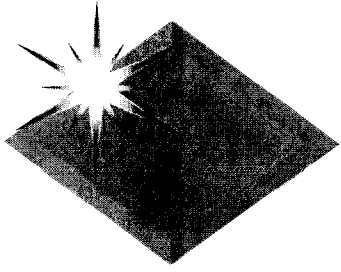


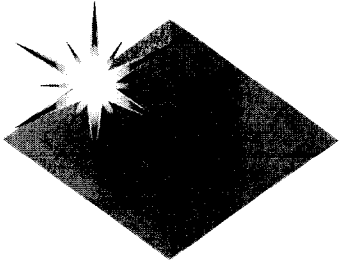
**Probing Cosmological Constant Through the Alcock-Paczynski  
Test Based on the Lyman-Alpha Forest**

Wen-Ching Lin, UCSD



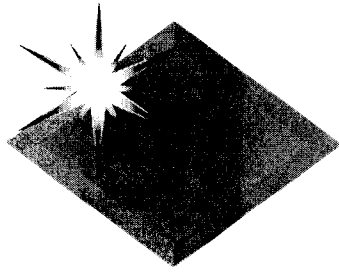
## Alcock-Paczynski Test :

- ◆ For a spherical object (which follows the Hubble flow) in the sky, its physical size along the line of sight and perpendicular to the line of sight should be the same.
- ◆ For a non-spherical structure like Lyman-Alpha forest, we use the correlation function instead of the physical size.
- ◆ **The cross-correlation function is most sensitive to  $\Omega_\Lambda$**

