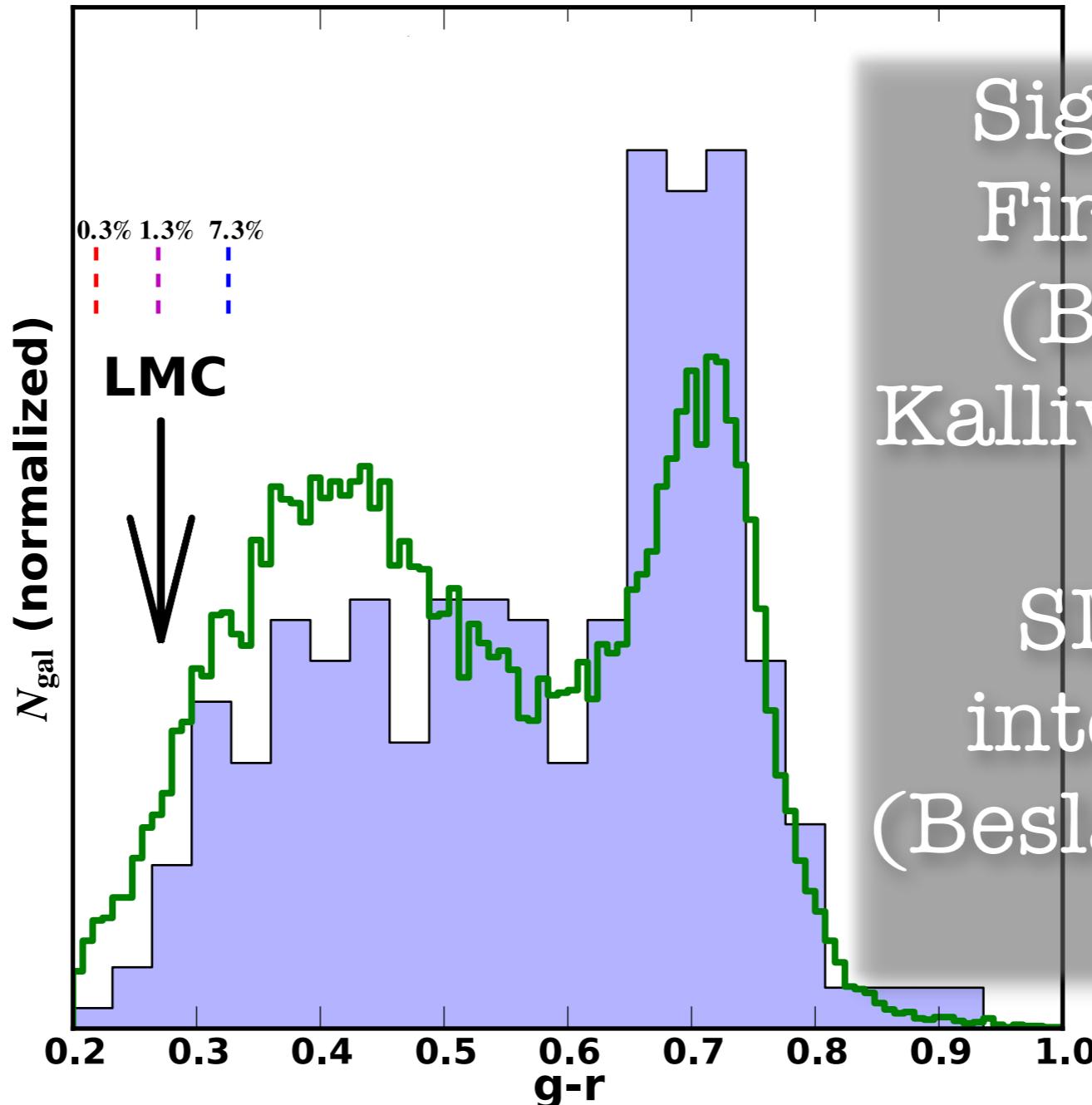
- RC3
- Eskew & Zaritsky II
- Bothun & Thompson
88



LMC is Very Blue

- RC3
- Eskew& Zaritsky II
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88

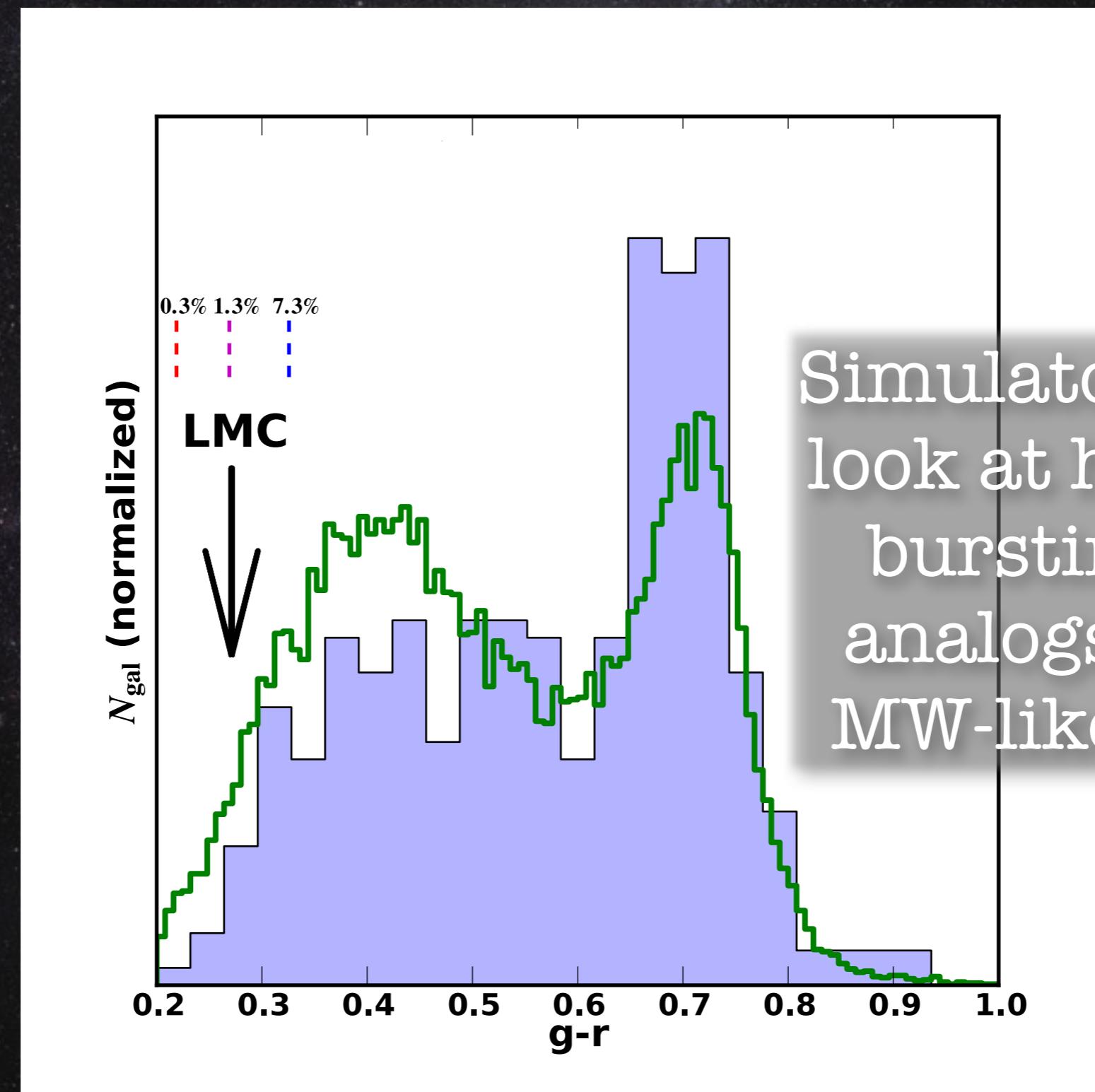


Signature of
First Infall?
(Besla+ 07,
Kallivayalil+ 06)

SMC/LMC
interaction?
(Besla+ 12, other
talks)

LMC is Very Blue

- RC3
 - Eskew & Zaritsky II
 - Bothun & Thompson
- 88



Summary

- ♦ LCDM+Abundance matching successfully **predicts basic observables** for bright satellite ($\sim 0.1 L^*$) of L^* hosts.
- ♦ Δv_{pair} matches between SDSS and simulation, implying host masses are correct
- ♦ Radial distribution of satellite galaxies match subhalos
- ♦ Empirical fraction of L^* galaxies w/ bright satellites is not tiny (LMC is not that weird)
- ♦ Empirical bright satellite properties:
 - ♦ Satellites of **isolated** L^* galaxies are **quenched**
 - ♦ LMC is very blue/starforming relative to similar objects (LMC *is weird*)