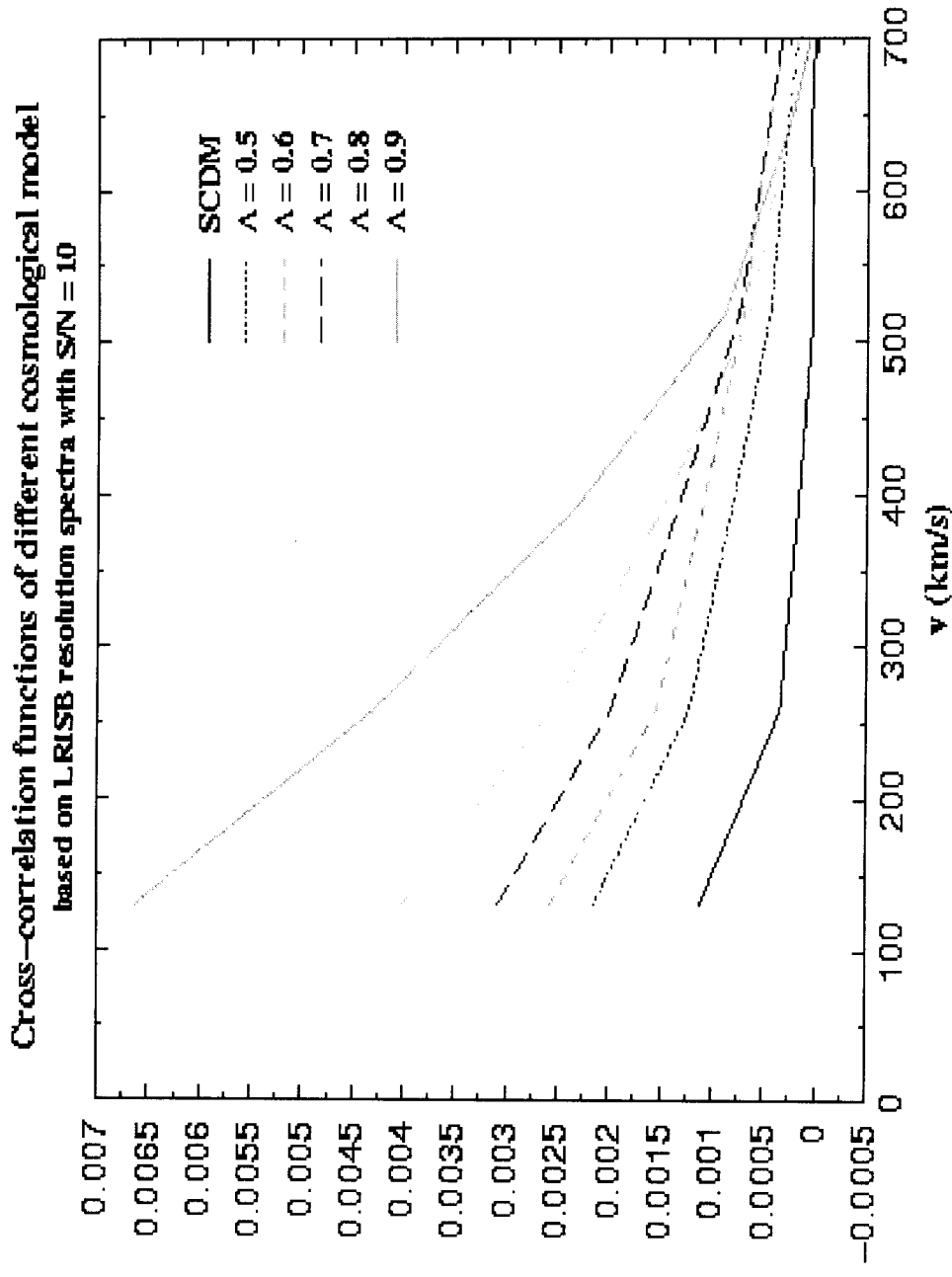


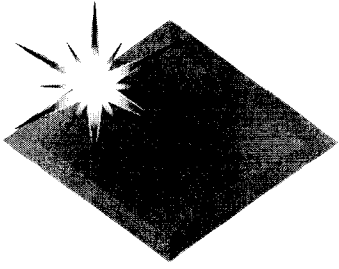
## Methodology :

- ◆ Use paired QSO spectra to provide real cross-correlation functions of our Universe.
- ◆ Use simulations to provide correlation-functions in different cosmological models (different  $\Omega_{\Lambda}$  ).
- ◆ Combine observational data and simulations to derive  $\Omega_{\Lambda}$ .



# Cross-correlation Functions For Different Cosmological Models





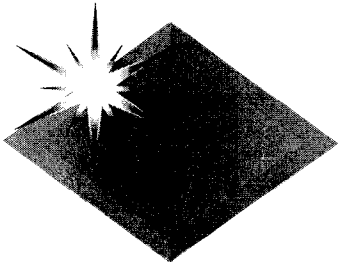
## **Major Difficulties and Solutions :**

- ◆ **Cosmic Variances :**

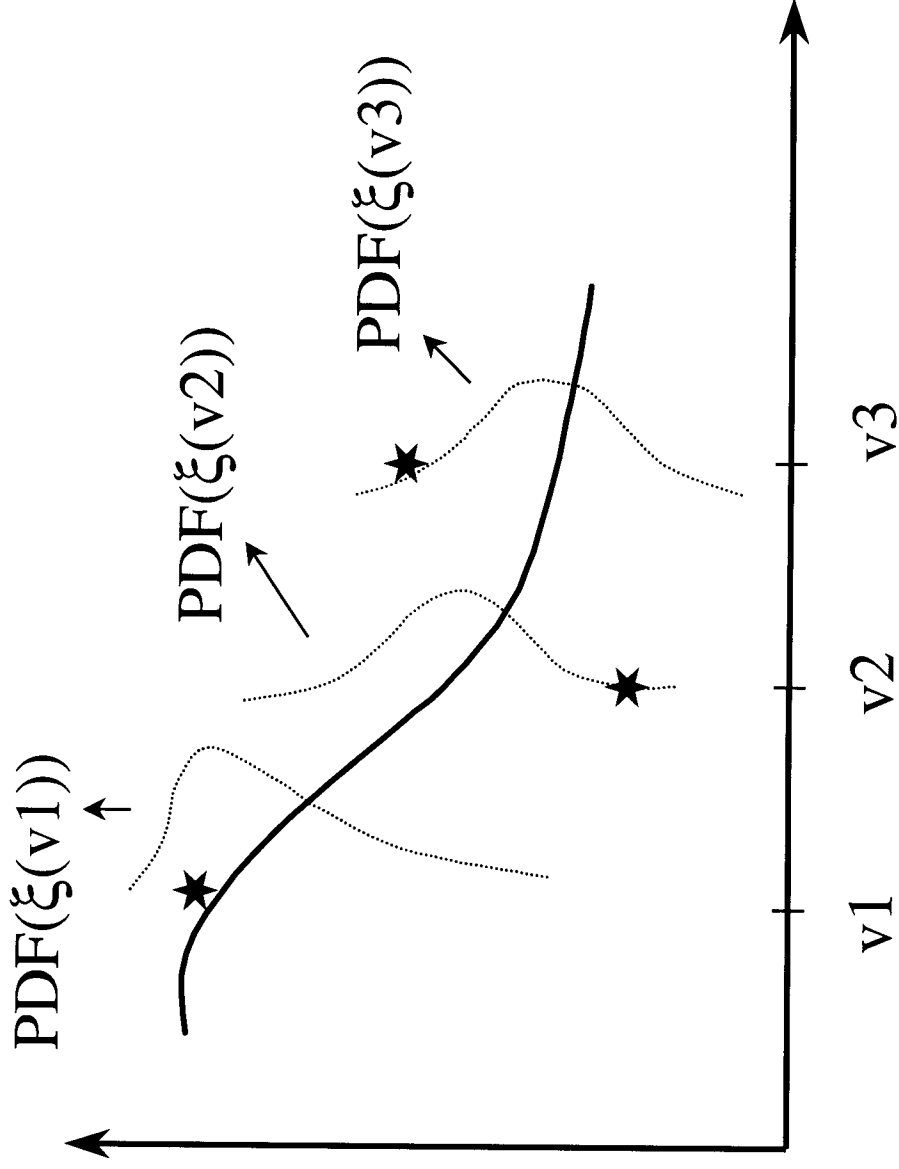
**Introduce the maximum likelihood estimation.**

- ◆ **Signal- to - noise (S/N) properties :**

**Generate artificial QSO spectra with different kinds of signal- to - noise properties for our study.**



# Maximum Likelihood Estimation :

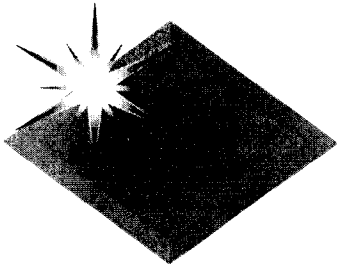


Cross-correlation

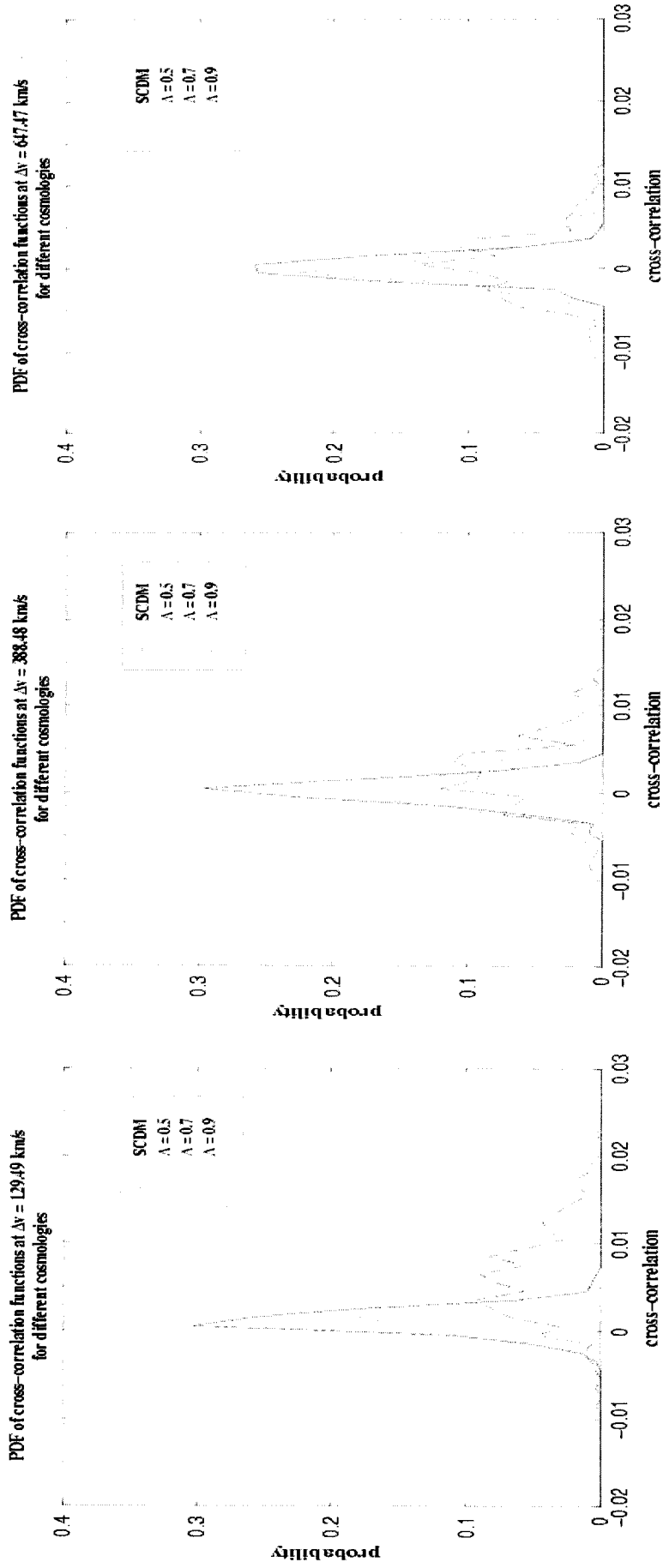
||

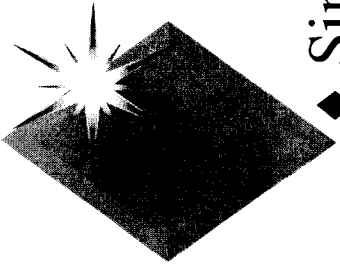
$\xi(v)$

★ : observational data point



# Probability Distribution Functions at Different velocity separations





## Procedures :

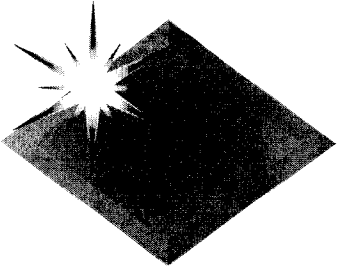
- ◆ Simulations provide PDF( $\xi(v) | \Omega_\Lambda$ ) at each  $v$  for different cosmological model.
- ◆ Take simulated QSO spectra in  $\Omega_\Lambda = 0.7$  simulations and pretended they were observation data.
- ◆ Then we can calculate the combined-likelihood based on all the data points we have :



where  $x_i$  are fake observation data points and the corresponding PDF are provided by our simulations.

- ◆ Maximizing the likelihood function  $L(\Omega_\Lambda)$  to determine  $\Omega_\Lambda$





## Conclusions :

- ◆ With 30 pairs of QSO spectra we have more than 60% probability to determine  $\Omega_{\Lambda}$  within 10% error and more than 90 % probability to obtain  $\Omega_{\Lambda}$  within 20% error.
- ◆ QSO spectra with  $S/N \geq 20$  are required to achieve the above results.

TABLE I  
PROBABILITY OF GETTING  $\Lambda$  WITHIN 10 % ERROR

Fail's/group	group #	Results of $\Omega_{\Lambda}$	Prob(20%)	Prob(10%)	Prob(5%)
10	30	$\Omega_{\Lambda} = 0.6 \pm 0.145$	55.97%	29.76%	15.11%
15	20	$\Omega_{\Lambda} = 0.6 \pm 0.104$	63.93%	33.55%	16.99%
20	15	$\Omega_{\Lambda} = 0.64 \pm 0.1$	76.54%	44.31%	23.03%
25	12	$\Omega_{\Lambda} = 0.7 \pm 0.11$	79.69%	47.55%	24.97%
30	10	$\Omega_{\Lambda} = 0.7 \pm 0.074$	94.15%	65.56%	36.38%